2 July 2019
'Beyond' conference - Warsaw

Dark Matter Indirect Searches

as of 2019

Marco Cirelli

(CNRS LPTHE Jussieu)

LPTHE
LABORATOIRE DE PHYSIQUE THEORIQUE ET HAUTES ENERGIES

cnrs dépasser les frontières

SORBONNE UNIVERSITÉ
CRÉATEURS DE FUTURS DEPUIS 1257

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DM detection

direct detection

production at colliders

indirect

- γ from annihil in galactic center or halo and from secondary emission
  Fermi, ICT, radio telescopes...
- e^+ from annihil in galactic halo or center
  PAMELA, Fermi, HESS, AMS,balloons...
- $\bar{p}$ from annihil in galactic halo or center
- $\bar{d}$ from annihil in galactic halo or center
  GAPS, AMS
- $\nu, \bar{\nu}$ from annihil in massive bodies
  SK, Icecube, Antares

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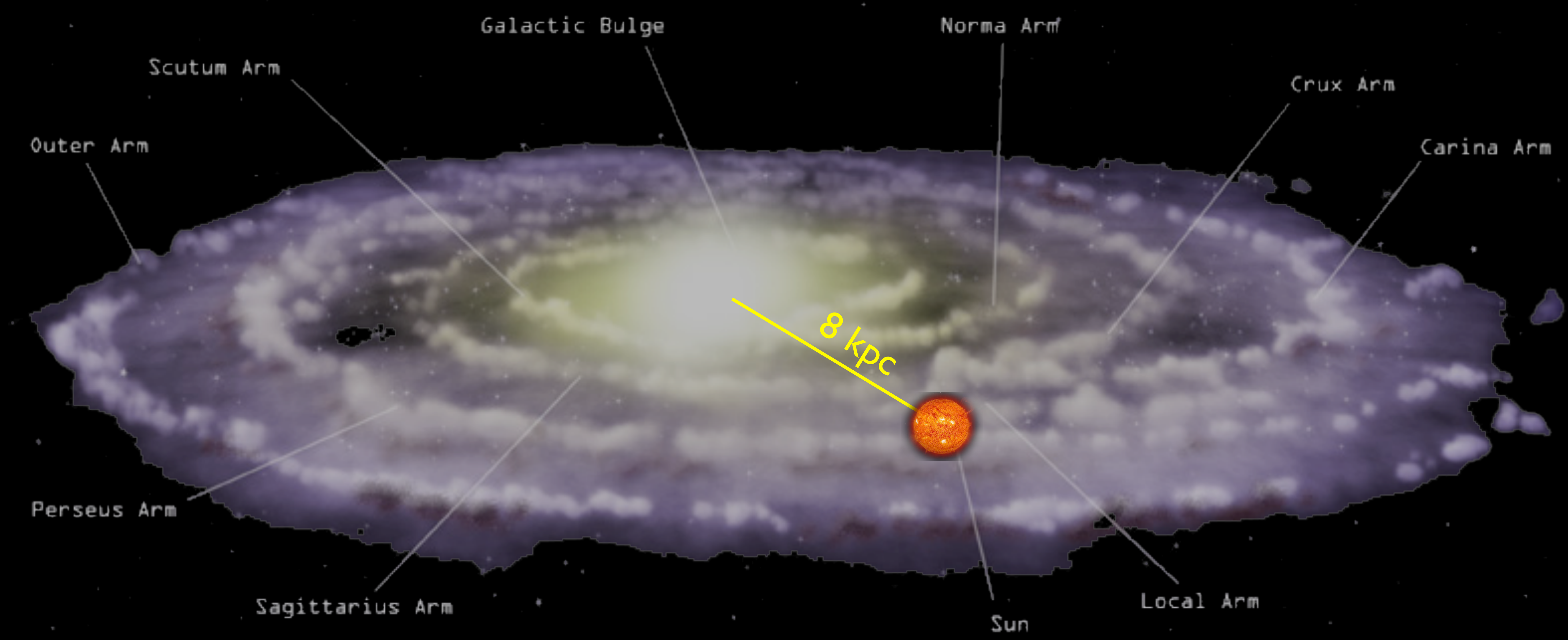
Indirect Detection: basics

$\bar{p}$ and e^+ from DM annihilations in halo



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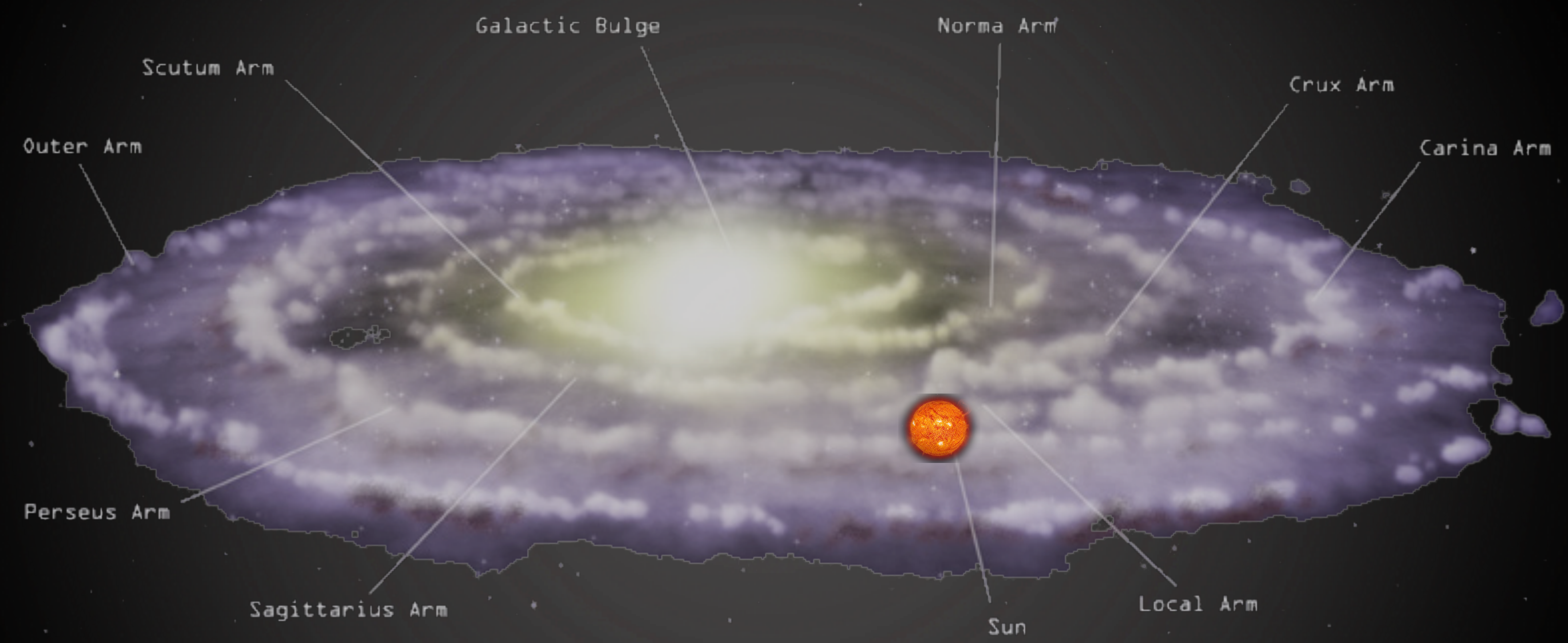

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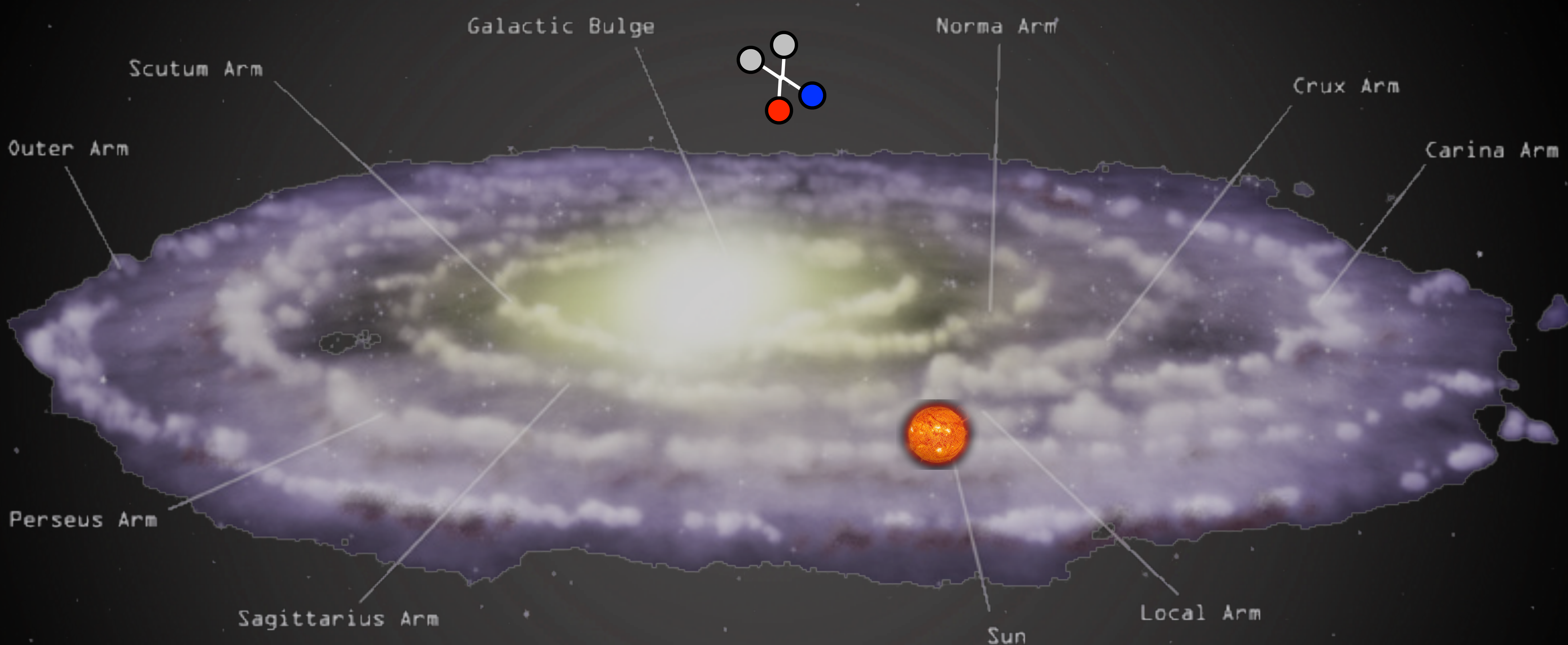

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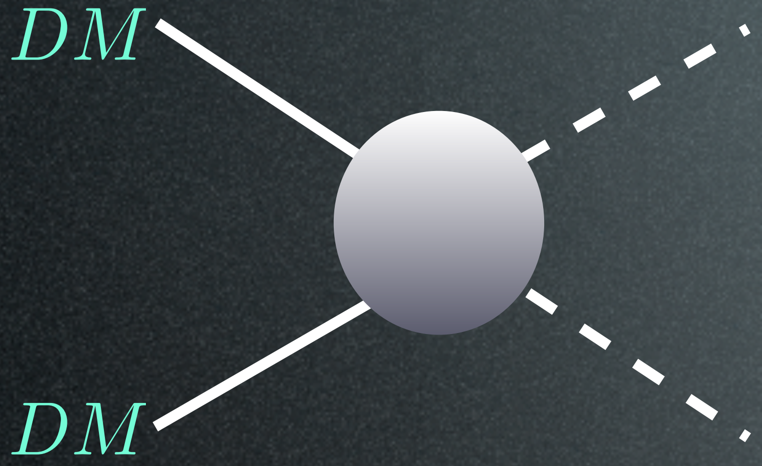
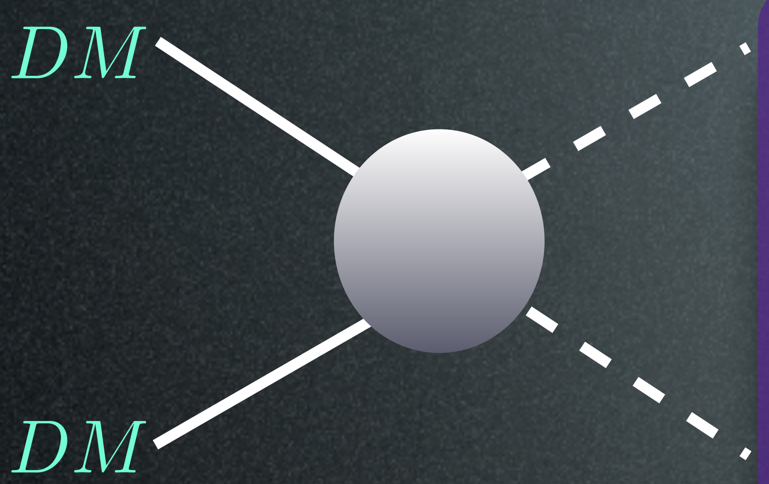


$$W^-, Z, b, \tau^-, t, h \ldots \rightsquigarrow e^\mp, \overset{(-)}{p}, \overset{(-)}{D} \ldots$$

$$W^+, Z, \bar{b}, \tau^+, \bar{t}, h \ldots \rightsquigarrow e^\pm, \overset{(-)}{p}, \overset{(-)}{D} \ldots$$

Indirect Detection: basics

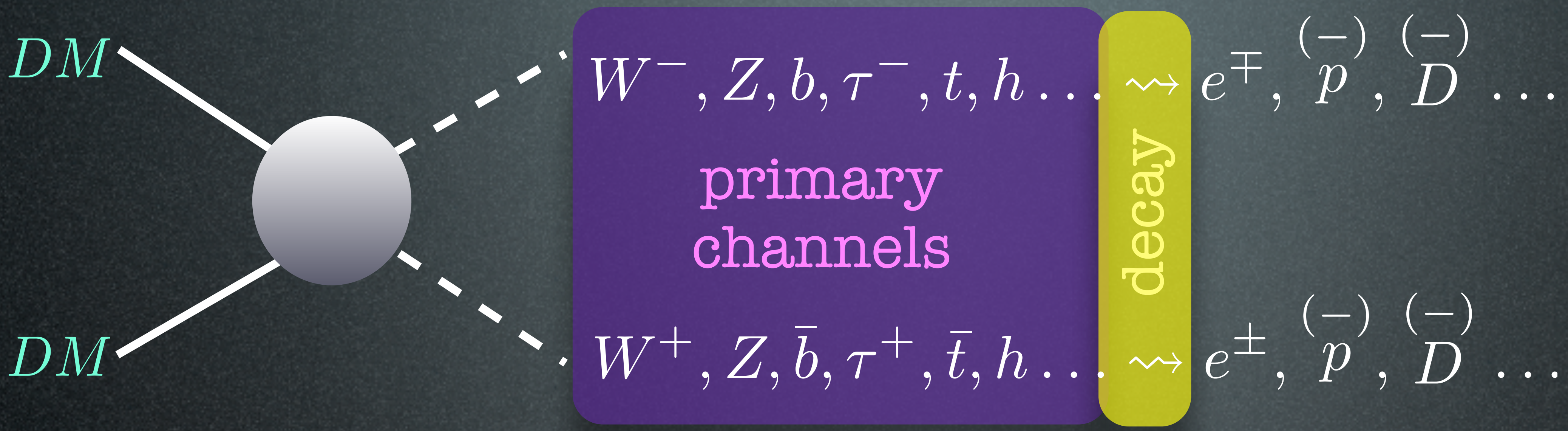


$W^-, Z, b, \tau^-, t, h \ldots \rightsquigarrow e^\mp, \overset{(-)}{p}, \overset{(-)}{D} \ldots$

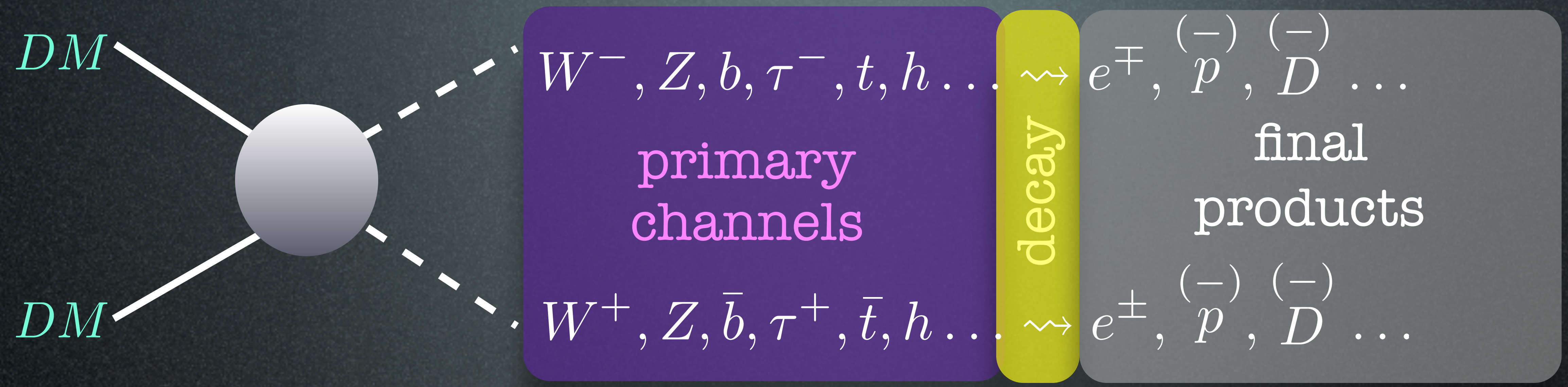
primary channels

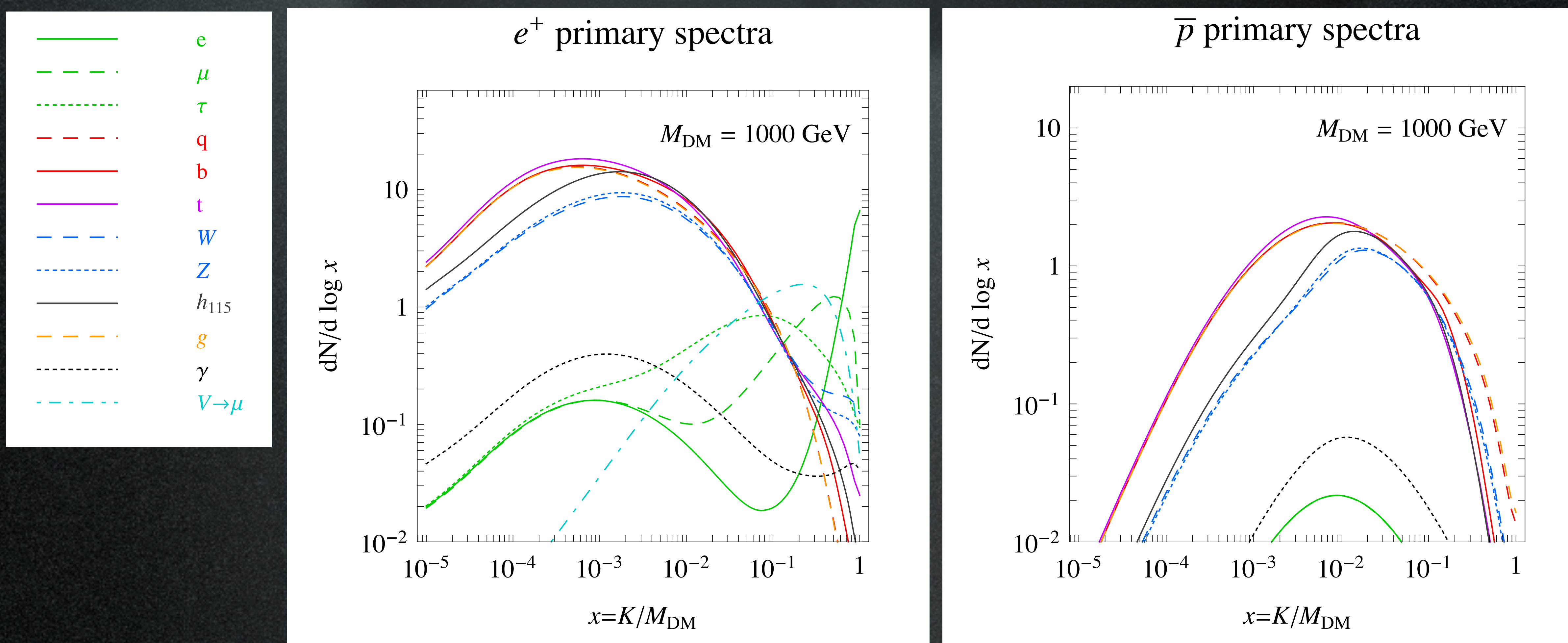
$W^+, Z, \bar{b}, \tau^+, \bar{t}, h \ldots \rightsquigarrow e^\pm, \overset{(-)}{p}, \overset{(-)}{D} \ldots$

Indirect Detection: basics

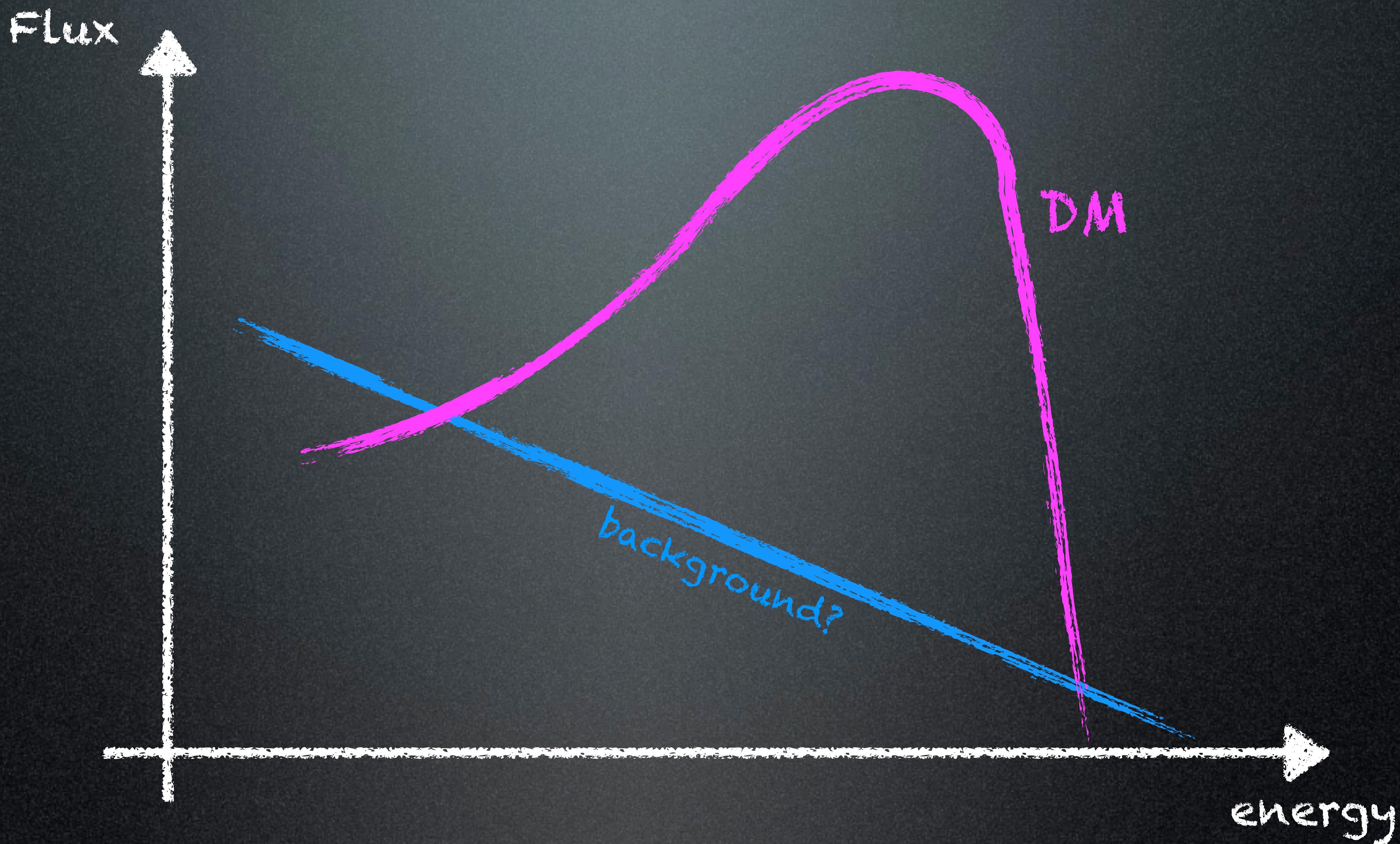


Indirect Detection: basics



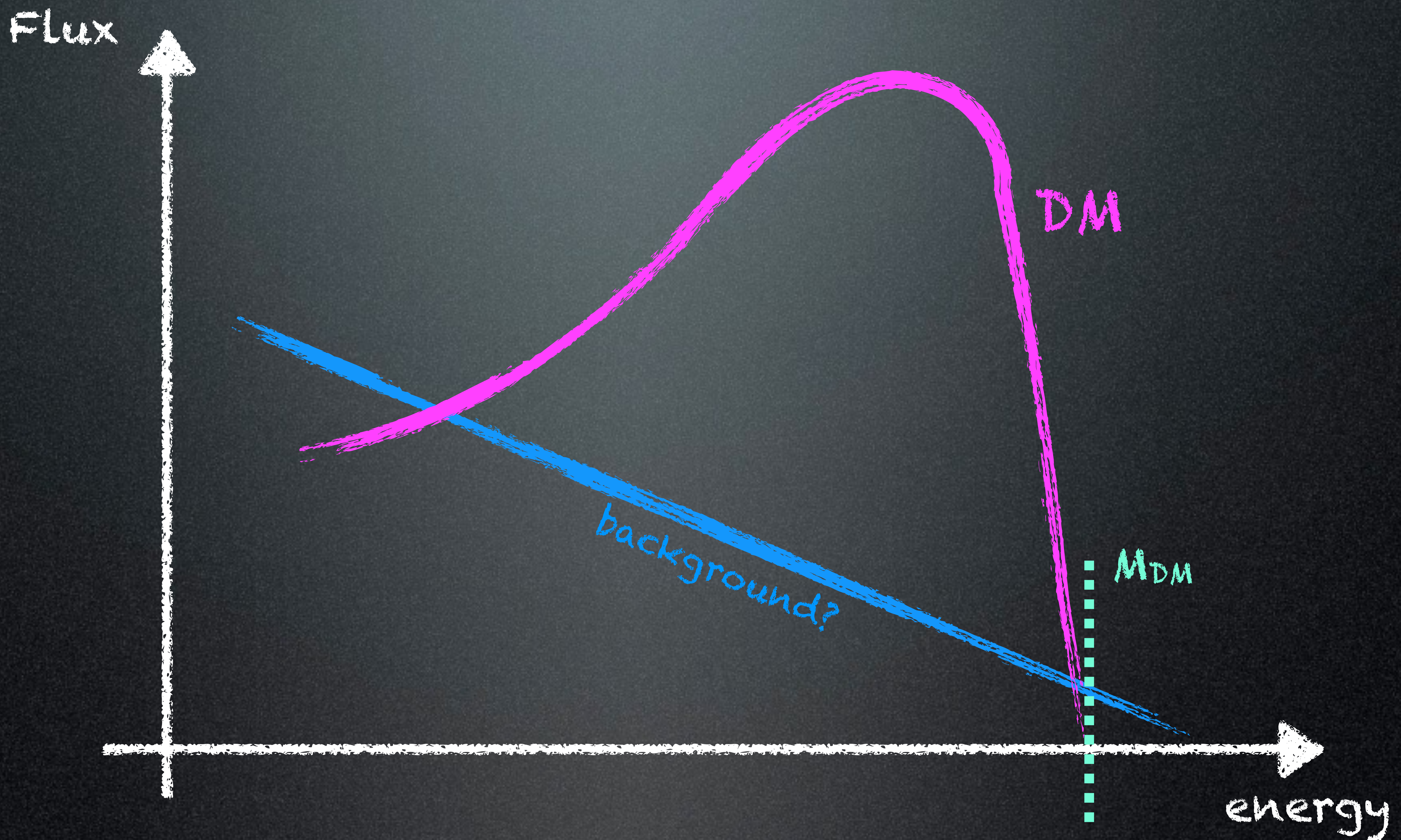


Fluxes at production



So what are the particle physics parameters?

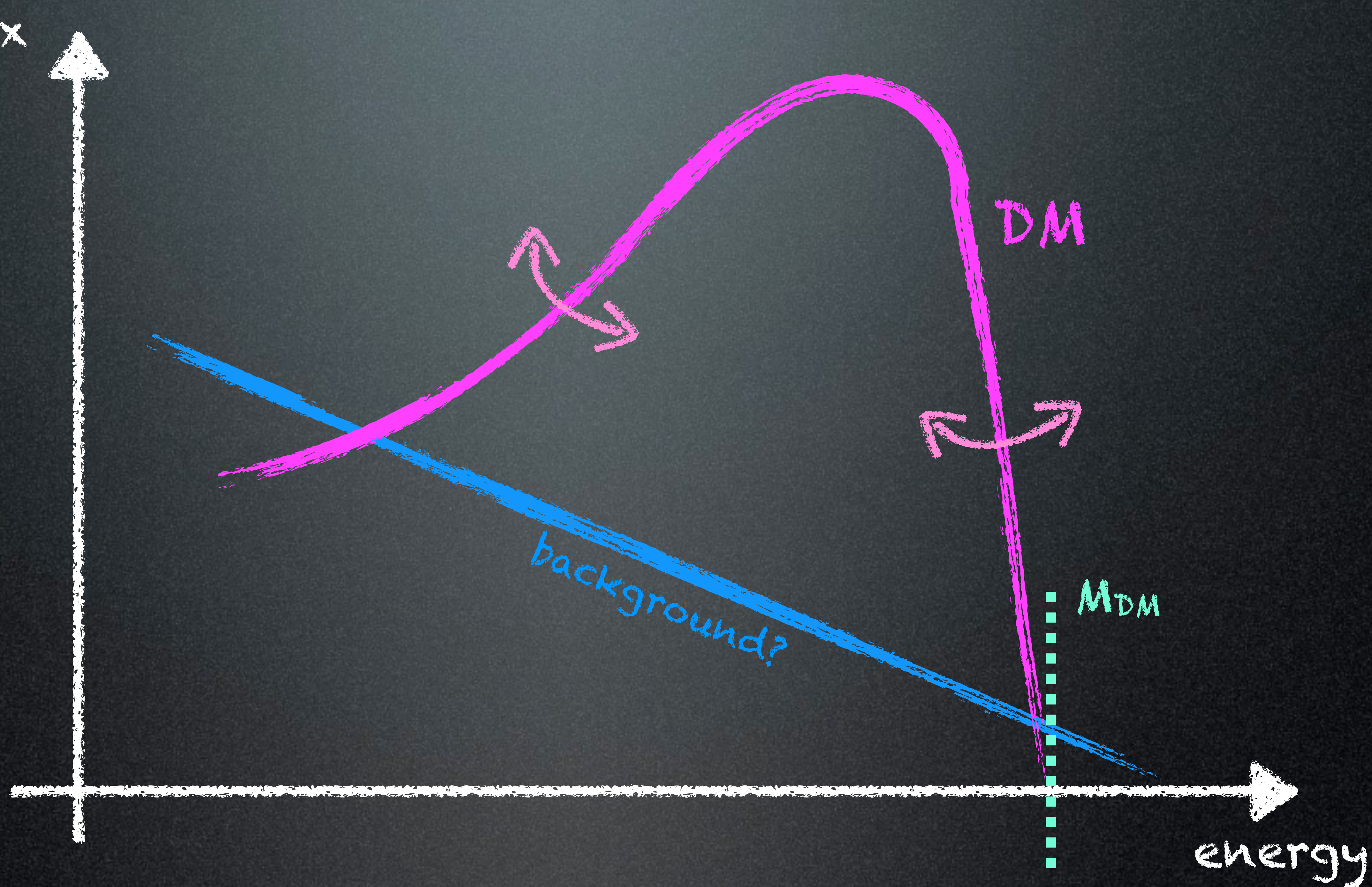
Fluxes at production



So what are the particle physics parameters?

1. Dark Matter mass

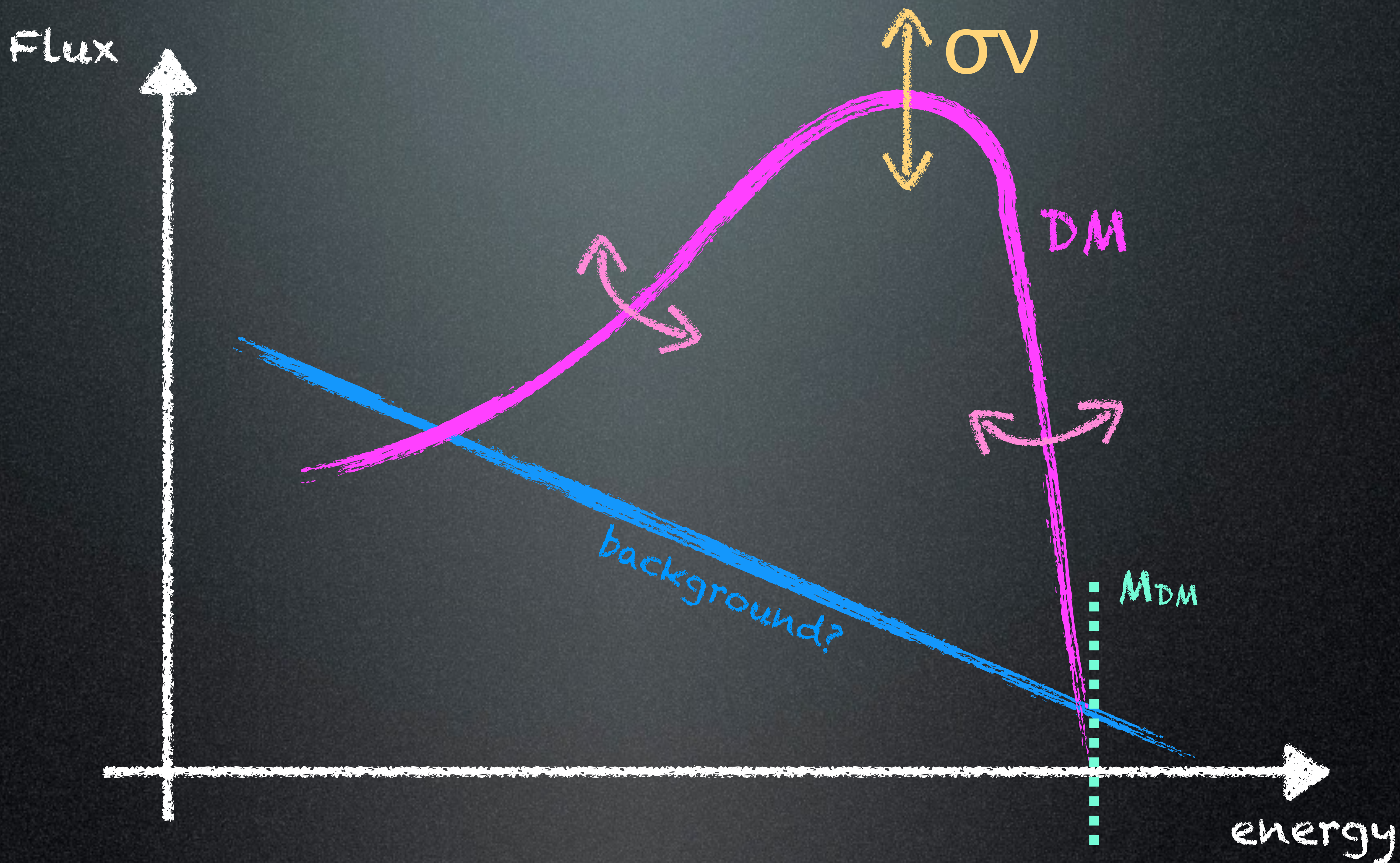
Fluxes at production



So what are the particle physics parameters?

1. Dark Matter mass
2. primary channel(s)

Fluxes at production



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1. Dark Matter mass
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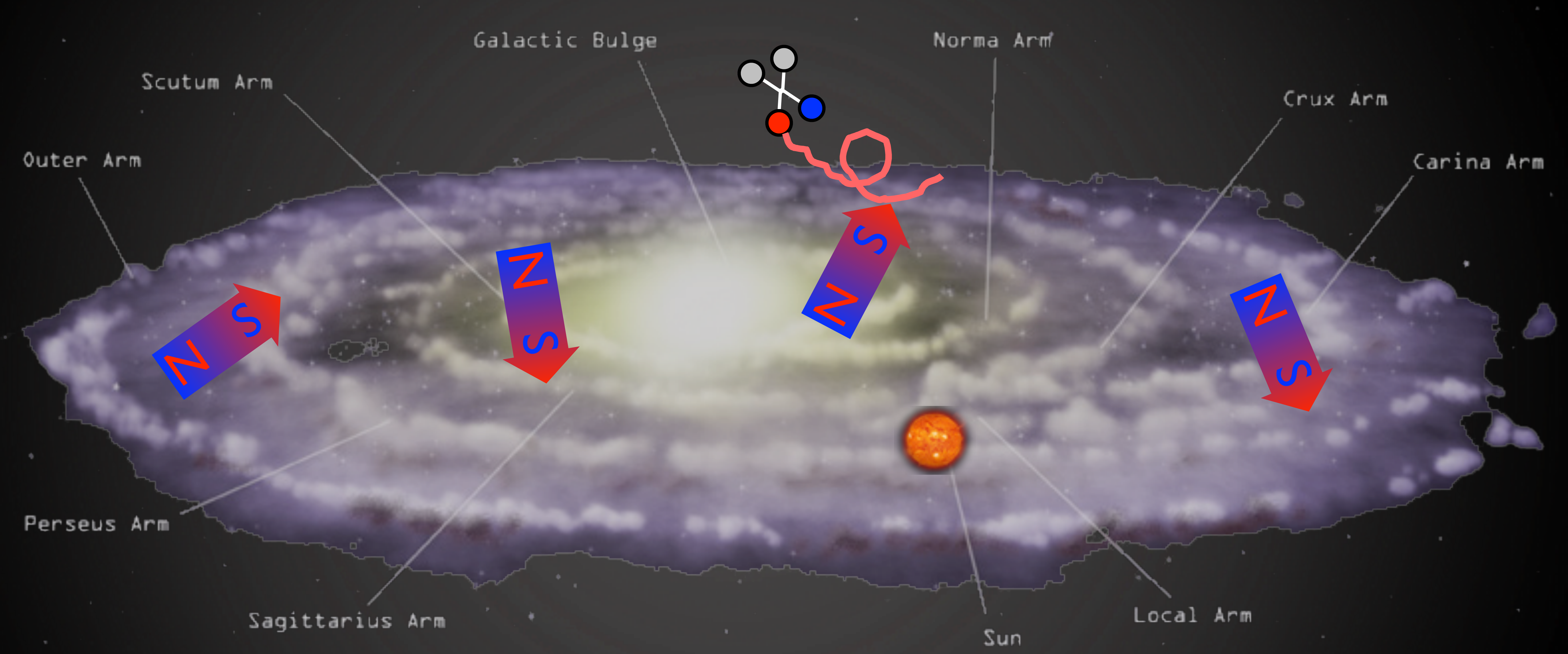
3. cross section

Indirect Detection: charged CRs

$\bar{p}$ and e^+ from DM annihilations in halo



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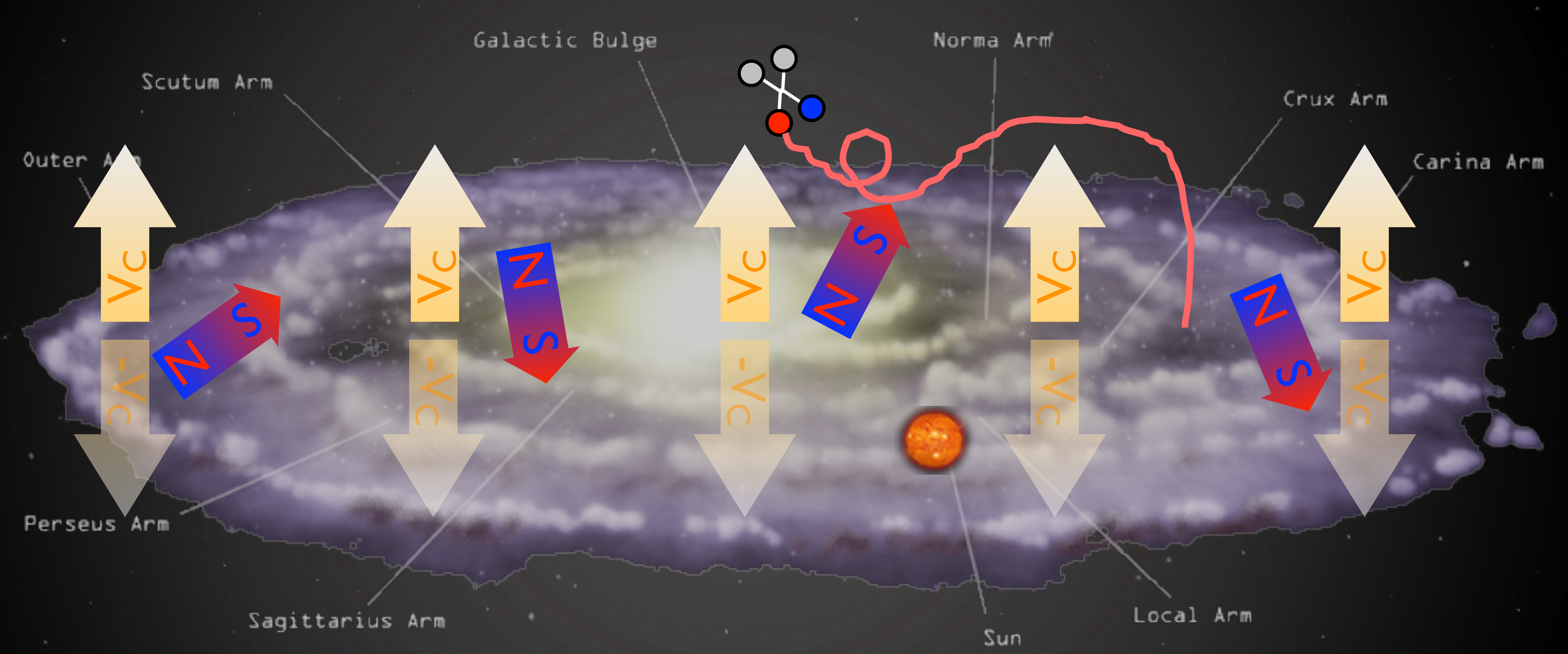

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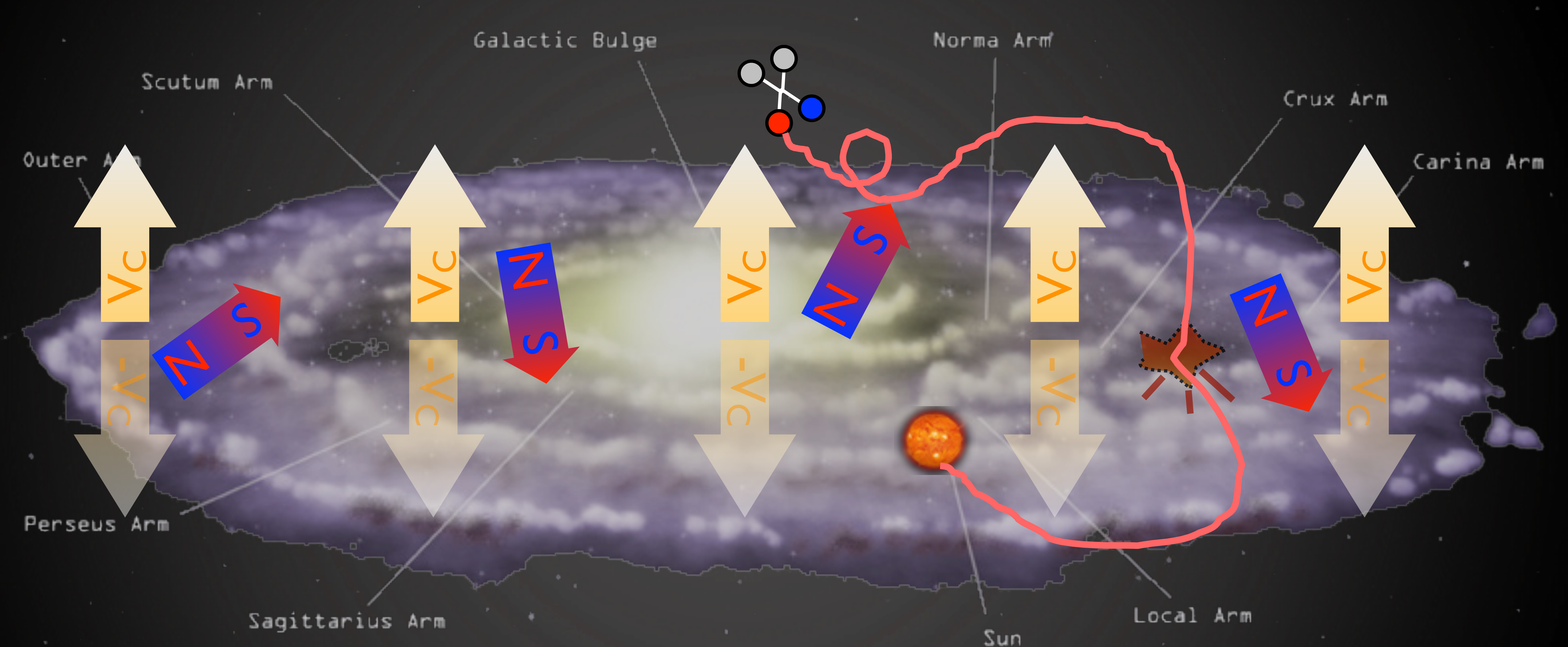

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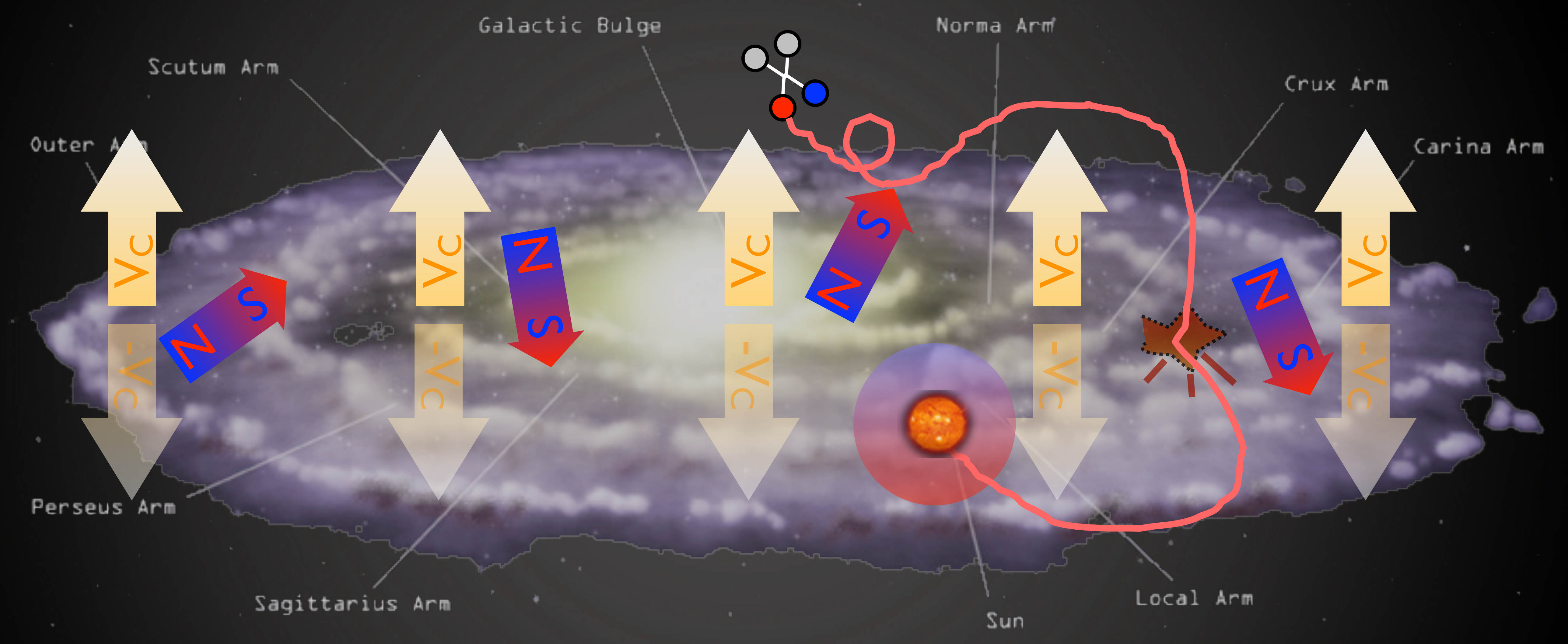

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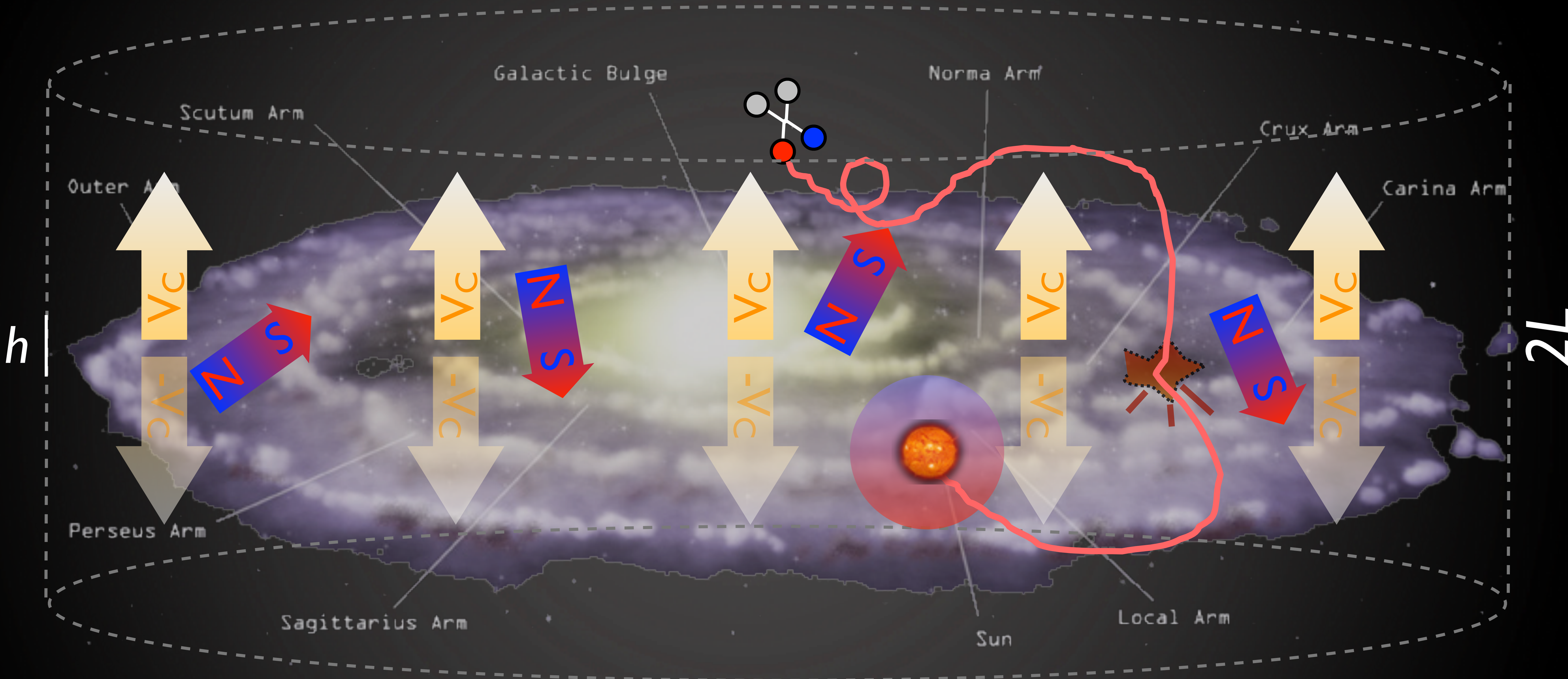


Salati, Chardonnay, Barrau, Donato, Taillet, Fornengo, Maurin, Brun... '90s, '00s

spectrum

$$\frac{\partial f}{\partial t} - K(E) \cdot \nabla^2 f - \frac{\partial}{\partial E}\left(b(E) f\right) + \frac{\partial}{\partial z}(V_c f) = Q_{\rm inj} - 2h\delta(z)\Gamma_{\rm spall} f$$

diffusion — energy loss — convective wind — source — spallations

[uncert]

Indirect Detection: charged CRs

$\bar{p}$ and e^+ from DM annihilations in halo

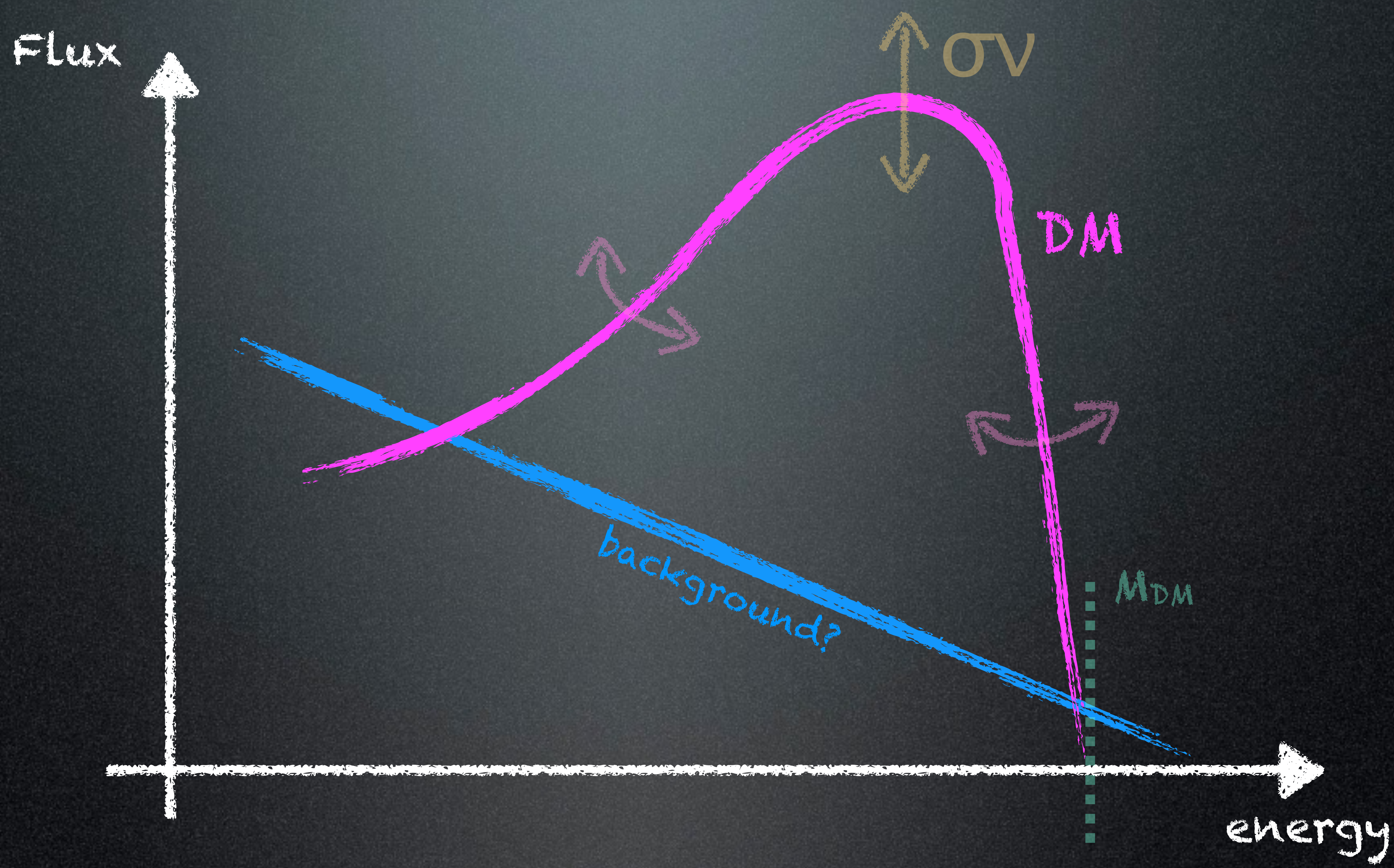
		KRA	KOL	CON	THK	THN	THN2	THN3
thickness	L [kpc]	4	4	4	10	0.5	2	3
diffusion	D_0 [10^{28} cm^2 s^{-1}]	2.64	4.46	0.97	4.75	0.31	1.35	1.98
diffusion	δ	0.50	0.33	0.6	0.50	0.50	0.50	0.50
diffusion	η	-0.39	1	1	-0.15	-0.27	-0.27	-0.27
diff. reacc.	v_A [km s^{-1}]	14.2	36	38.1	14.1	11.6	11.6	11.6
p index	γ	2.35	1.78/2.45	1.62/2.35	2.35	2.35	2.35	2.35
convection	dv_c/dz [km s^{-1} kpc^{-1}]	0	0	50	0	0	0	0
solar mod.	ϕ_F^p [GV]	0.650	0.335	0.282	0.687	0.704	0.626	0.623
	χ^2_{min}/dof (p in [25])	0.462	0.761	1.602	0.516	0.639	0.343	0.339

Cirelli, Gaggero, Giesen, Taoso, Urbano 1407.2173
cfr. Evoli, Cholis, Grasso, Maccione, Ullio, 1108.0664

	Electrons or positrons		Antiprotons (and antideuterons)			
Model	δ	$\mathcal{K}_0$ [kpc^2/Myr]	δ	$\mathcal{K}_0$ [kpc^2/Myr]	V_{conv} [km/s]	L [kpc]
MIN	0.55	0.00595	0.85	0.0016	13.5	1
MED	0.70	0.0112	0.70	0.0112	12	4
MAX	0.46	0.0765	0.46	0.0765	5	15

Donato et al., 2003+

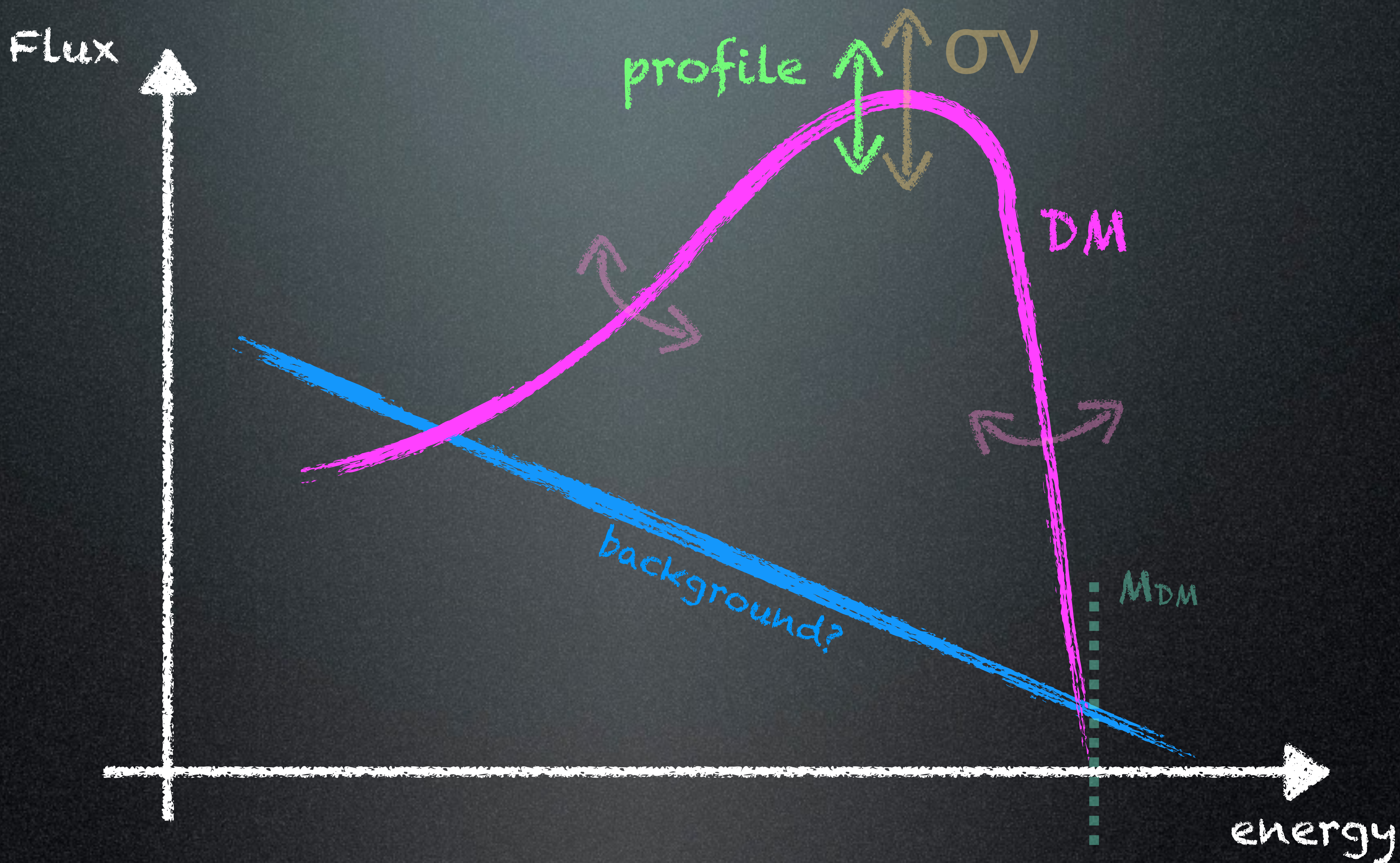
Fluxes at detection



So what are the astrophysics parameters?

1. Dark Matter mass
2. primary channel(s)
3. cross section

Fluxes at detection



So what are the **astrophysics** parameters?

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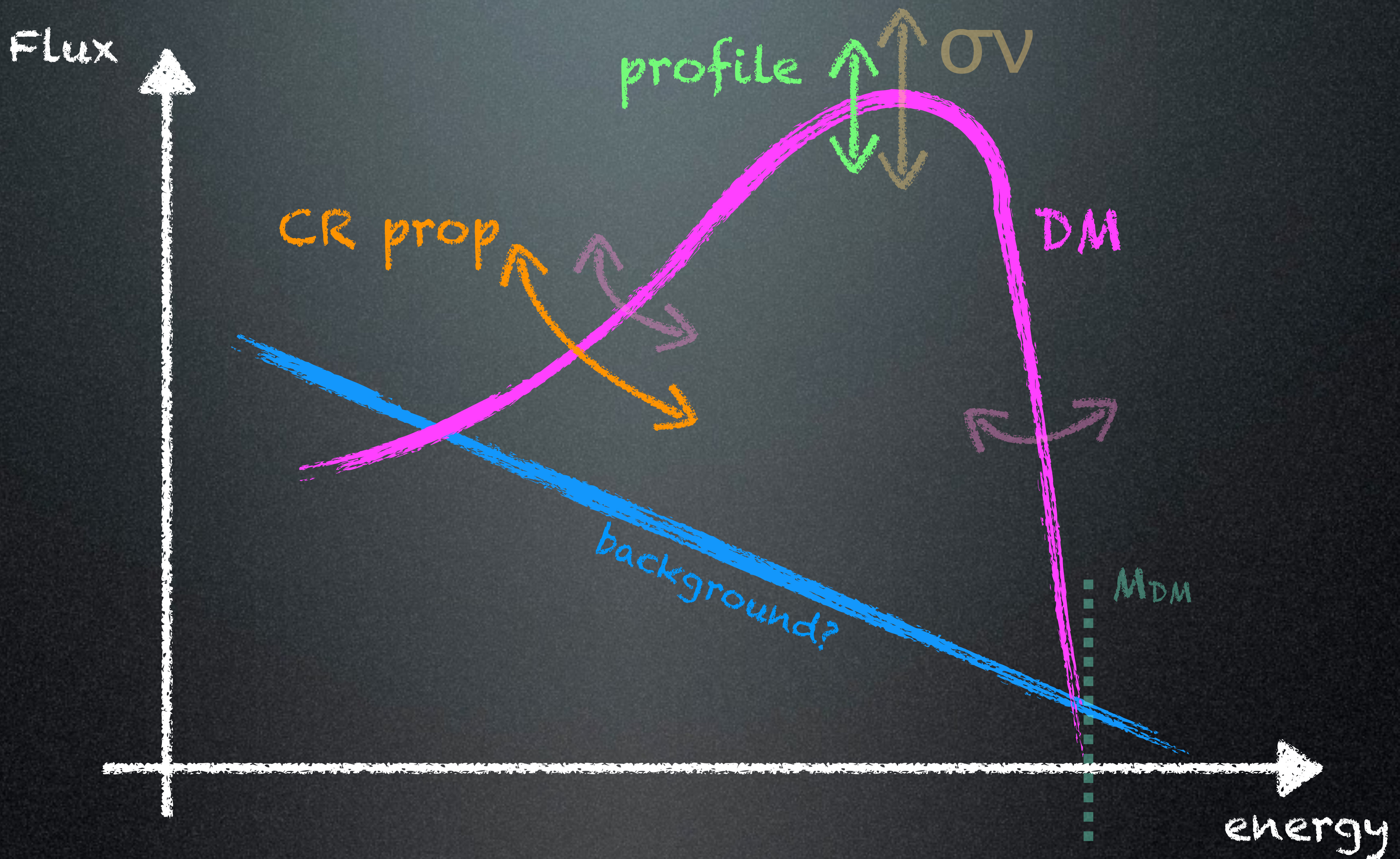
3. cross section

1. DM abundance/profile

Fluxes at detection



So what are the astrophysics parameters?

1. Dark Matter mass
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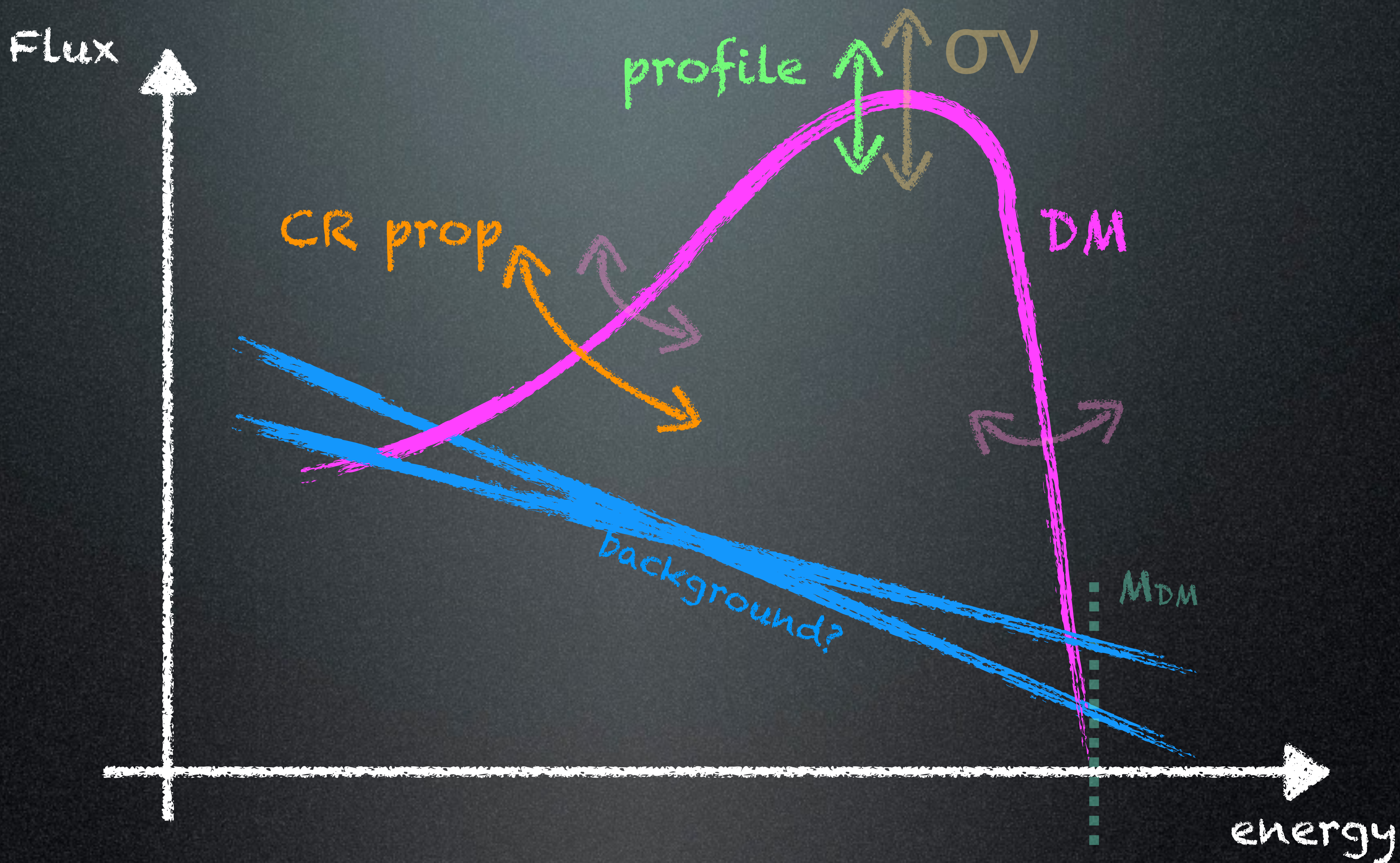
1. DM abundance/profile
2. propagation

Fluxes at detection



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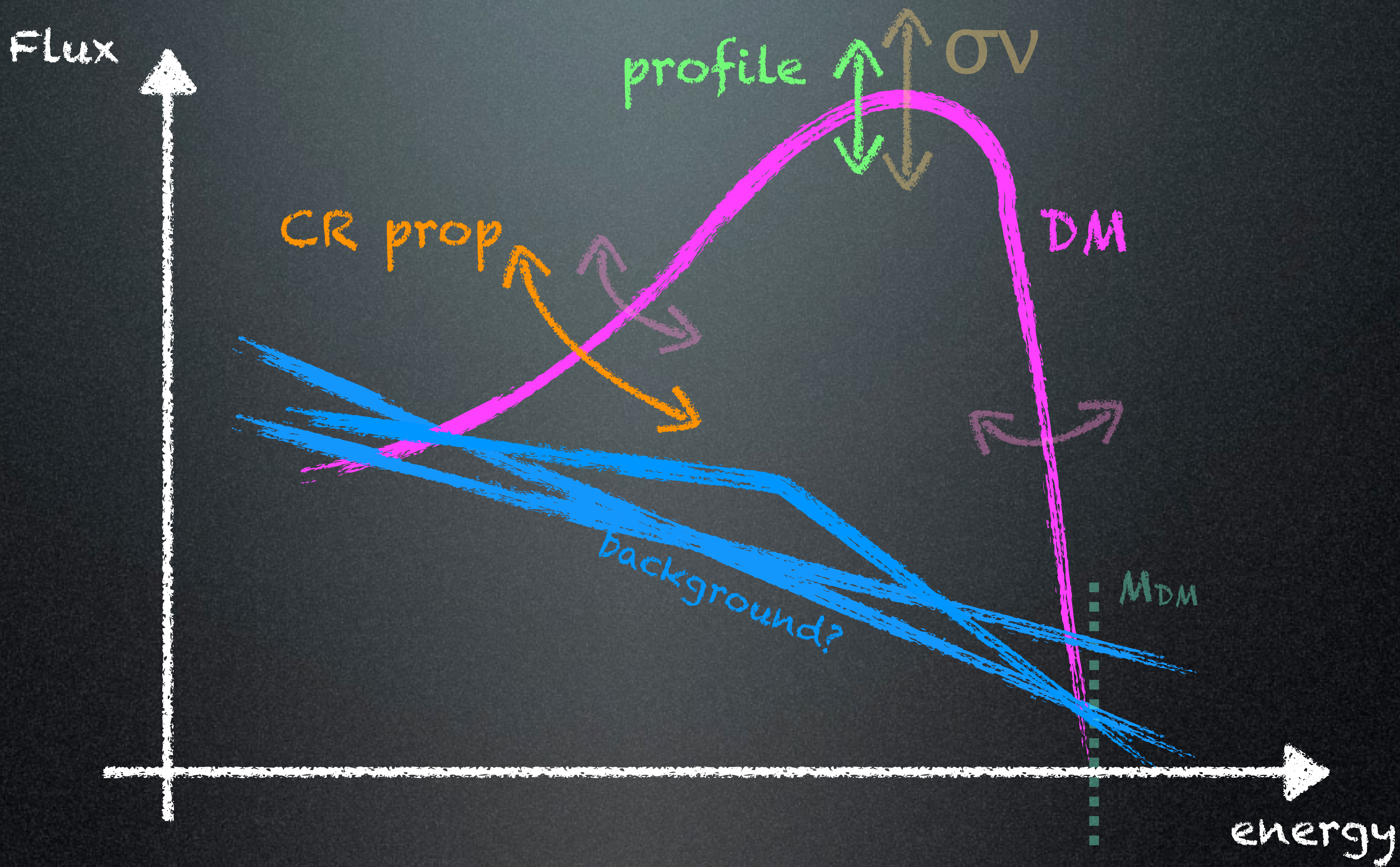
3. background

Fluxes at detection



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Antiprotons

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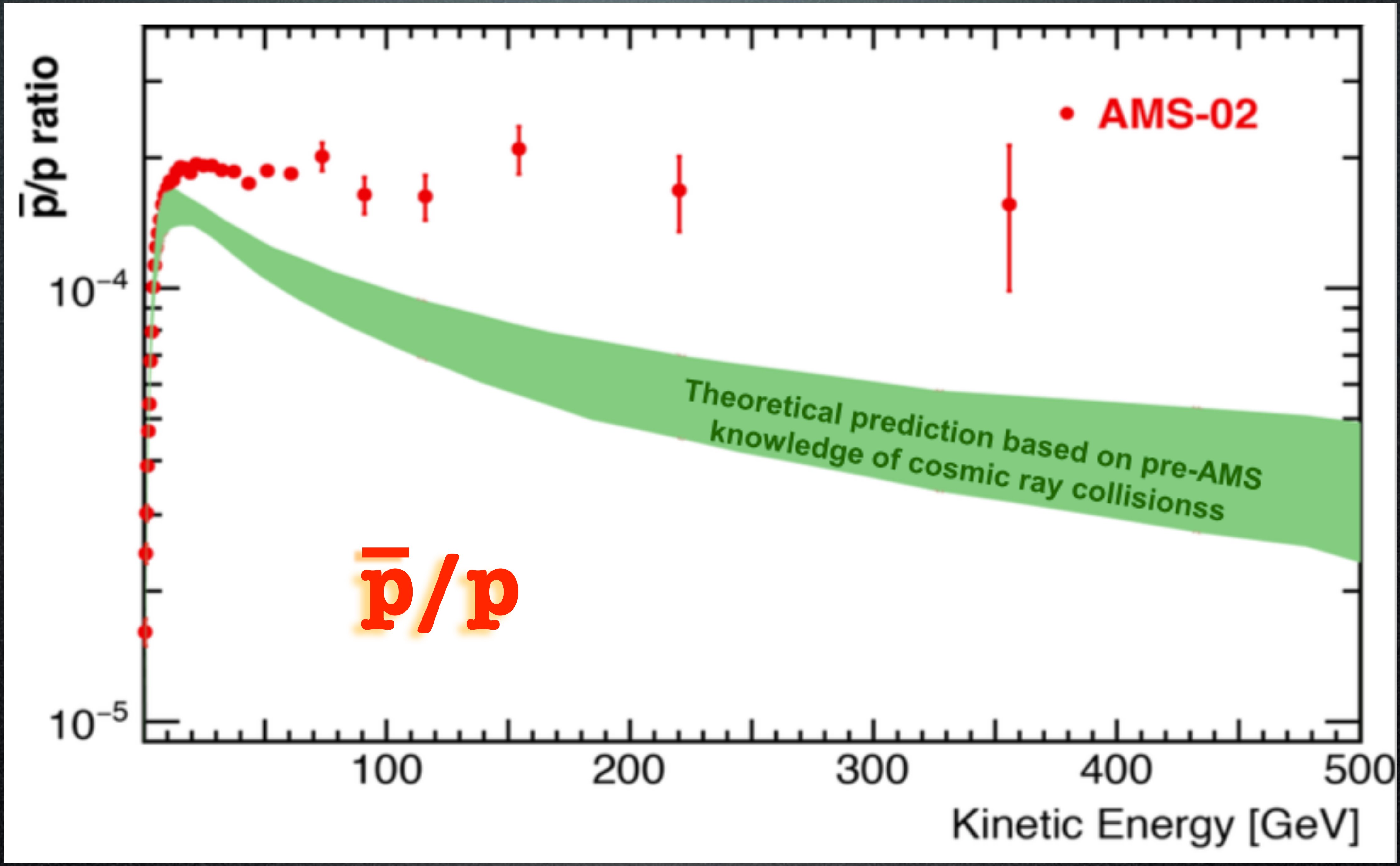
  SK, Icecube, Km3Net

Data: antiprotons

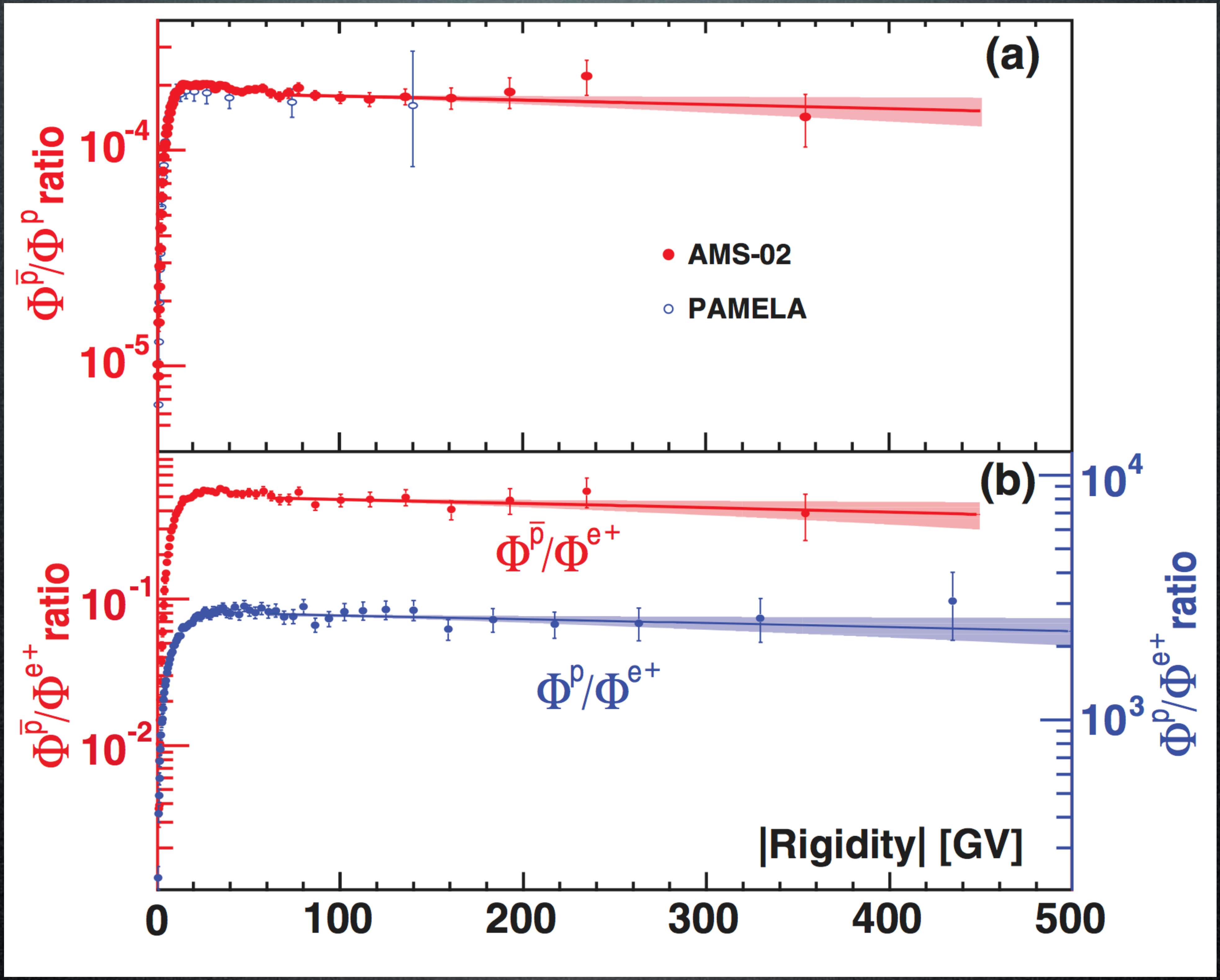
AMS-02



S. Ting - AMS days @ CERN apr 2015
A. Kounine - AMS days @ CERN apr 2015

Data: antiprotons

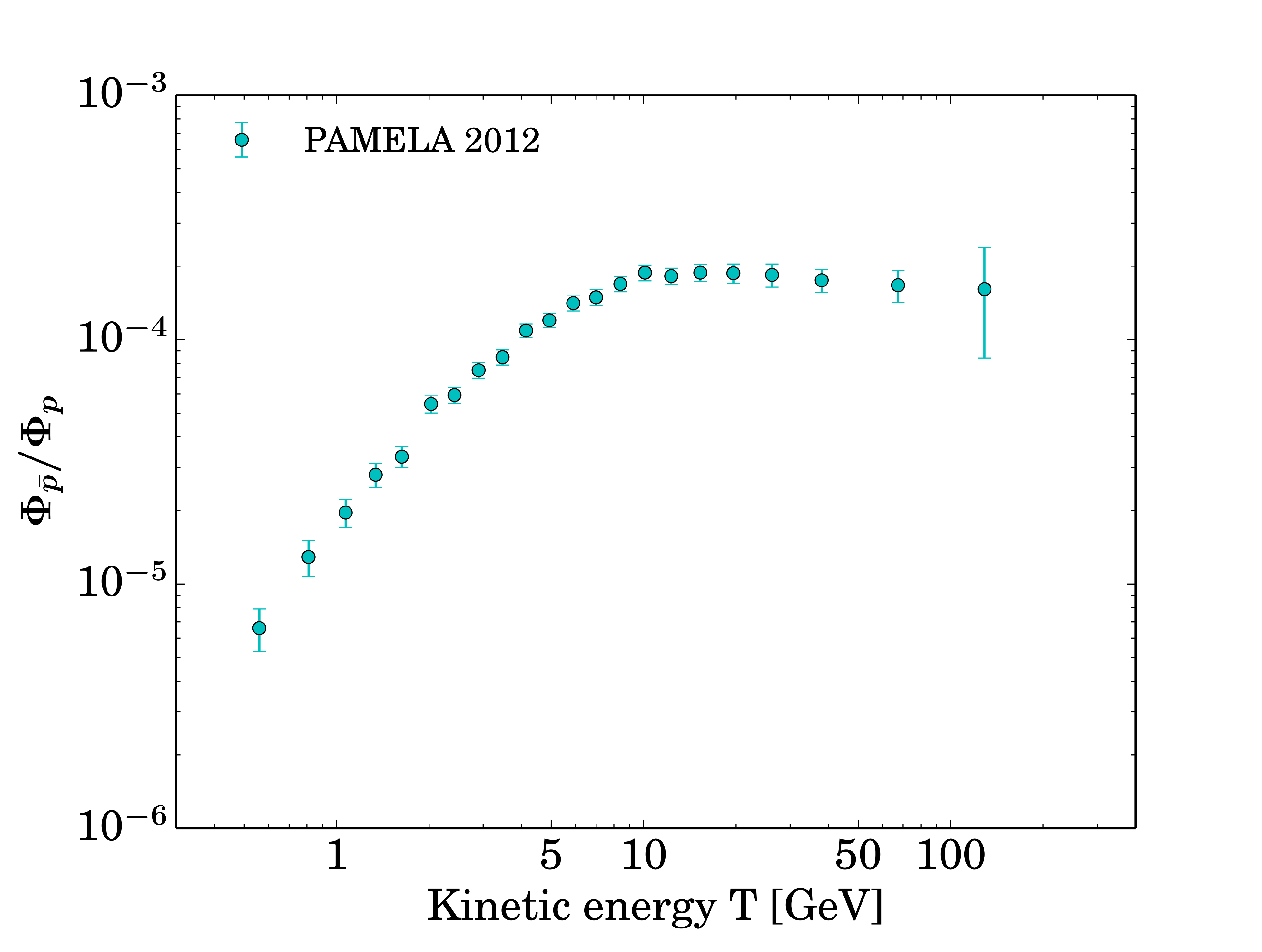
AMS-02



AMS coll.,PRL 117, 091103 (2016)

Antiprotons

Antiproton data vis-à-vis the secondaries:



Giesen, Boudaud, Génolini, Poulin, Cirelli, Salati, Serpico 1504.04276

Antiprotons

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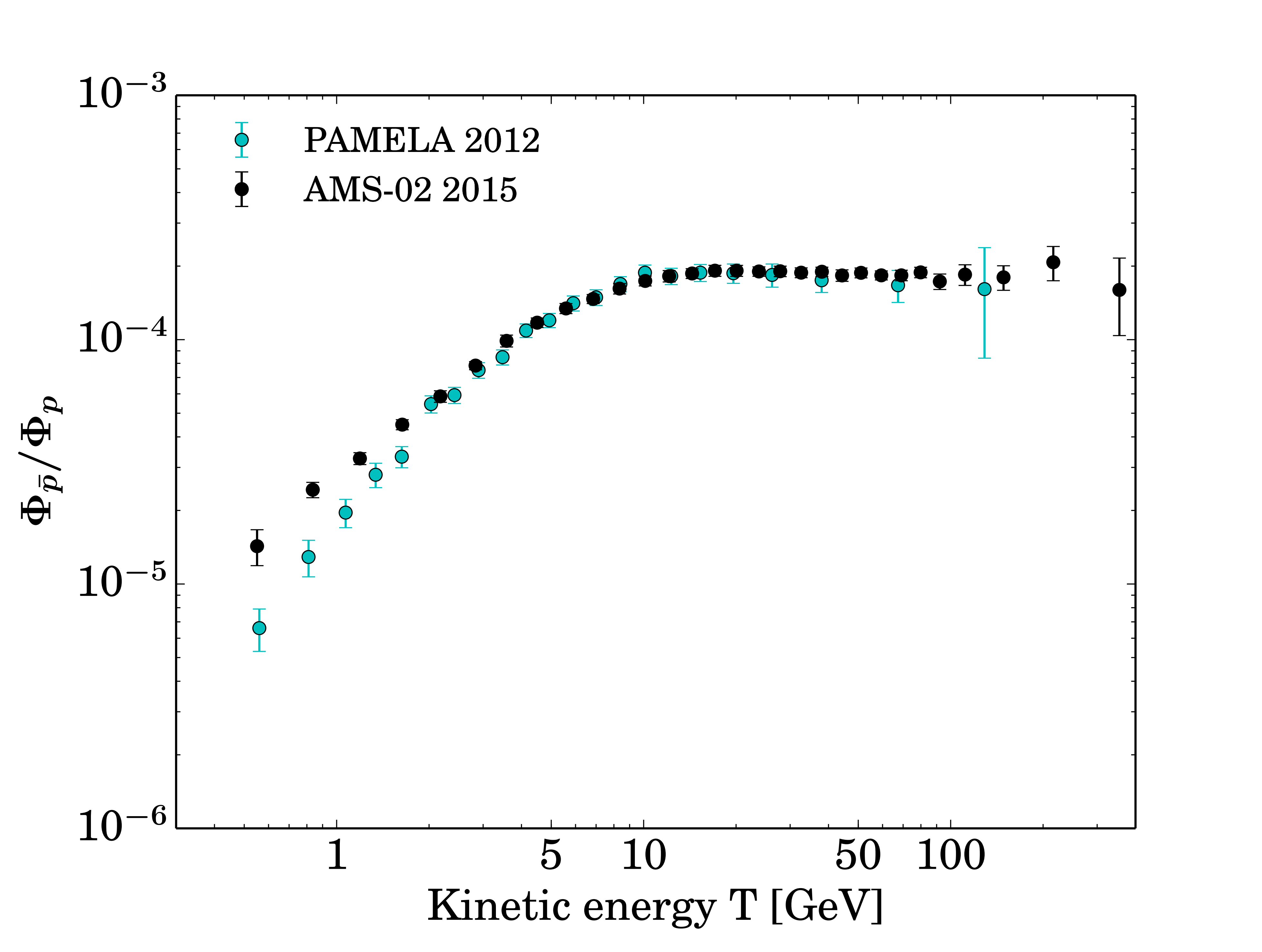

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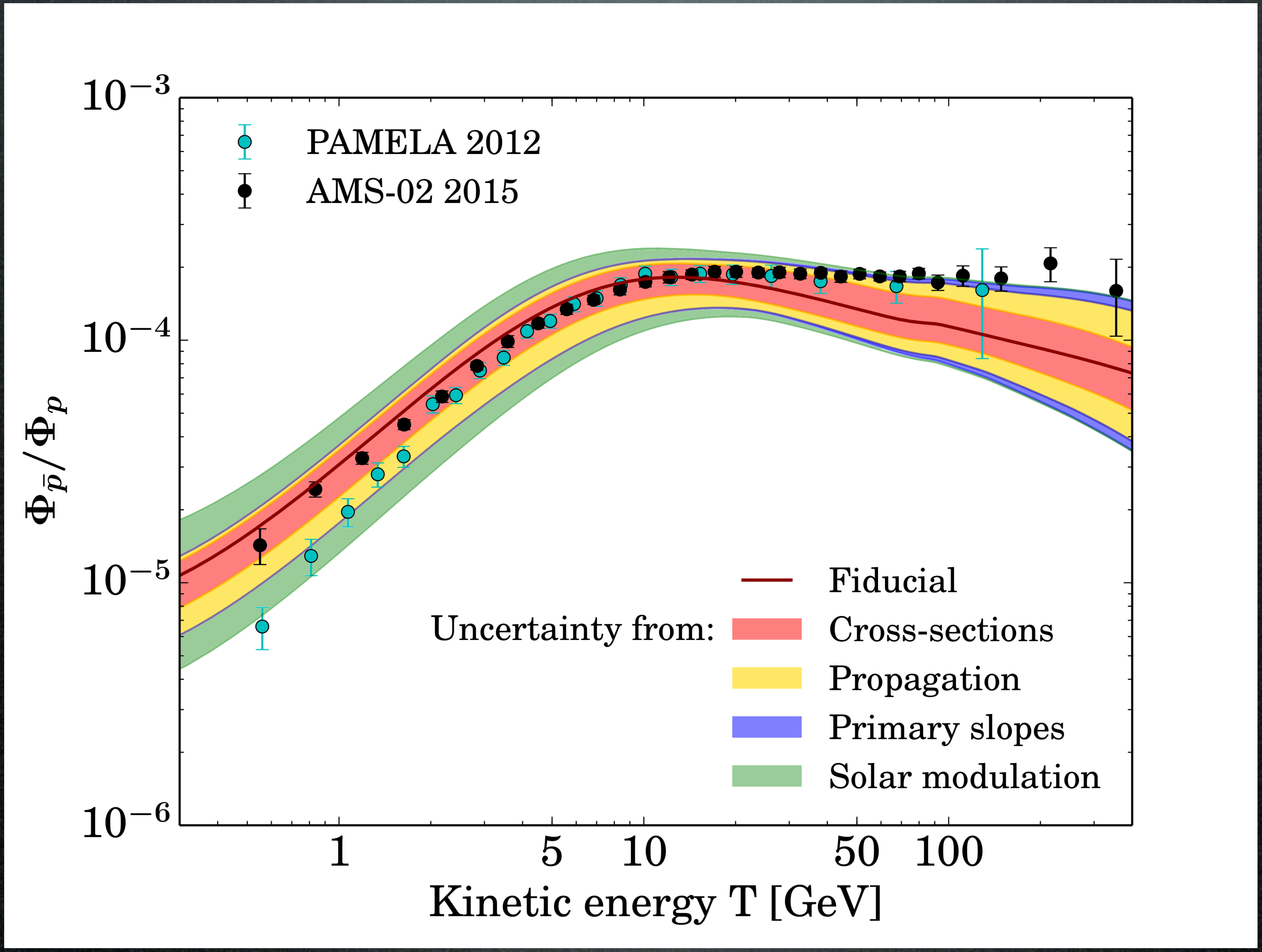

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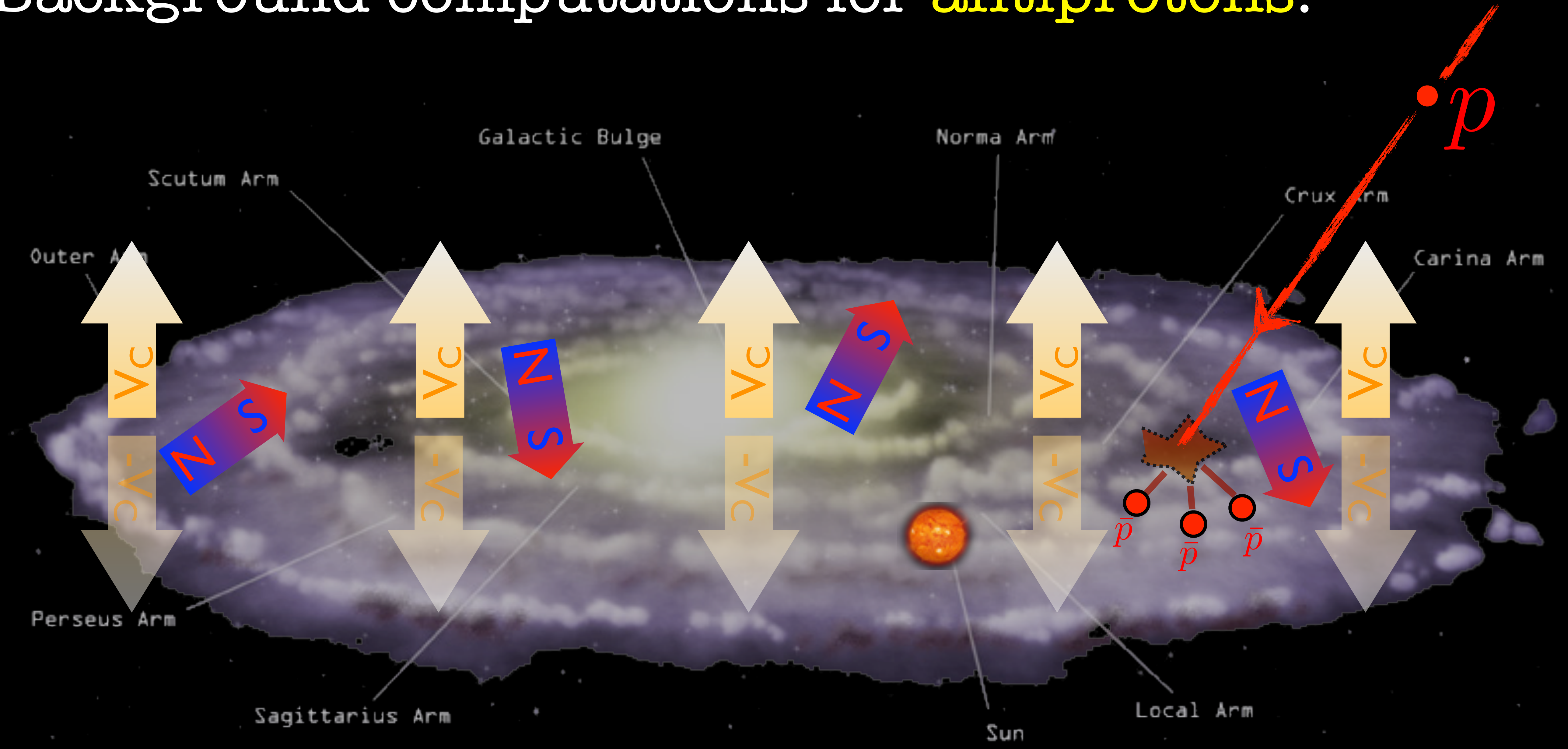
Indirect Detection

Background computations for antiprotons:



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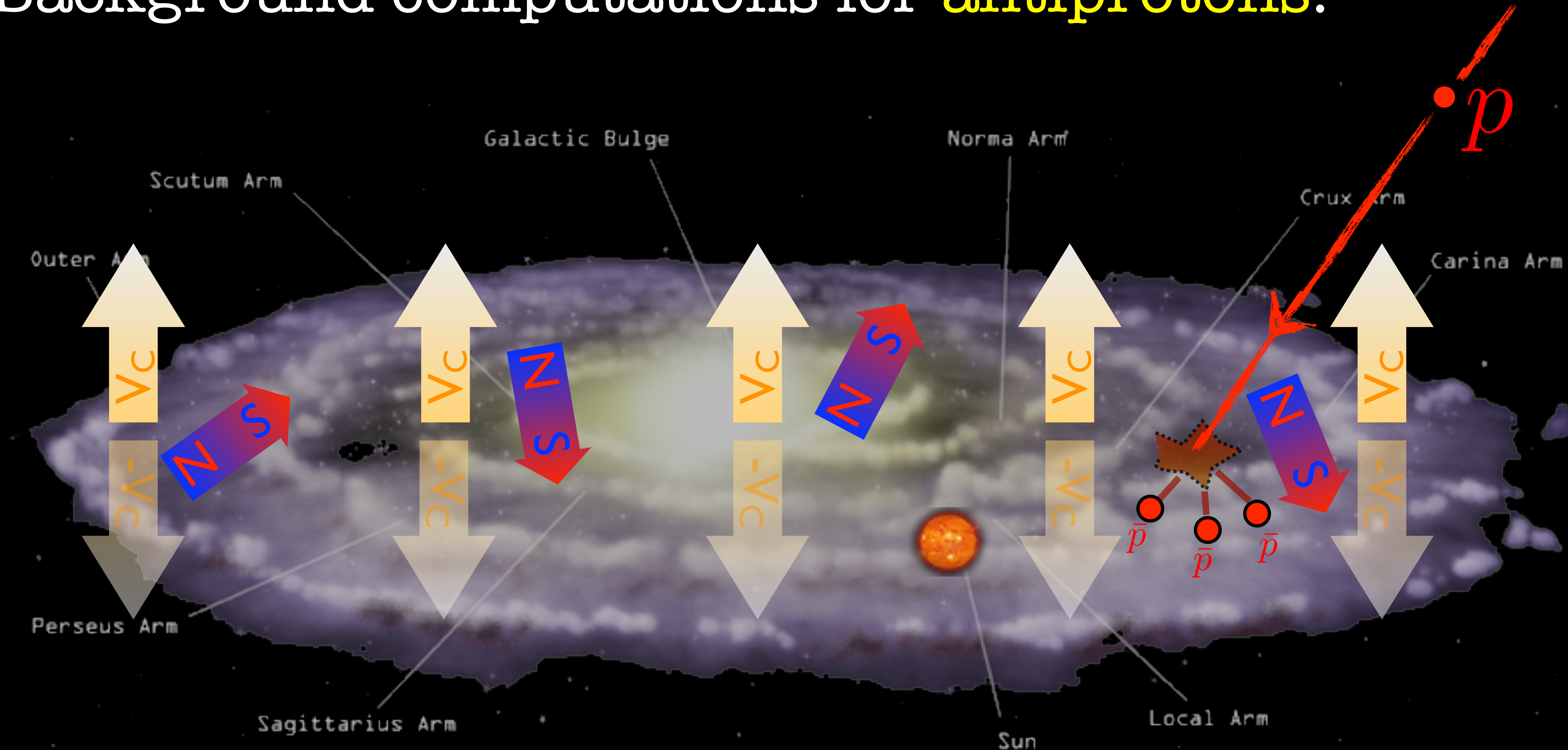
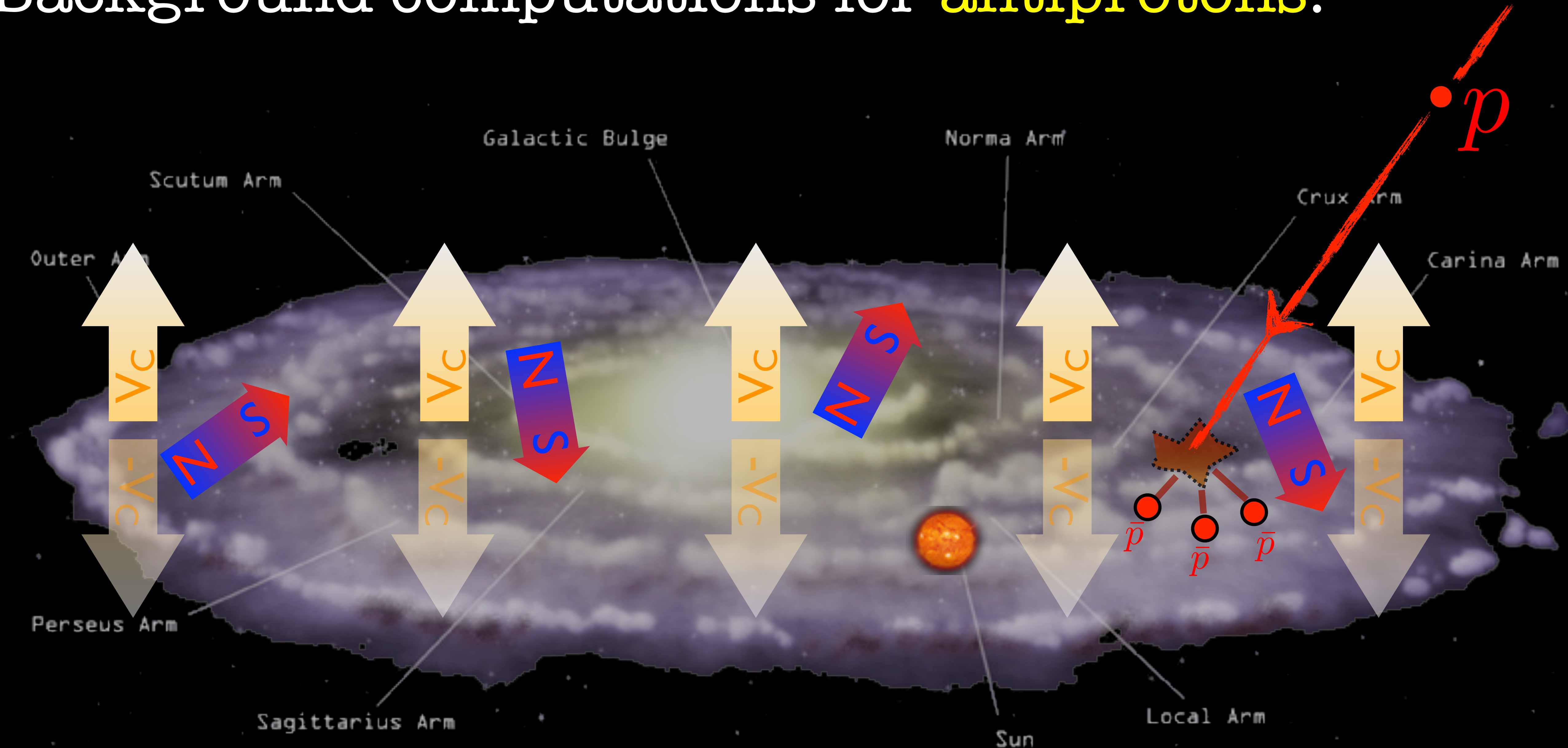


Main ingredients:

- primary p (and He)
- spallation cross-sections $\sigma_{p\mathrm{H}\to\bar{p}\mathrm{X}}, \sigma_{p\mathrm{He}\to\bar{p}\mathrm{X}}, \sigma_{\mathrm{HeH}\to\bar{p}\mathrm{X}}, \sigma_{\mathrm{HeHe}\to\bar{p}\mathrm{X}}$
- propagation
- solar modulation

Indirect Detection

Background computations for antiprotons:

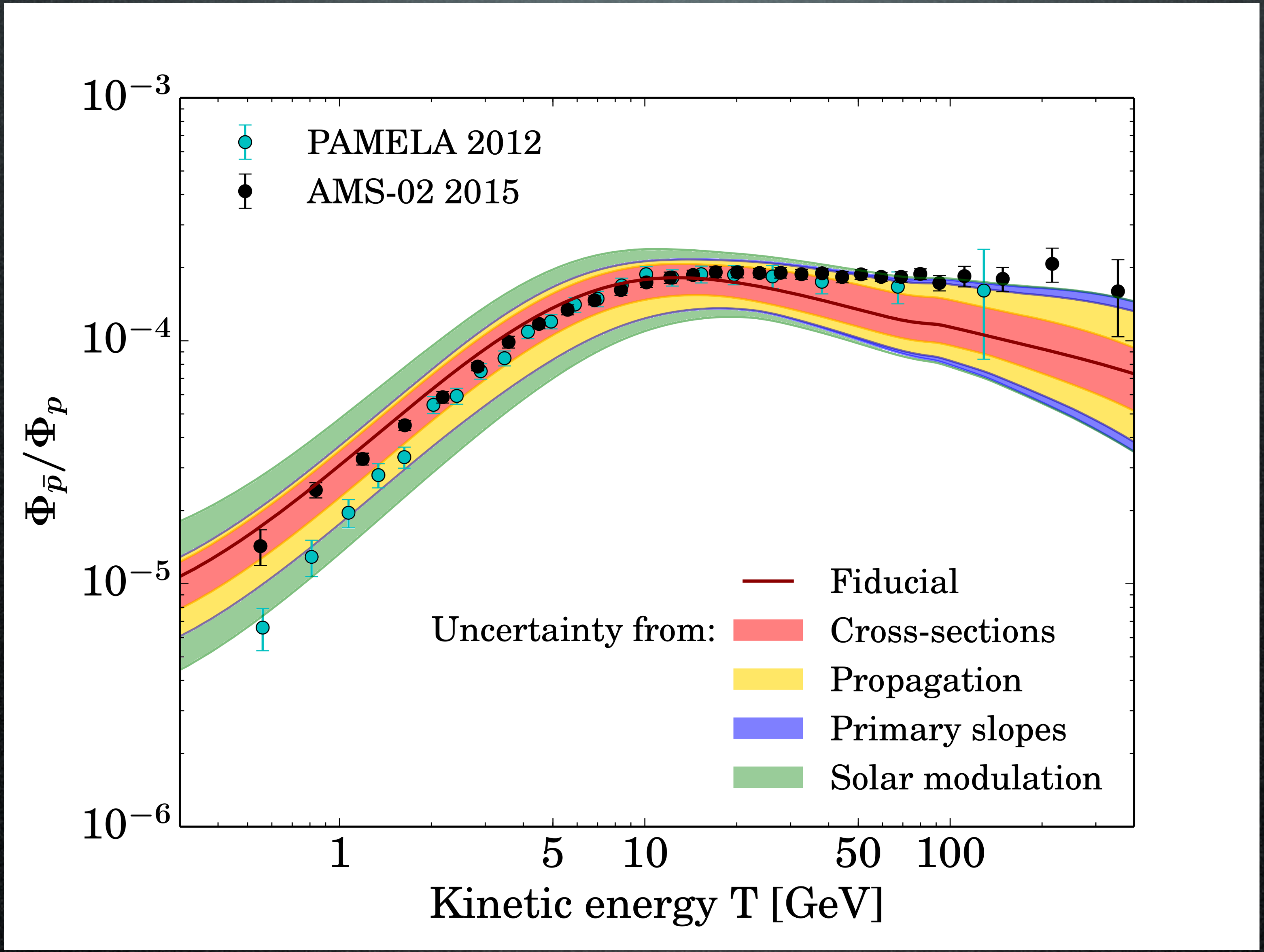


Main ingredients:

- primary p (and He) **New!**
- spallation cross-sections $\sigma_{pH\to\bar{p}X}$, $\sigma_{pHe\to\bar{p}X}$, $\sigma_{HeH\to\bar{p}X}$, $\sigma_{HeHe\to\bar{p}X}$ **New!**
- propagation
- solar modulation

Antiprotons

Antiproton data vis-à-vis the secondaries:



No evident **excess**

Giesen, Boudaud, Génolini, Poulin, Cirelli, Salati, Serpico
1504.04276

Antiprotons

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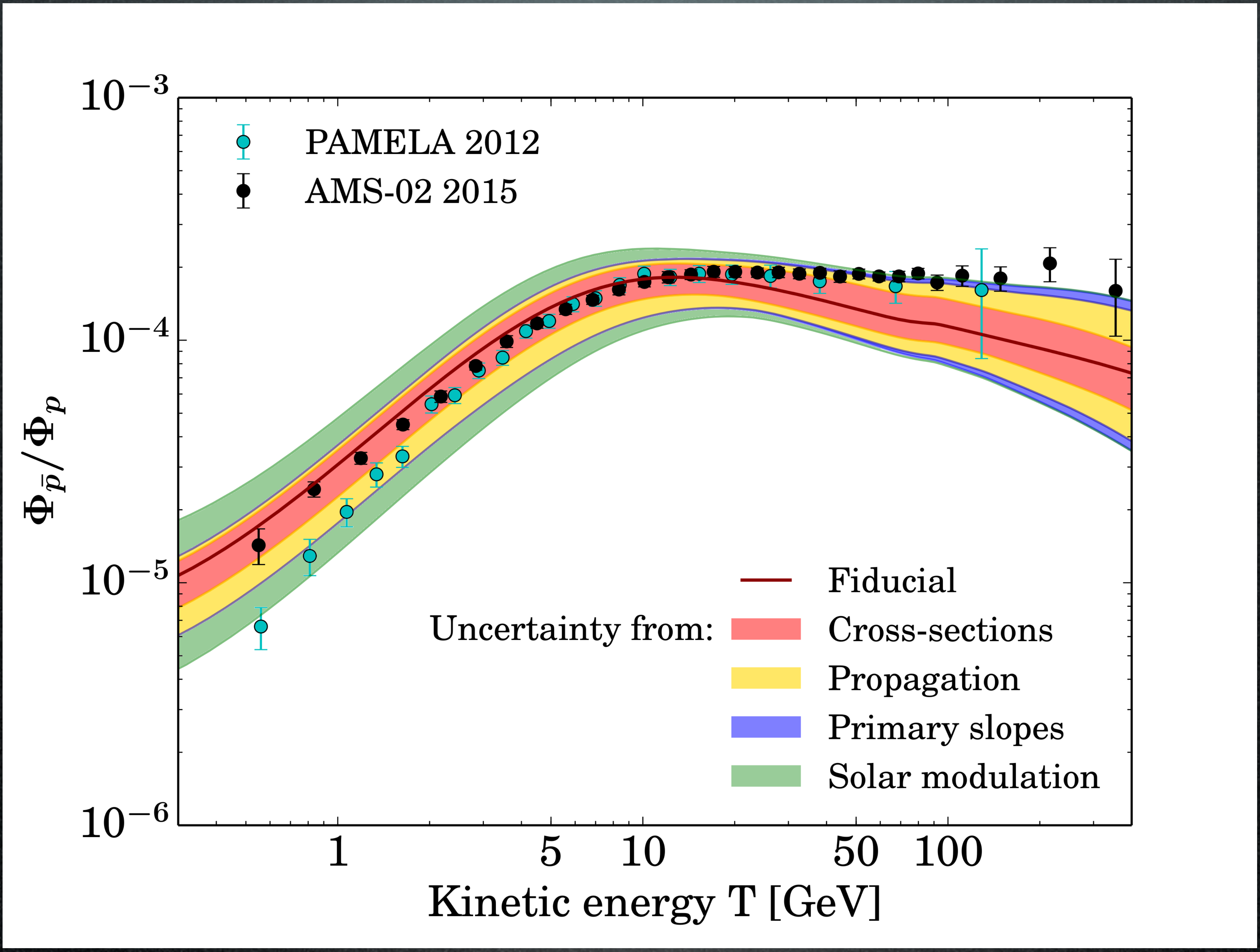


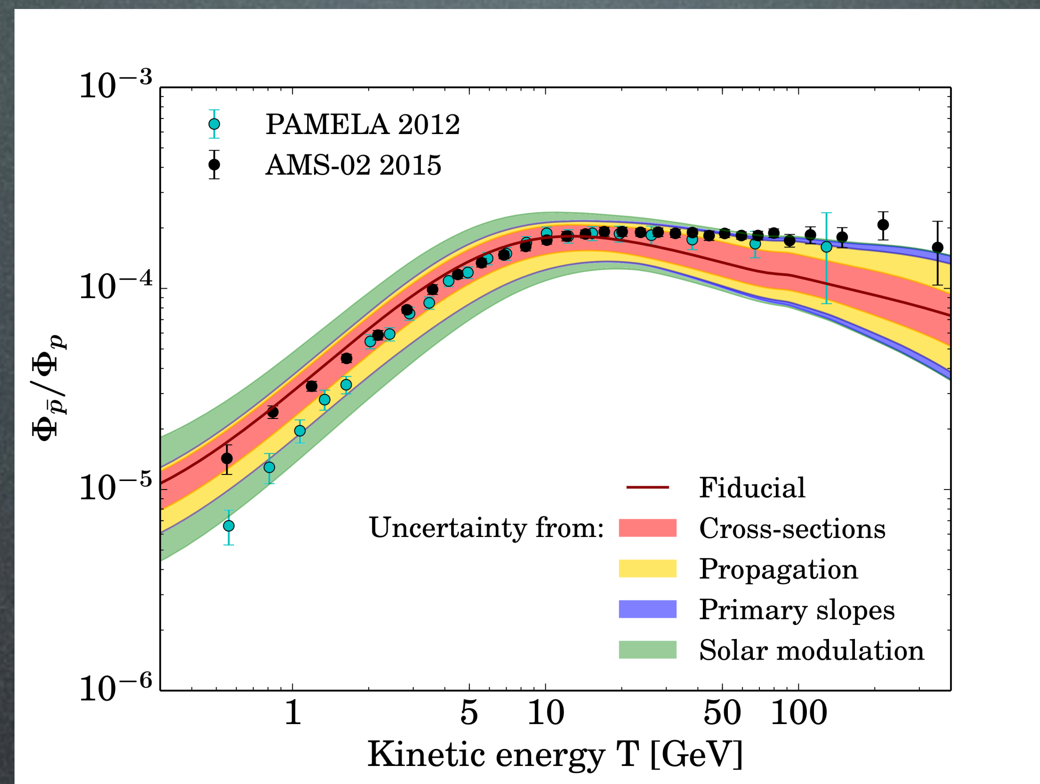
No evident **excess**

Some preference for flatness

Giesen, Boudaud, Génolini, Poulin, Cirelli, Salati, Serpico 1504.04276

Antiprotons

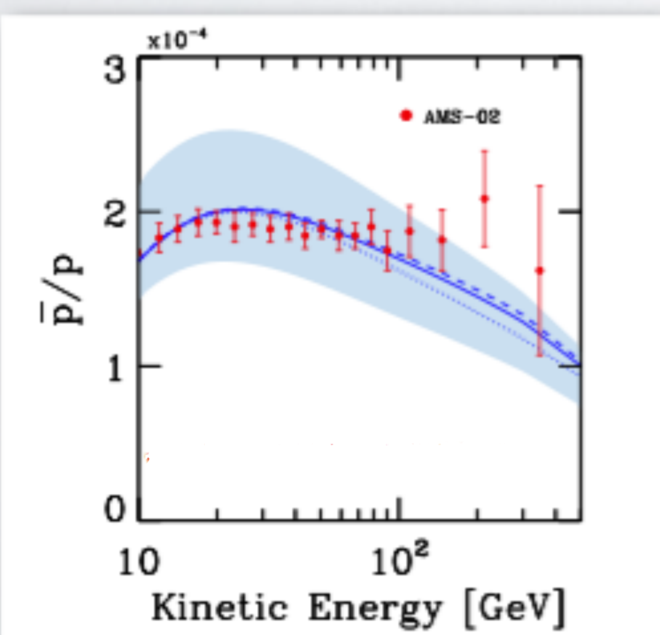
Antiproton data vis-à-vis the secondaries:



Giesen, Boudaud, Génolini, Poulin, Cirelli, Salati, Serpico 1504.04276

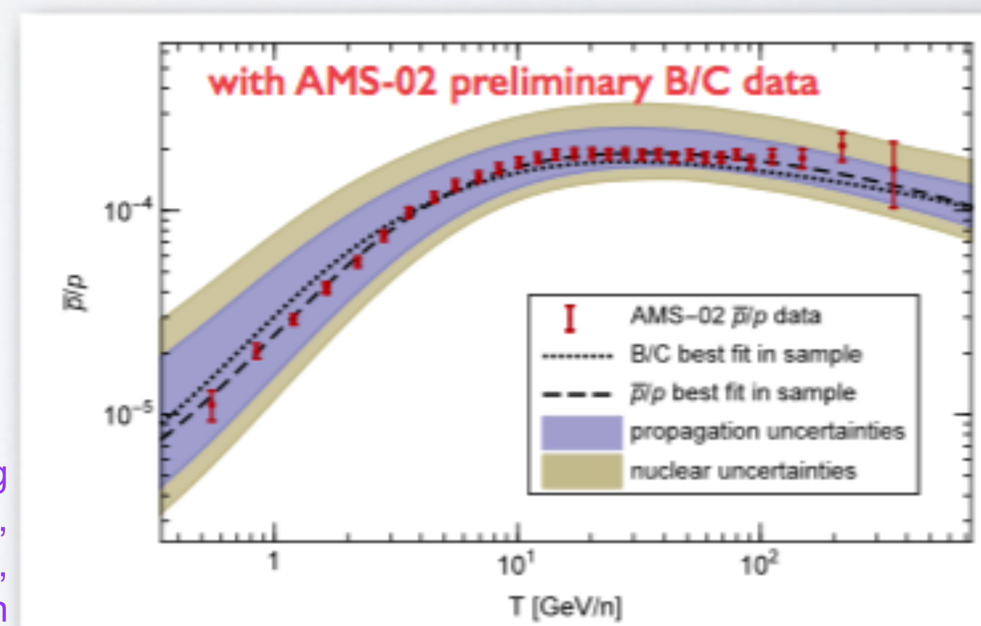
using
p, He by AMS-02,
B/C by PAMELA

C. Evoli, D. Gaggero and D. Grasso, arXiv:1504.05175



using
p, He, B/C
by AMS-02

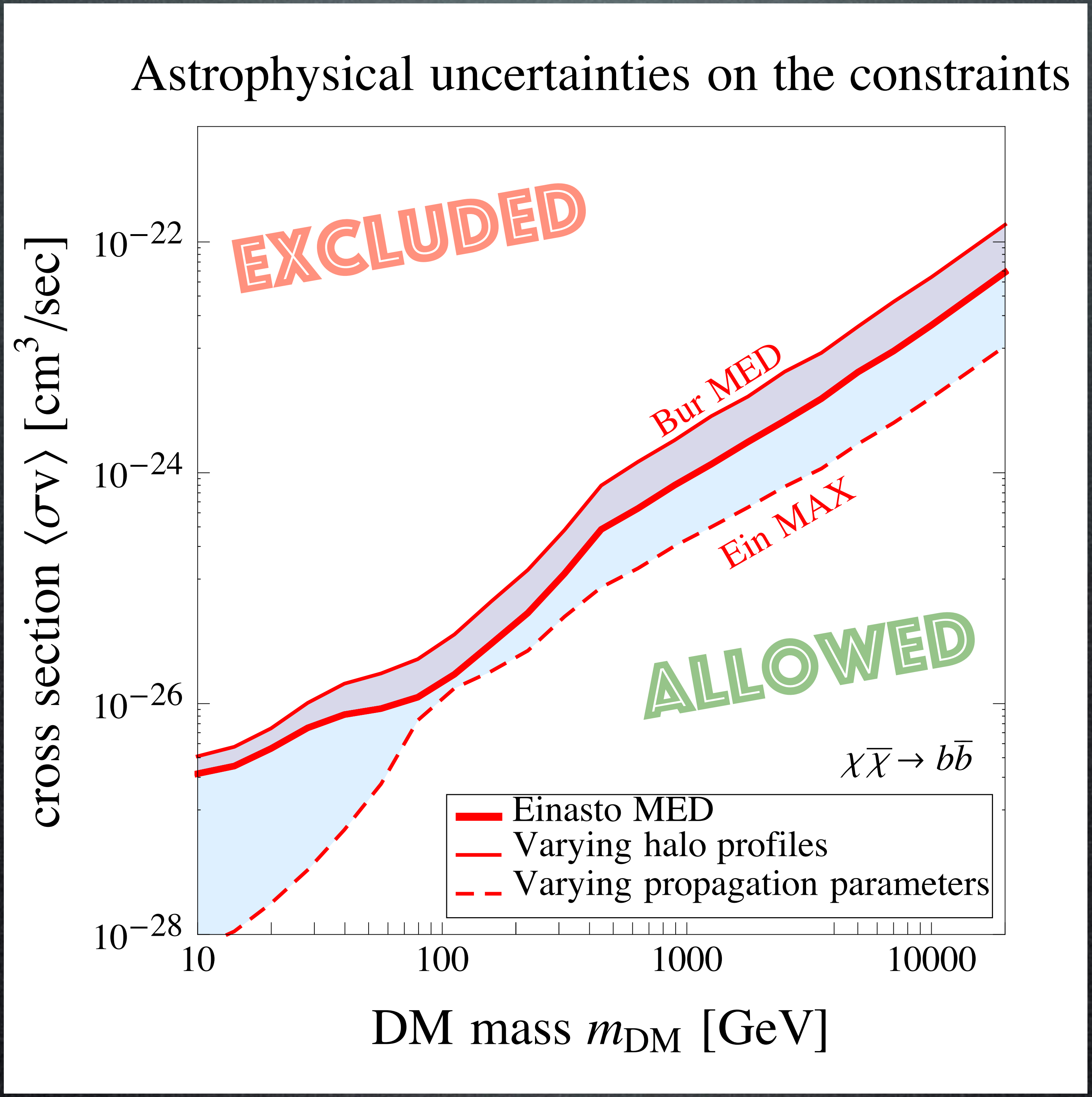
R. Kappl, A. Reinert and M.W. Winkler, arXiv:1506.04145



using
p, He by AMS-02 and CREAM,
B/C by AMS-02,
heavier nuclei by compilation

Dark Matter interpretation

Based on AMS-02 $\bar{p}/p$ data (april 2015)

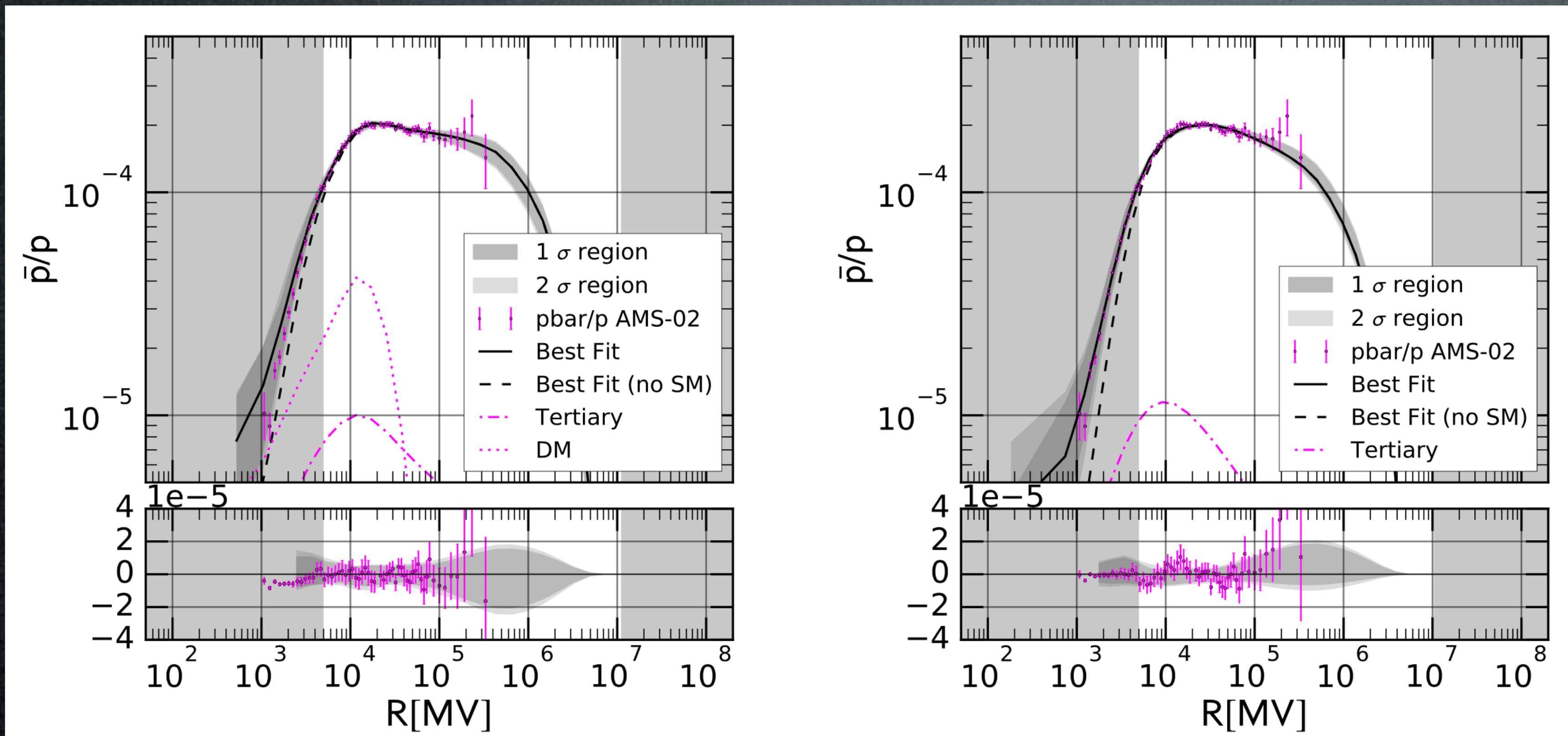


Giesen, Boudaud, Génolini, Poulin, Cirelli, Salati, Serpico 1504.04276

Antiprotons

Recent developments

Cuoco, Krämer, Korsmeier 1610.03071



finds a possible excess

(formally ~4.5σ)

m_{DM} = 80 GeV, bb, thermal cross-section

similarly:

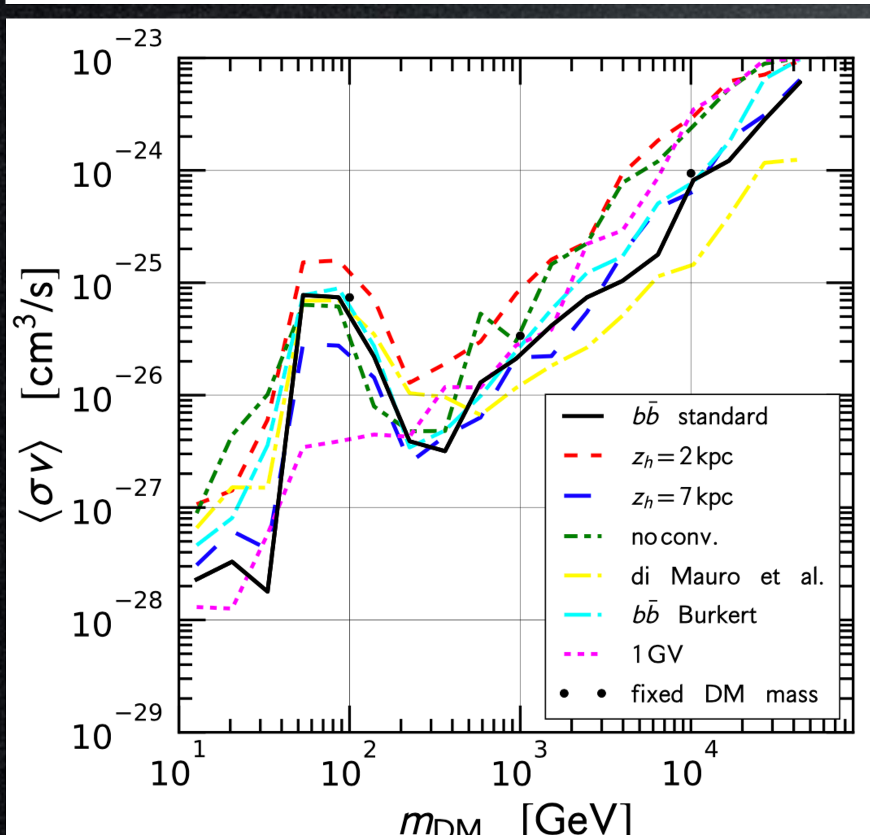
Cui, Yuan, Tsai, Fang 1610.03840

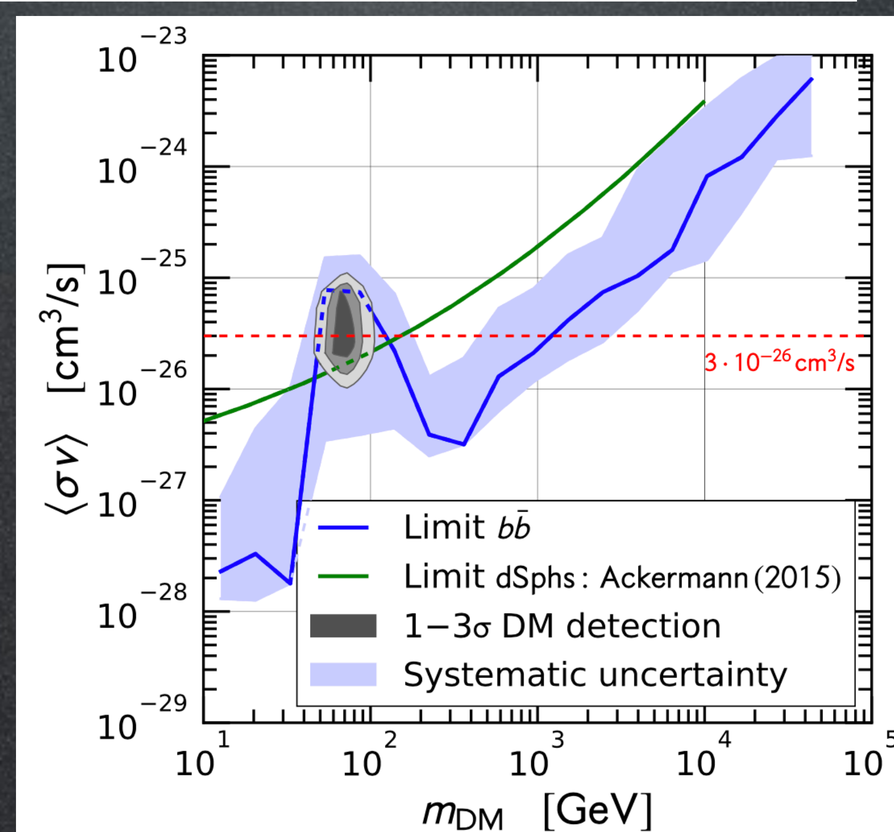
Huang + 1611.01983 (light mediators)

Feng, Zhang 1701.02263

Cuoco, Heisig, Krämer, Korsmeier 1704.08258

Boschini+ (Galprop) 1704.06337 (but only 1σ)





reiterated:

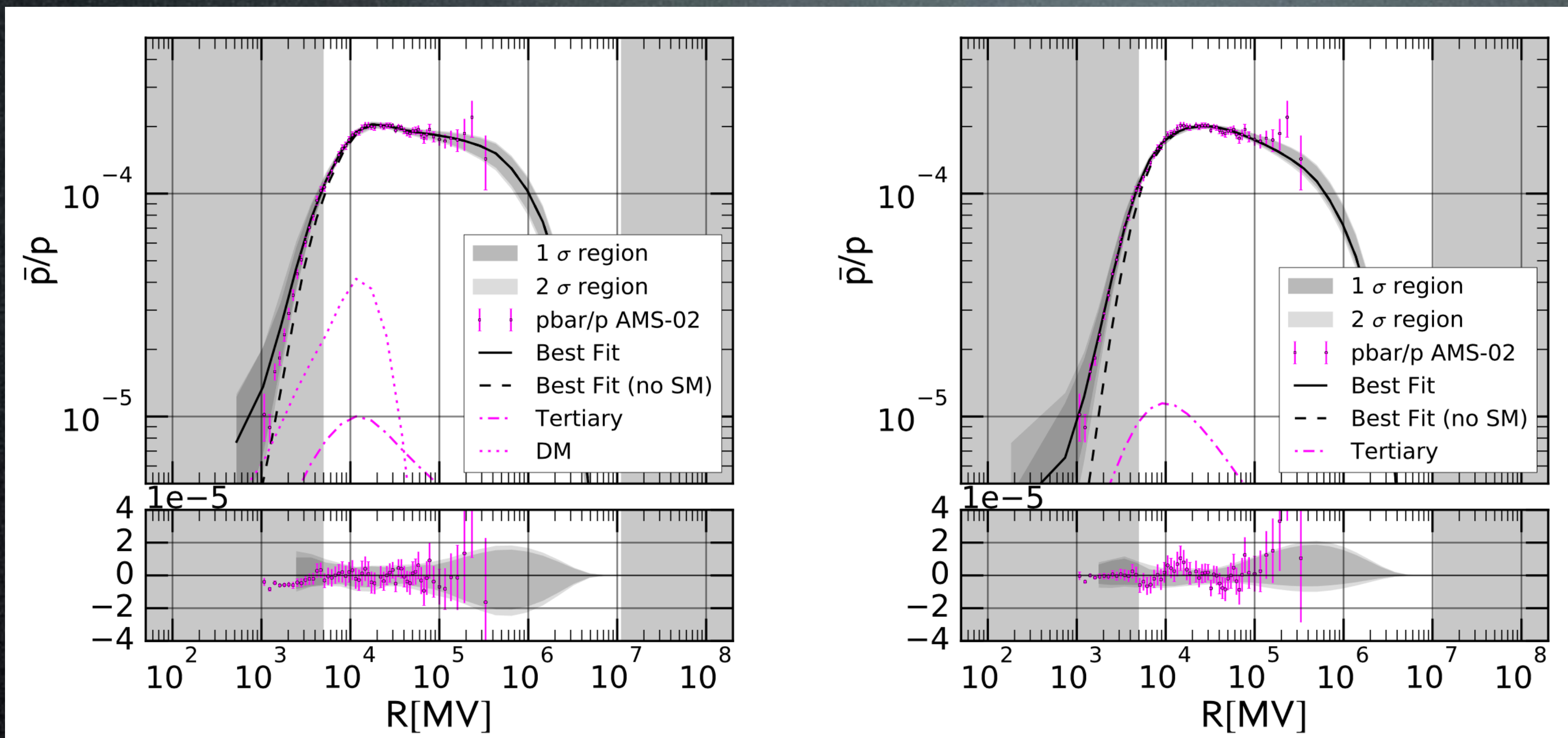
Cuoco, Heisig, K³ 1903.01472

Cholis, Linden, Hooper 1903.02549

Antiprotons

Recent developments

Cuoco, Krämer, Korsmeier 1610.03071



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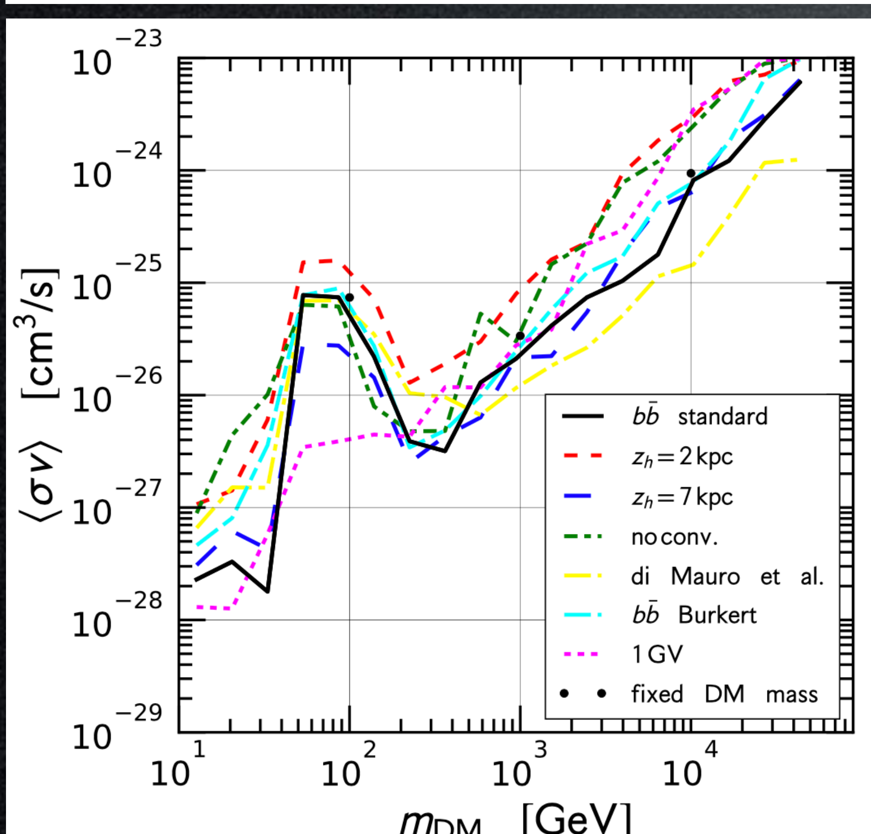
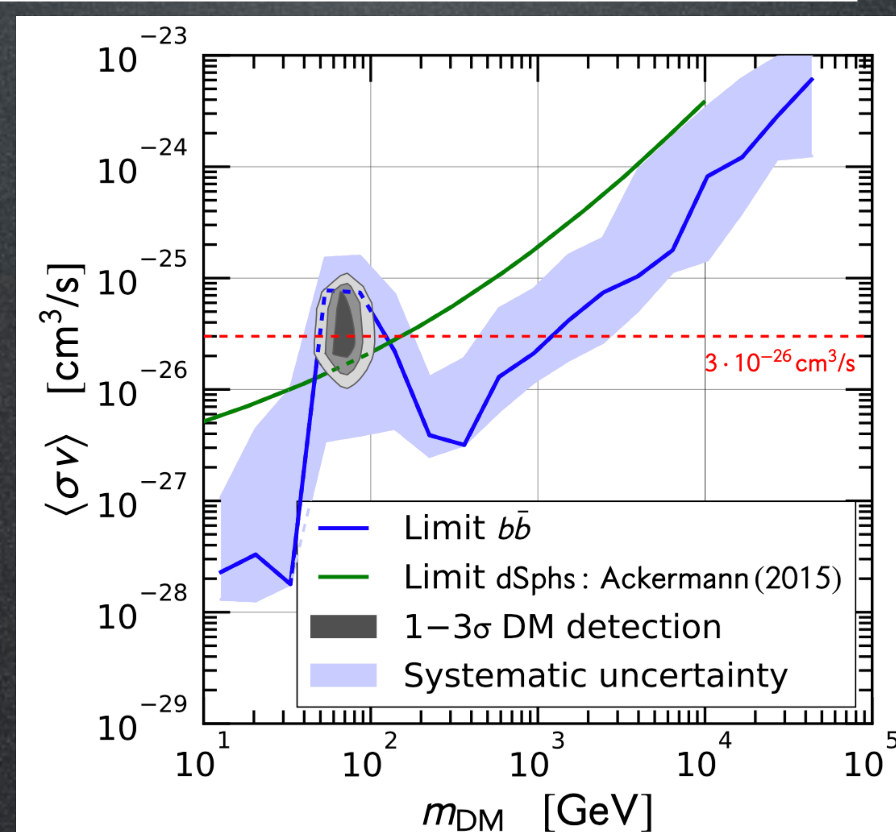




criticisms:

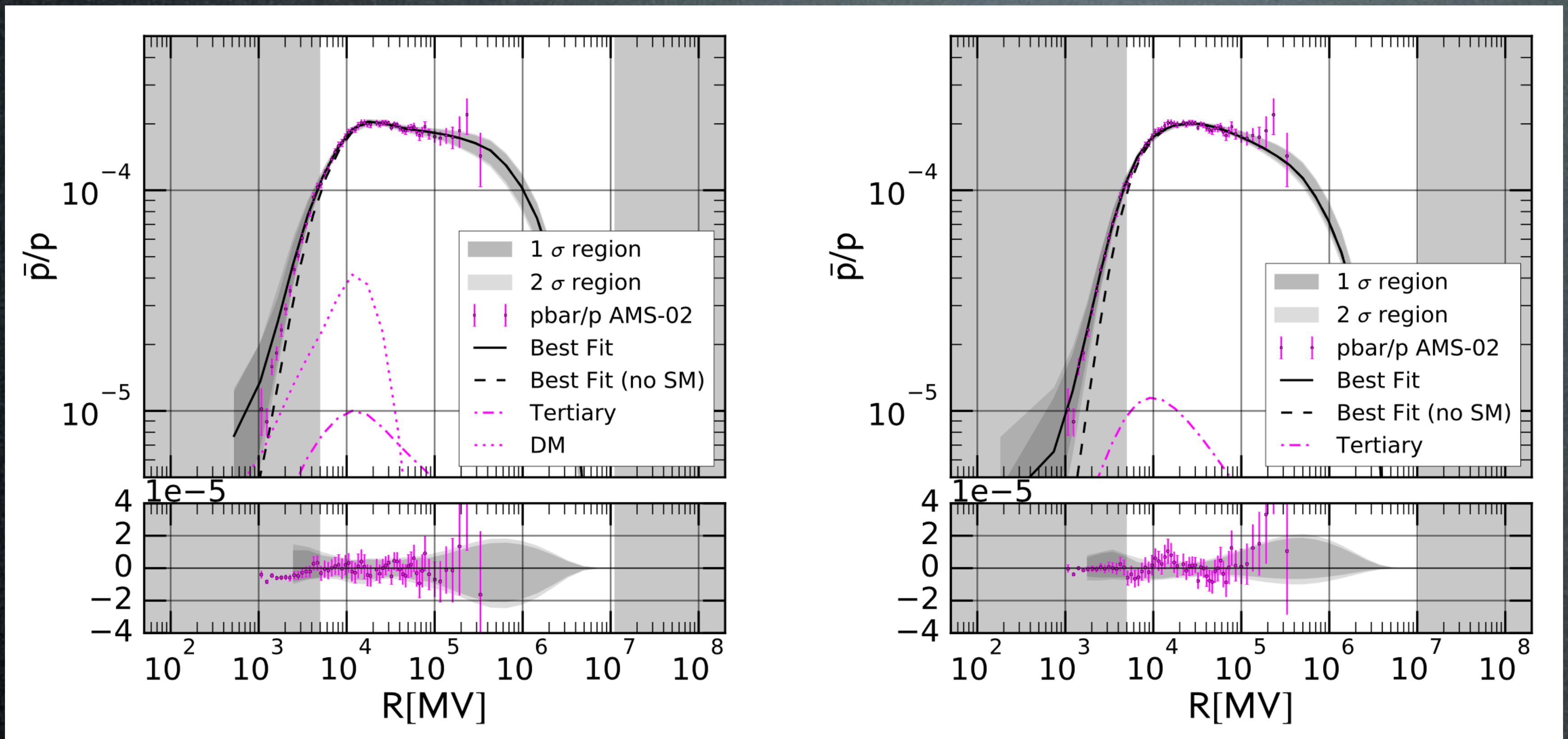
propagation parameters determined with p, He data only, w/o B/C

excess evaporates including low energies

Antiprotons

Recent developments

Cuoco, Krämer, Korsmeier 1610.03071



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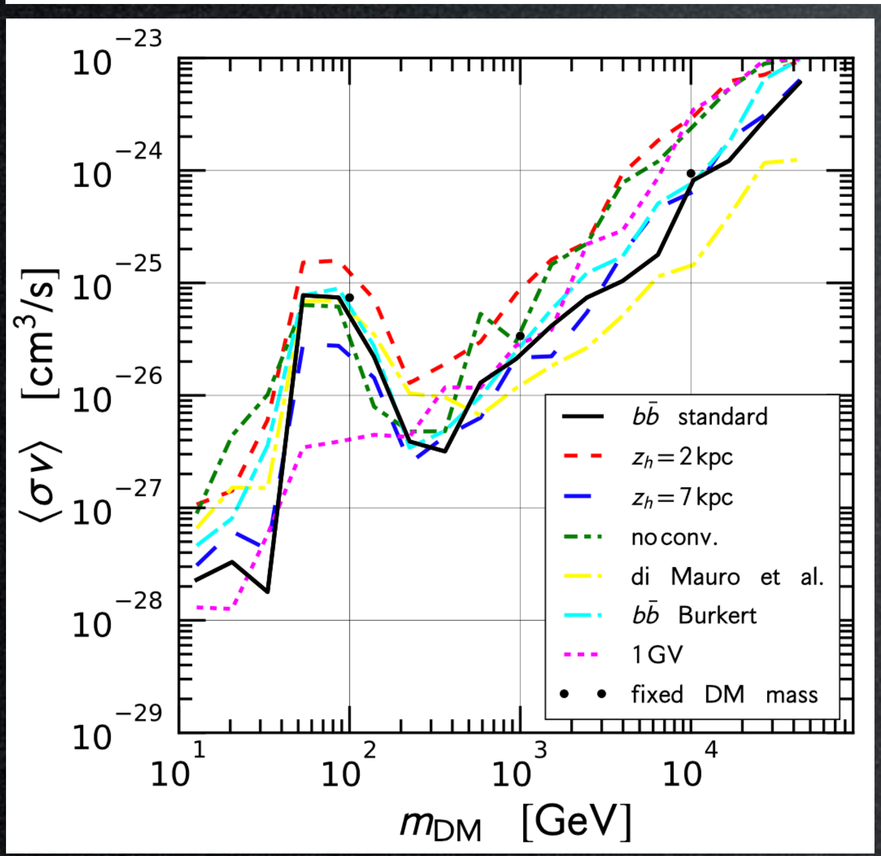
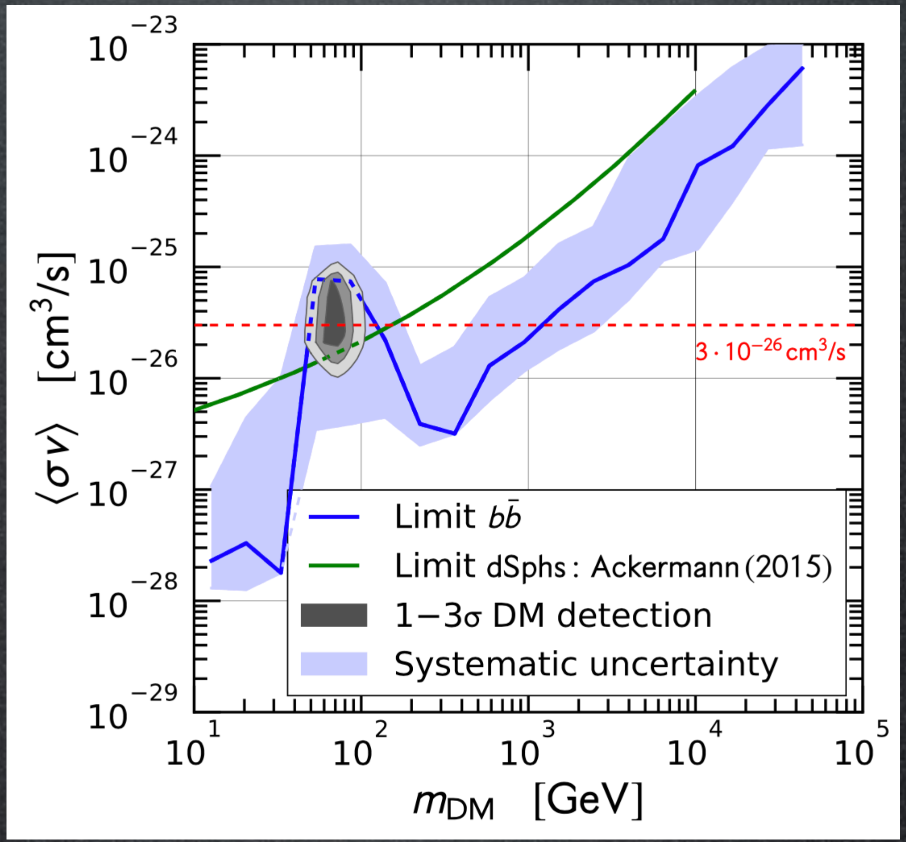




on the other hand:

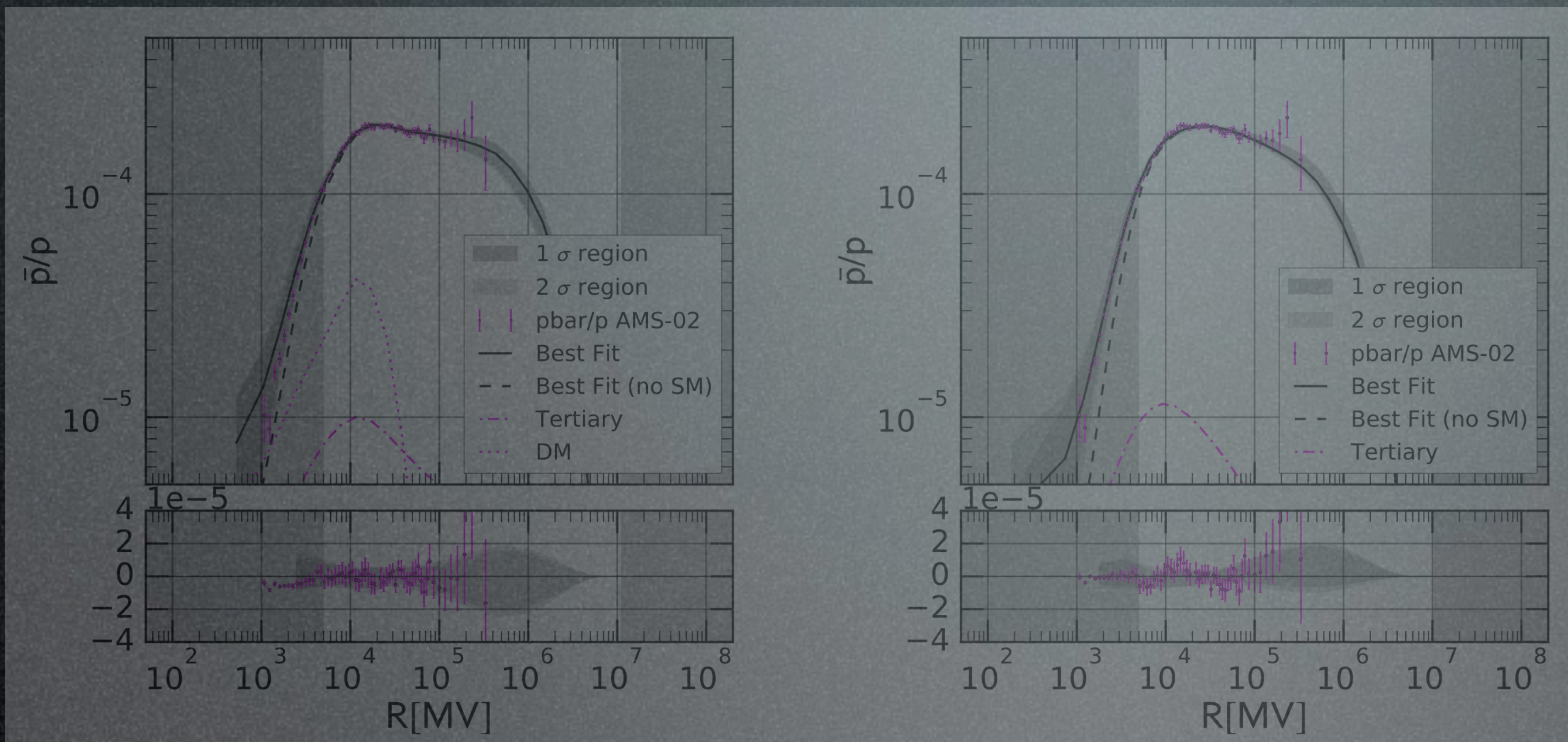
B/C and p probably probe different regions

it's a very tricky region, cool things can hide there

Antiprotons

Recent developments

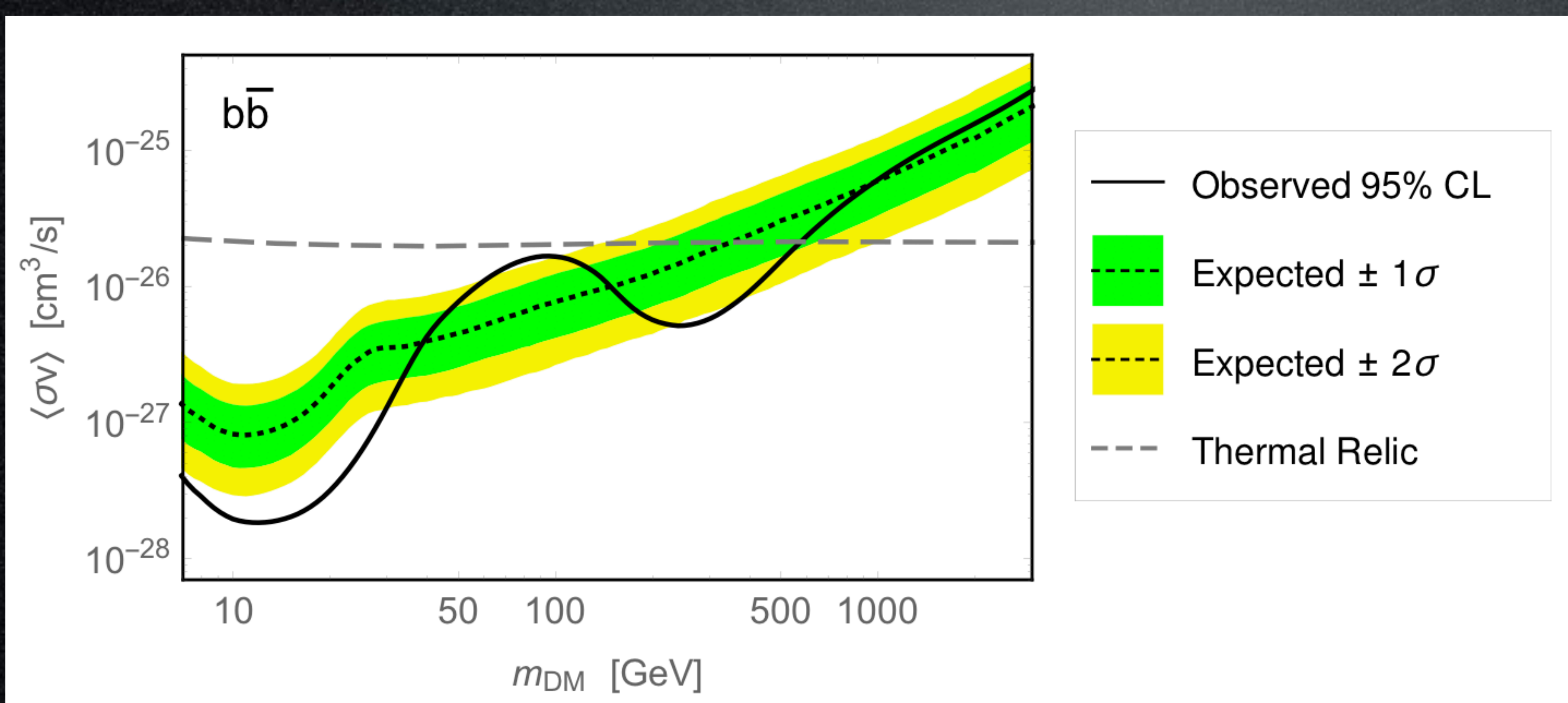
Cuoco, Krämer, Korsmeier 1610.03071



finds a possible excess

m_{DM} = 80 GeV, bb,
thermal cross-section

similarly:
Cui, Yuan, Tsai, Fang 1610.03840
Huang + 1611.01983 (light mediators)
Feng, Zhang 1701.02263
Cuoco, Heisig, Krämer, Korsmeier 1704.08258
Boschini+ (Galprop) 1704.06337 (but only 1σ)



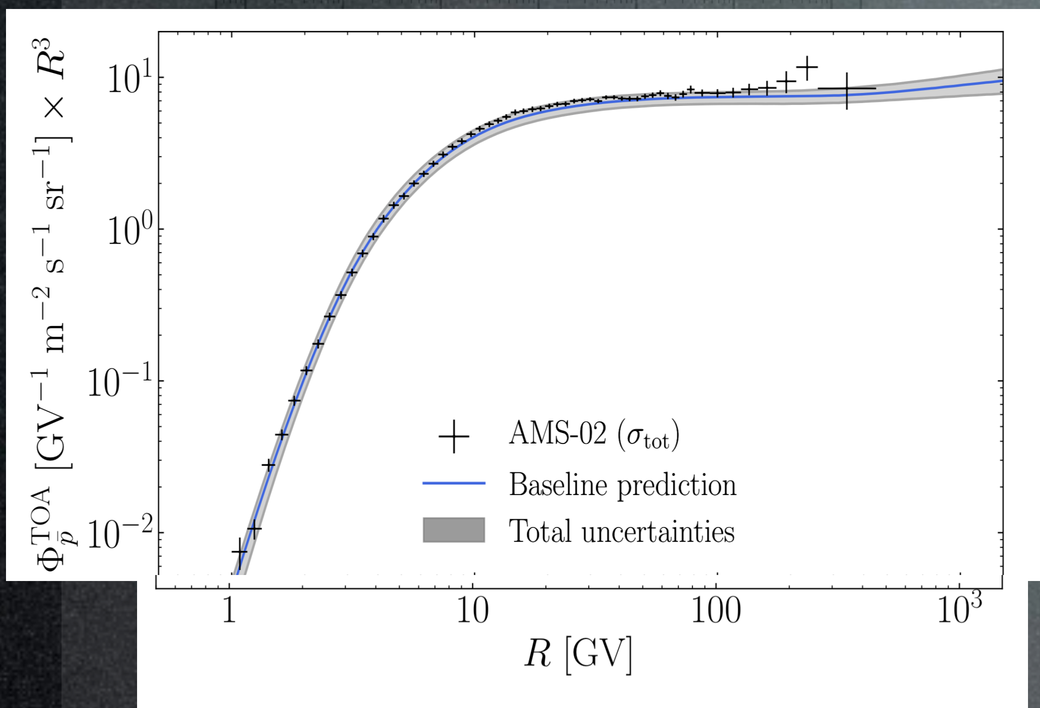
Reinert, Winkler 1712.00002

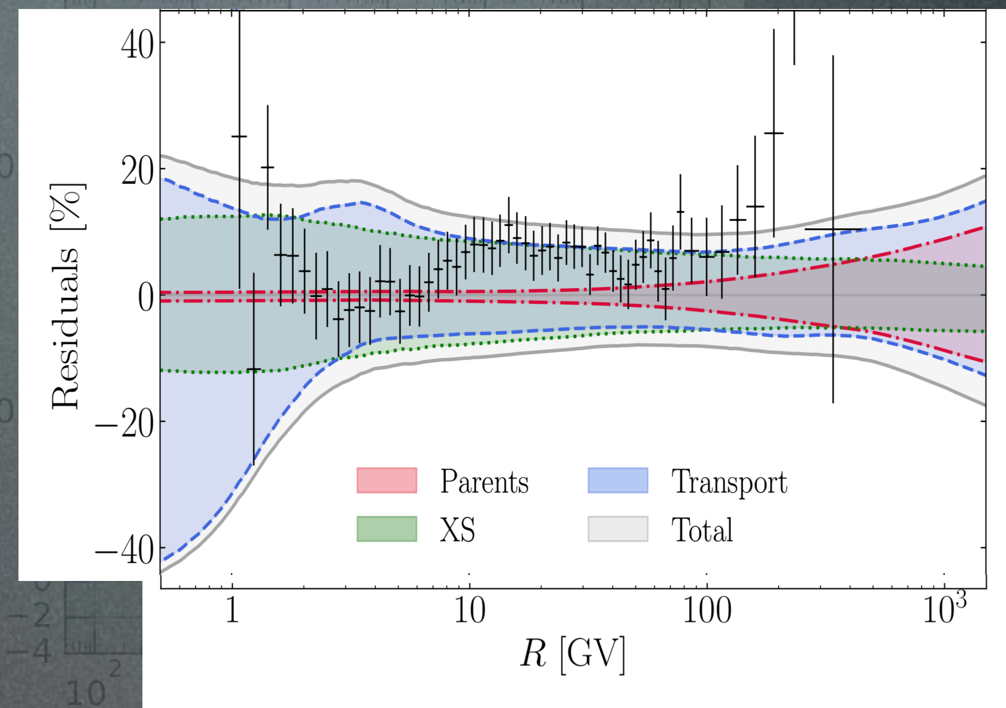
excess exists

but significance ~1σ,
given all uncertainties

Antiprotons

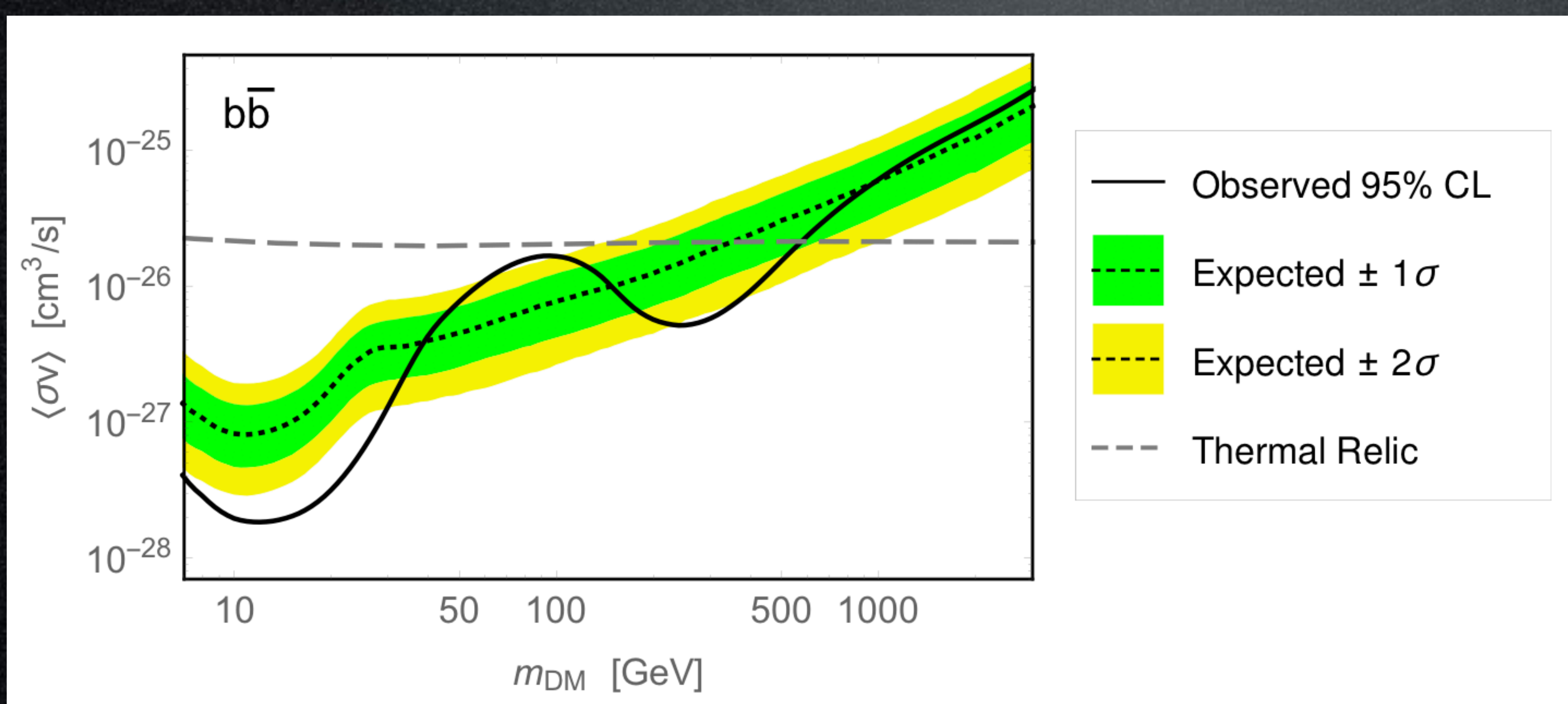
Recent developments





Boudaud et al. 1906.0719

"antiprotons are consistent with a secondary astrophysical origin"



Reinert, Winkler 1712.00002

excess exists

but significance ~1σ, given all uncertainties

Positrons (and electrons)

direct detection

production at colliders

indirect

γ from annihil in galactic center or halo
and from secondary emission

Fermi, ICT, radio telescopes...

e^+ from annihil in galactic halo or center

PAMELA, Fermi, HESS, AMS,balloons...

$\bar{p}$ from annihil in galactic halo or center

$\bar{d}$ from annihil in galactic halo or center

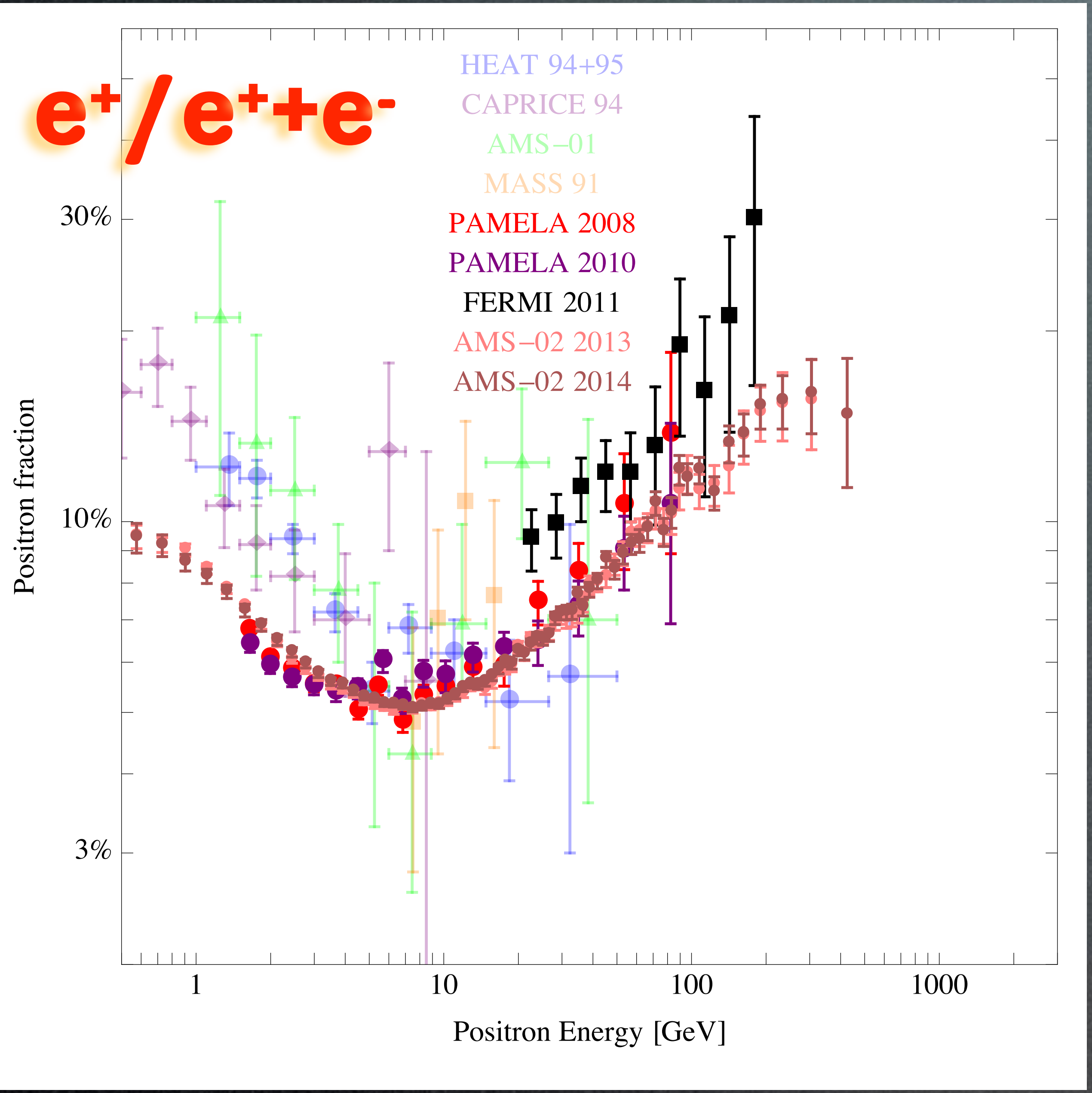
GAPS, AMS

$\nu, \bar{\nu}$ from annihil in massive bodies

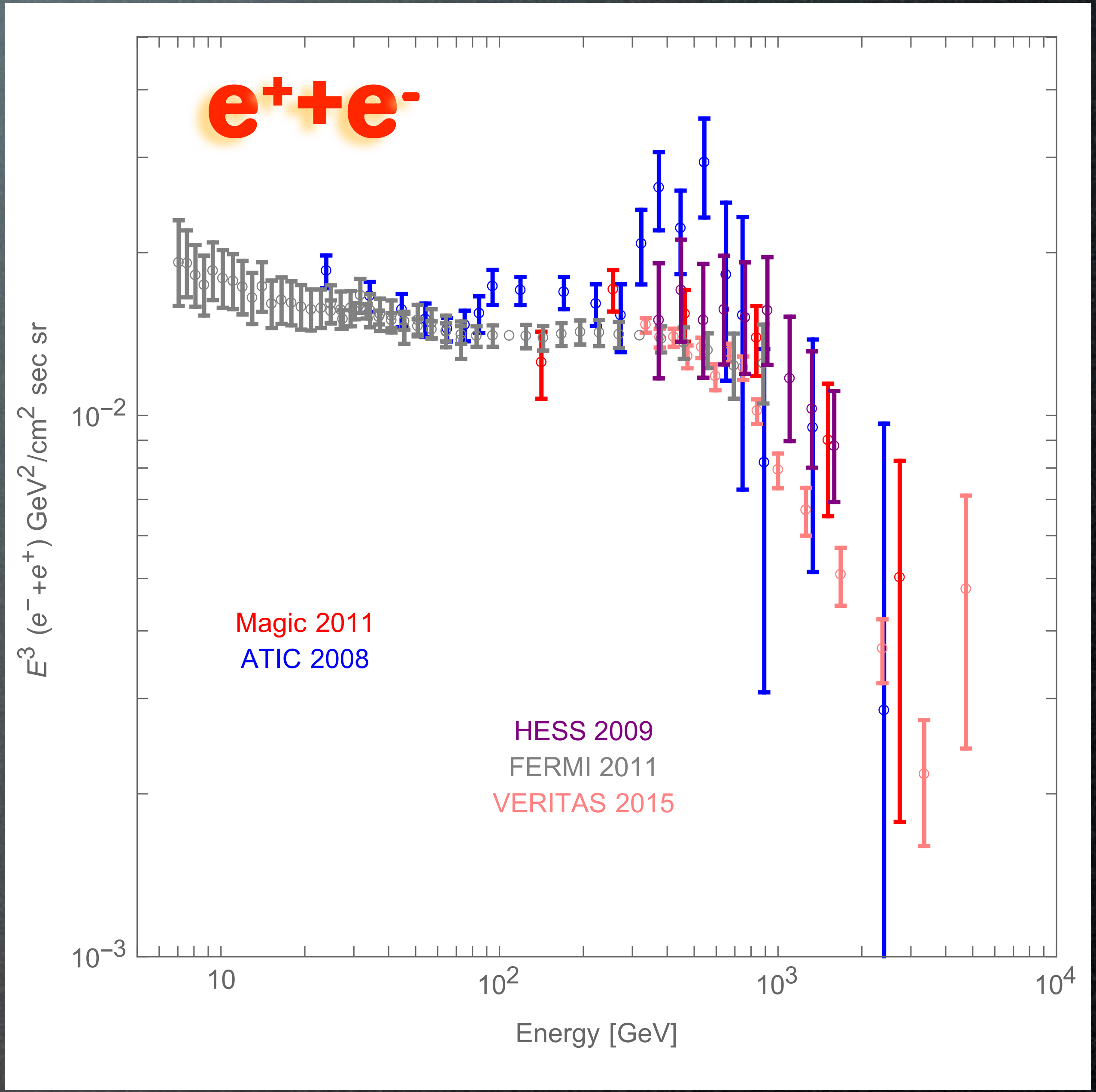
SK, Icecube, Km3Net

Data: leptons

high energy



M. Cirelli - compilation ICRC 2015



M. Cirelli - compilation

Data: leptons

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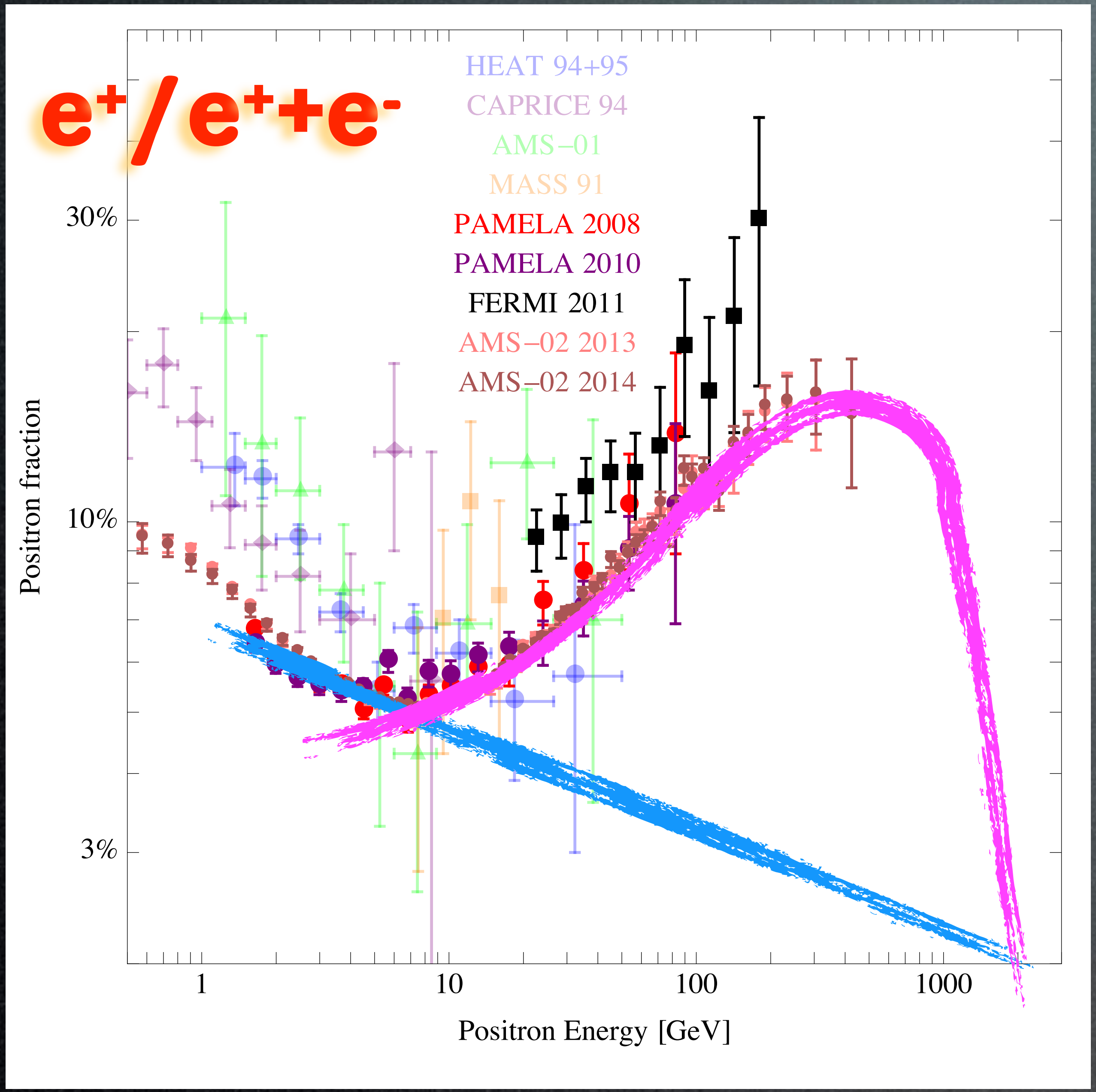

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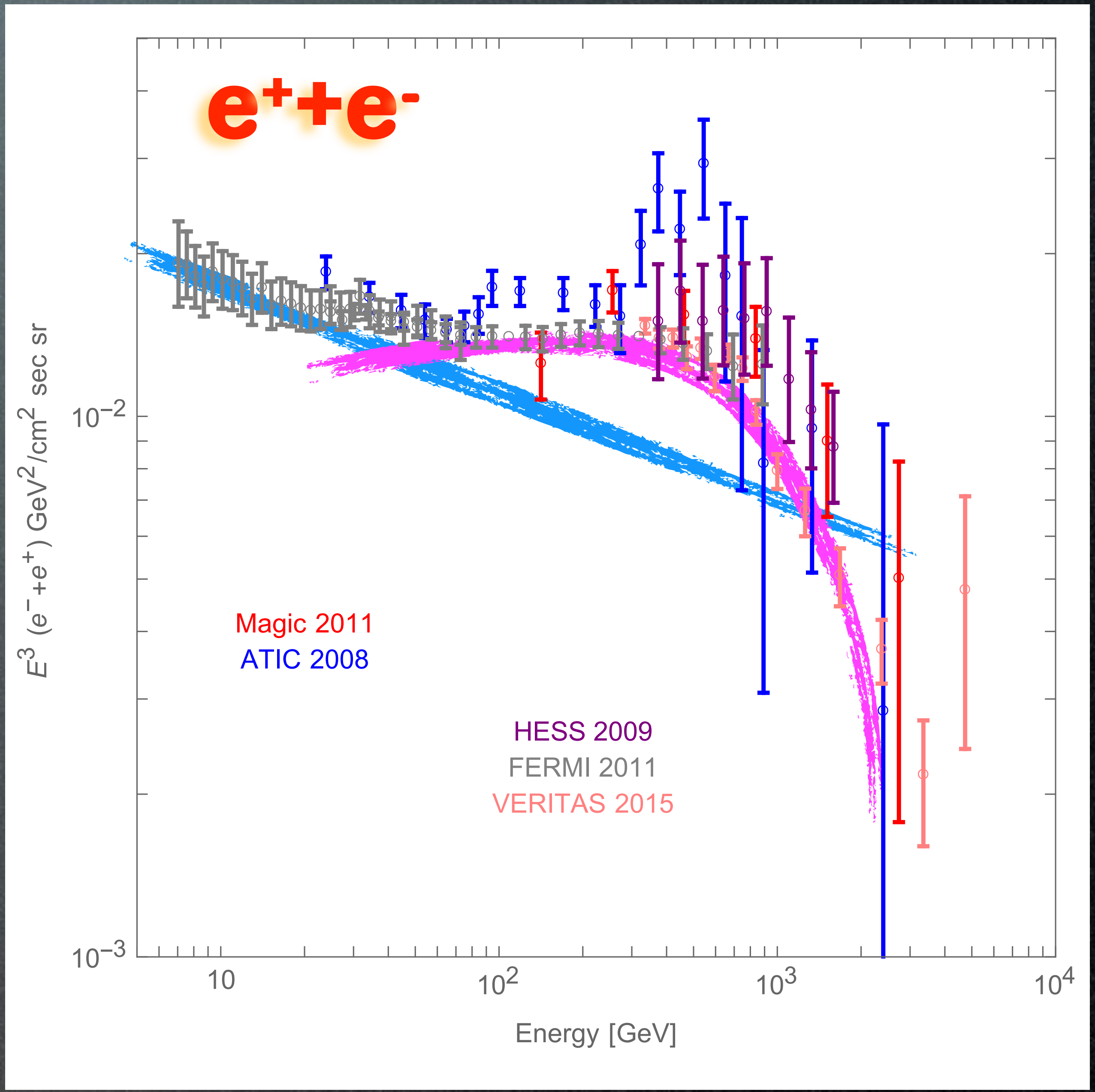
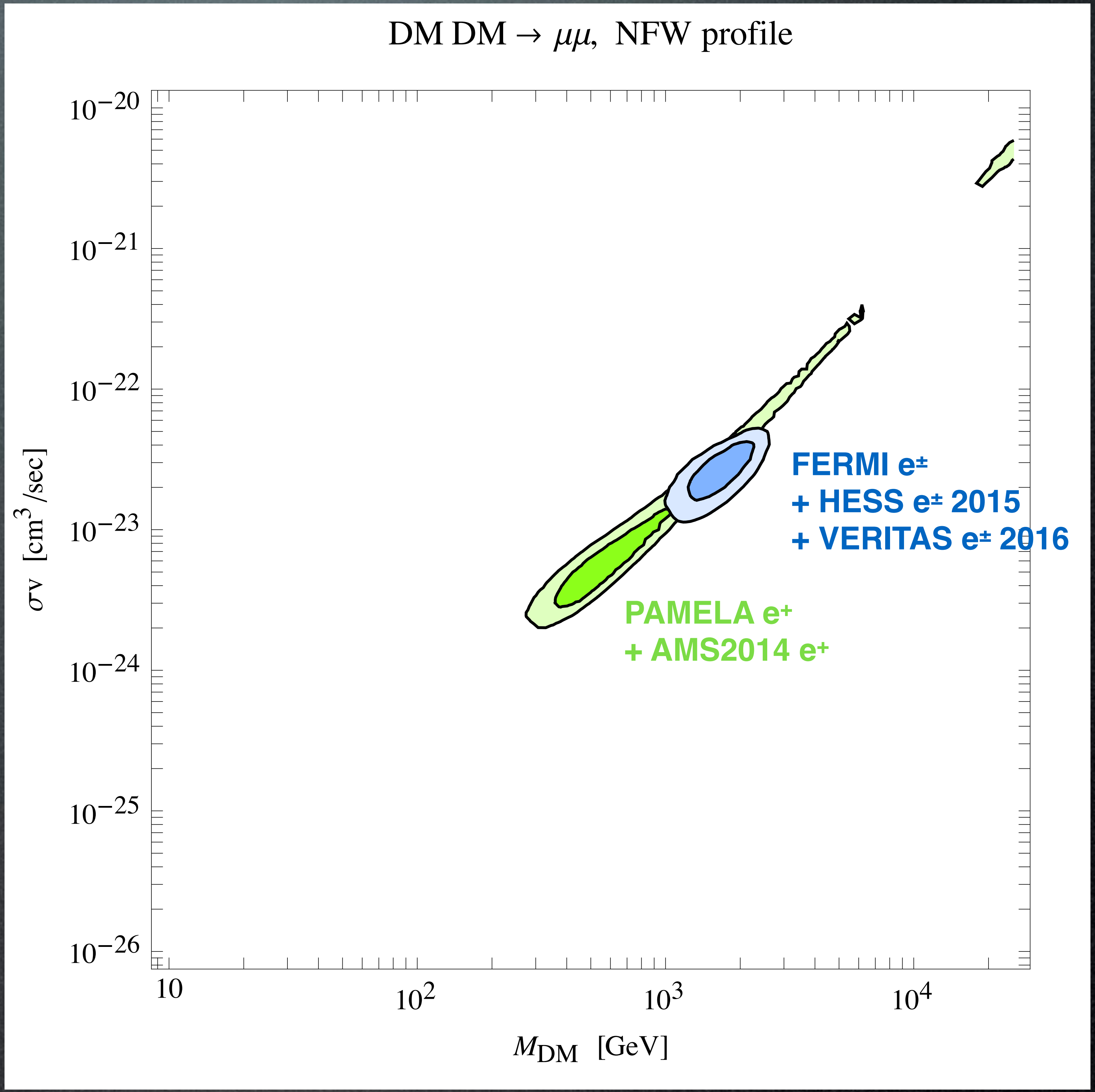


M. Cirelli - compilation ICRC 2015

Dark Matter interpretation

- leptophilic
- $m_{DM} \sim 1$ TeV
- huge annihilation cross section



M. Cirelli - compilation ICRC 2015

Dark Matter interpretation

However:

Dark Matter interpretation

However:

- increased precision brings increased tension

  "The improved accuracy of AMS-02 [...]
  now excludes channels previously allowed."

  M. Boudaud et al., 1410.3799

Dark Matter interpretation

However:

- increased precision brings increased tension

  "The improved accuracy of AMS-02 [...]
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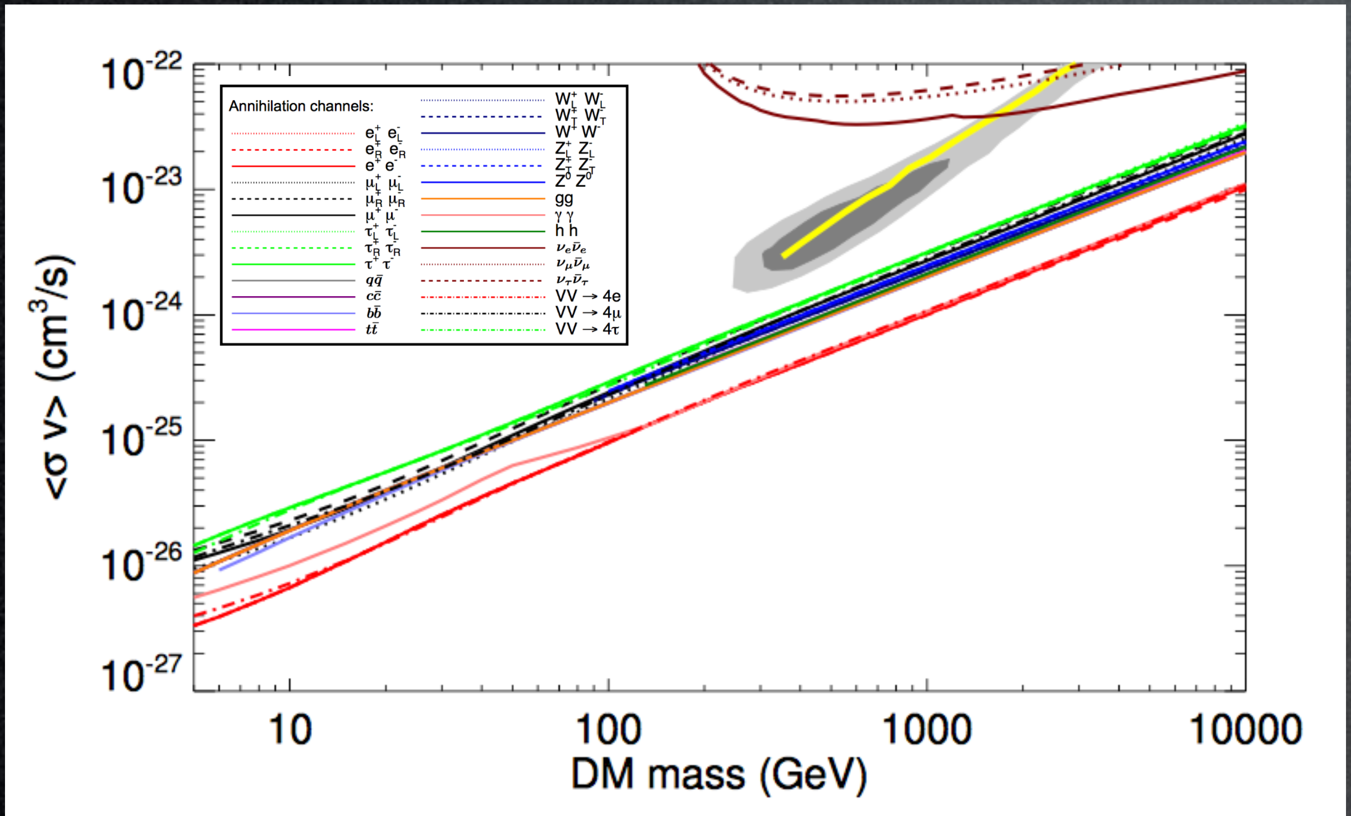
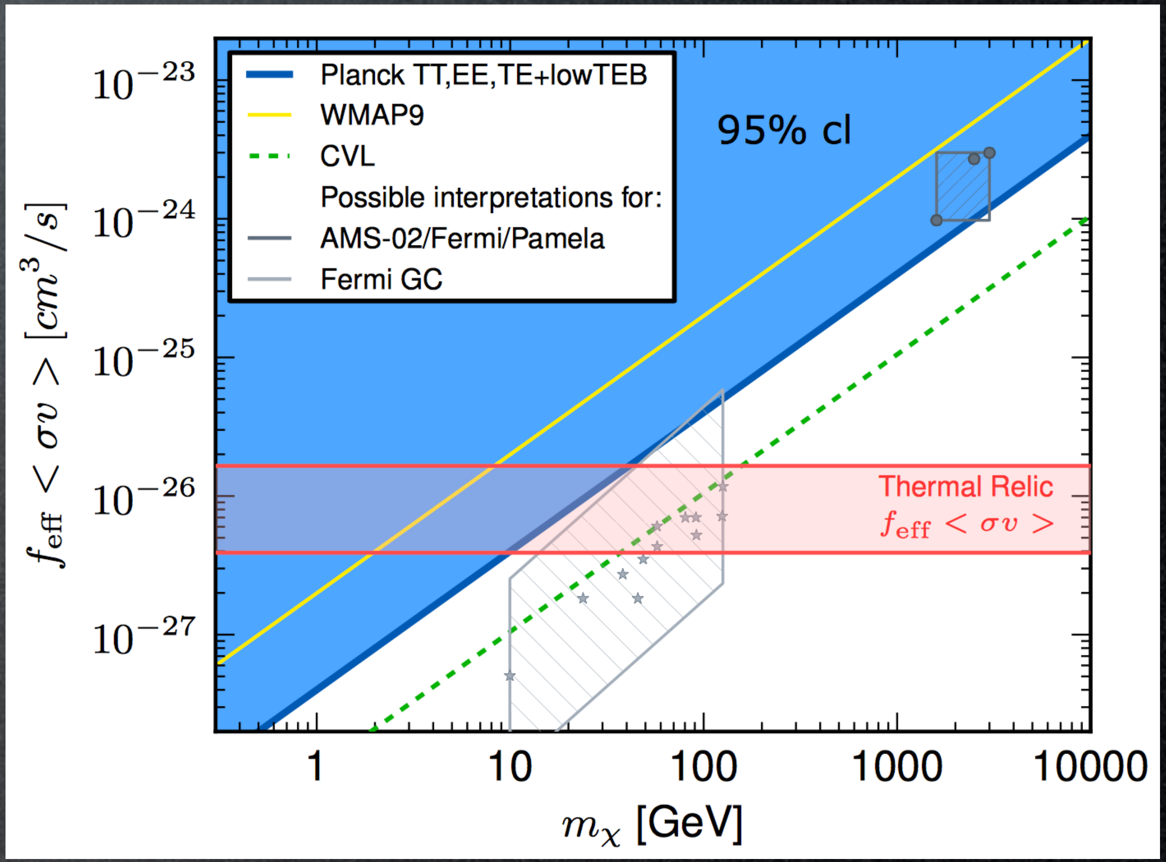
  M. Boudaud et al., 1410.3799

- constraints: gamma rays, neutrinos, CMB...

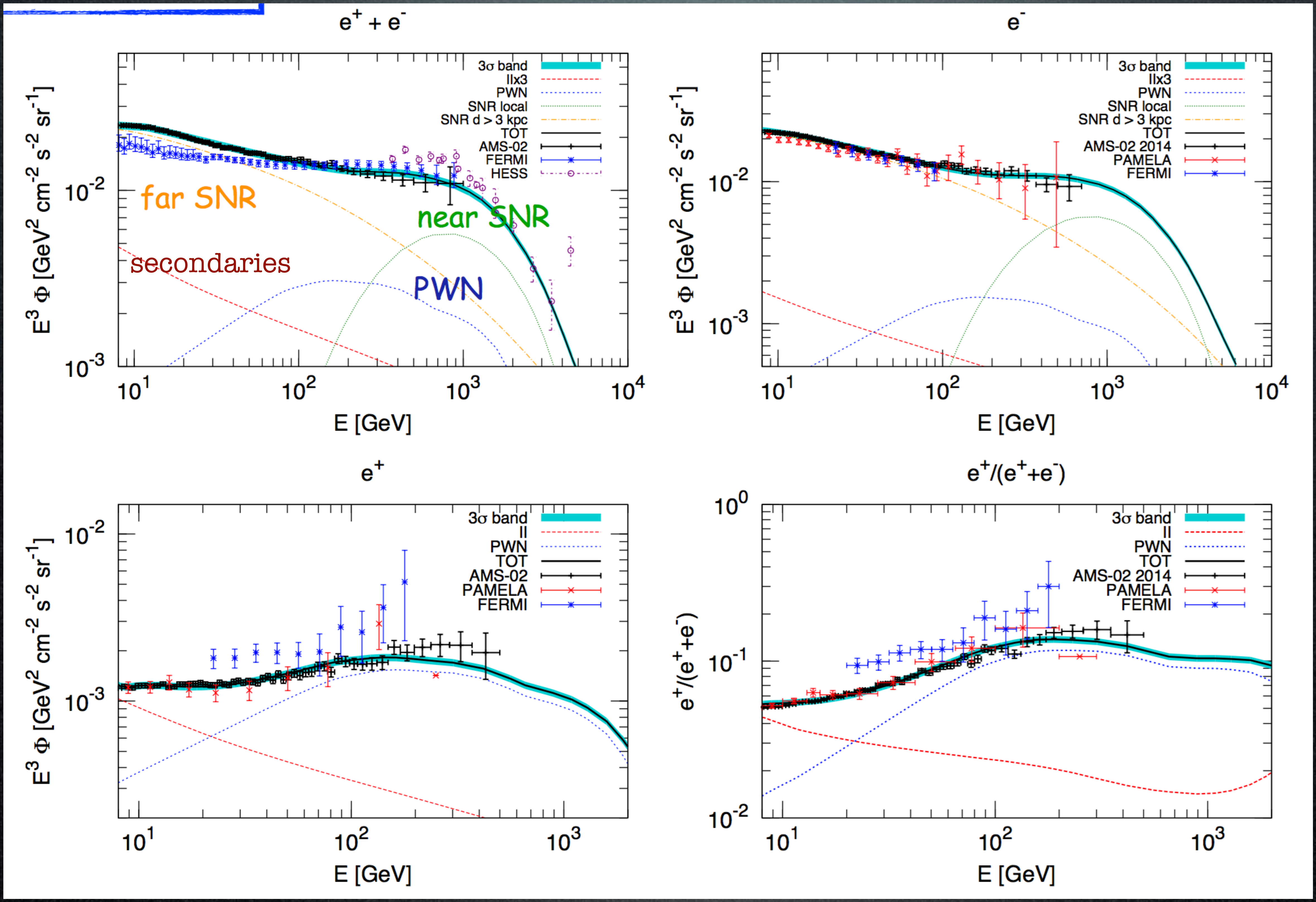


T.Slatyer 1506.03811



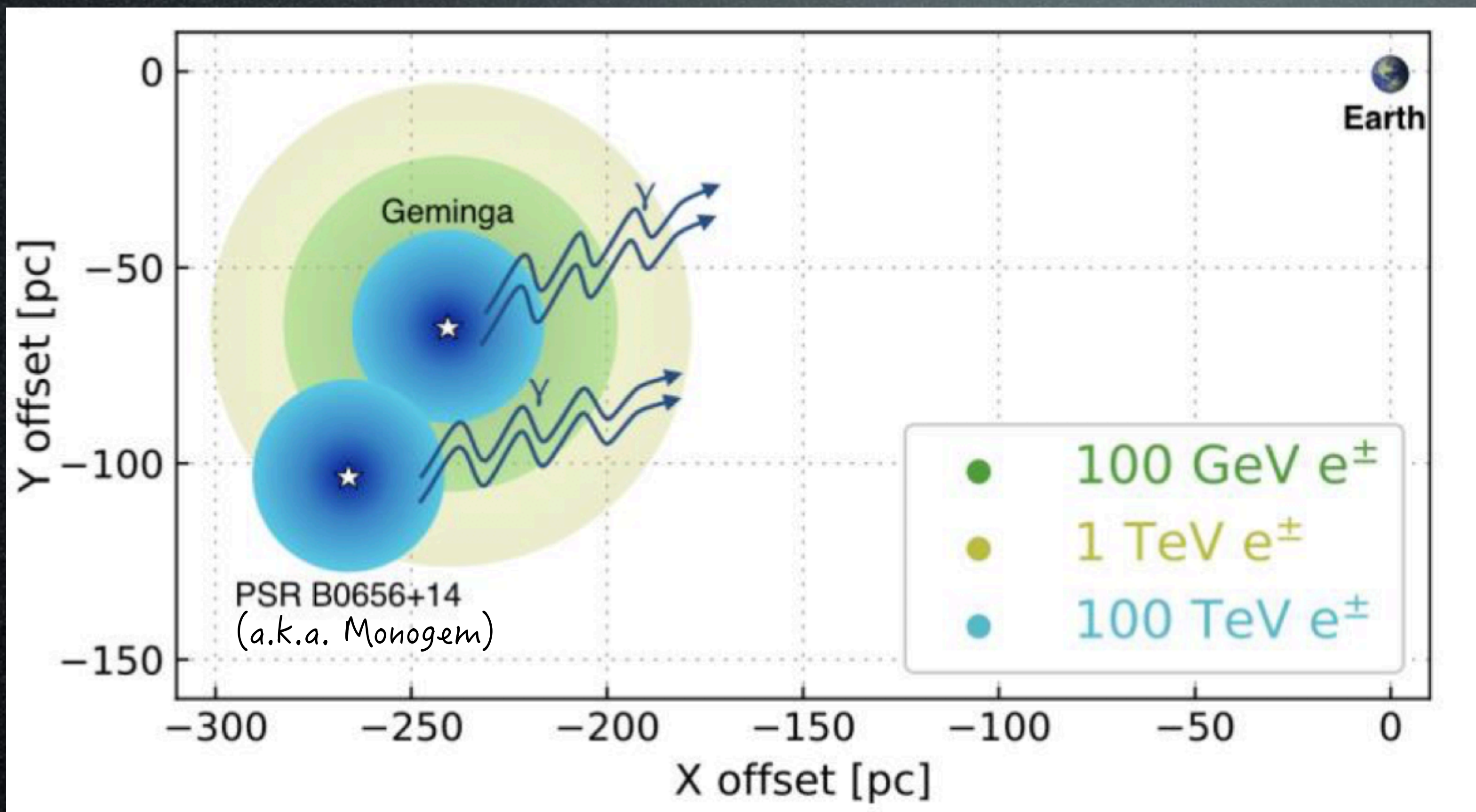
Planck 2015 (1502.01589)

Astro interpretation



M. Di Mauro et al. 1507.07001

Dark Matter interpretation: the come back?



HAWC Coll., Science 359 (2017) 911 - 1711.06223

HAWC sees ICS TeV γ-rays from ~100 TeV e^+e^- from Geminga and Monogem

e^+ are 'very trapped' around these pulsars (diffusion is very slow)

e^+ cannot reach Earth to explain 100 GeV excesses, must be stg else (DM?)

Geminga and PSR B0656+14 are the oldest pulsars for which a tera-electron volt nebula has so far been detected. Under our assumption of isotropic and homogeneous diffusion, the dominant source of the positron flux above 10 GeV cannot be either Geminga or PSR B0656+14. Under the unlikely situation that the field is nearly aligned along the direction between Earth and the nearby tera-electron volt nebulae, the local positron flux can be increased; however, the tera-electron volt morphology of the sources matches our isotropic diffusion model. We therefore favor the explanation that instead of these two pulsars, the origin of the local positron flux must be explained by other processes, such as different assumptions about secondary production [although that has been questioned (33; 34)], other pulsars, other types of cosmic accelerators such as micro-quasars (35) and supernova remnants (34), or the annihilation or decay of dark matter particles (9).

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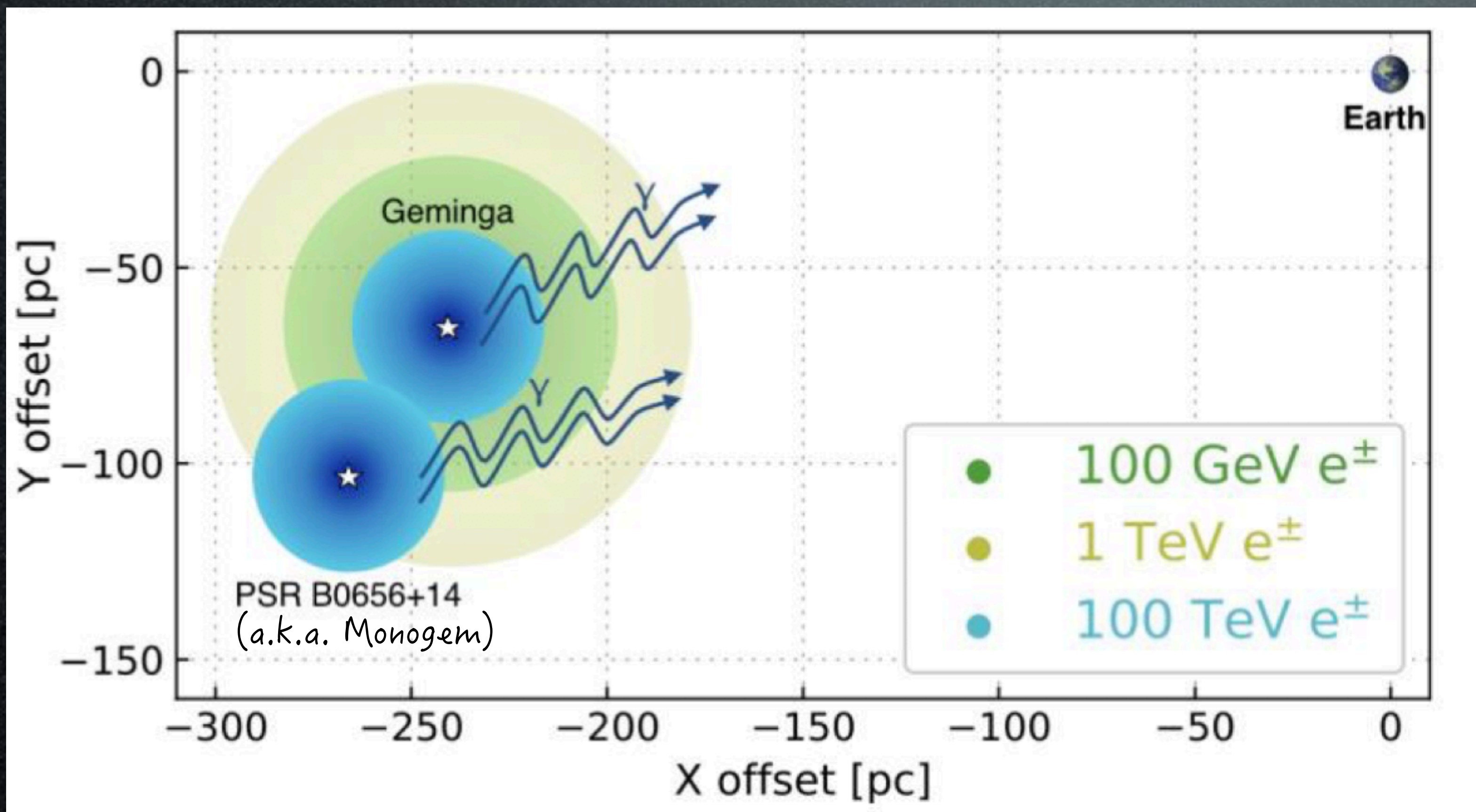

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Criticisms:

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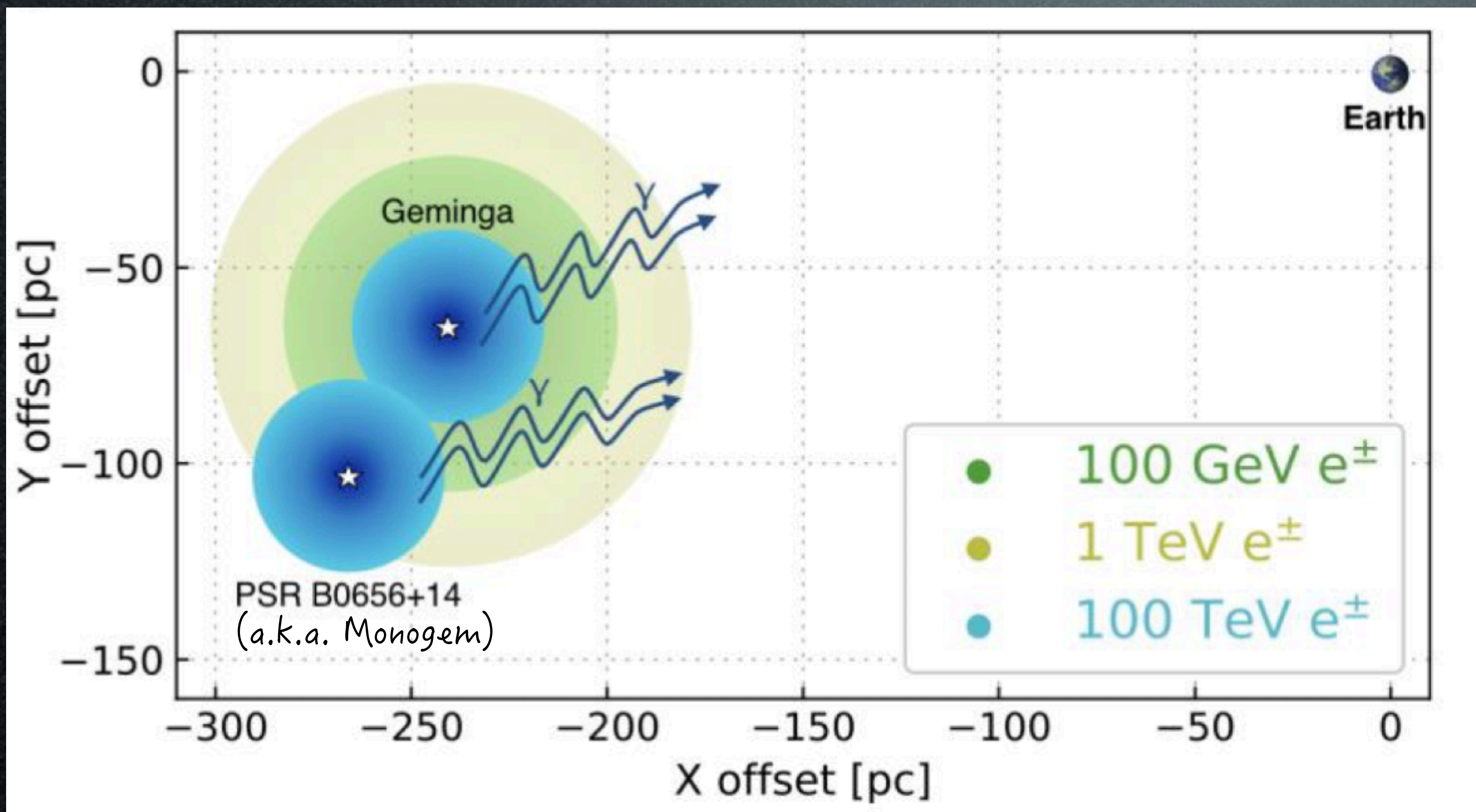

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Criticisms:

- space-dep diffusion: local ≠ global
- E-dep diffusion: 100 TeV vs 100 GeV

usually $\mathcal{K}(E) = \mathcal{K}_0\,(E/\mathrm{GeV})^\delta$ so E factored out,
but cannot exclude residual dependance

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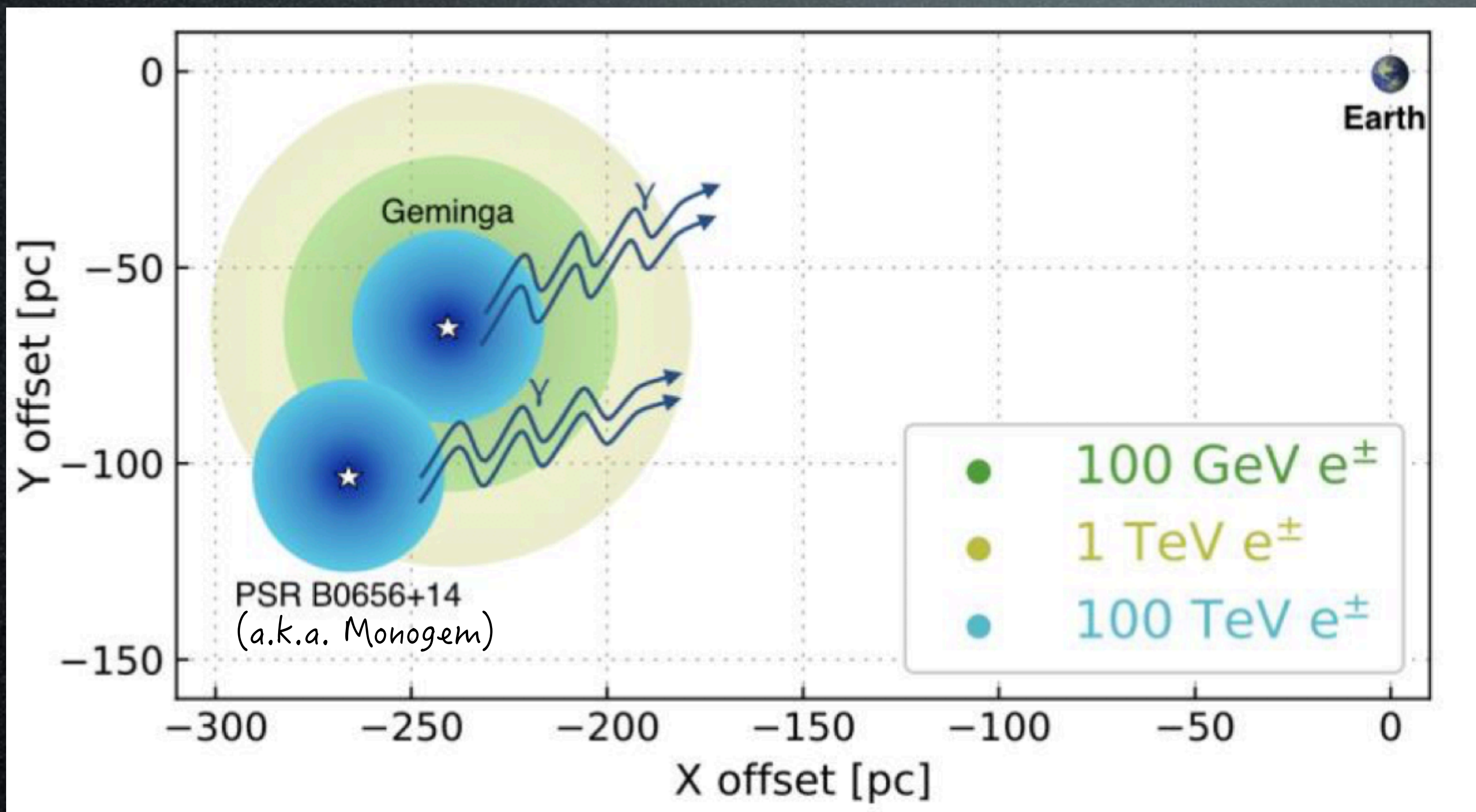

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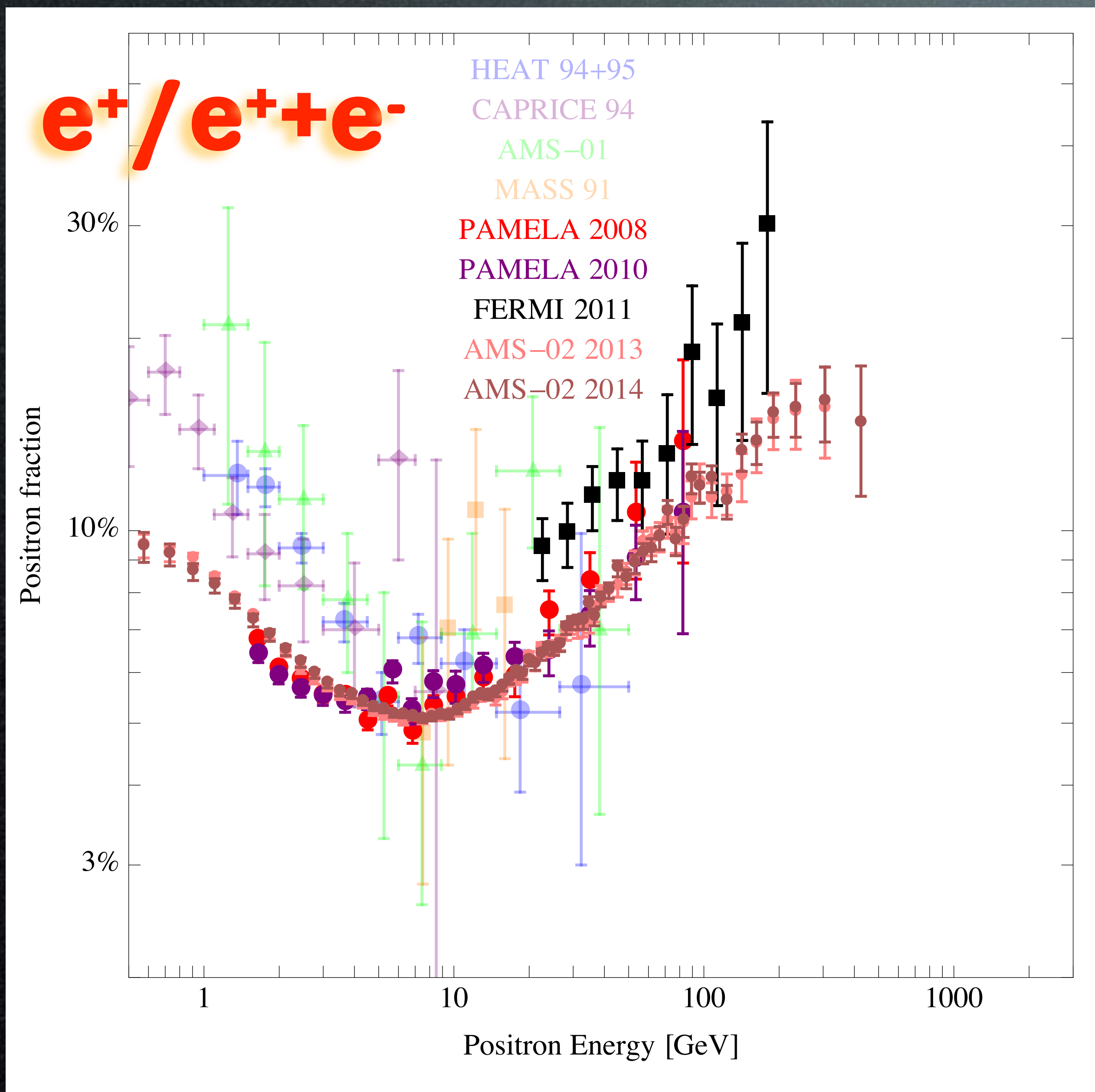
Criticisms:

- space-dep diffusion: local ≠ global
- E-dep diffusion: 100 TeV vs 100 GeV
- t-dep: γ-rays today, but e^+ 10^4 yrs ago

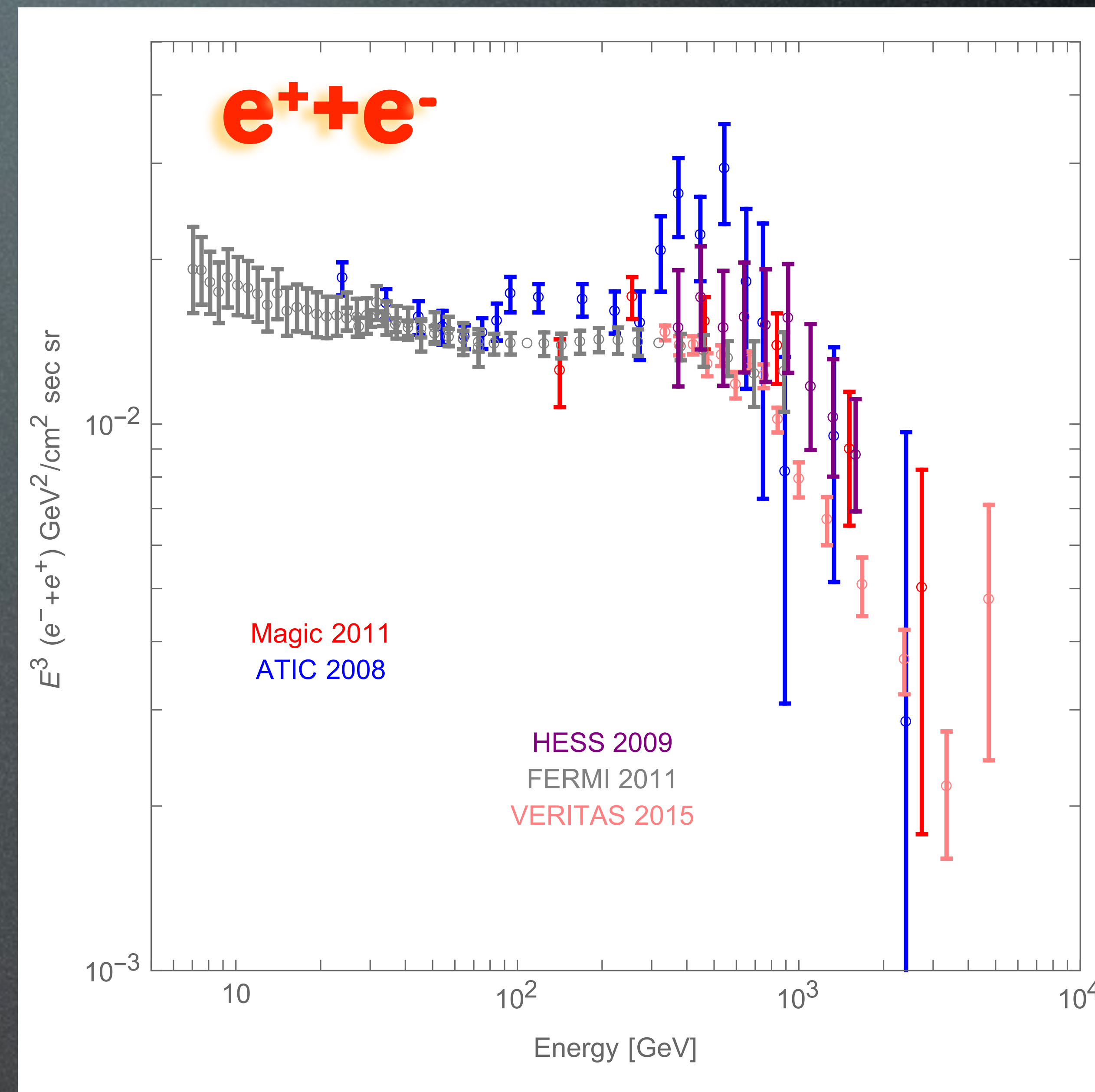
Credit: Mathieu Boudaud

Data: leptons

high energy



M. Cirelli - compilation ICRC 2015



M. Cirelli - compilation

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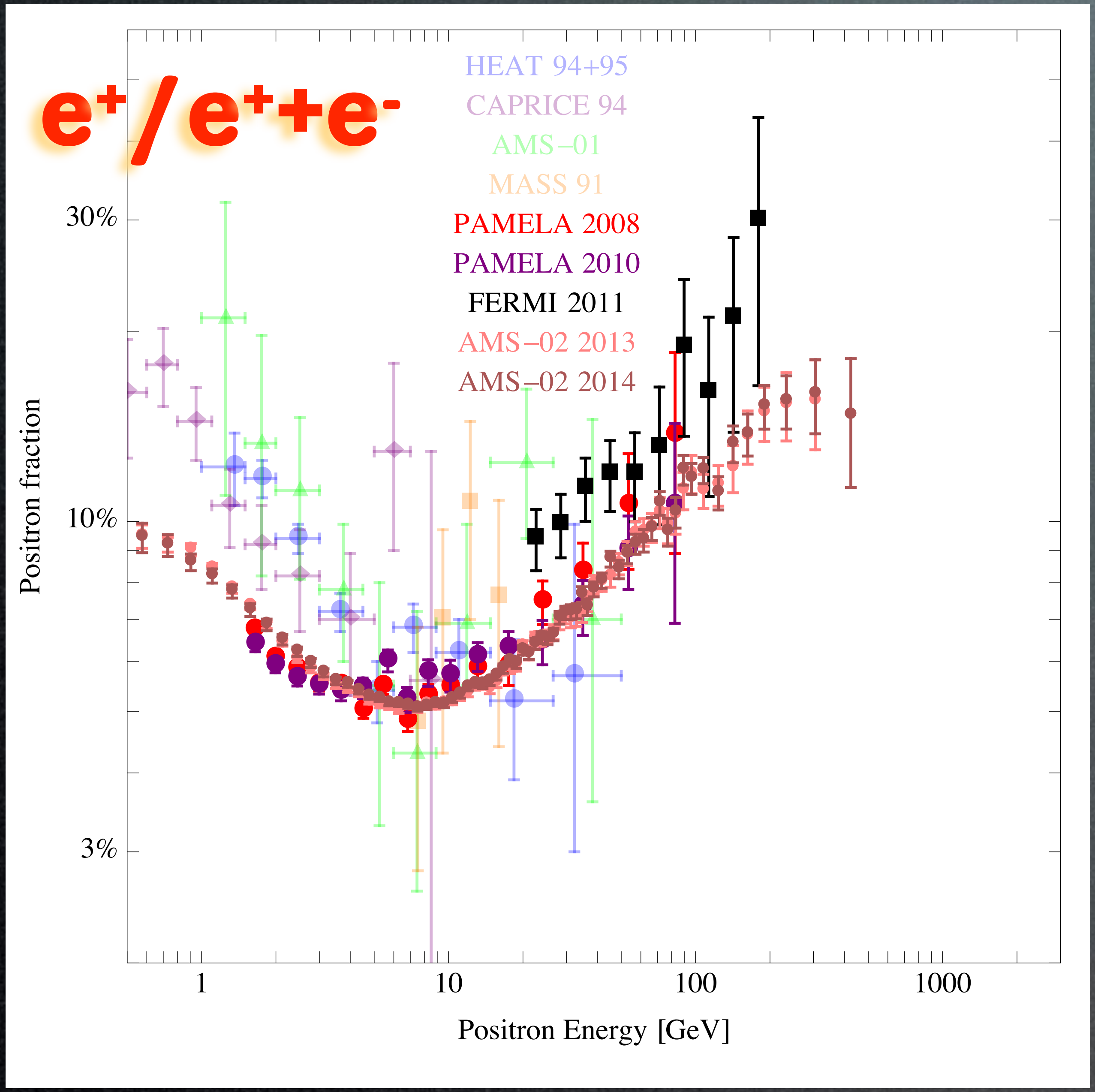
high energy



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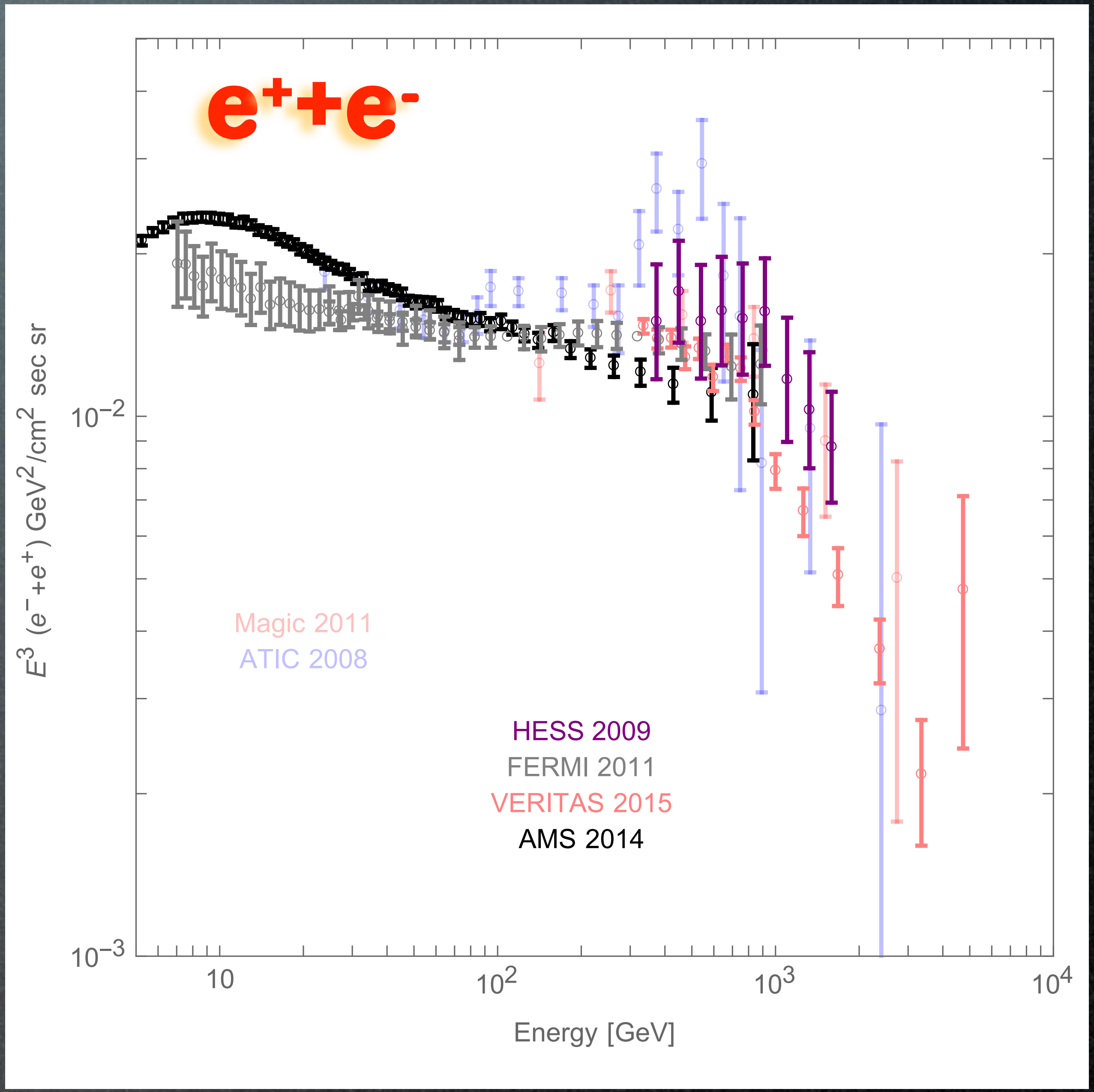

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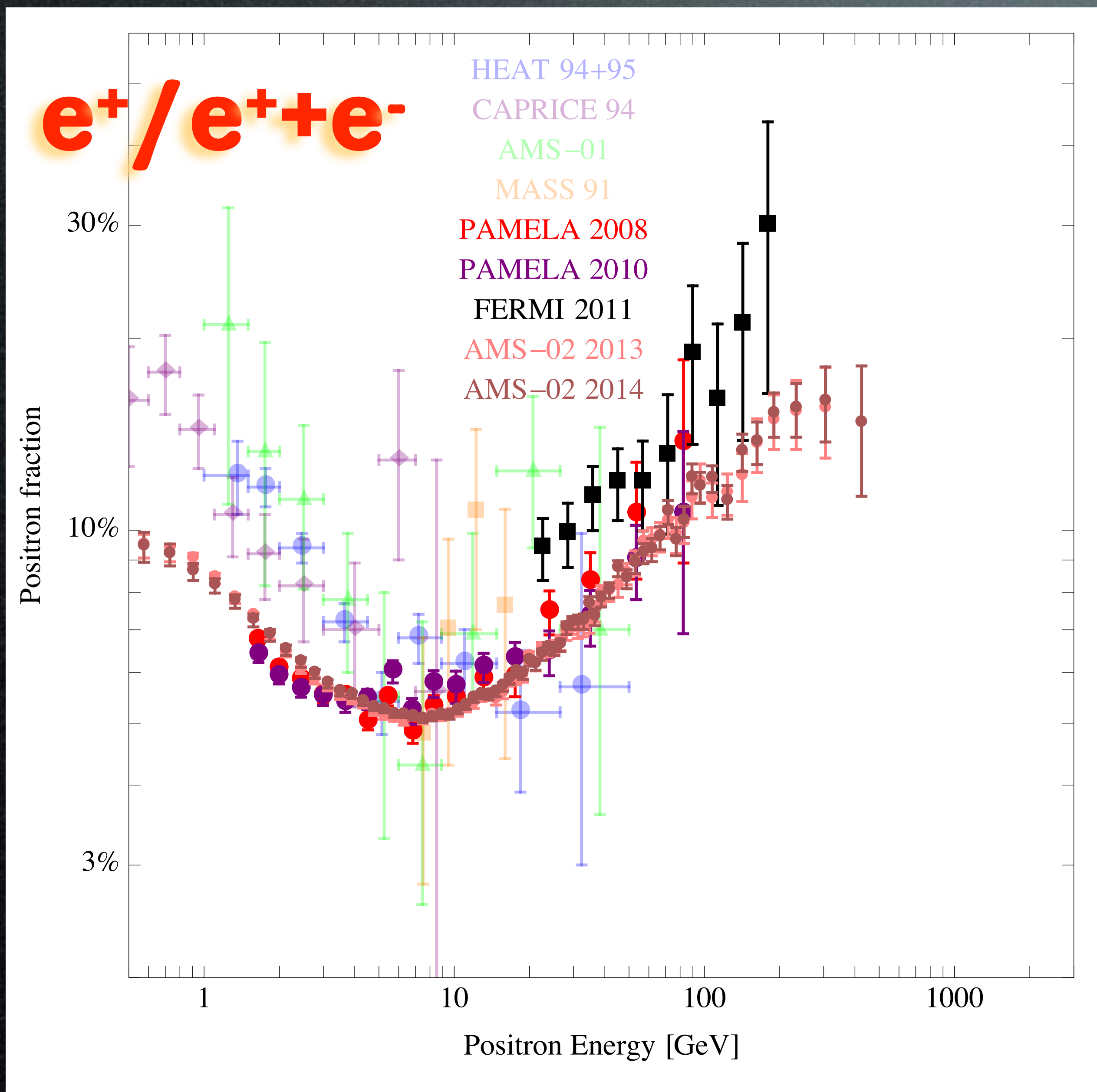

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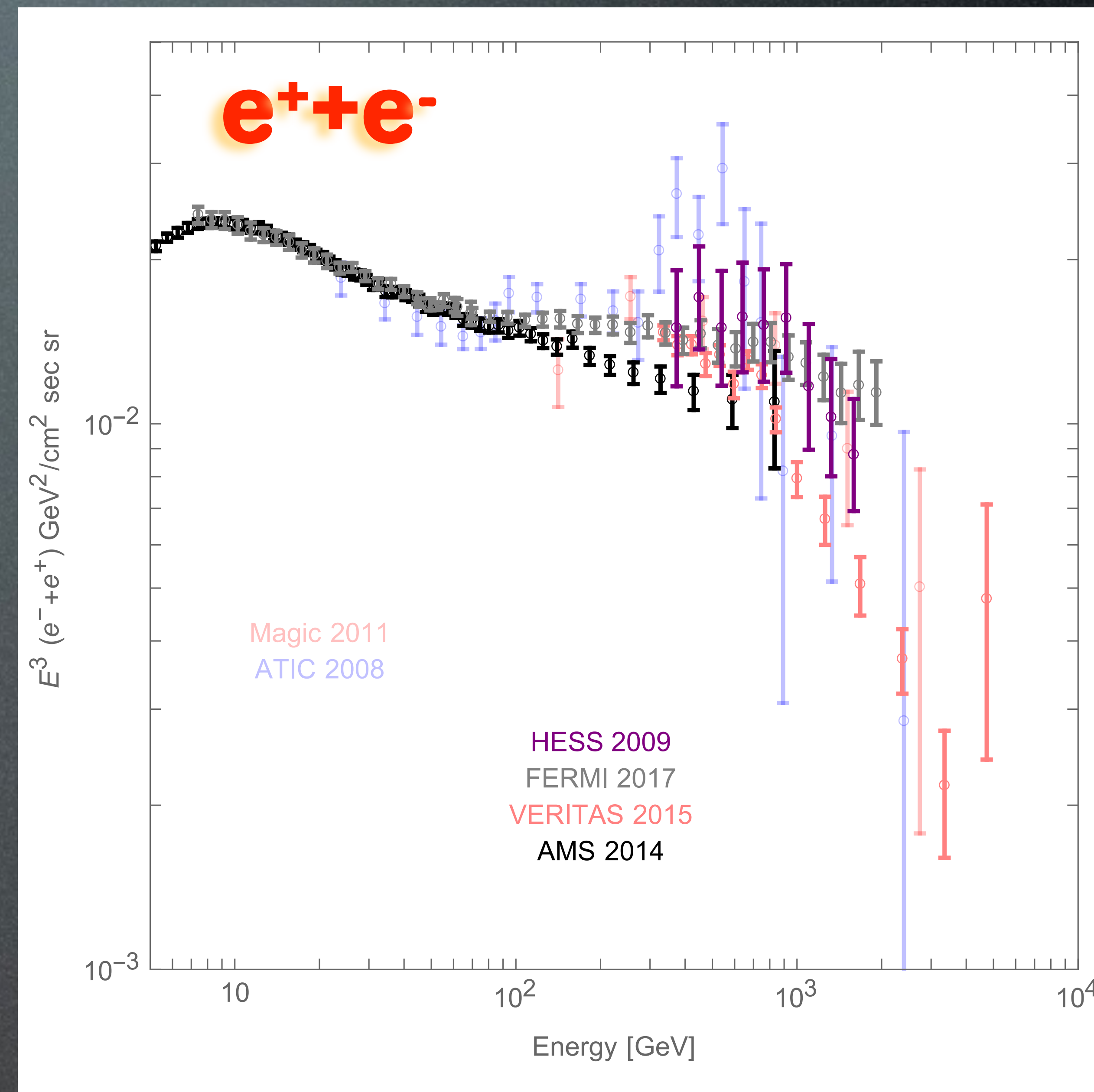


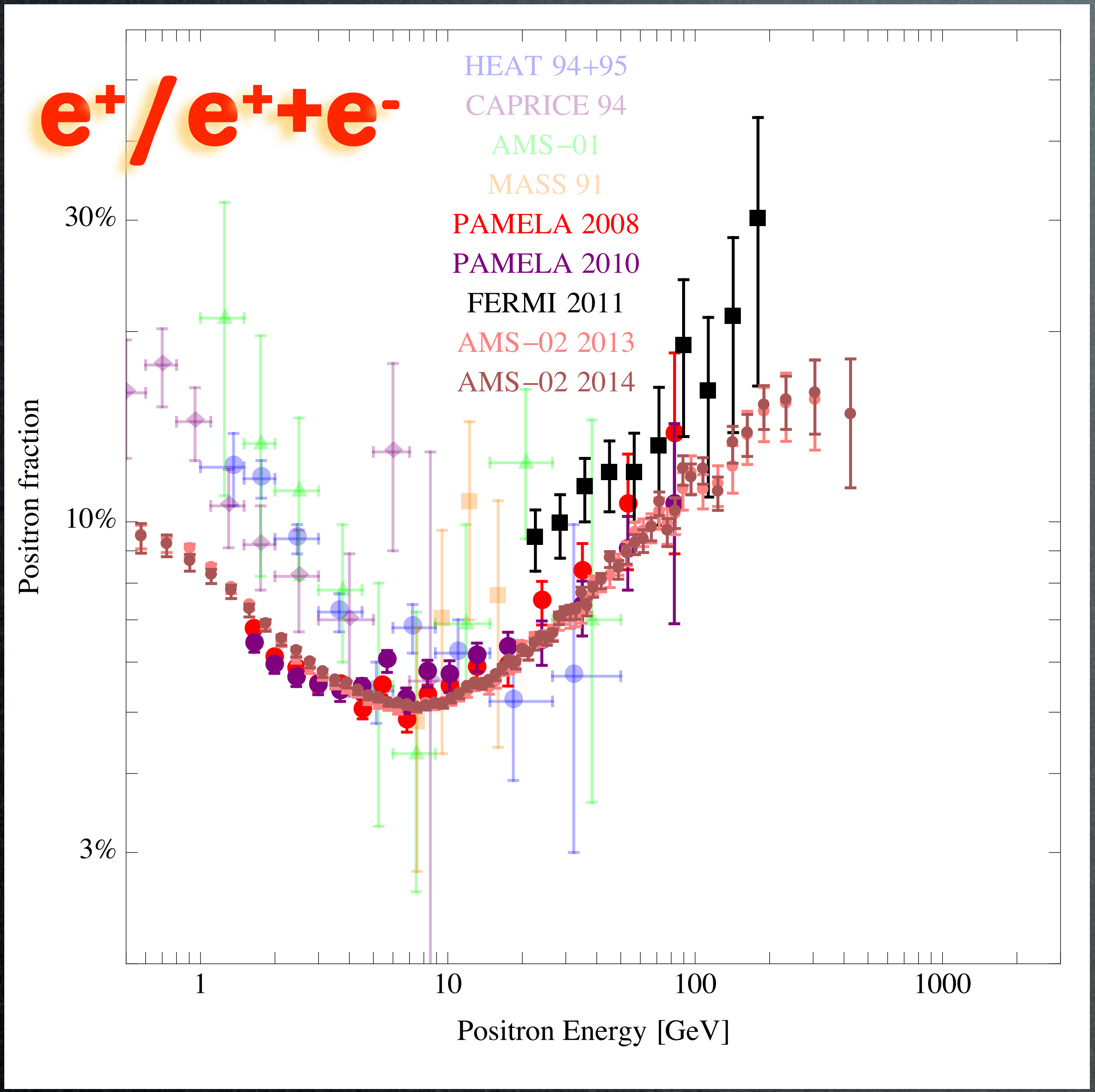
M. Cirelli - compilation

Fermi Coll.
1704.07195
(PRD 96)

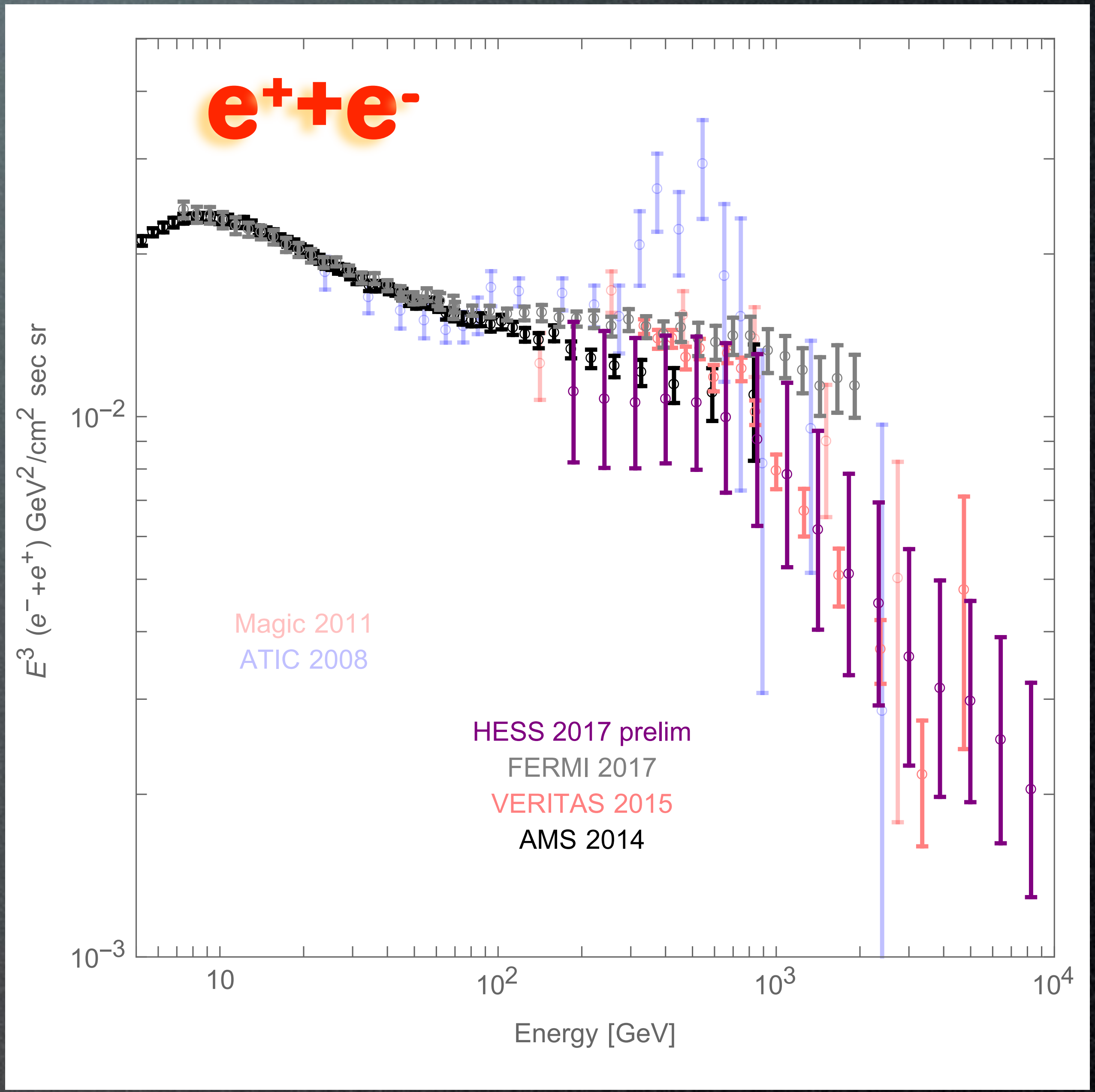
Below 100 GeV, the new LAT measurement differs from the previous one by 10–30%, as can be seen in Fig. 13. A large part of this difference below 30 GeV is due to the lack of correction in the previous analysis for the loss of CREs above the geomagnetic energy cutoff. After applying this correction, the remaining difference is 10–15% and is due to imperfections in the simulation that was used in the previous analysis (remnants of electronic signals from out-of-time particles were not simulated [34]).

Data: leptons

high energy



M. Cirelli - compilation ICRC 2015

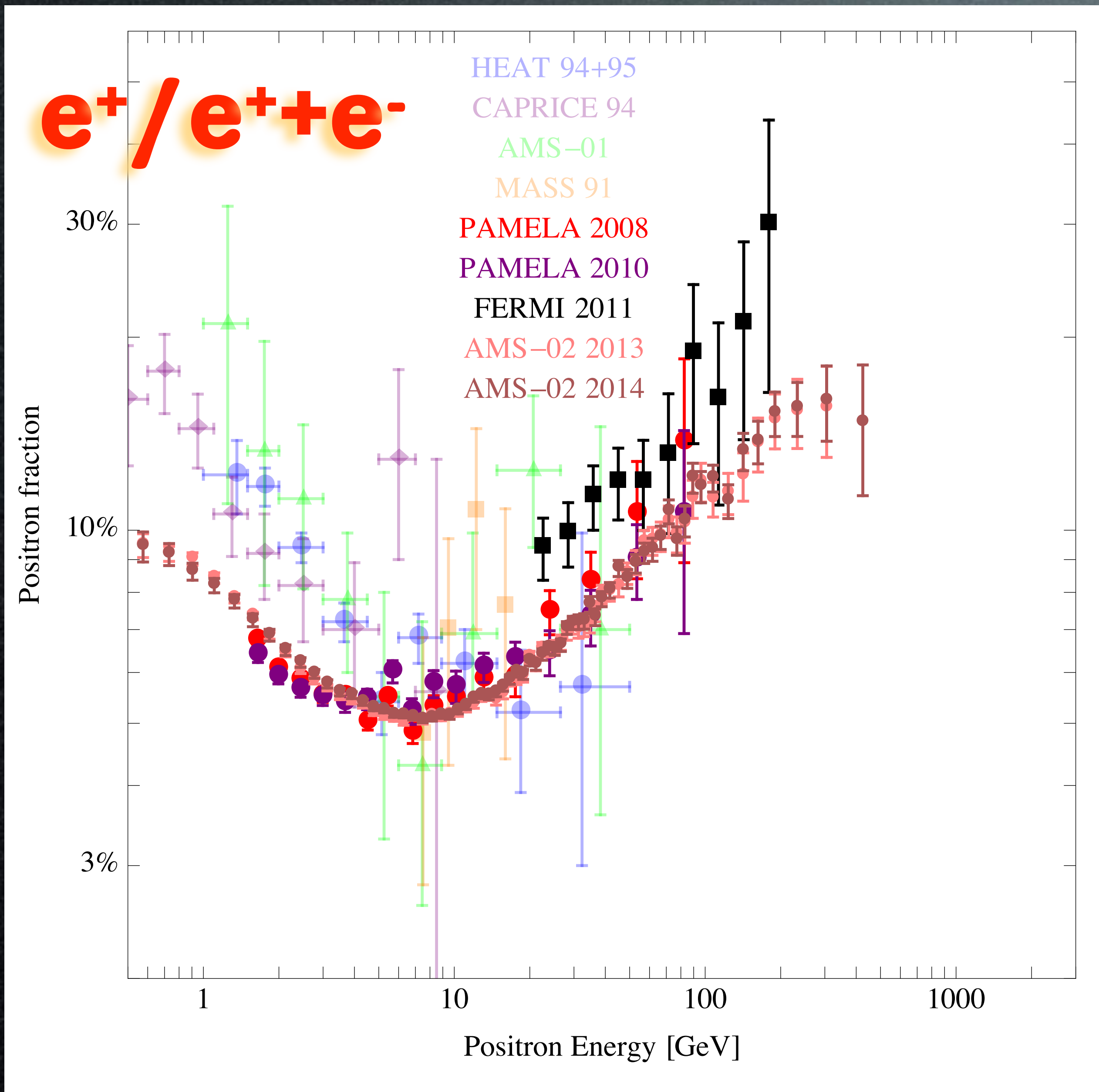


M. Cirelli - compilation

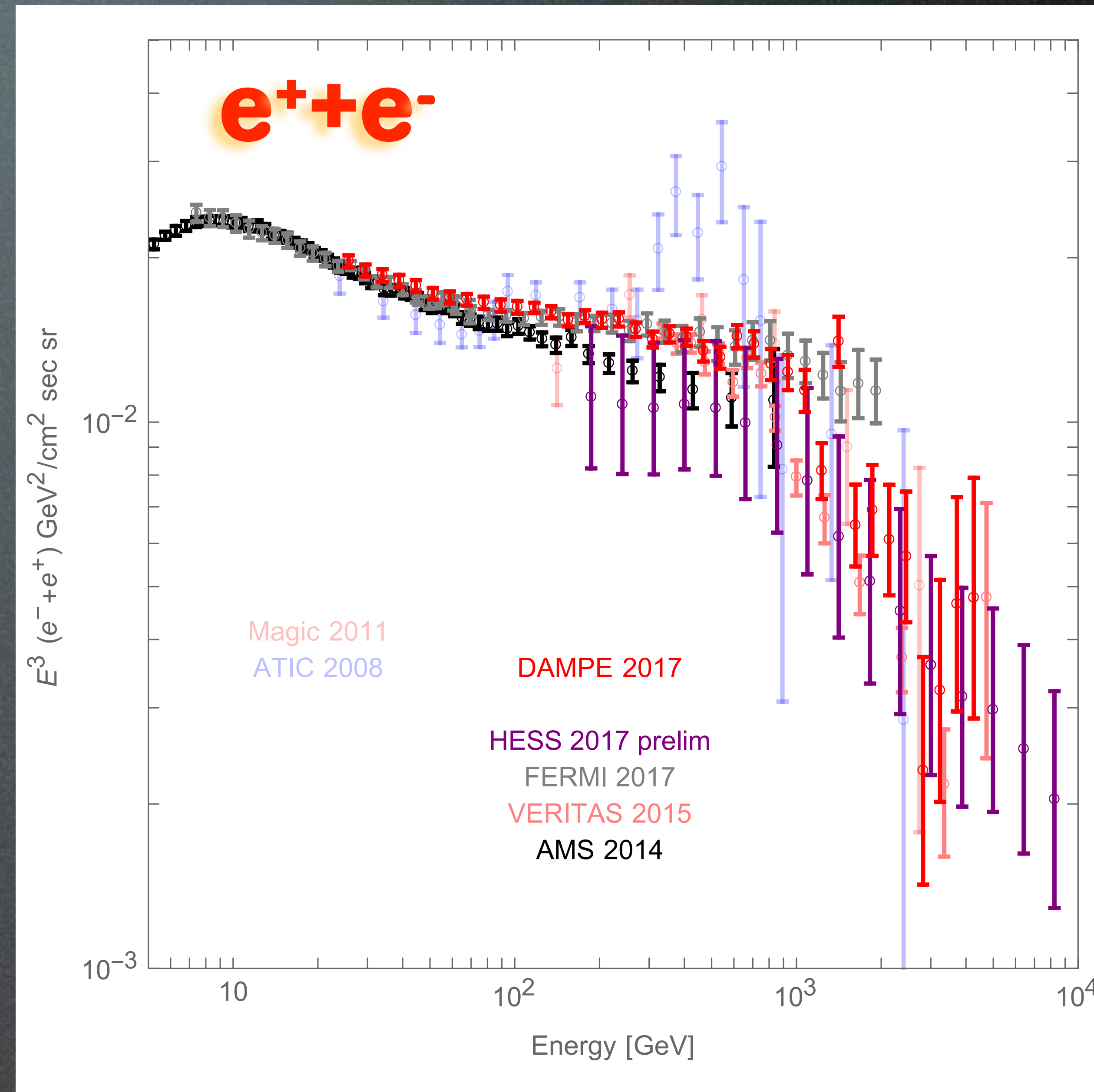
HESS Coll.
ICRC 2017
(D. Kerszberg)
no paper nor proceeding yet

Data: leptons

high energy



M. Cirelli - compilation ICRC 2015



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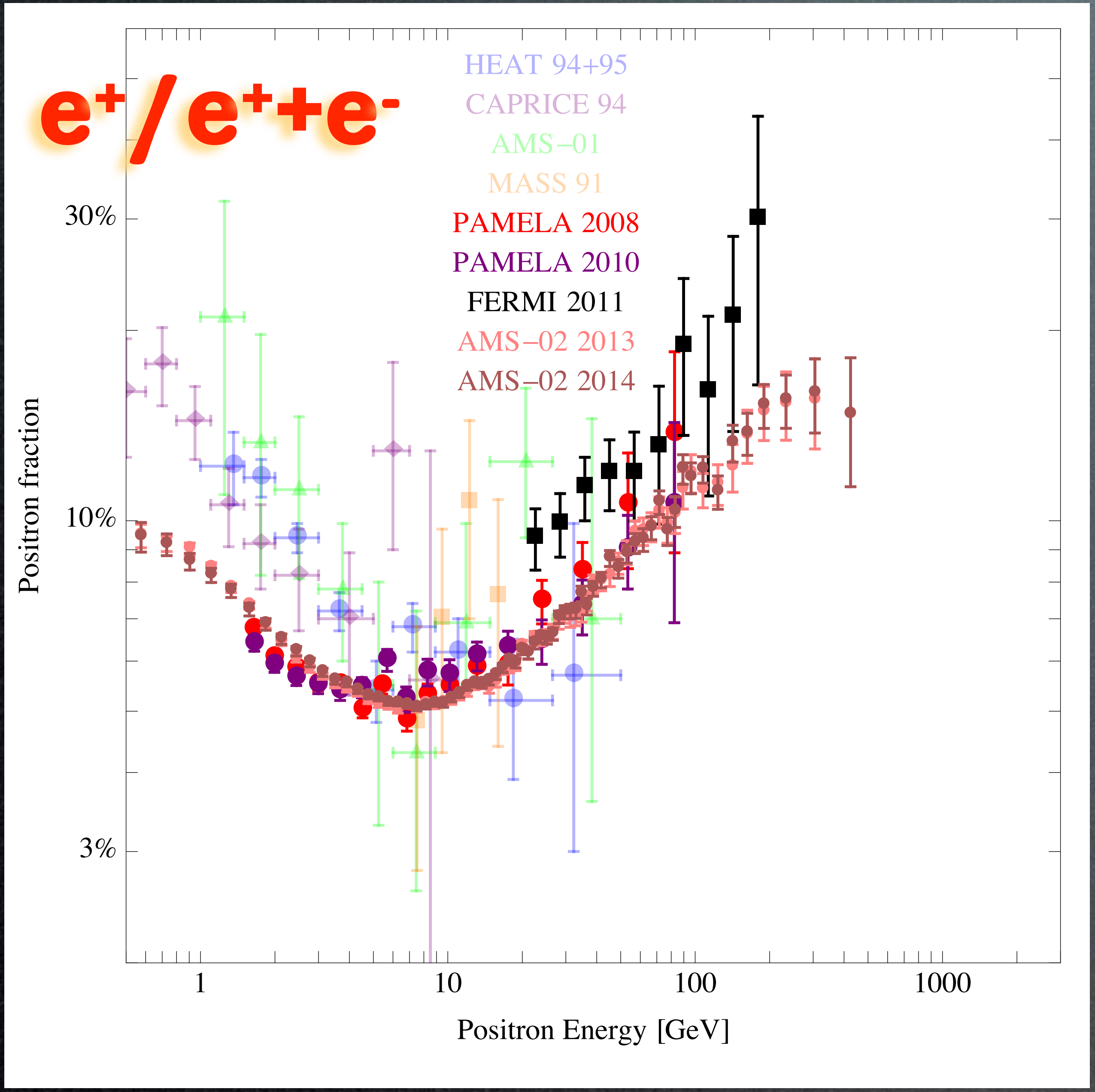

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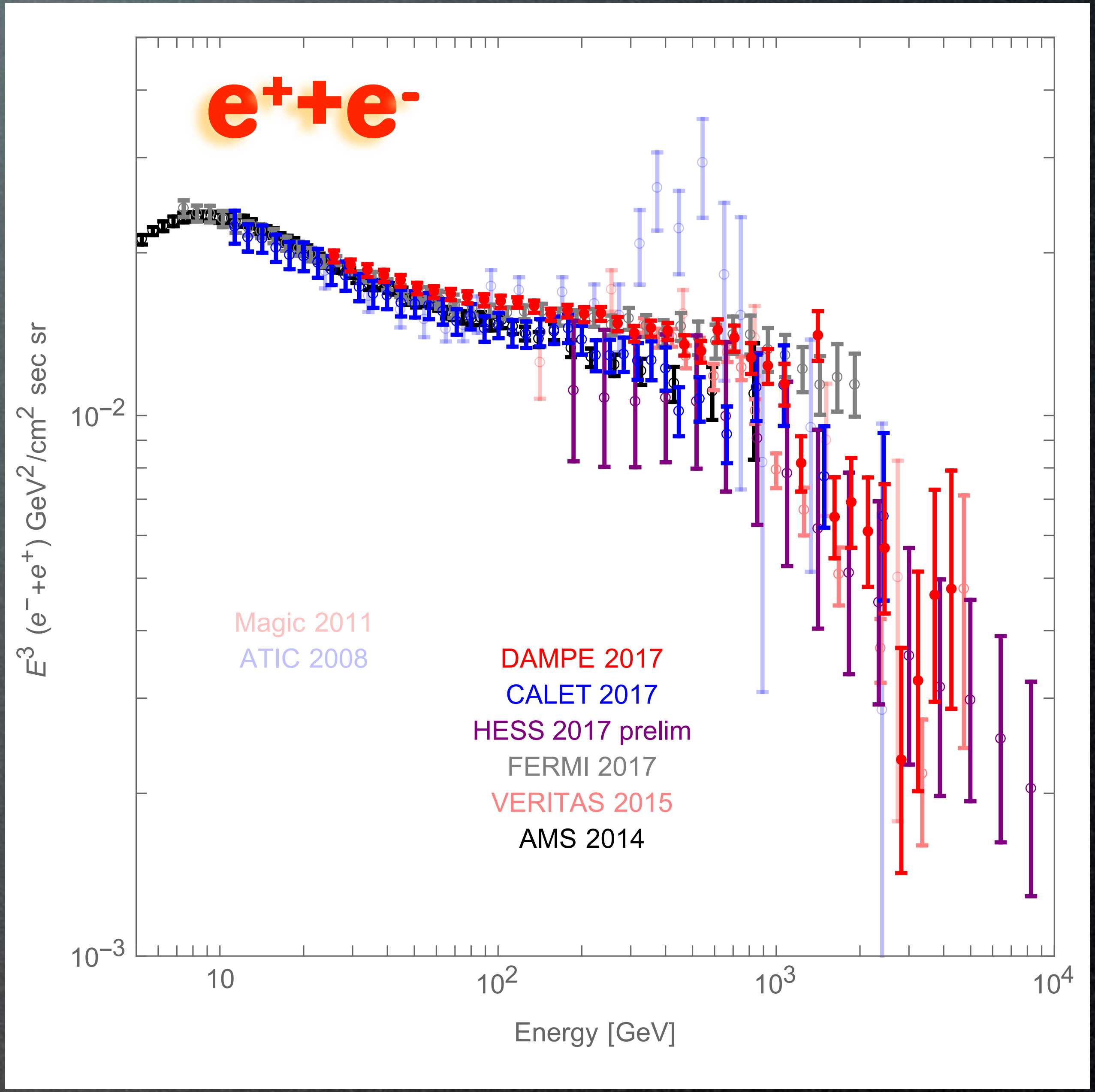

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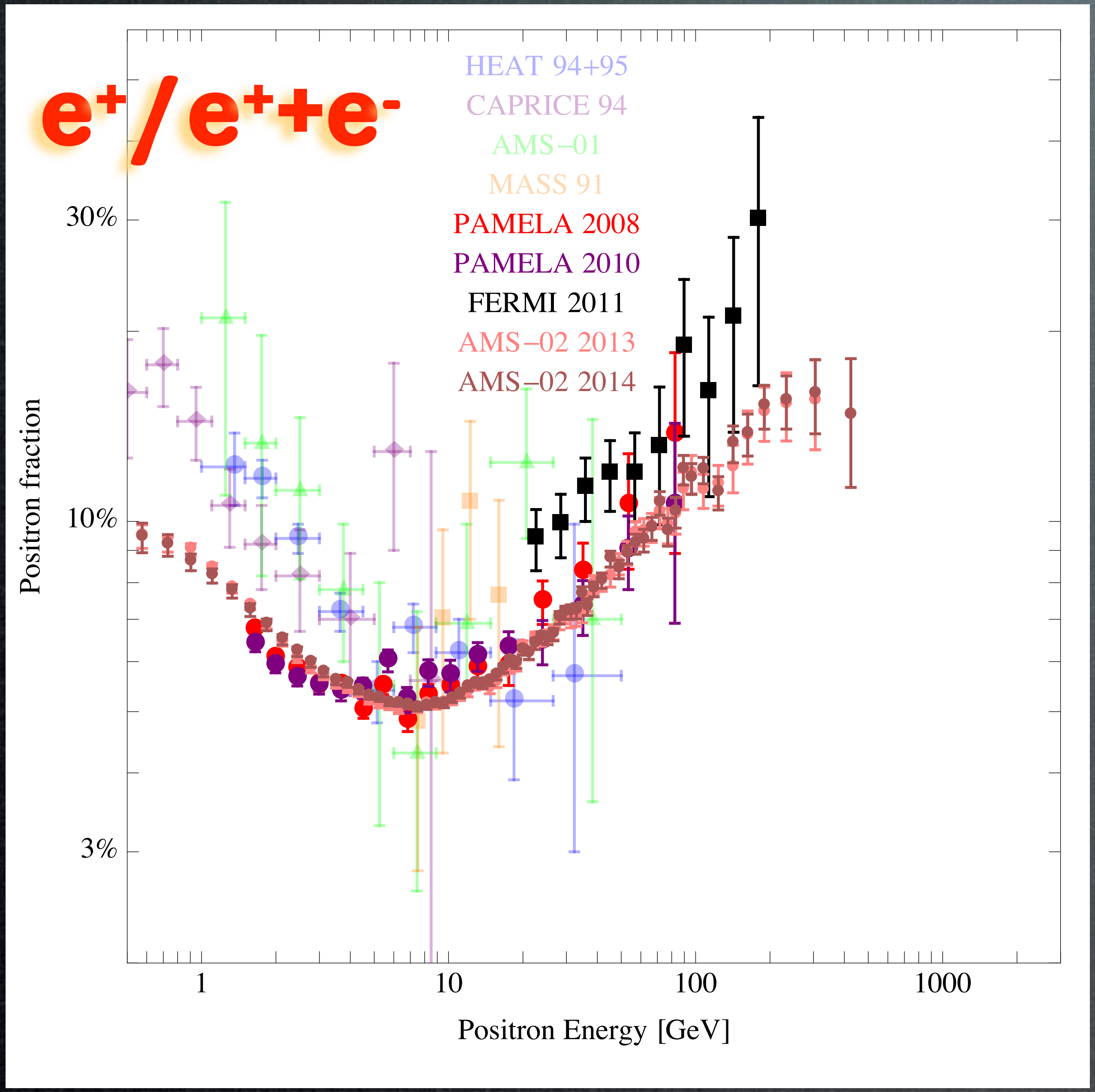

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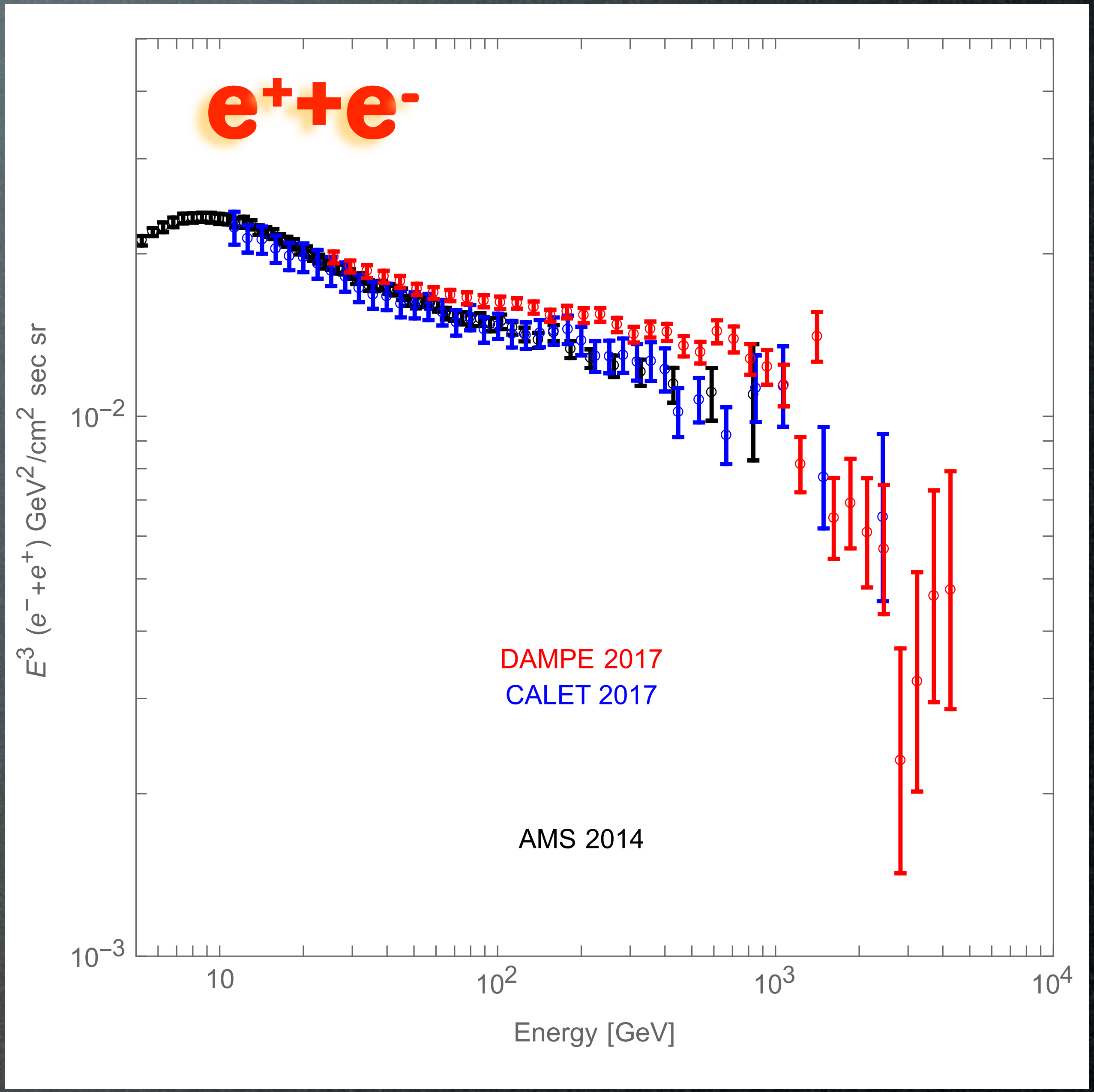

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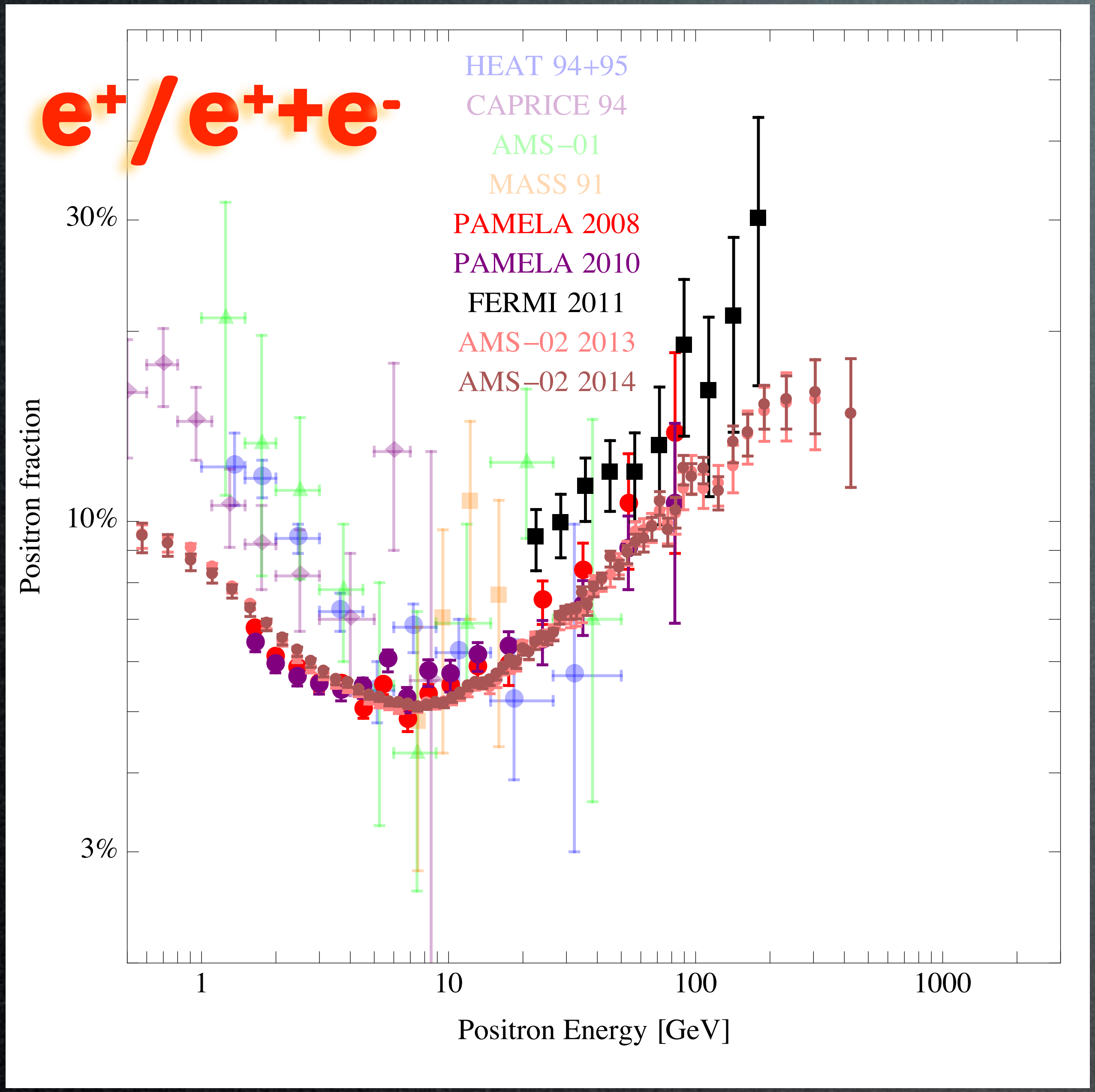
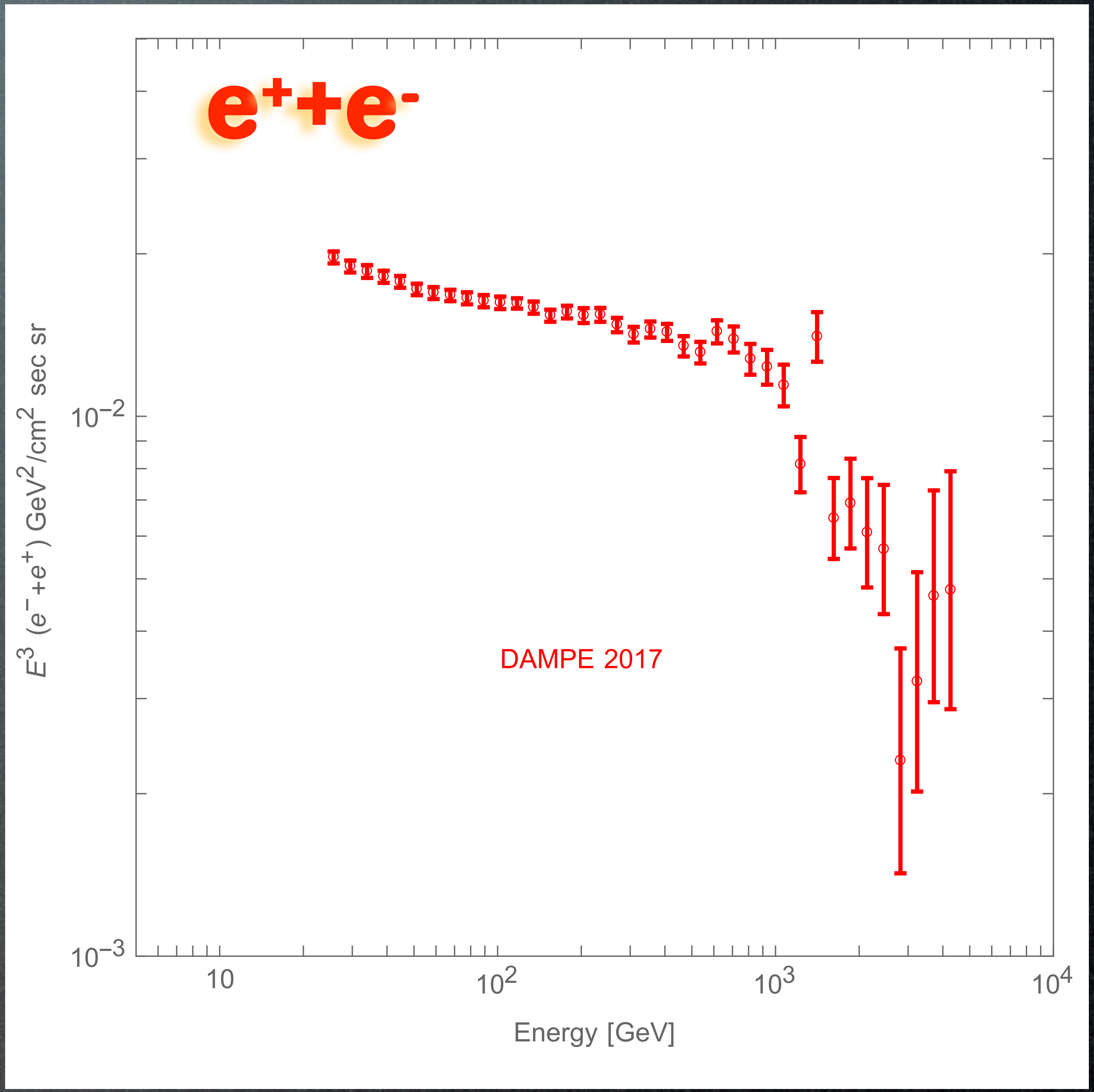


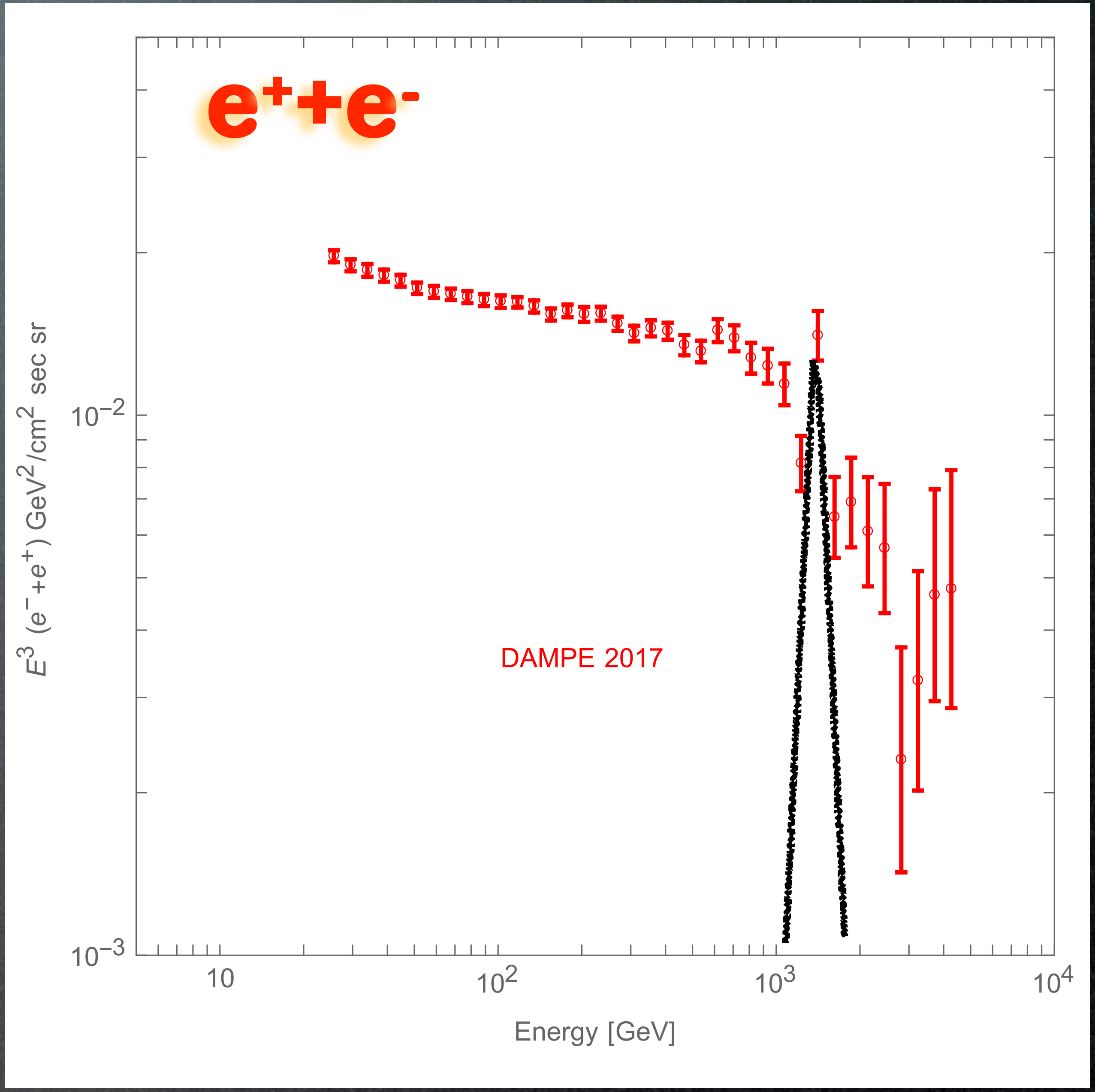


M. Cirelli - compilation ICRC 2015

M. Cirelli - compilation

Data: leptons
high energy

frenetic activity in December 2017
(38 papers / 29 days)



M. Cirelli - compilation

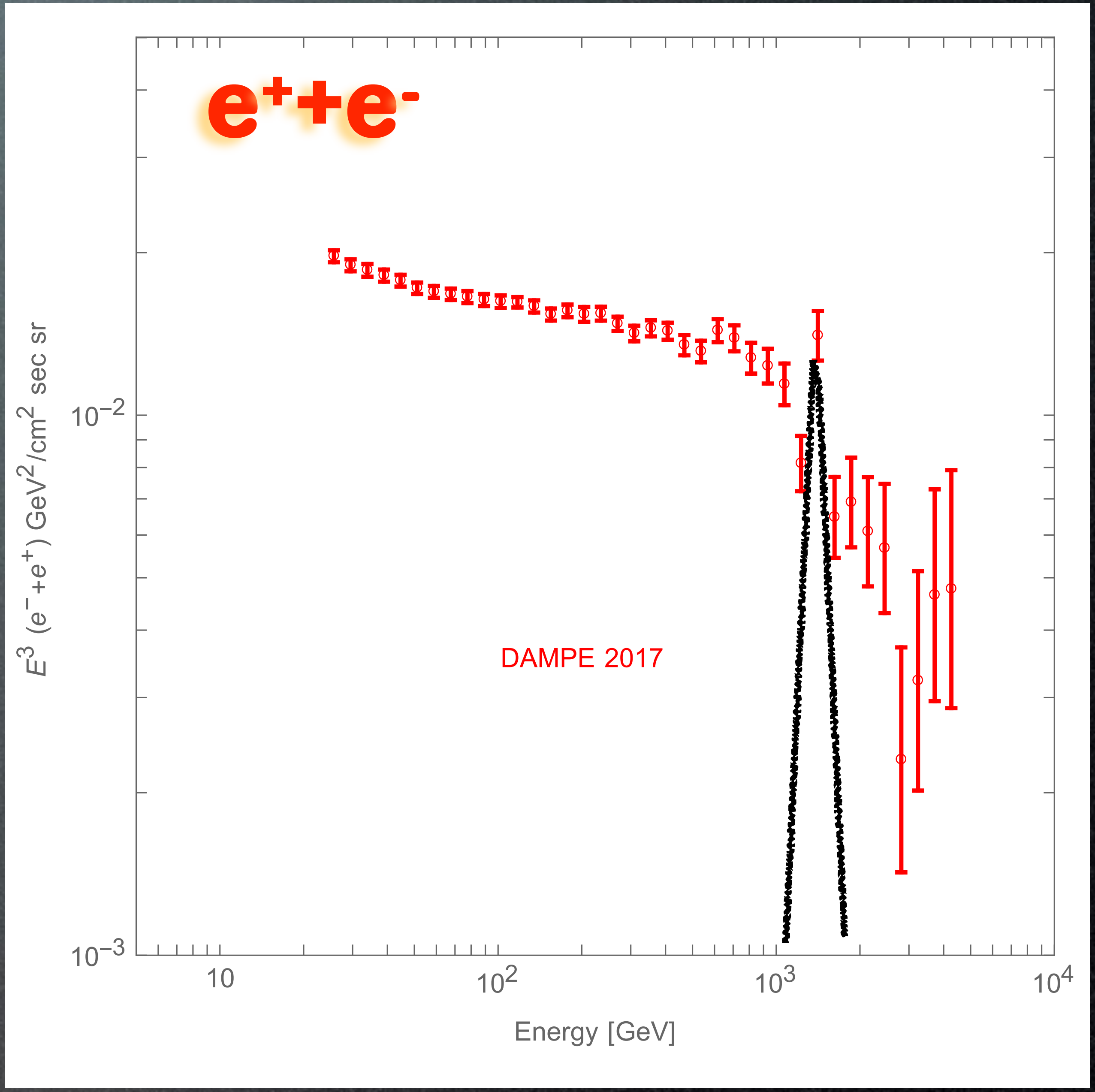
Data: leptons

high energy

frenetic activity in December 2017
(38 papers / 29 days)

- leptonic channel (e^+e^- or $\mu^+\mu^-$)

- nearby (0.2 kpc) huge (10^8 M_{sun}) DM clump
  - for large flux
  - for peaked spectrum

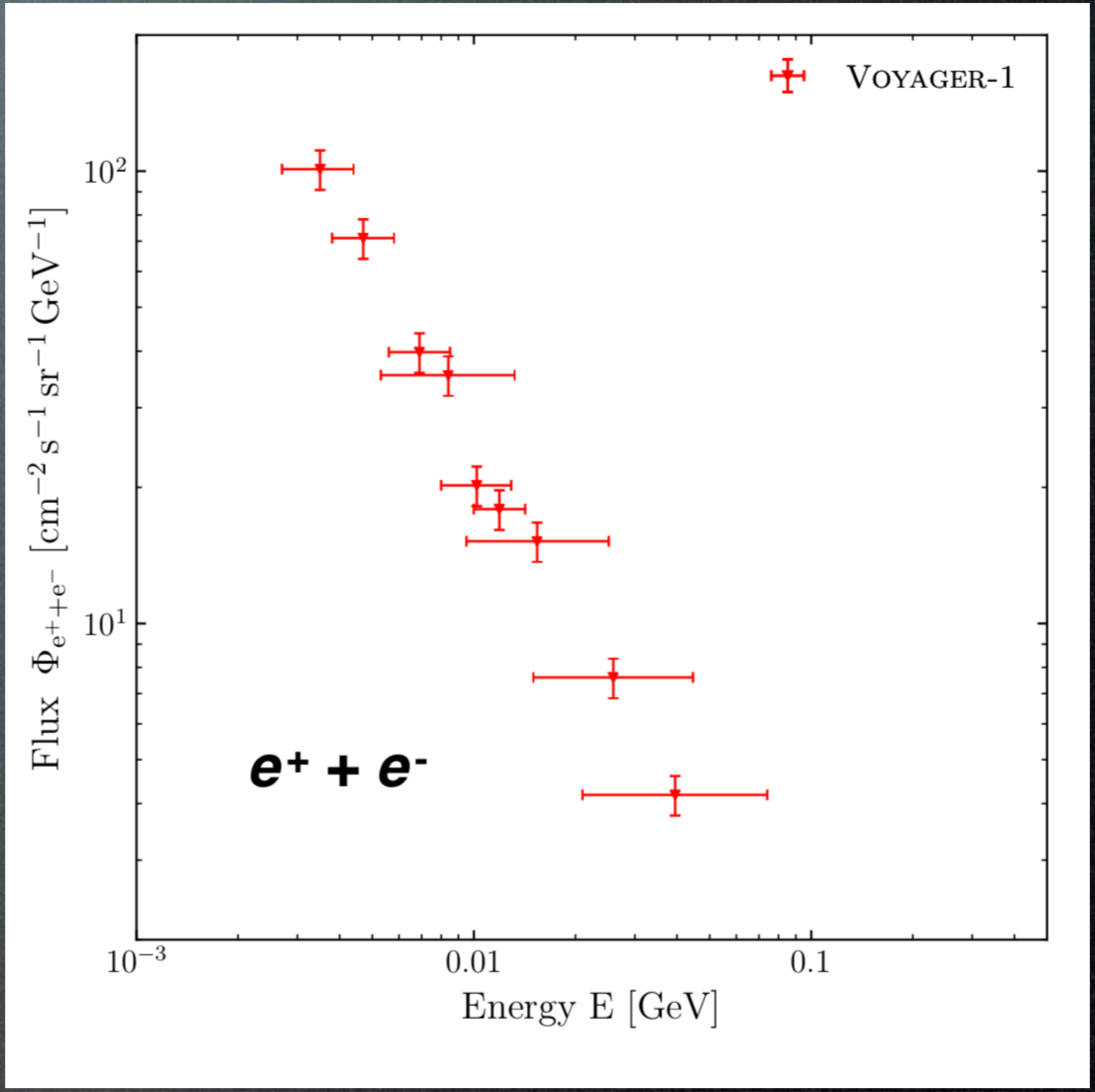


M. Cirelli - compilation

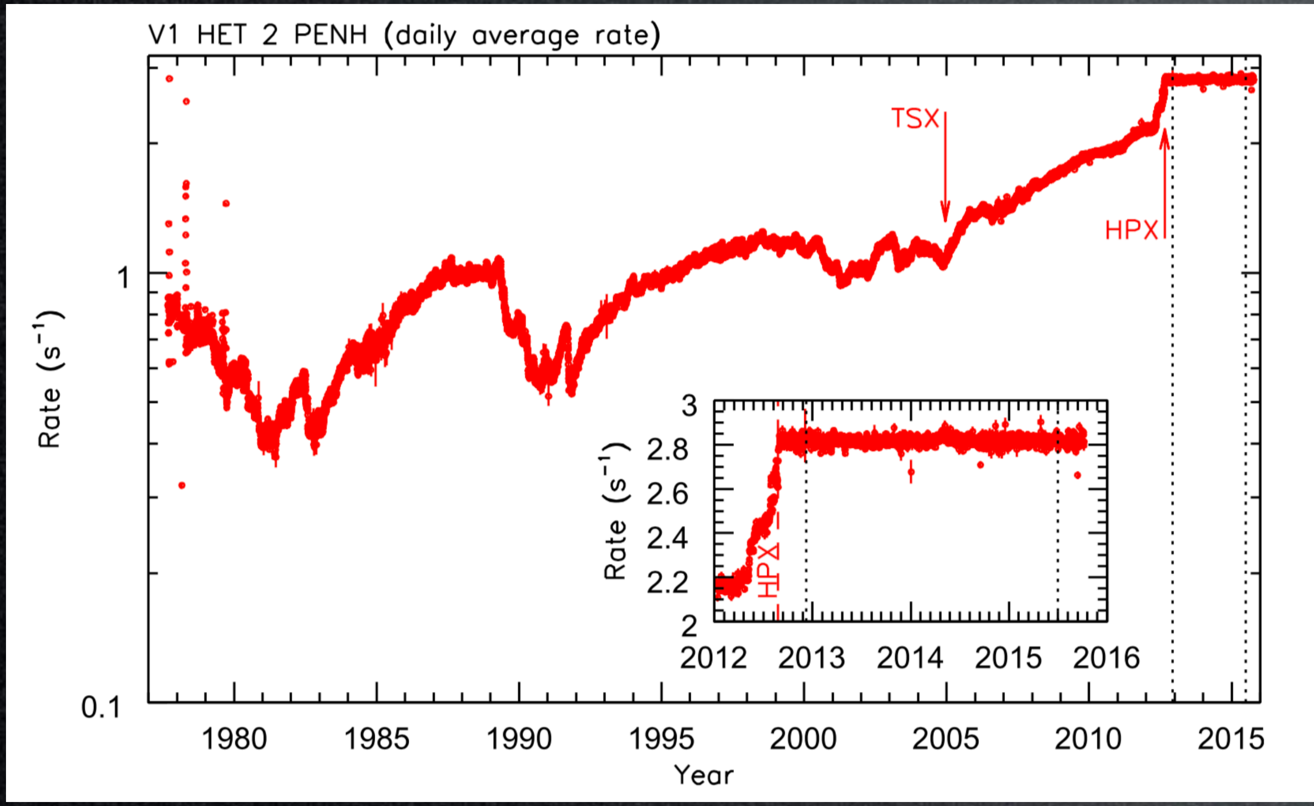
Data: leptons
low energy

Voyager-1 left the heliosphere in 08.2012

First ever measurement of sub-GeV $e^{+}+e^{-}$



courtesy of M. Boudaud, based on Cummings+ (Voyager-1 coll.), The Astrophysical Journal, 831:18, 2016

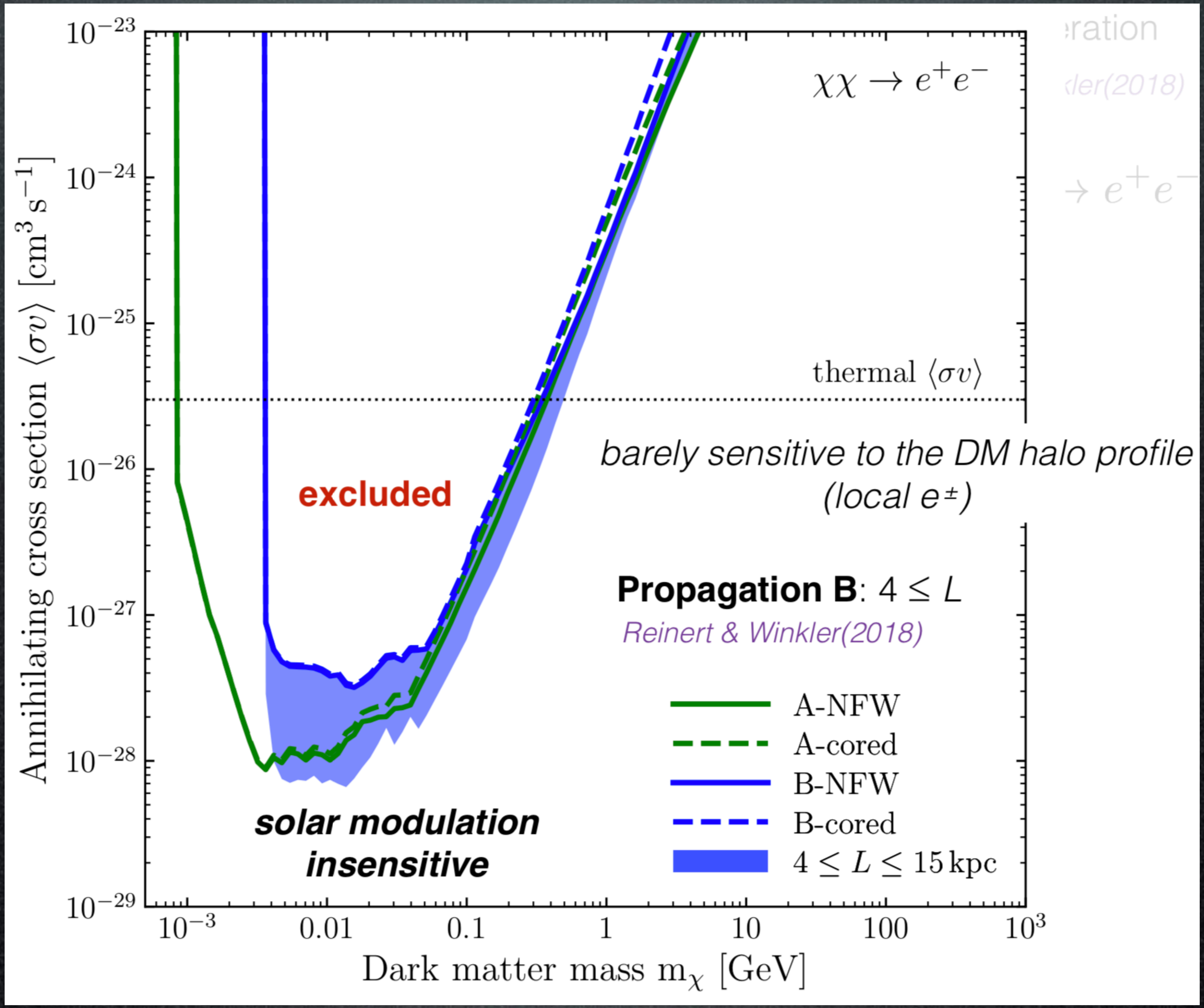


Cummings+ (Voyager-1 coll.), The Astrophysical Journal, 831:18, 2016

Dark Matter interpretation

low energy

Constraints on sub-GeV DM



Boudaud, Lavalle, Salati 1612.07698

Dark Matter interpretation

low energy

Constraints on Primordial Black Holes

DM could consist of PBHs

huge range of sizes:

$$M \simeq 10^{15}(t/10^{-23}\ \mathrm{sec})\ \mathrm{g}$$

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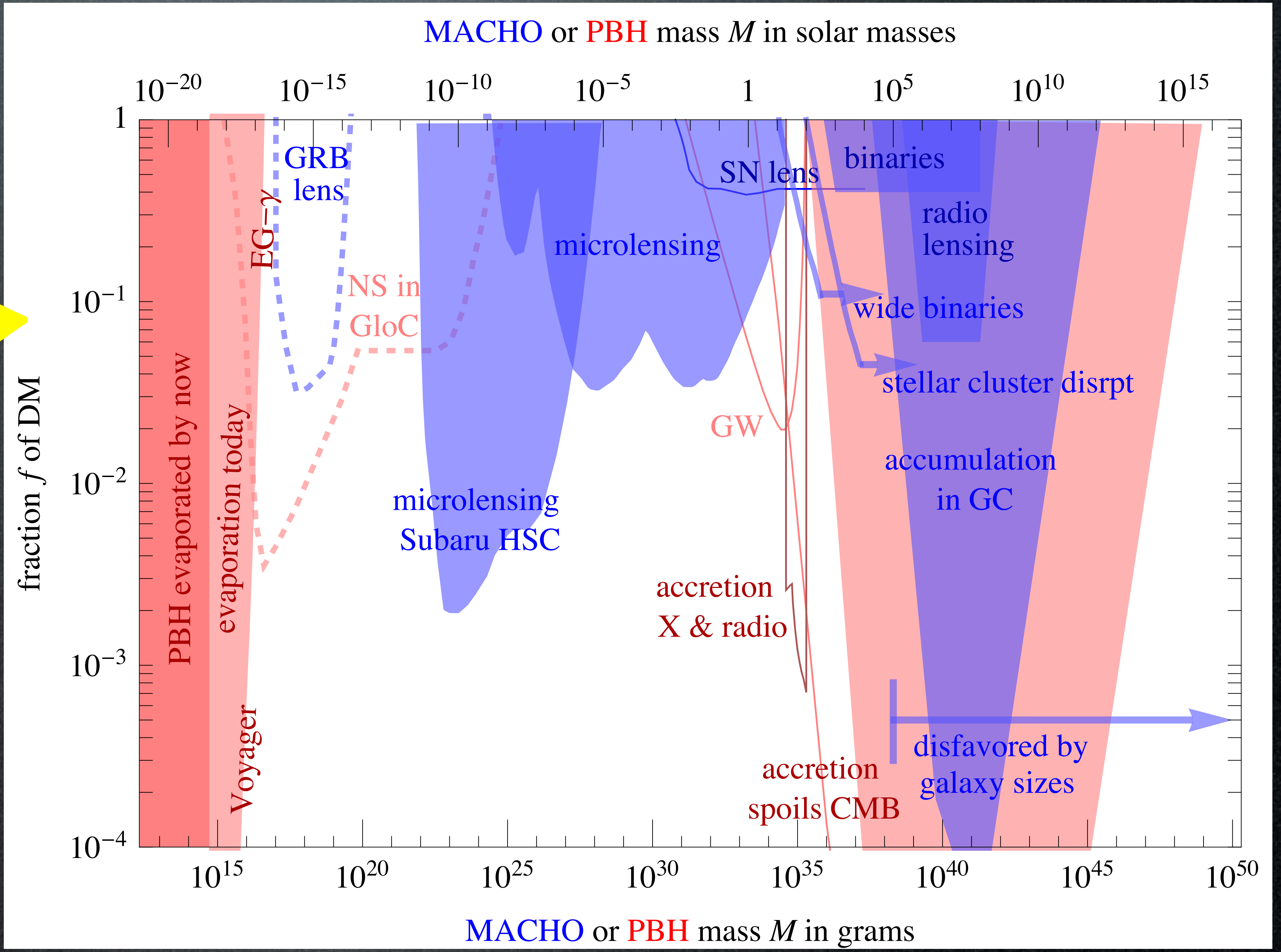

M. Cirelli, A. Strumia, J. Zupan to appear

Dark Matter interpretation

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'small' PBHs emit today by Hawking evaporation

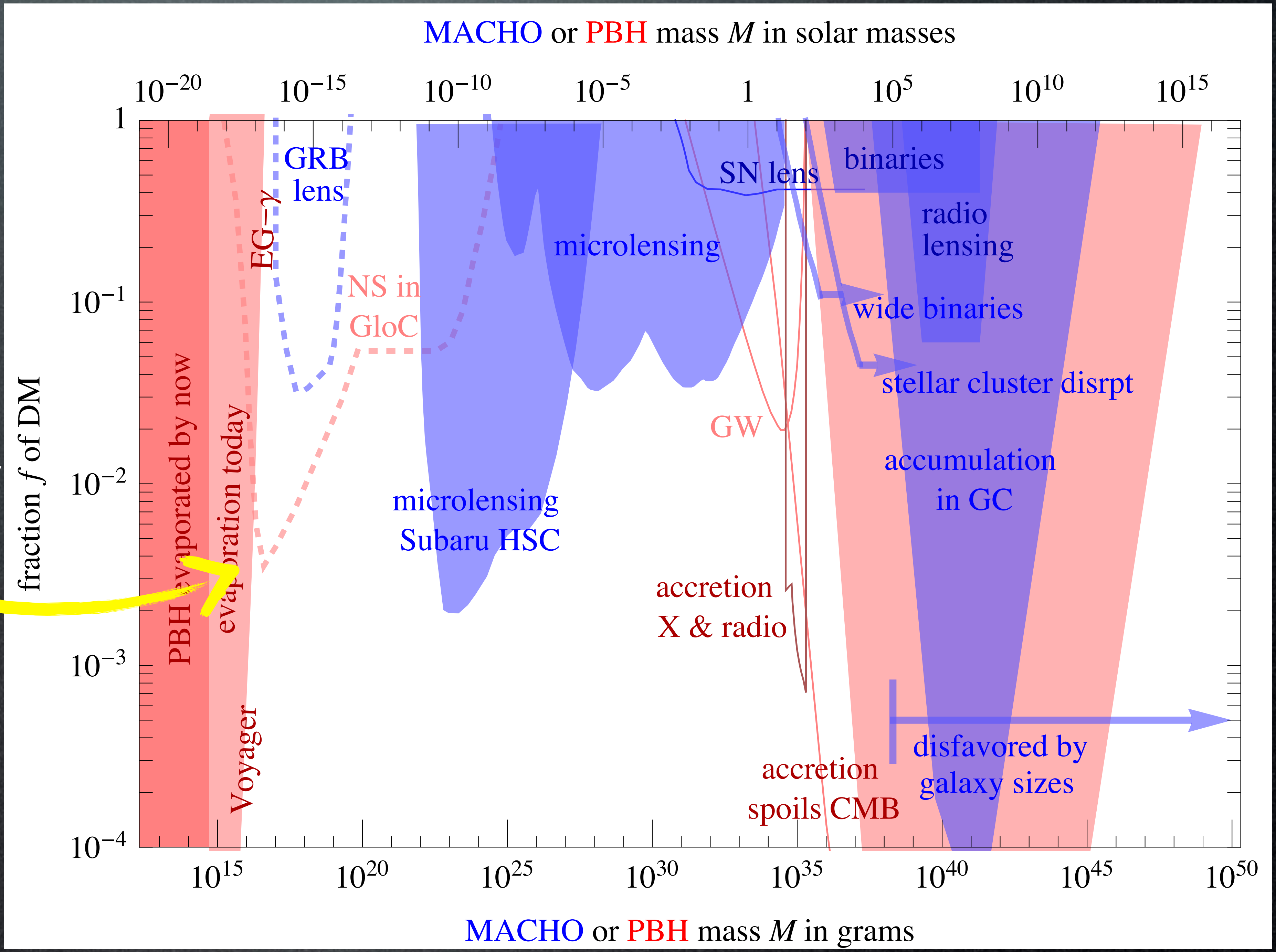
$$T = \frac{1}{8\pi G_N M}$$

rate

$$\frac{dM}{dt} \simeq -5\times 10^{25} f(M) \left(\frac{\text{g}}{M}\right)^2 \text{g/s}$$

spectrum

$$\frac{dN}{dt\,dE} = \frac{27}{2\pi}\frac{G^2M^2E^2}{e^{E/T}+1}$$



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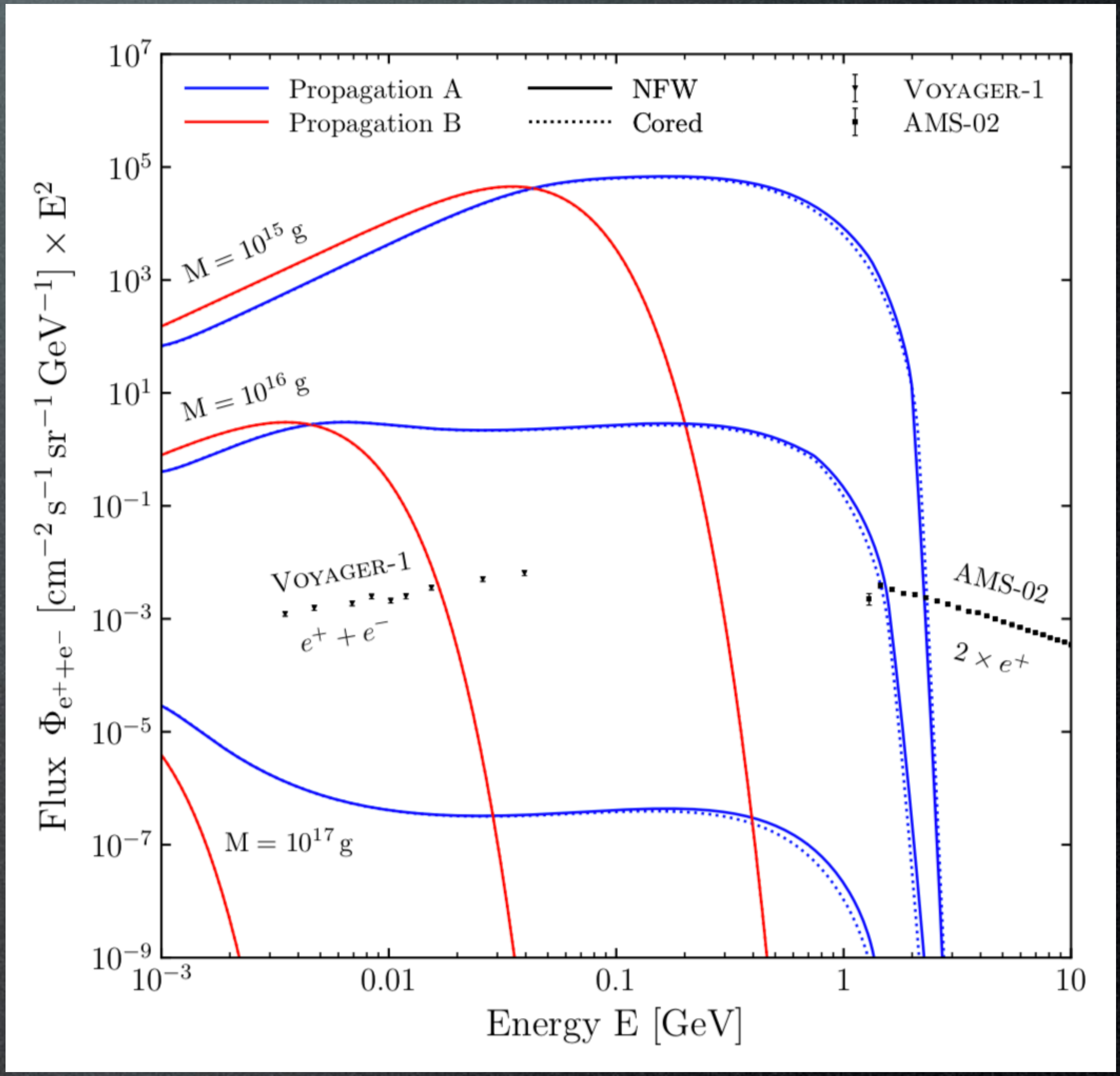

Boudaud, Cirelli 1807.03075, PRL 122 (2019)

Dark Matter interpretation
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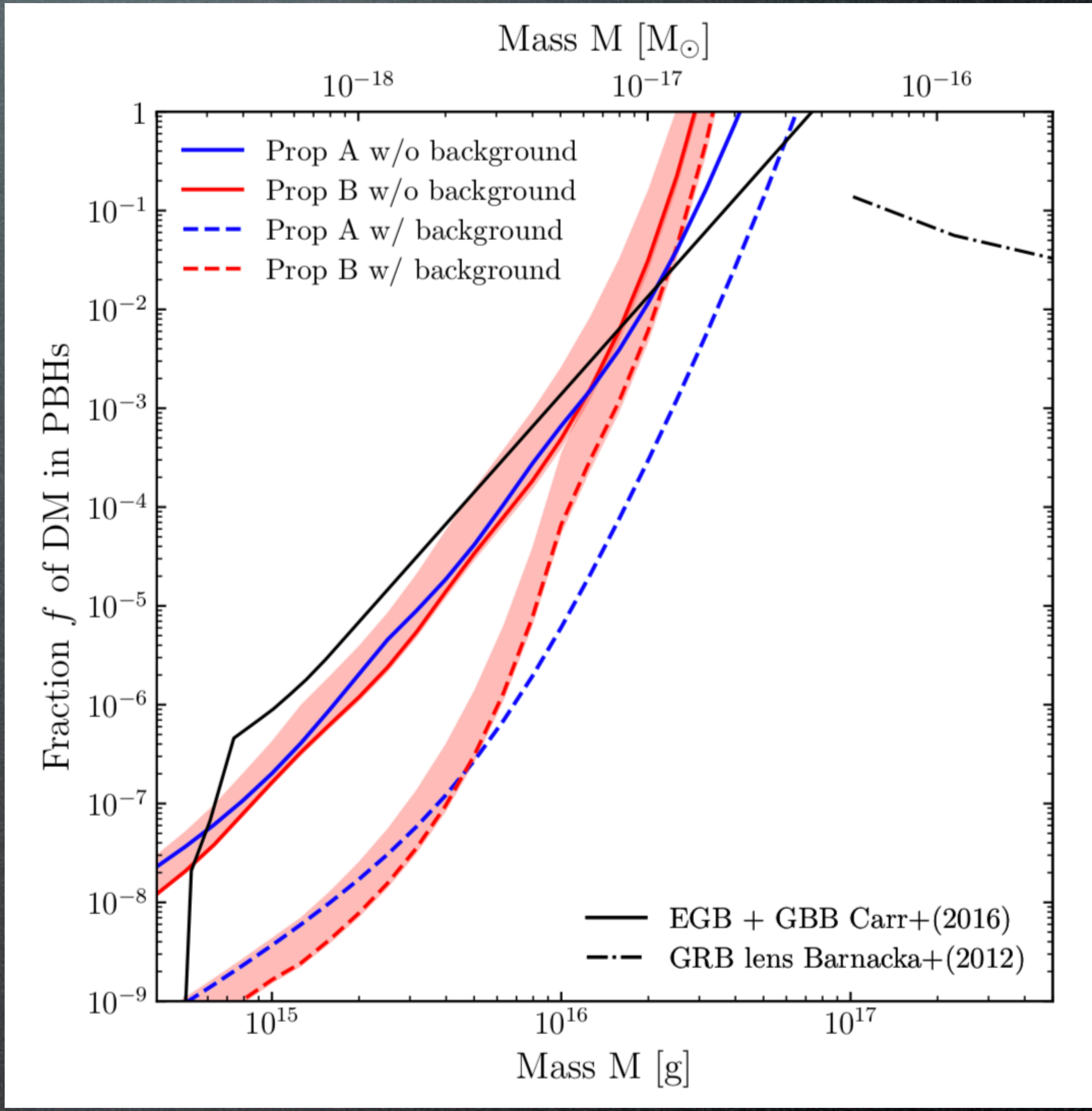


Boudaud, Cirelli 1807.03075, PRL 122 (2019)

Dark Matter interpretation

low energy

Constraints on Primordial Black Holes

An illustration of Voyager 1, now 21.7 billion kilometers away JPL CALTECH/NASA

Aging Voyager 1 spacecraft undermines idea that dark matter is tiny black holes

By Adrian Cho | Jan. 9, 2019 , 2:25 PM

Forbes

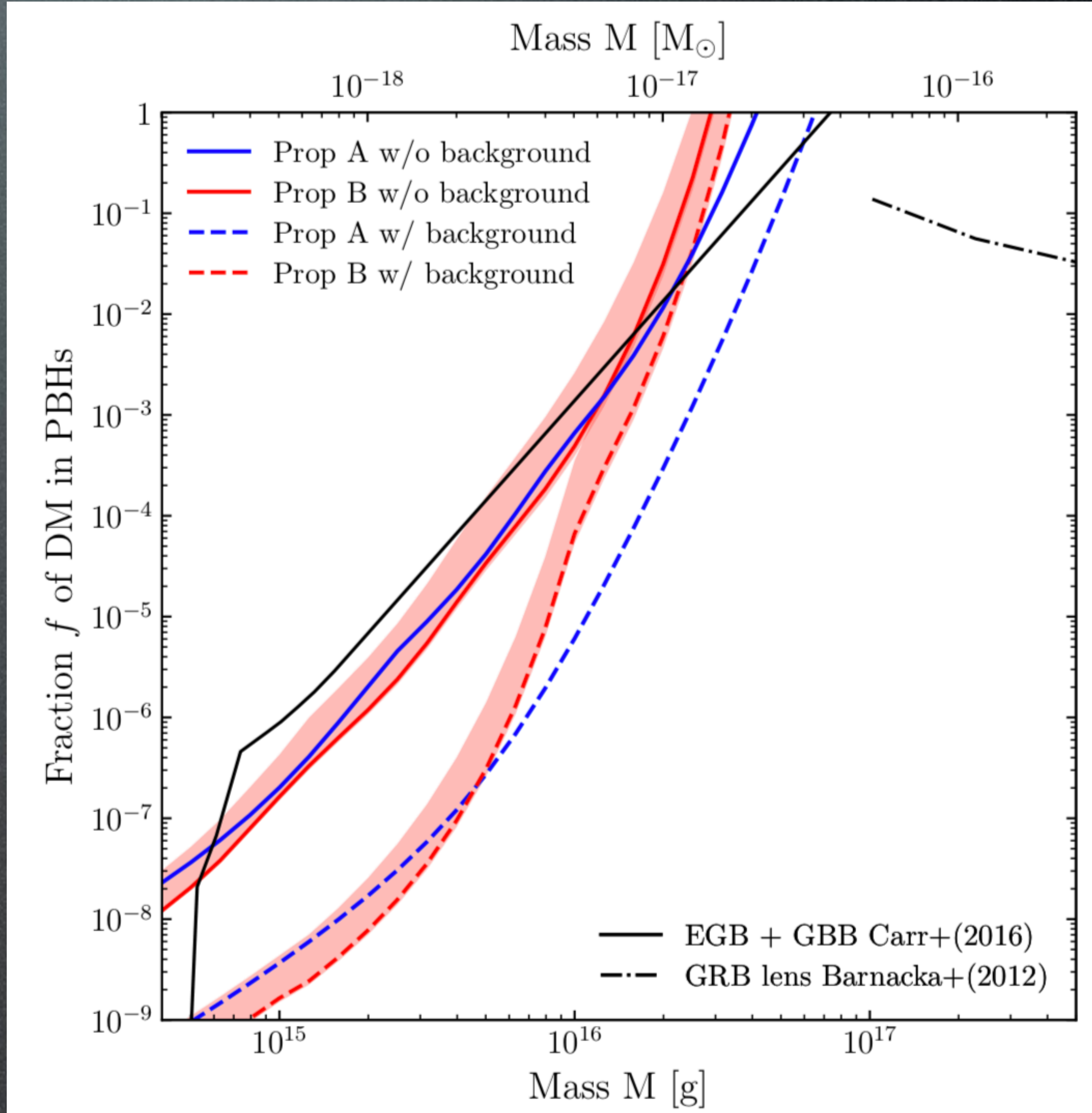
25,121 views | Jul 10, 2018, 05:59pm

NASA's Voyager-1 Spacecraft Opens Door On New Way To Look For Dark Matter

Bruce Dorminey Contributor ⓘ

Science

I cover over-the-horizon technology, aerospace and astronomy.



Boudaud, Cirelli 1807.03075, PRL 122 (2019)

Gamma rays

direct detection

production at colliders

indirect

γ from annihil in galactic center or halo
and from secondary emission

Fermi, ICT, radio telescopes...

e^+ from annihil in galactic halo or center

$\bar{p}$ from annihil in galactic halo or center

$\bar{d}$ from annihil in galactic halo or center

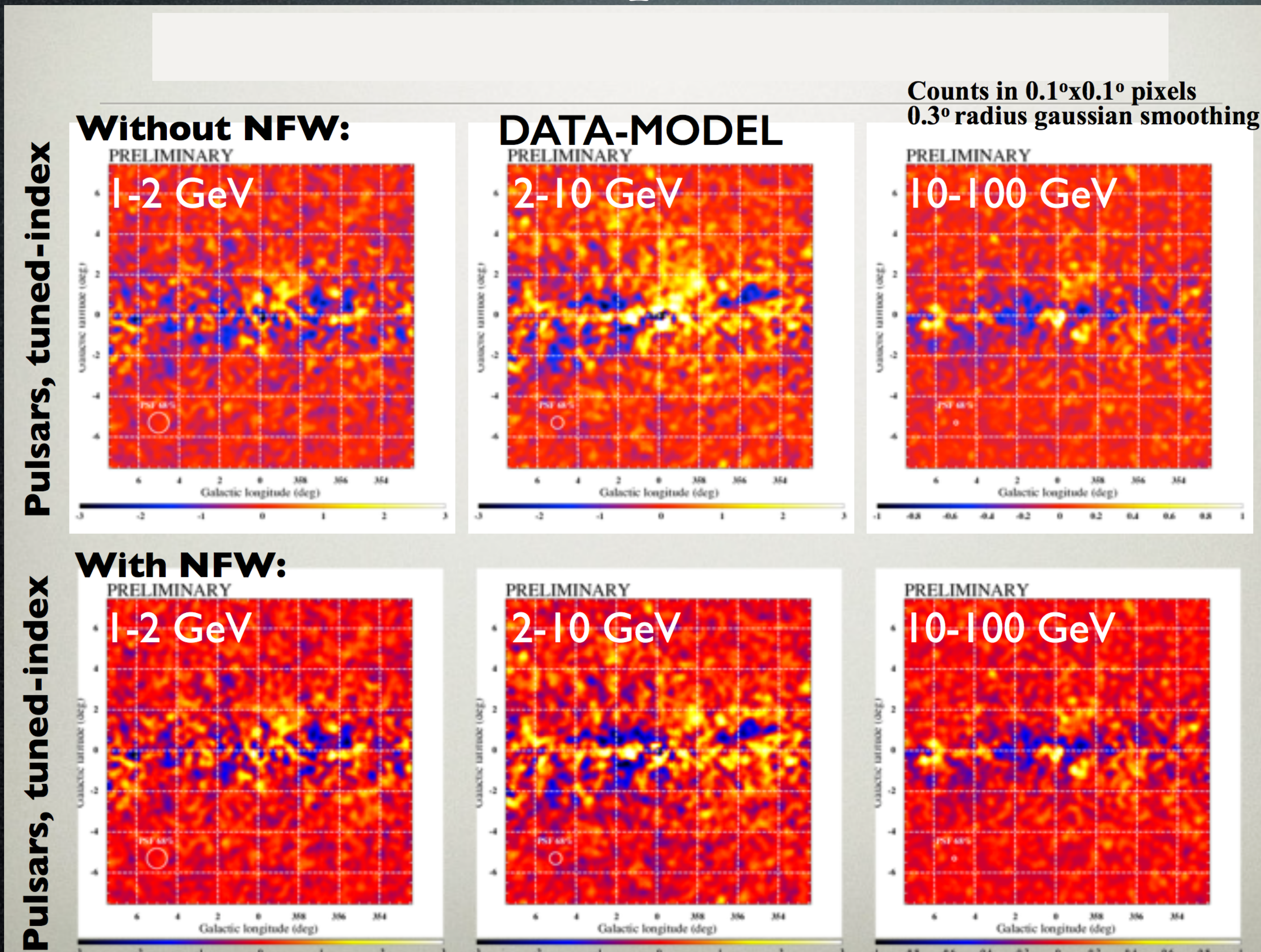
GAPS, AMS

$\nu, \bar{\nu}$ from annihil in massive bodies

SK, Icecube, Antares

GC GeV excess

Dark Matter interpretation:



S. Murgia for FERMI-LAT - ICRC 2015
T. Porter for FERMI-LAT - ICRC 2015 #815
Fermi coll. 1511.02938

GC GeV excess

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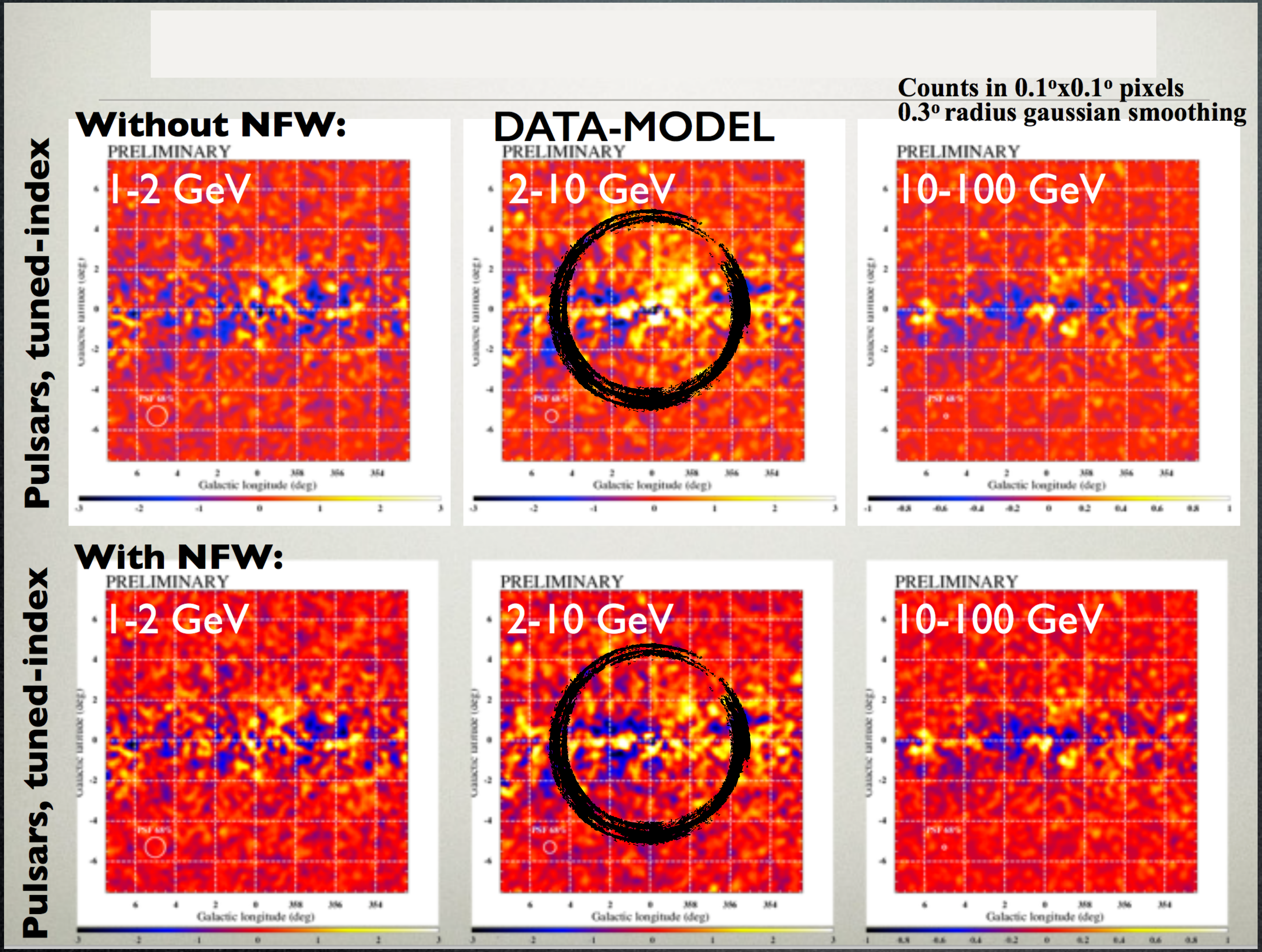

S. Murgia for FERMI-LAT - ICRC 2015
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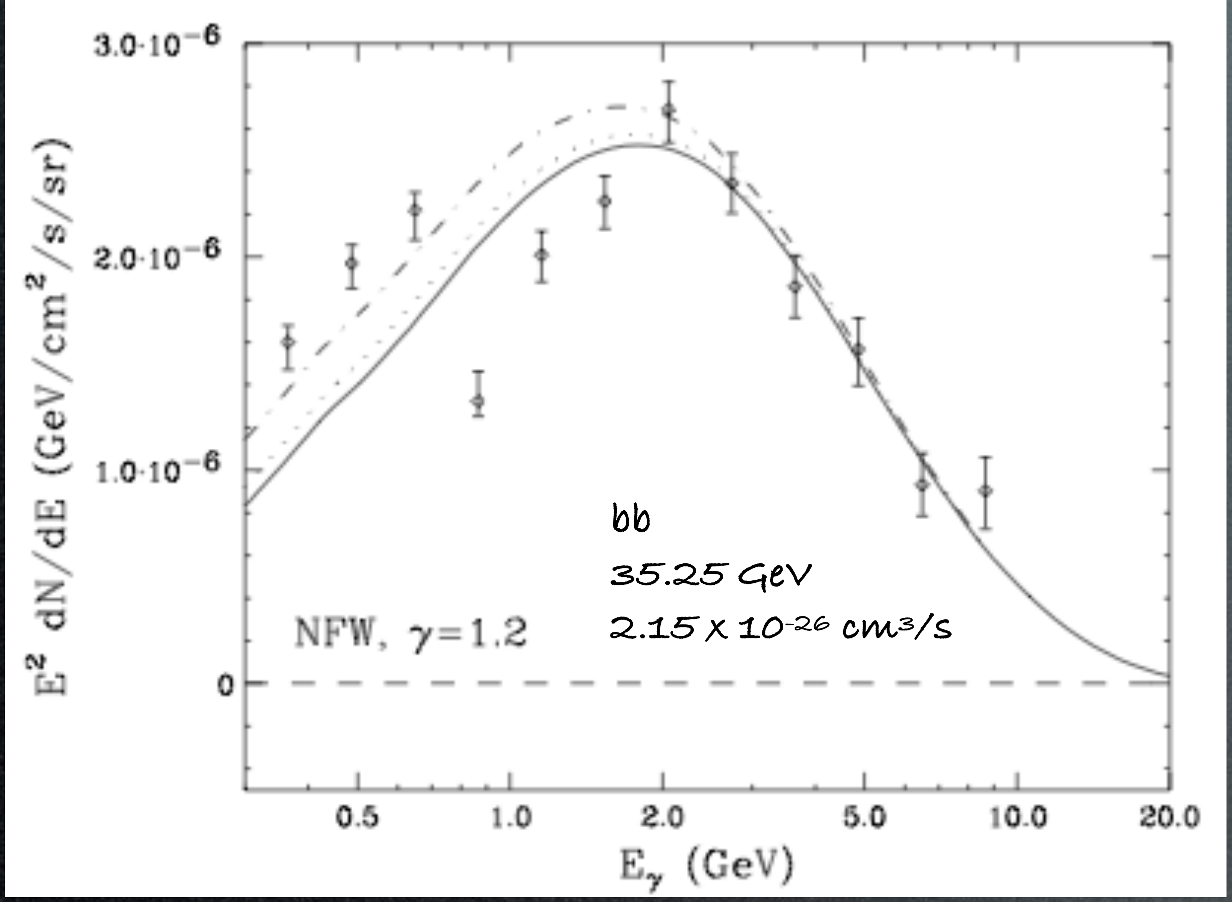
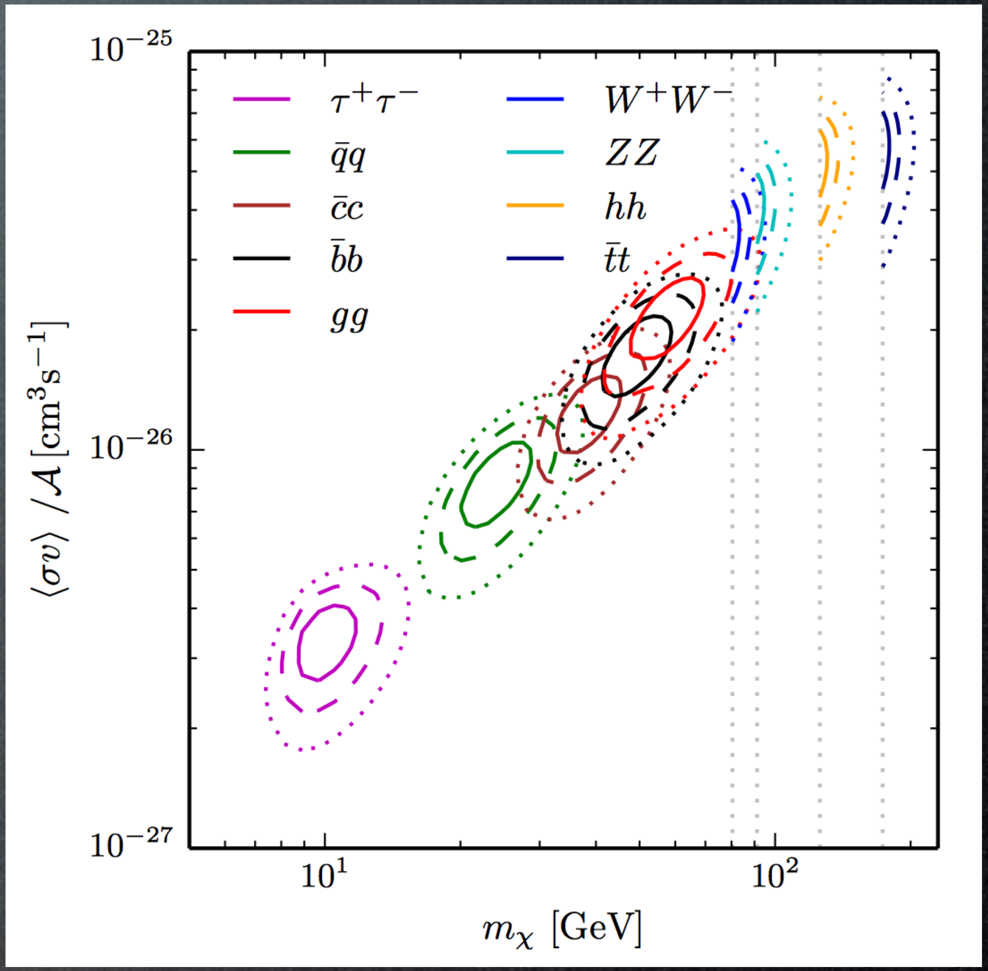
GC GeV excess

Dark Matter interpretation:

Best fit:
~35 GeV, quarks, ~thermal σv





F. Calore et al. 1411.4647

A compelling case for annihilating DM

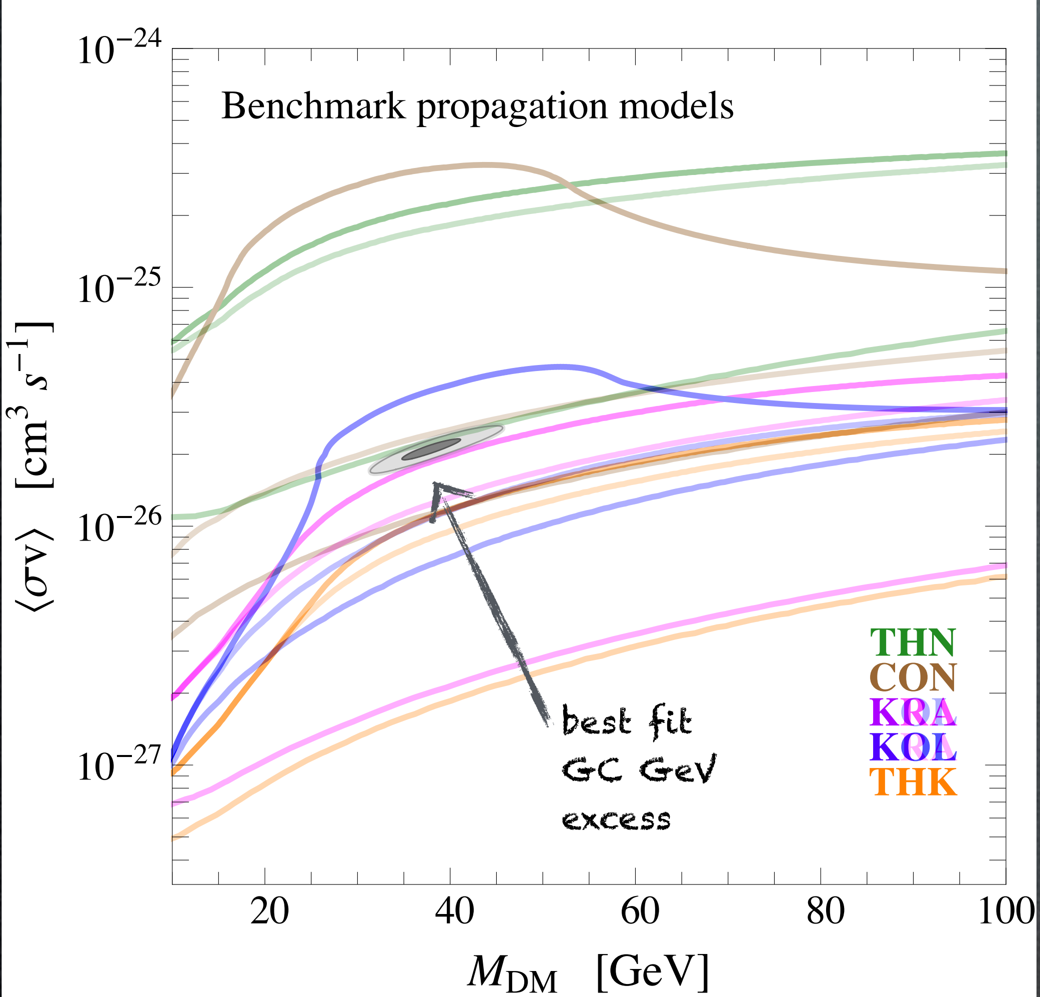
Daylan, Finkbeiner, Hooper, Linden, Portillo, Rodd, Slatyer 1402.6703

...as good as it can get.

GC GeV excess

Dark Matter interpretation:

Antiproton constraints
are not conclusive



Cirelli, Gaggero,
Giesen, Taoso,
Urbano 1407.2173

Also:

Bringmann, Vollmann,
Weniger 1406.6027

Hooper, Linden, Mertsch
1410.1527

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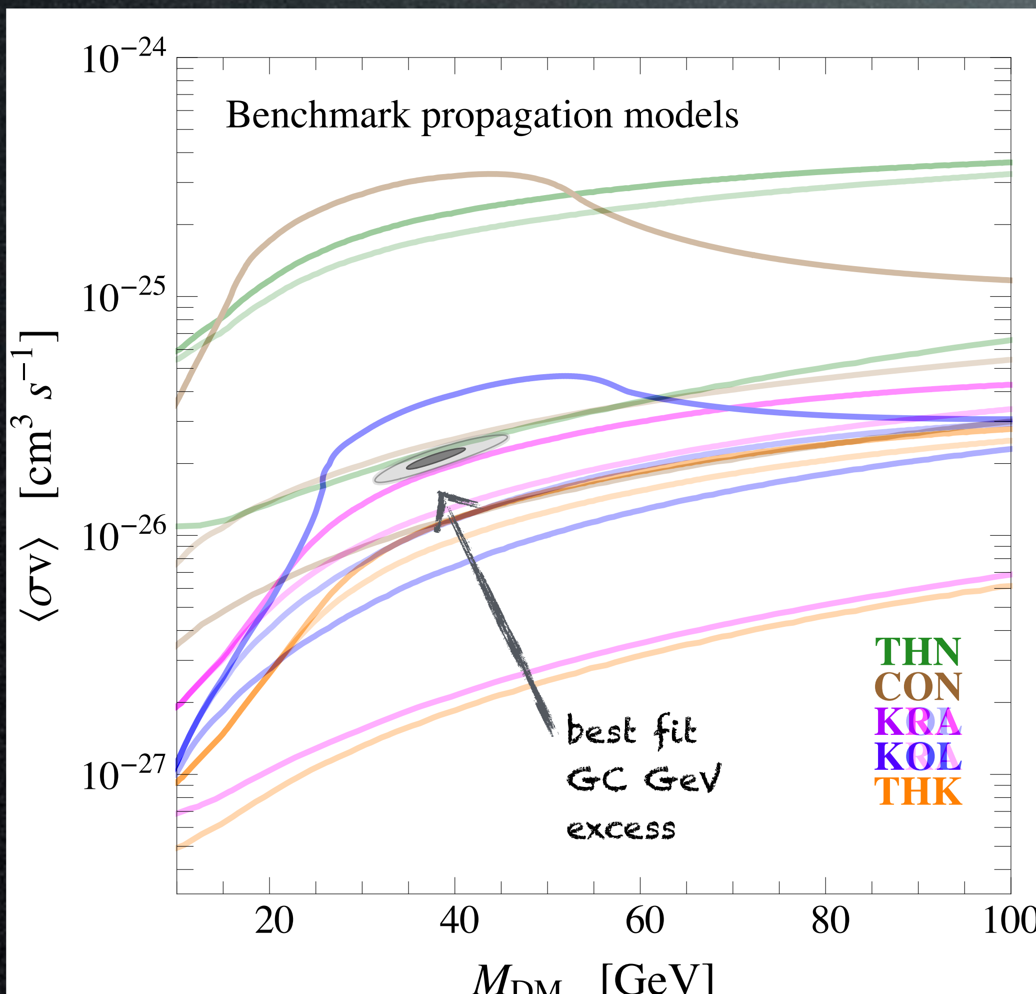

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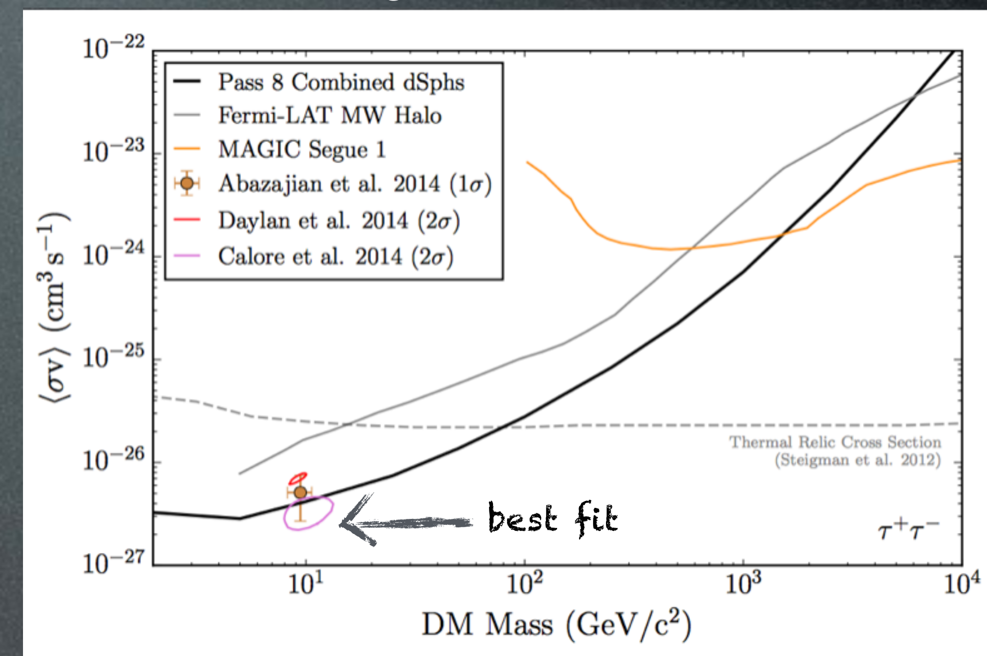
Gamma ray ones neither



FERMI 1503.02641

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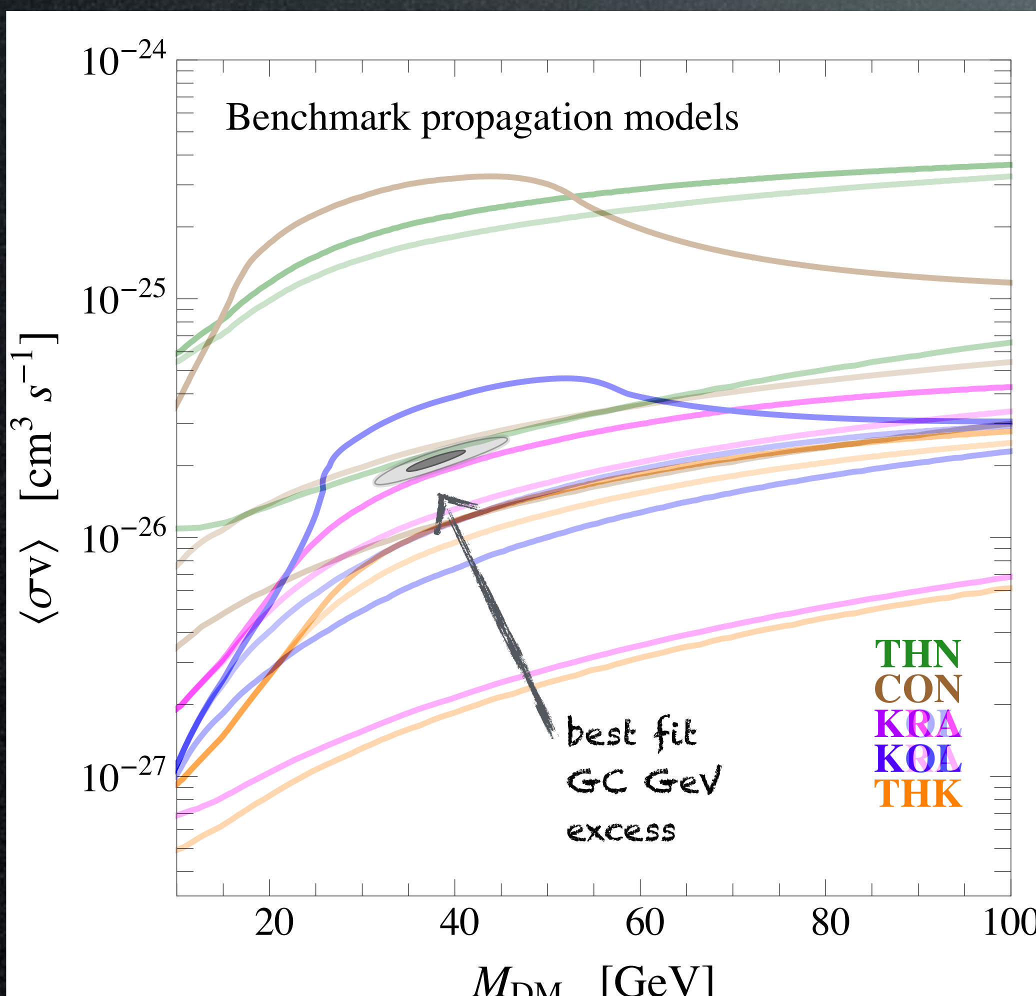

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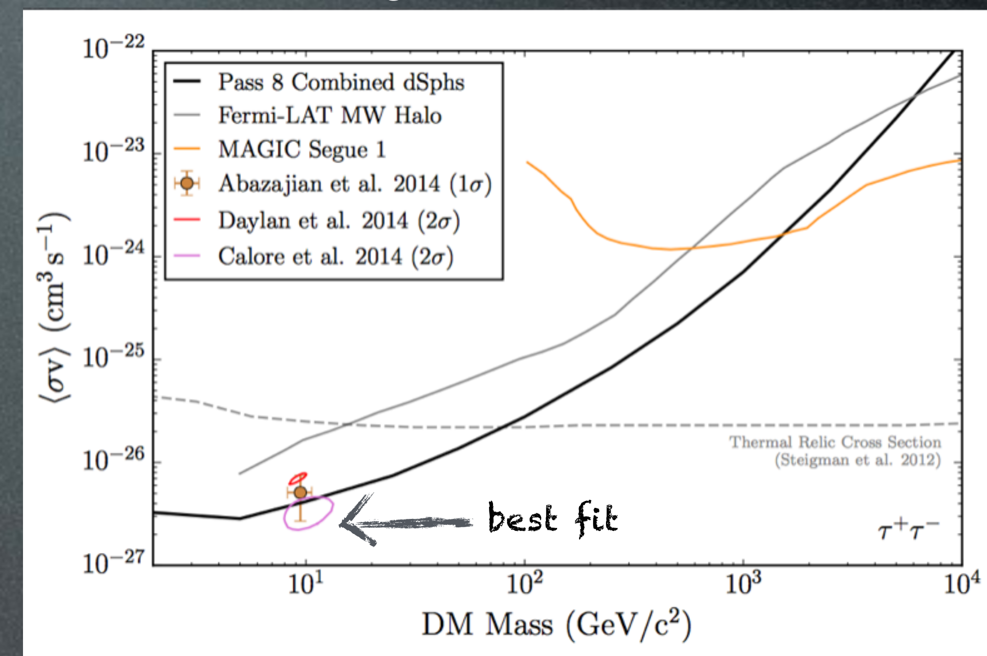
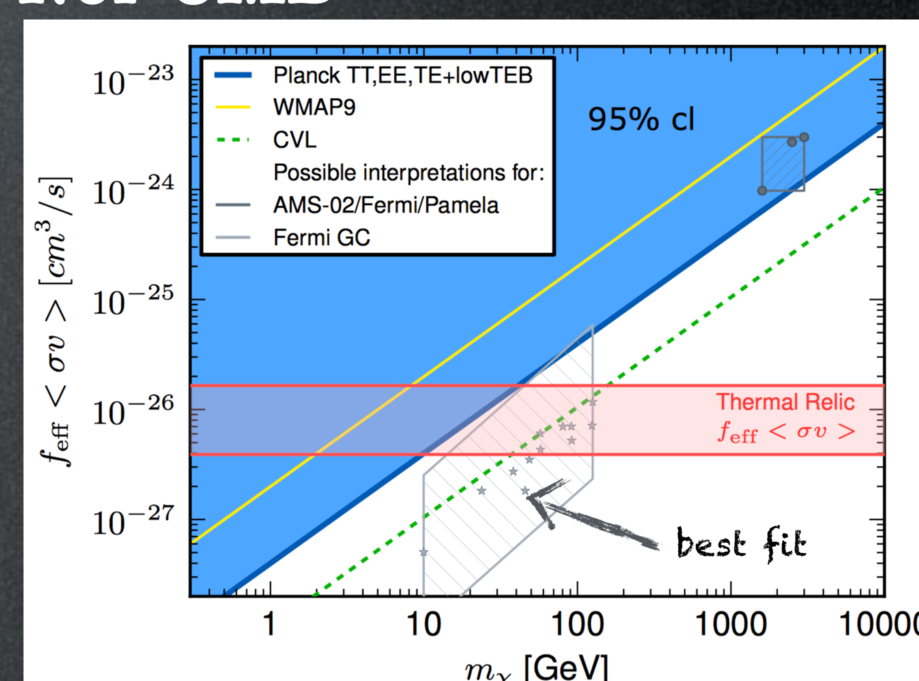


FERMI 1503.02641

Nor CMB



Planck
2015

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Bringmann, Vollmann,
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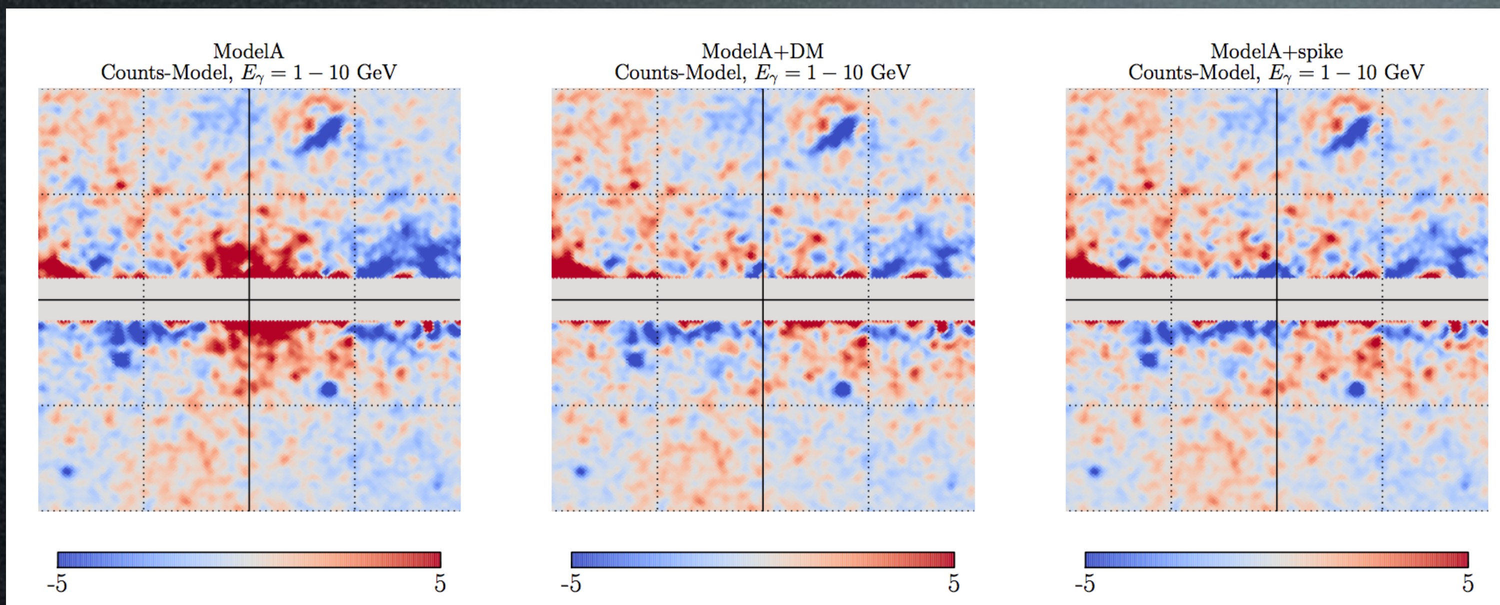
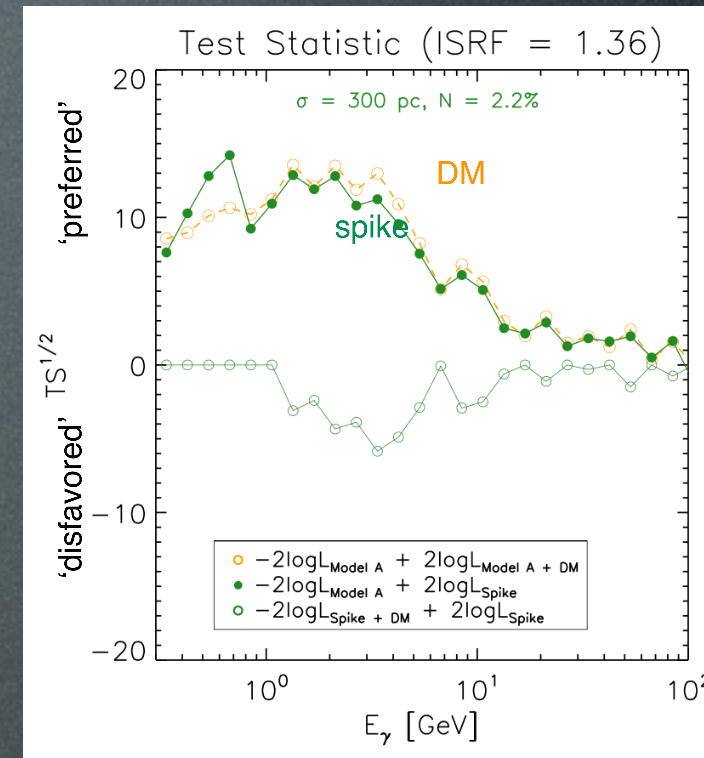
1410.1527

GC GeV excess

'Astro' interpretation(s):

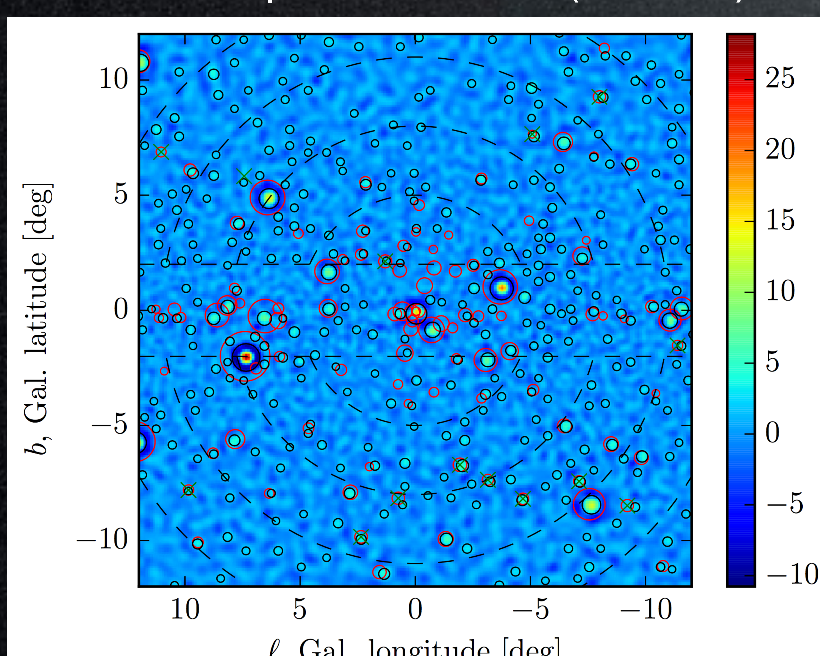




An additional steady-source spike of CRs (from SNRs?) that emit via ICS

D. Gaggero et al 1507.06129

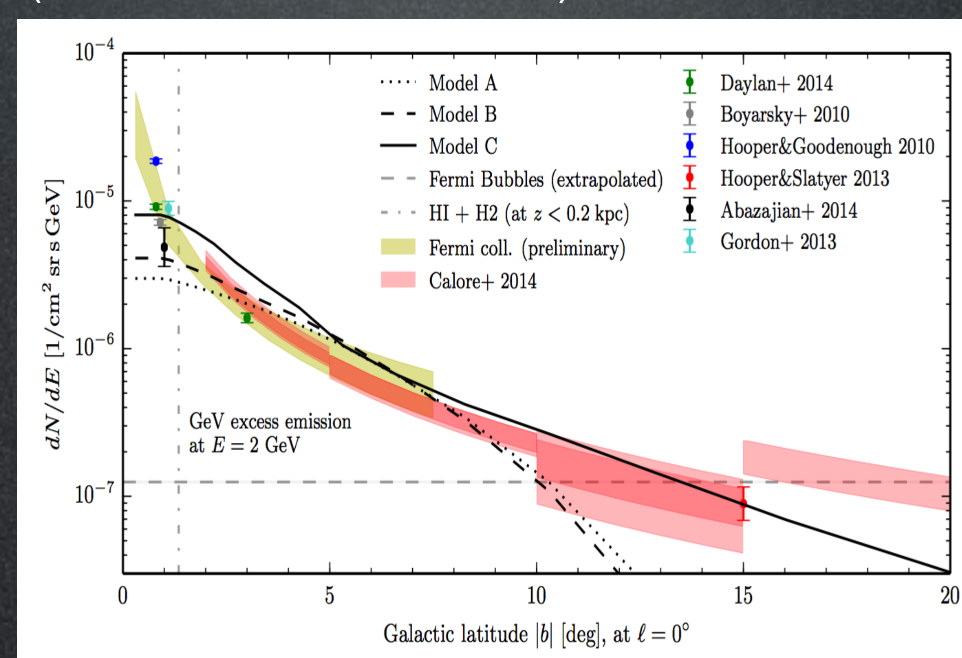
Unresolved point sources (MSPs?)



Bartels…Weniger 1506.05104

Lee, Lisanti…Slatyer 1506.05124

Leptonic outbursts: old + young (1 + 0.1 Myr)
(but even this is not ideal)



F. Calore 1506.05119

What does the FERMI coll. say?

Unclear…

- Excess exists (1511.02938), adding DM improves the fit.
- Excesses elsewhere in the GP, the GC one not significant (1704.03910).
- We found point sources! DM 'strongly disfavored' (1705.00009v1).
- *Sure? (Bartels et al., 1710.10266)*
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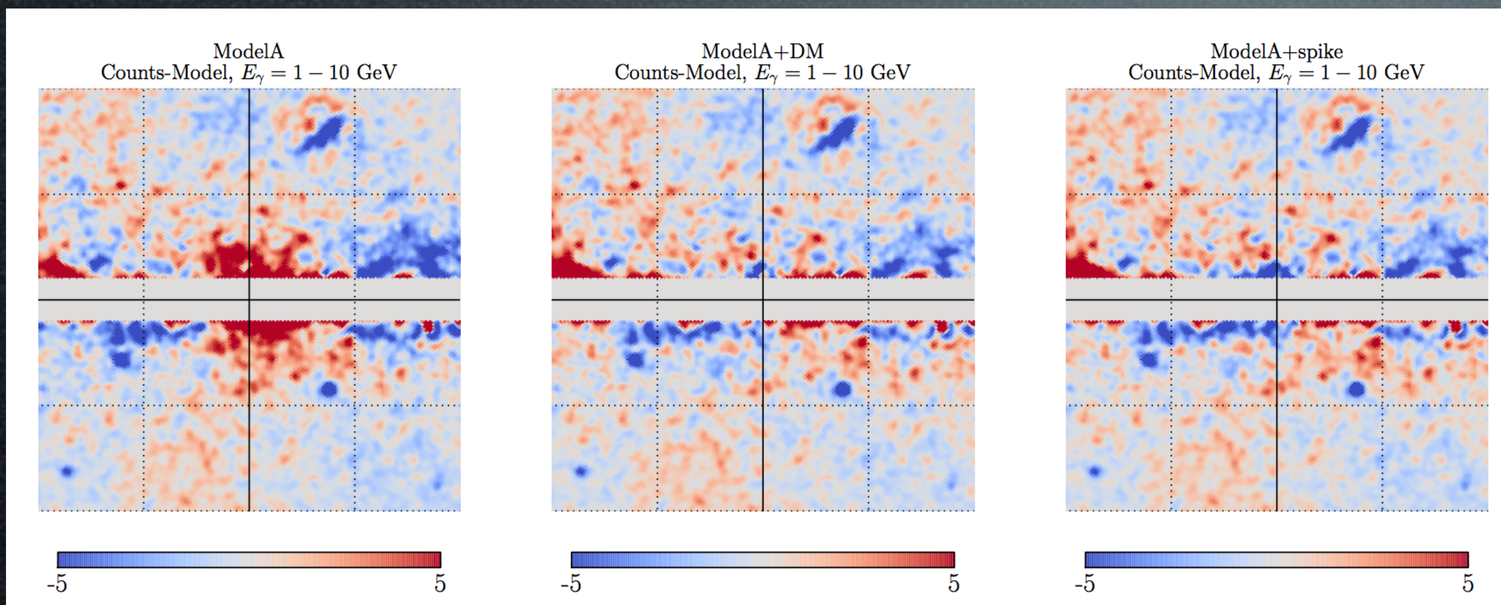

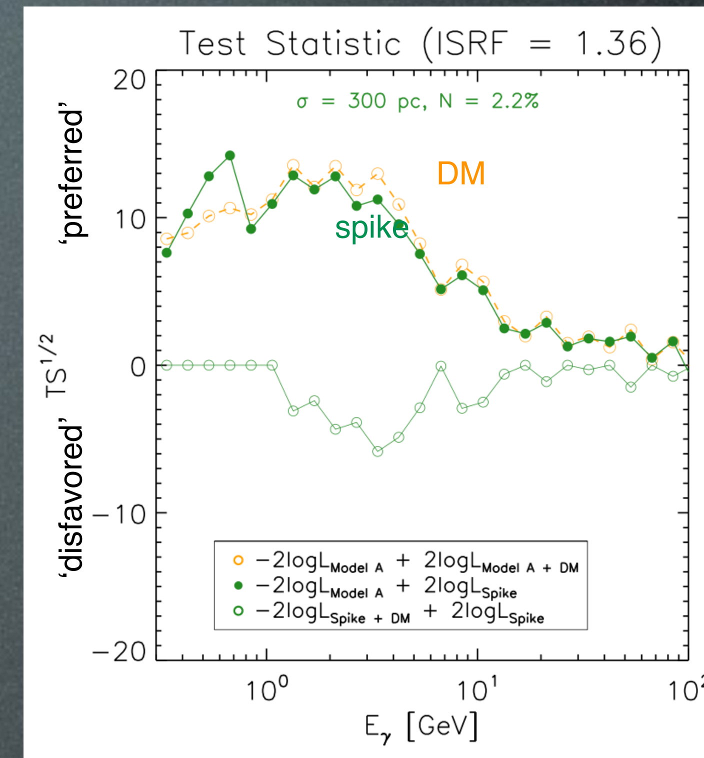

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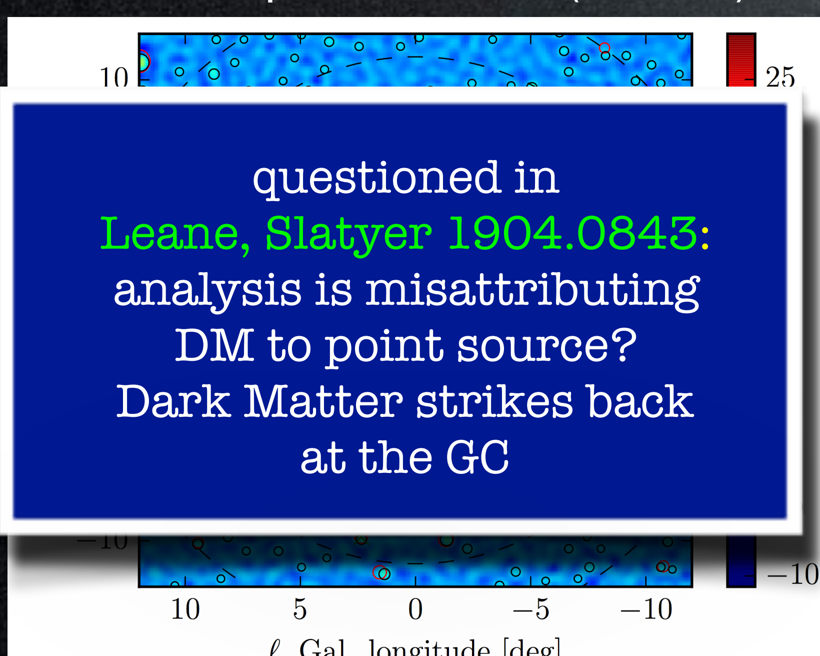

questioned in
Leane, Slatyer 1904.0843:
analysis is misattributing
DM to point source?
Dark Matter strikes back
at the GC

Bartels…Weniger 1506.05104
Lee, Lisanti…Slatyer 1506.05124

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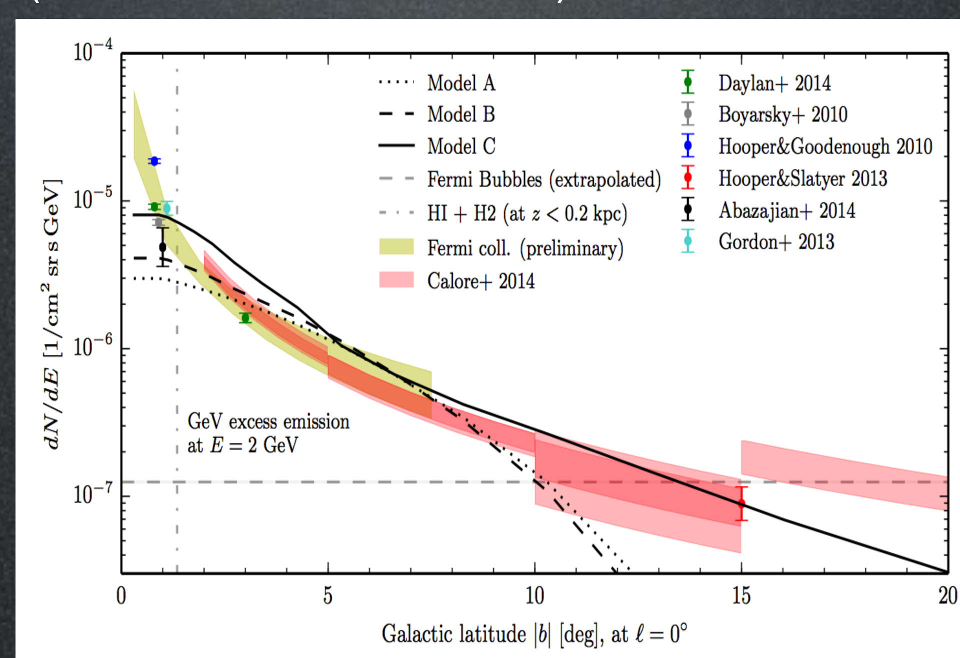

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PS: excesses in Andromeda?

Fermi has also observed M31 (and M33):
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Ackermann et al, FERMI Coll. 1702.08602

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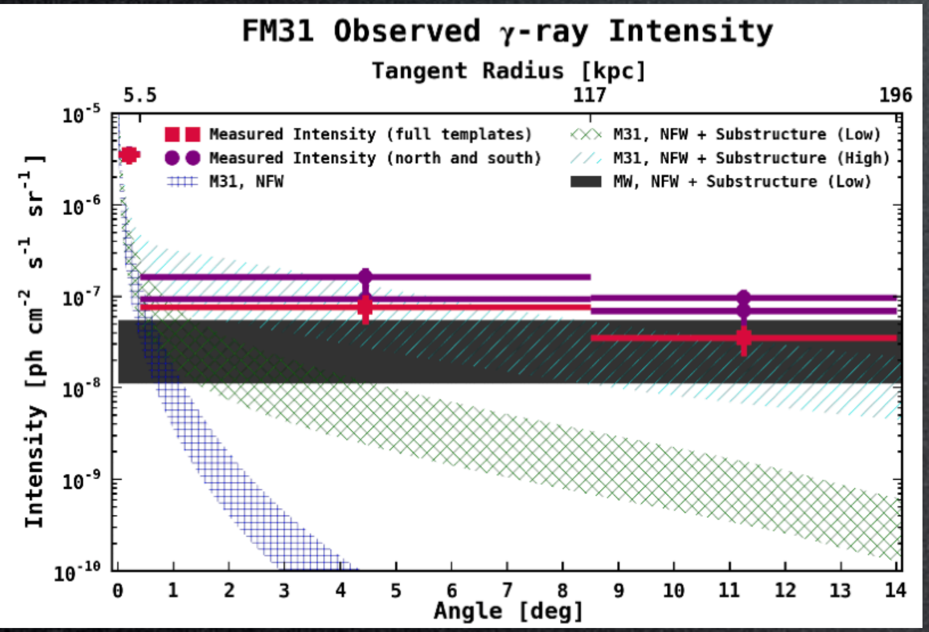
Ackermann et al, FERMI Coll. 1702.08602

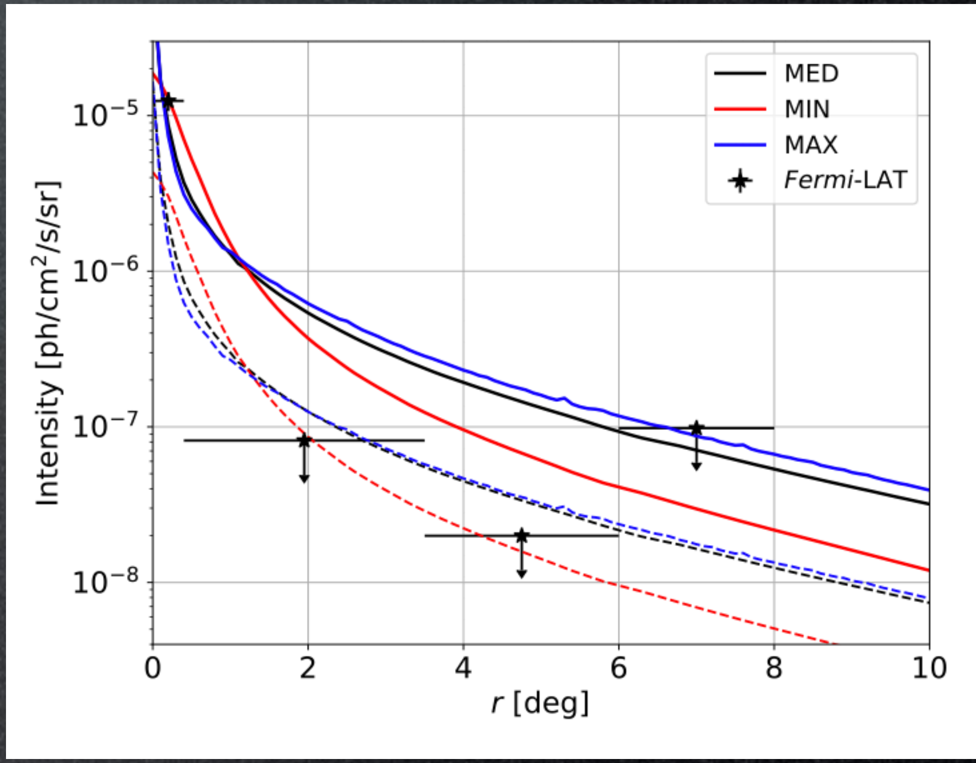
an excess from the center of M31 **but:** intensity = 5 x GC GeV excess

Murgia, Moskalenko et al 1903.10533
excess from the outer halo
spectrum agrees w GC GeV excess



but: fitting with DM requires huge subhalo boost, and in any case MW DM emission a.l.o.s. contributes at least as much, or more

Di Mauro, Zaharijas, Charles et al 1904.10977
no excess
upper limits only



DM detection

direct detection

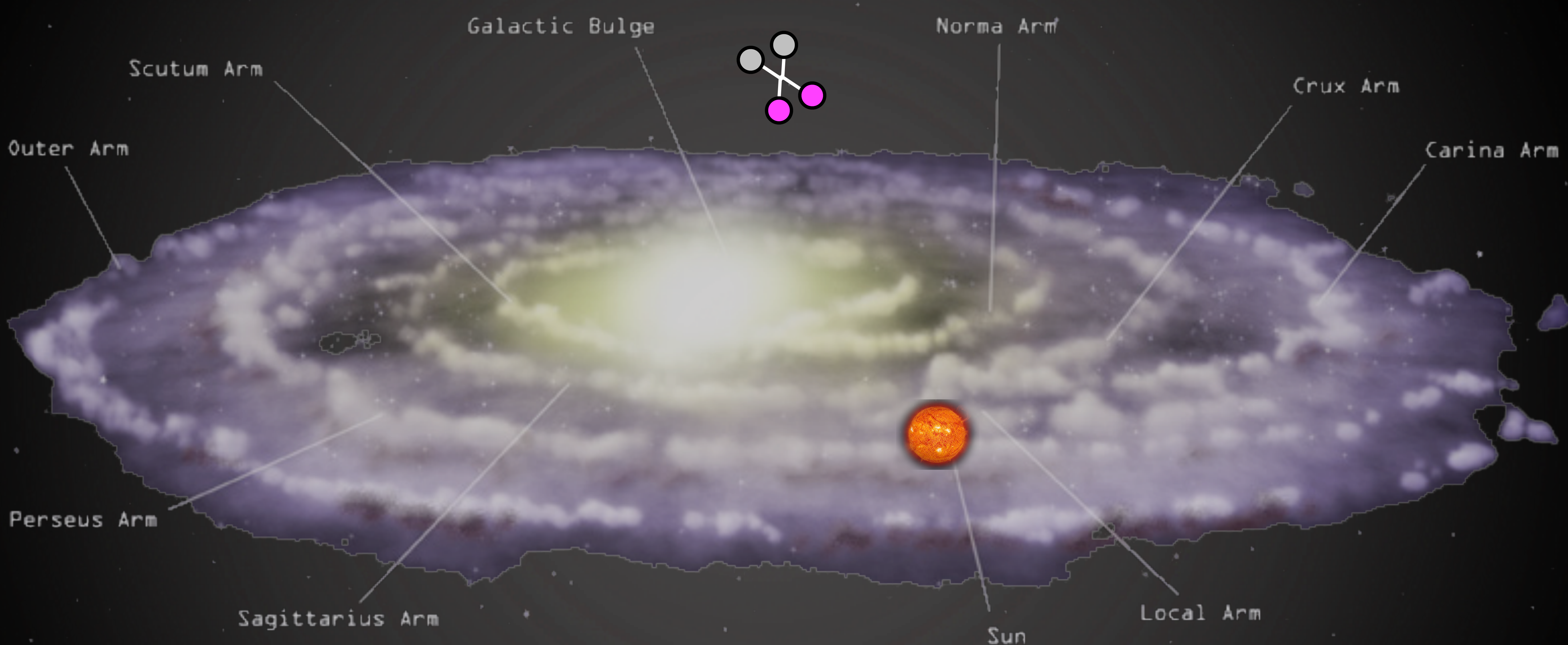
production at colliders

indirect

- γ from annihil in galactic center or halo and from secondary emission
  Fermi, ICT, radio telescopes...
- e^+ from annihil in galactic halo or center
  PAMELA, Fermi, HESS, AMS,balloons...
- $\bar{p}$ from annihil in galactic halo or center
- $\bar{d}$ from annihil in galactic halo or center
  GAPS, AMS
- $\nu, \bar{\nu}$ from annihil in massive bodies
  SK, Icecube, Km3Net

Indirect Detection

$\bar{d}$ from DM annihilations in halo



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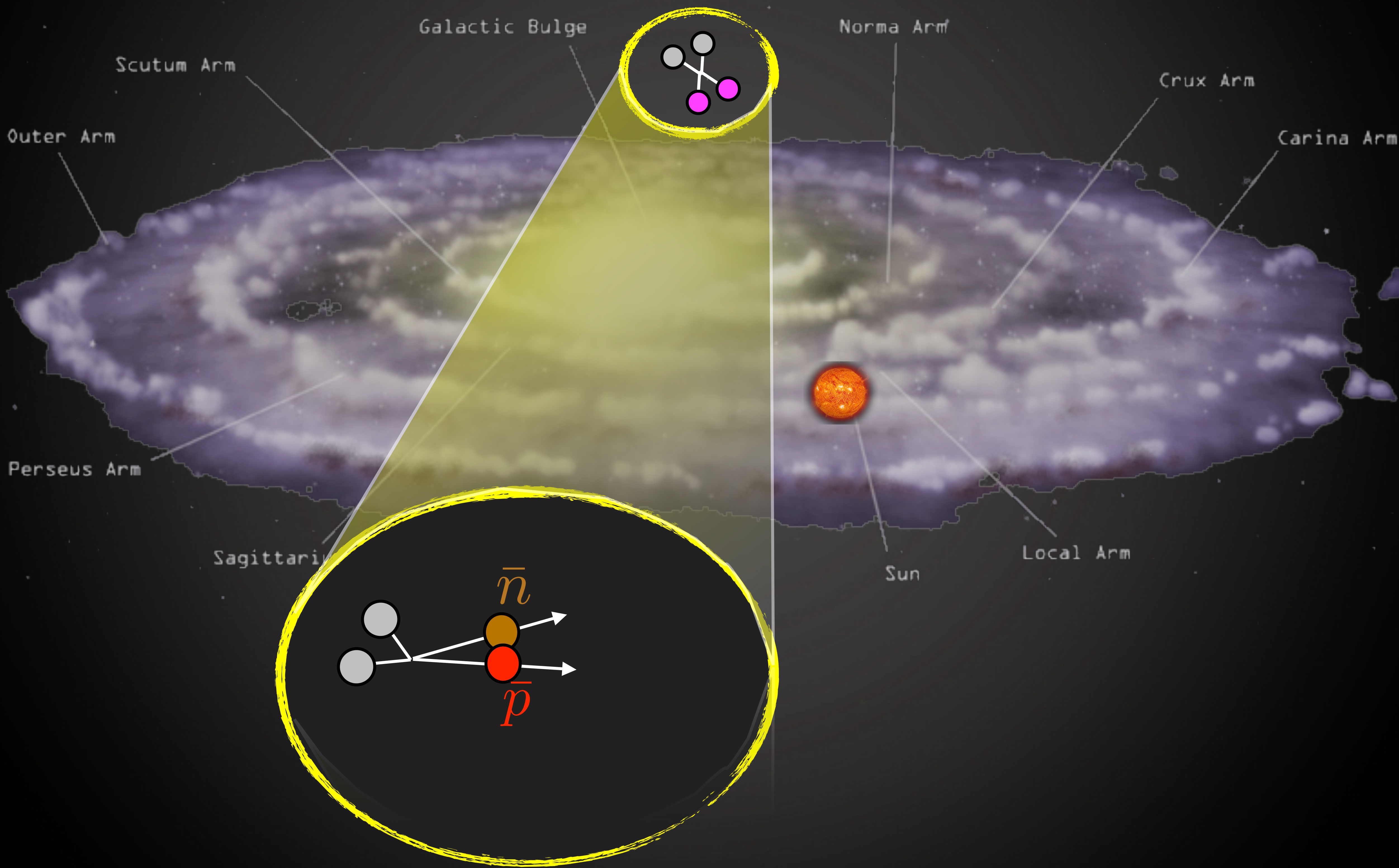

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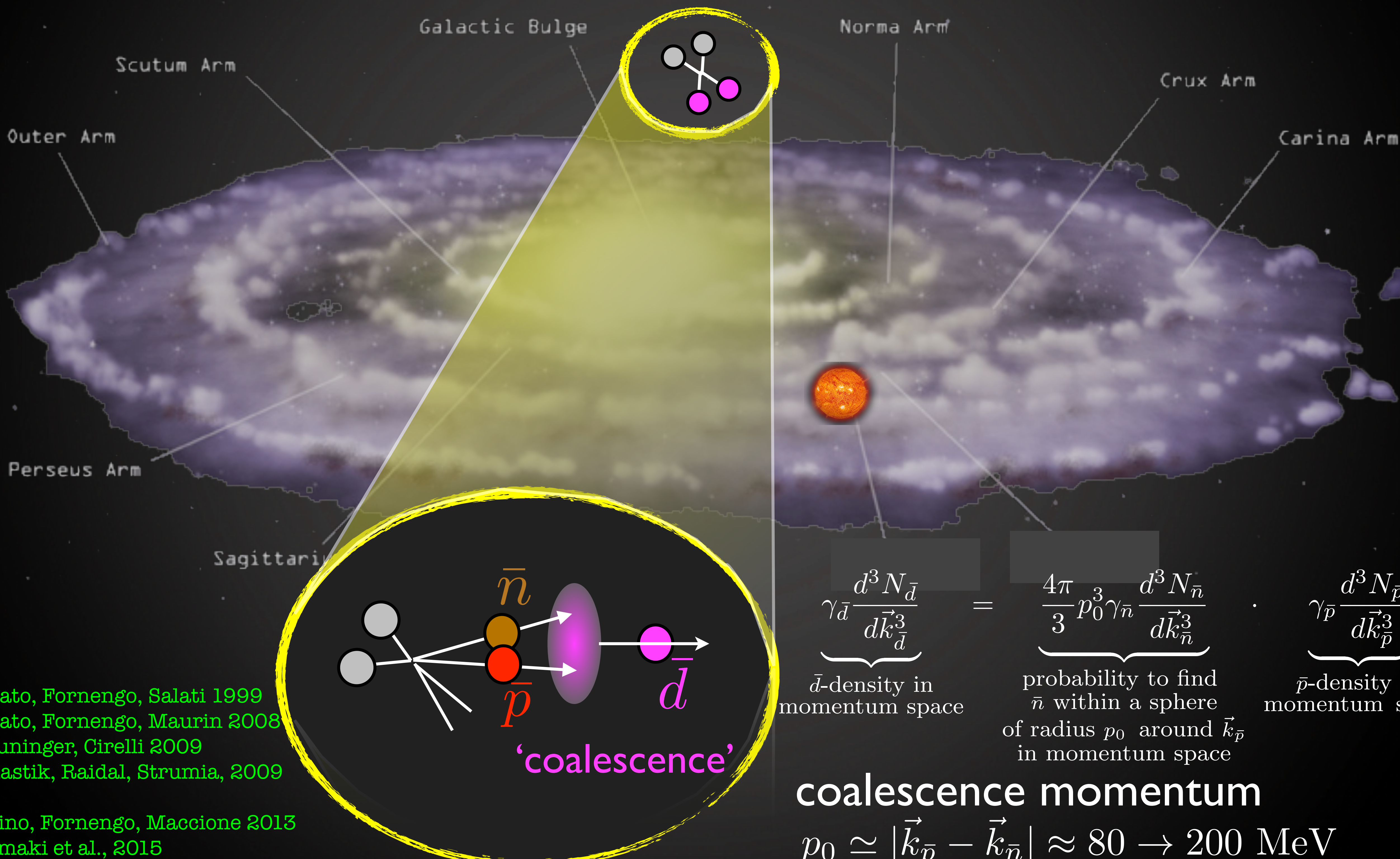

$$\underbrace{\gamma_{\bar{d}} \frac{d^3 N_{\bar{d}}}{d\vec{k}^3_{\bar{d}}}}_{\bar{d}\text{-density in momentum space}} = \underbrace{\frac{4\pi}{3} p_0^3 \gamma_{\bar{n}} \frac{d^3 N_{\bar{n}}}{d\vec{k}^3_{\bar{n}}}}_{\text{probability to find } \bar{n} \text{ within a sphere of radius } p_0 \text{ around } \vec{k}_{\bar{p}} \text{ in momentum space}} \cdot \underbrace{\gamma_{\bar{p}} \frac{d^3 N_{\bar{p}}}{d\vec{k}^3_{\bar{p}}}}_{\bar{p}\text{-density in momentum space}}$$

Donato, Fornengo, Salati 1999
Donato, Fornengo, Maurin 2008
Bräuninger, Cirelli 2009
Kadastik, Raidal, Strumia, 2009
...
Vittino, Fornengo, Maccione 2013
Aramaki et al., 2015

coalescence momentum

$$p_0 \simeq |\vec{k}_{\bar{p}} - \vec{k}_{\bar{n}}| \approx 80 \to 200 \text{ MeV}$$

Indirect Detection

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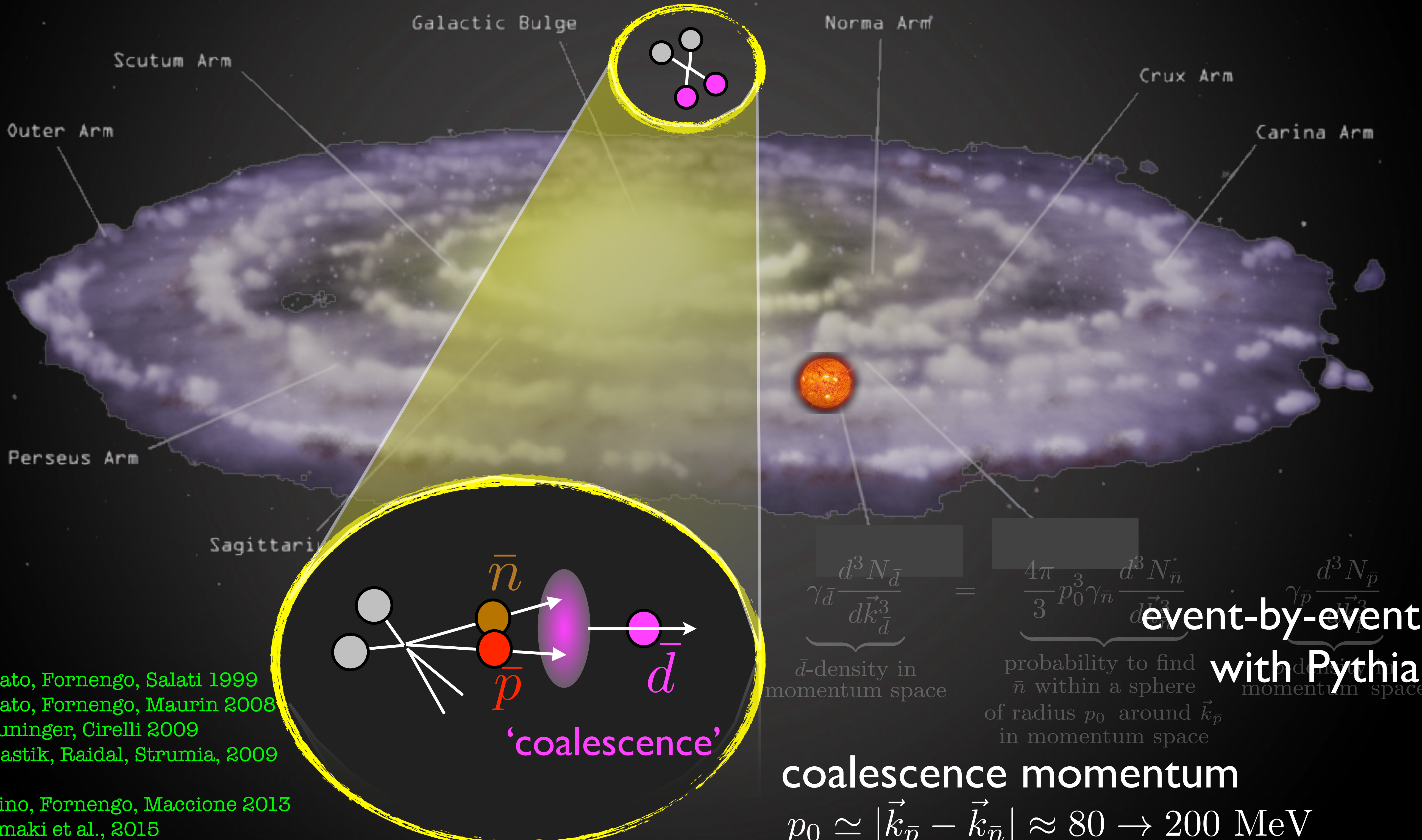

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event-by-event
with Pythia

coalescence momentum

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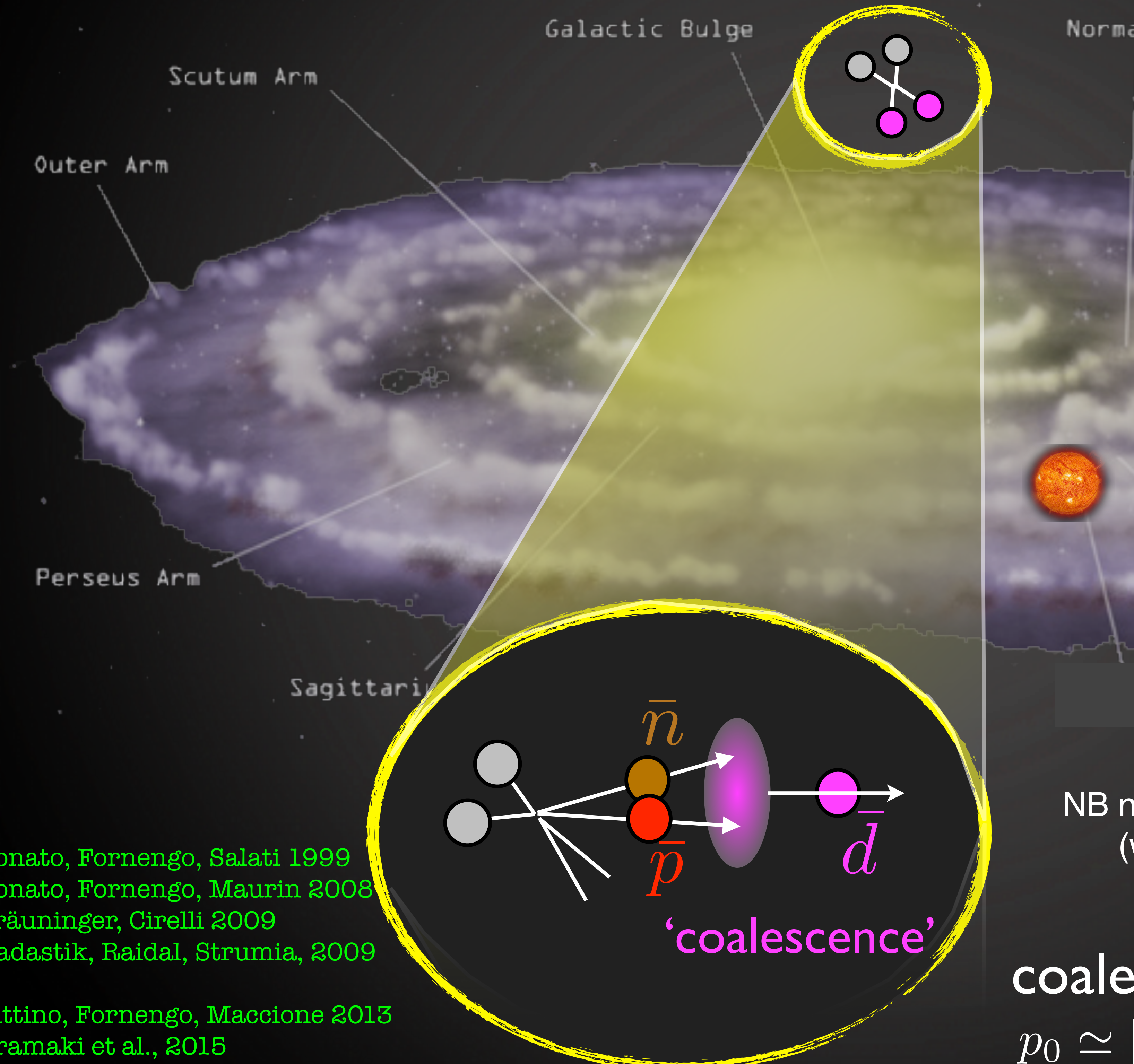
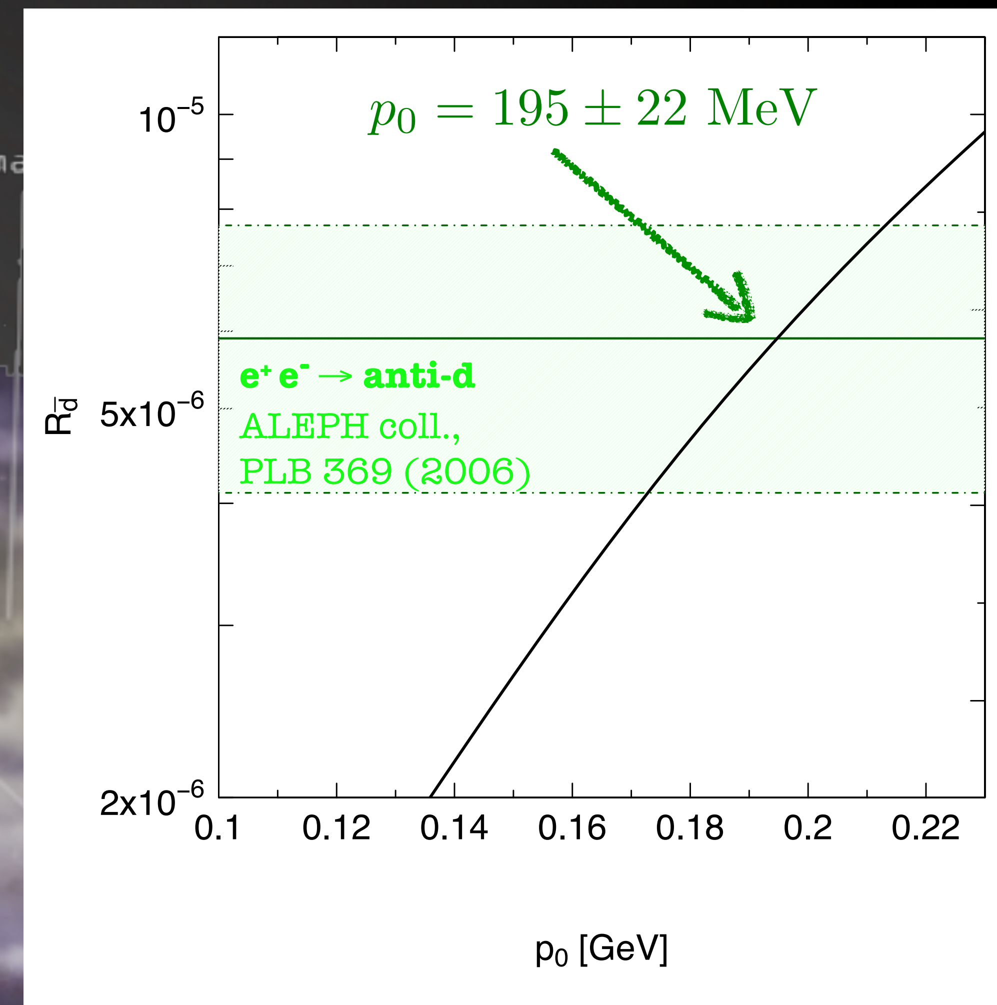
Indirect Detection

$\bar{d}$ from DM annihilations in halo





Vittino, Fornengo, Maccione 2013

NB naïve guess would be $p_0 = \sqrt{E_b\, m_p} = 47$ MeV (with E_b the d binding energy): not too far…

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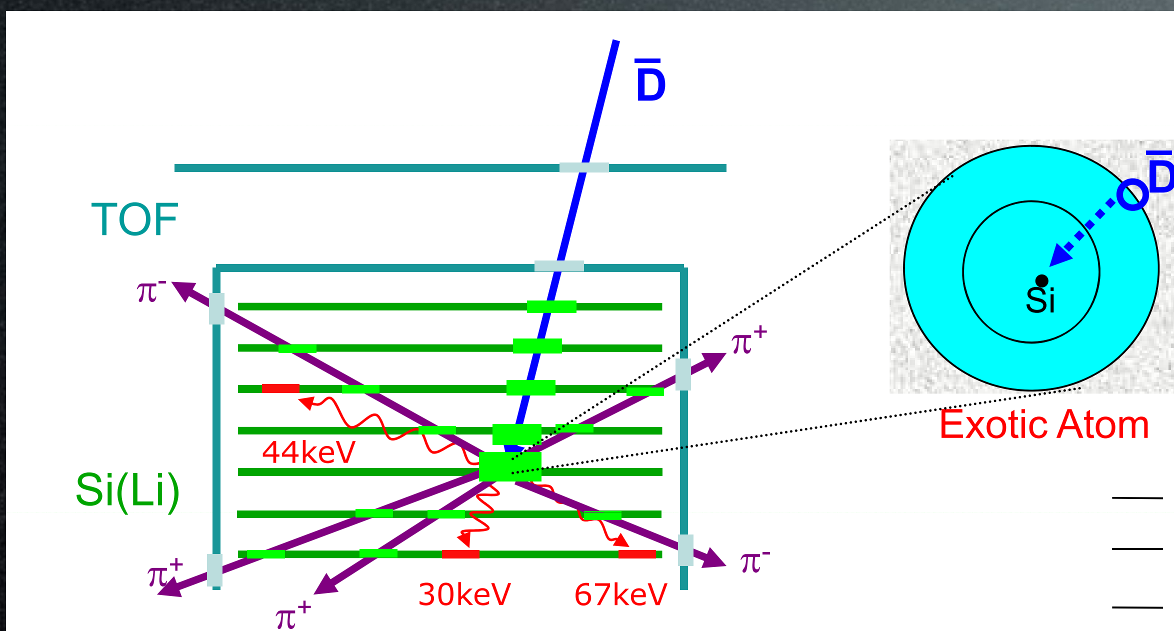
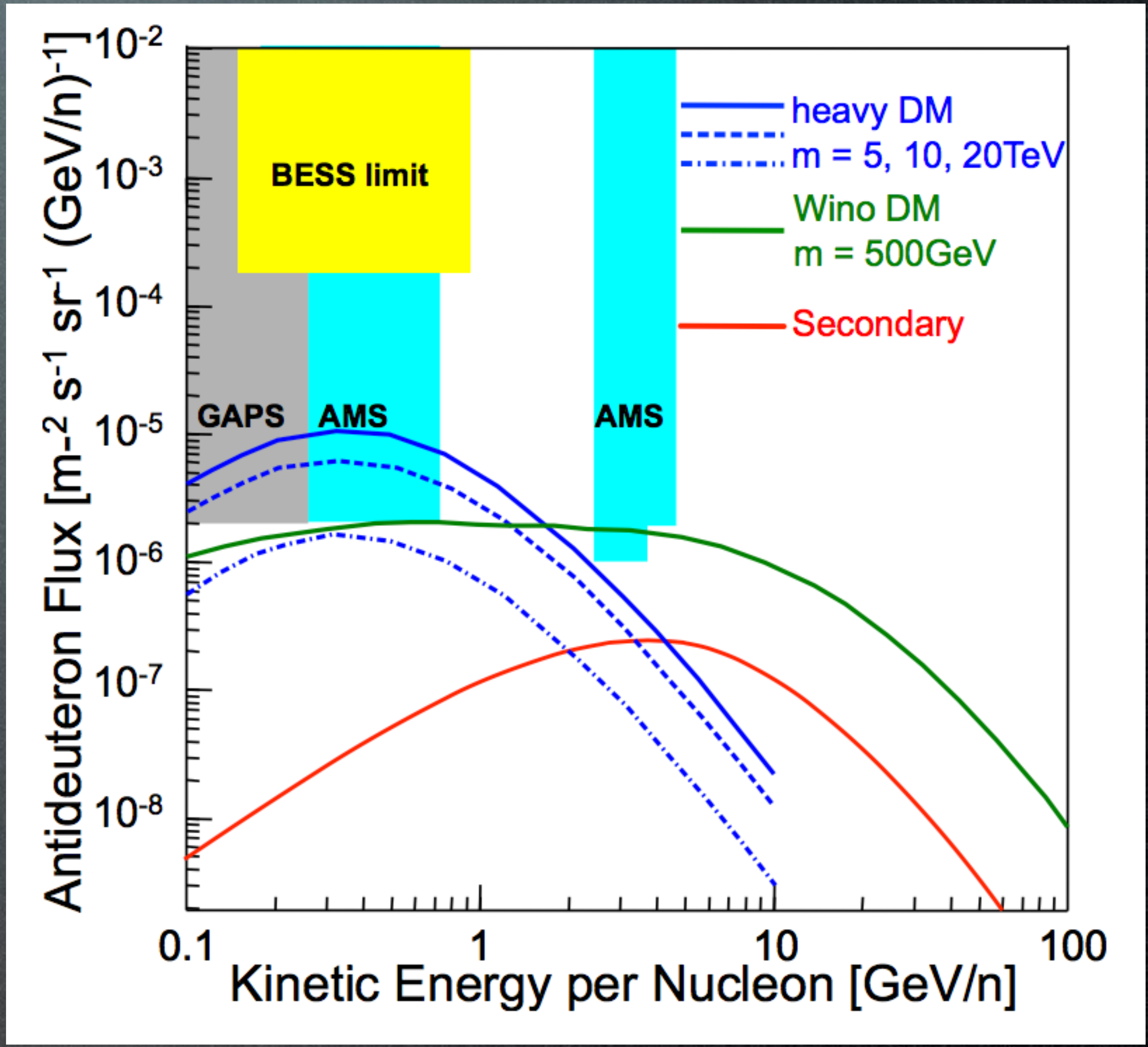
Indirect Detection

$\bar{d}$ from DM annihilations in halo

GAPS detection principle



$\bar{d}$ is slowed down,
captured (exotic atom),
annihilates w distinctive emissions



P. von Doetinchem et al., 2015

DM signal in the reach of GAPS and AMS-02

Neutrinos

direct detection

production at colliders

indirect

γ from annihil in galactic center or halo and from secondary emission

Fermi, ICT, radio telescopes...

e^+ from annihil in galactic halo or center

$\bar{p}$ from annihil in galactic halo or center

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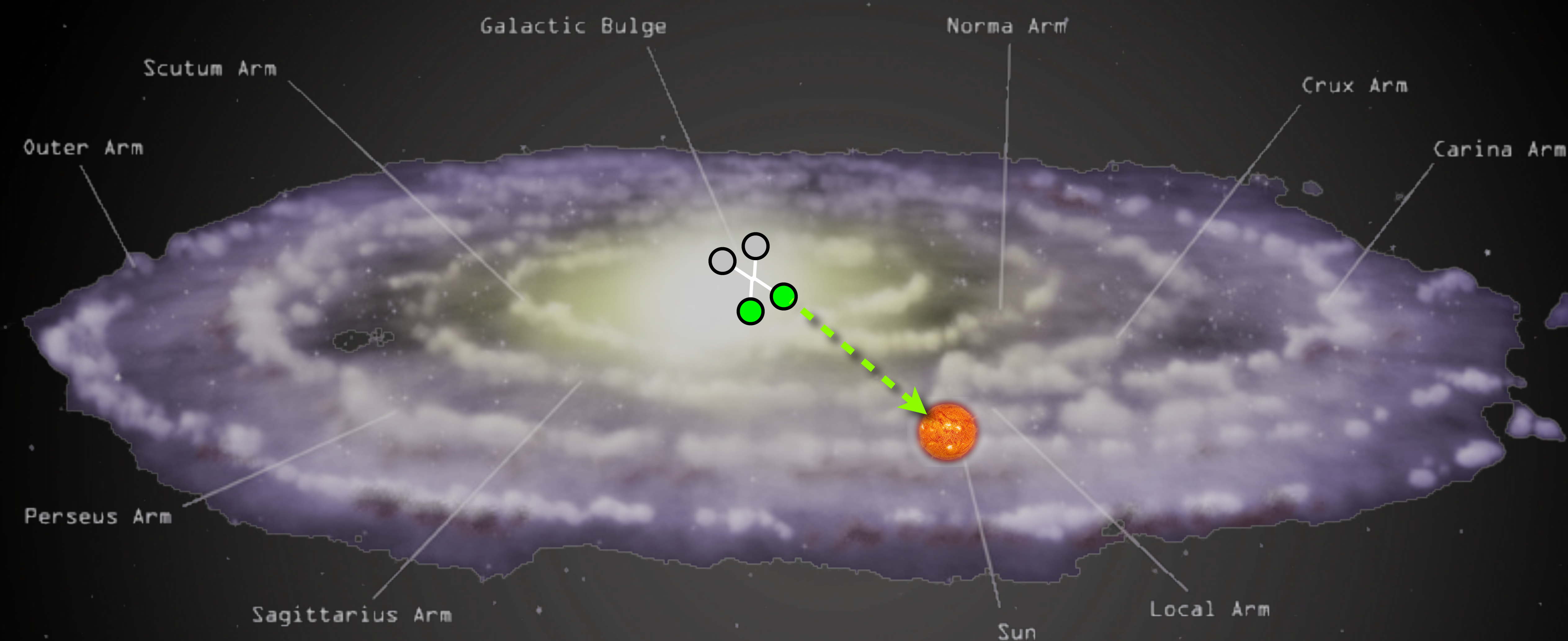
GAPS, AMS

$\nu, \bar{\nu}$ from annihil in massive bodies

SK, Icecube, Antares

ID with neutrinos

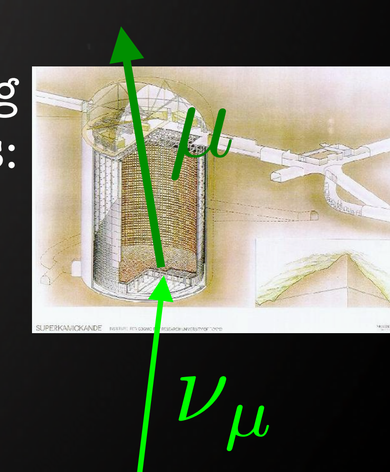
ν from DM annihilations in galactic center



$DM\, DM \to W^-, Z, b, \tau^-, t, h \ldots \rightsquigarrow e^\mp, \overset{(-)}{p}, \overset{(-)}{D} \ldots$ and ν

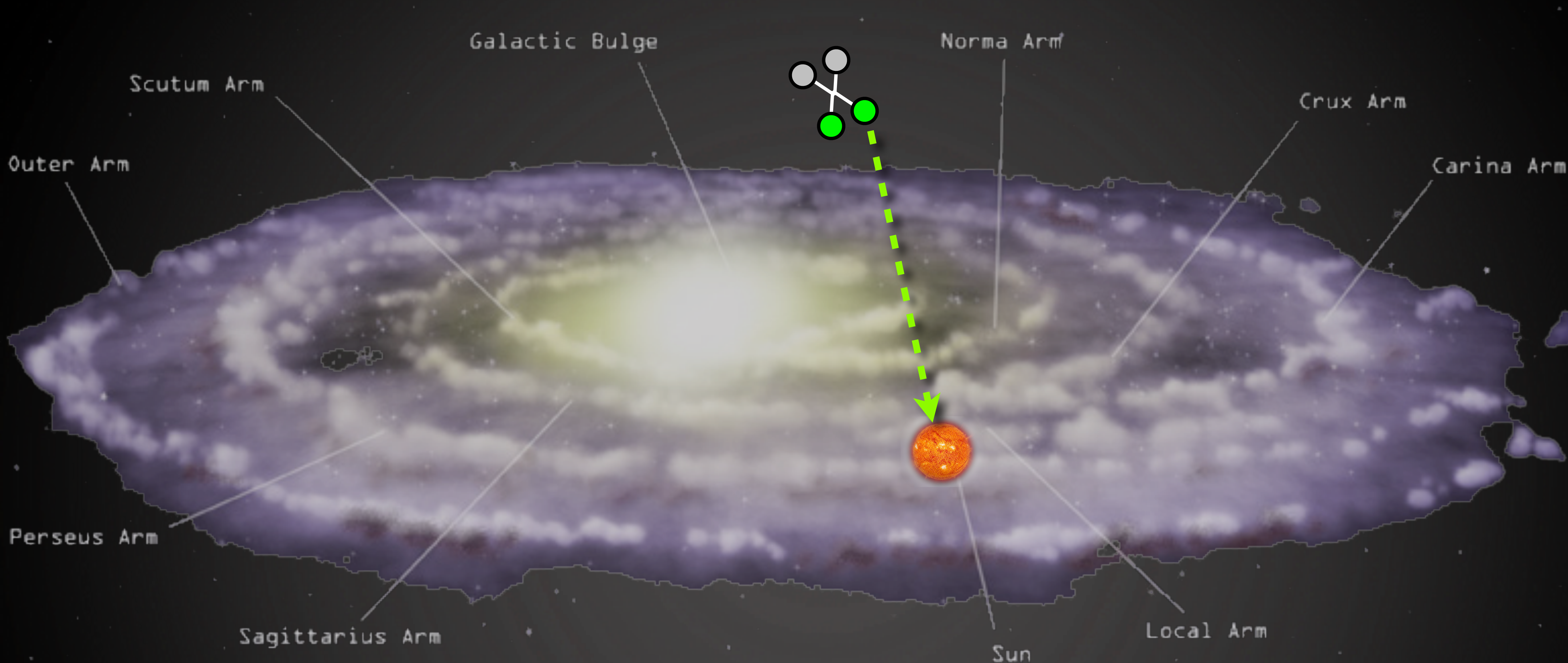
$W^+, Z, \bar{b}, \tau^+, \bar{t}, h \ldots \rightsquigarrow e^\pm, \overset{(-)}{p}, \overset{(-)}{D} \ldots$ and ν

up-going muons:



ID with neutrinos

ν from DM annihilations in galactic halo



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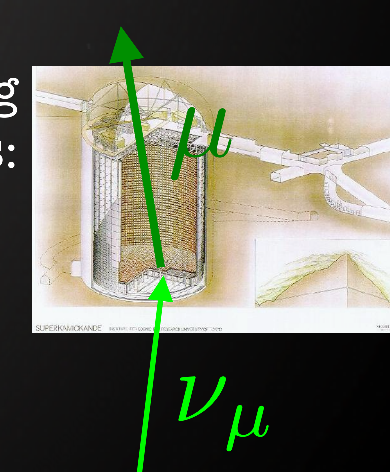
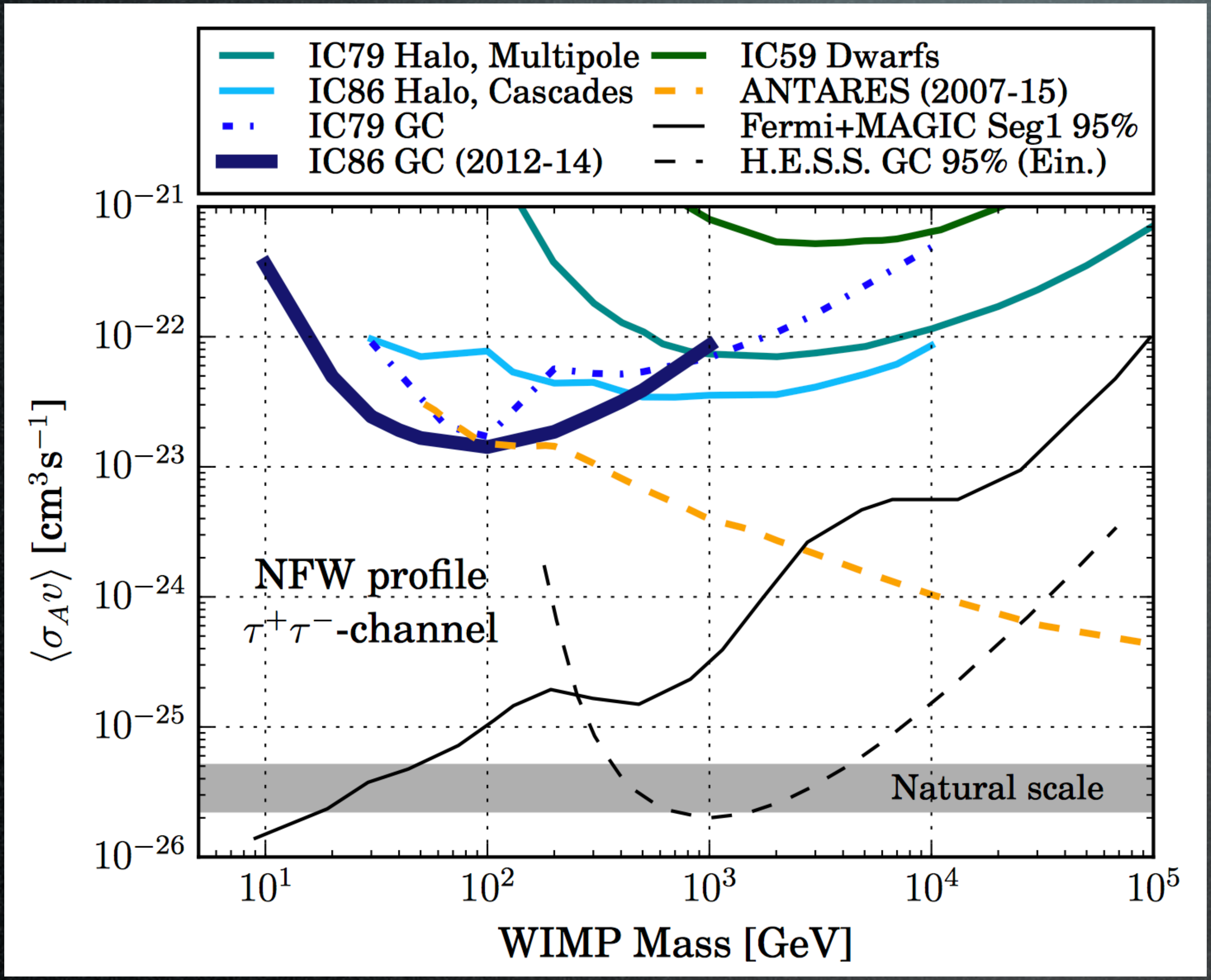


ID with neutrinos

ν from DM annihilations in galactic center/halo

ICECUBE & Antares

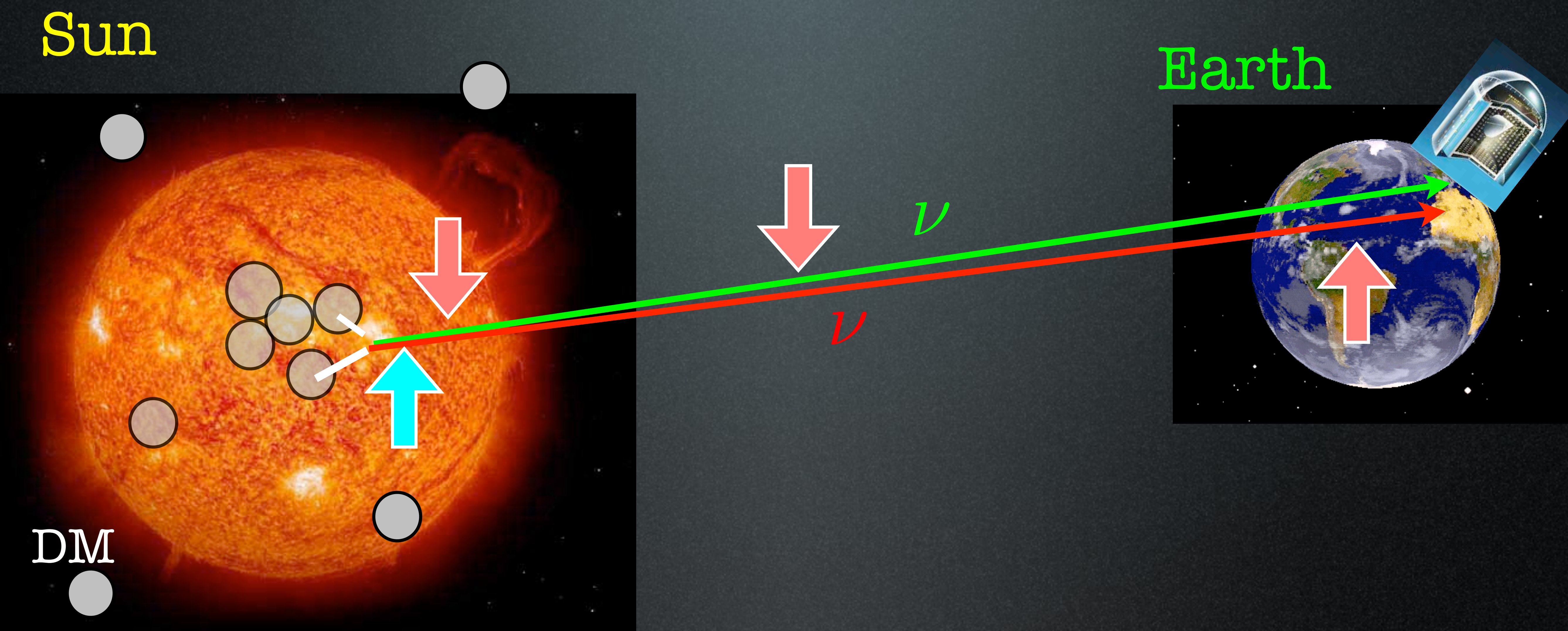


Icecube Coll., 1705.08103

Competitive constraints
(especially for large m_{DM})

ID with neutrinos

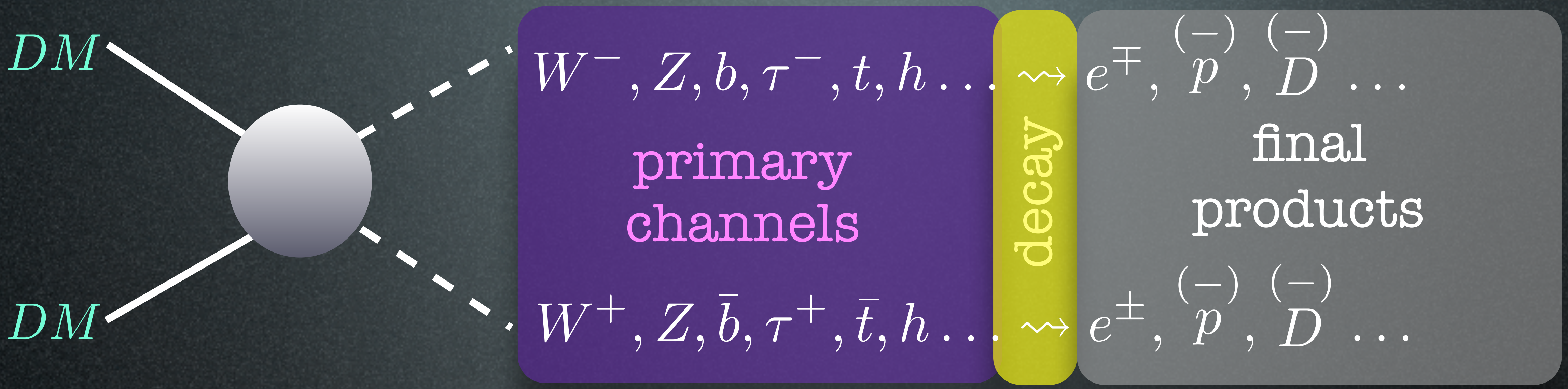
ν from DM annihilations in the Sun



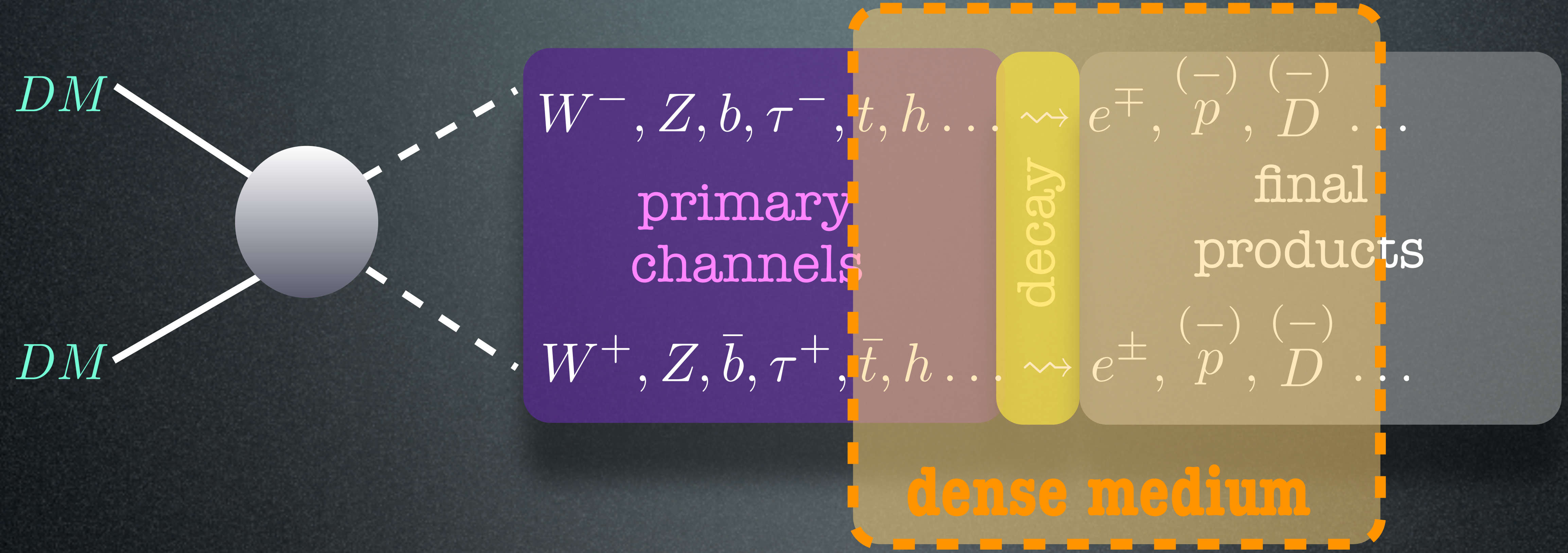
Include oscillations + interactions:

- reshuffling of the 3 flavors
- distortions the spectra
- attenuations of the fluxes

ID with neutrinos



ID with neutrinos

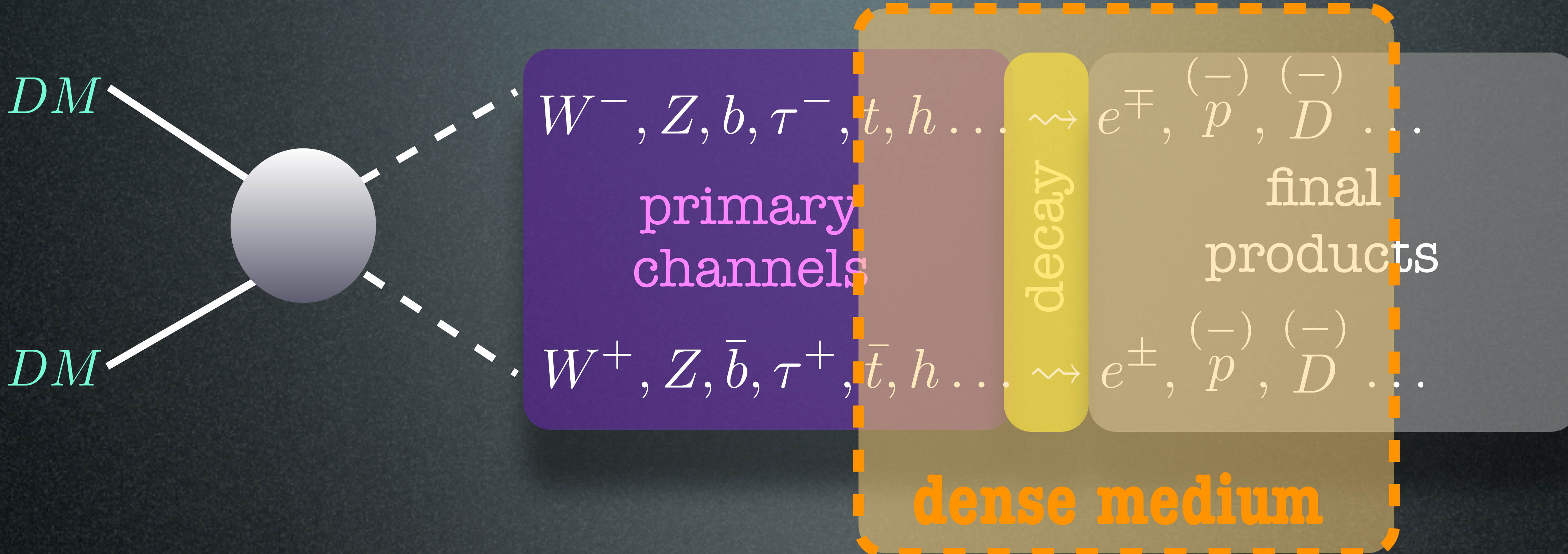


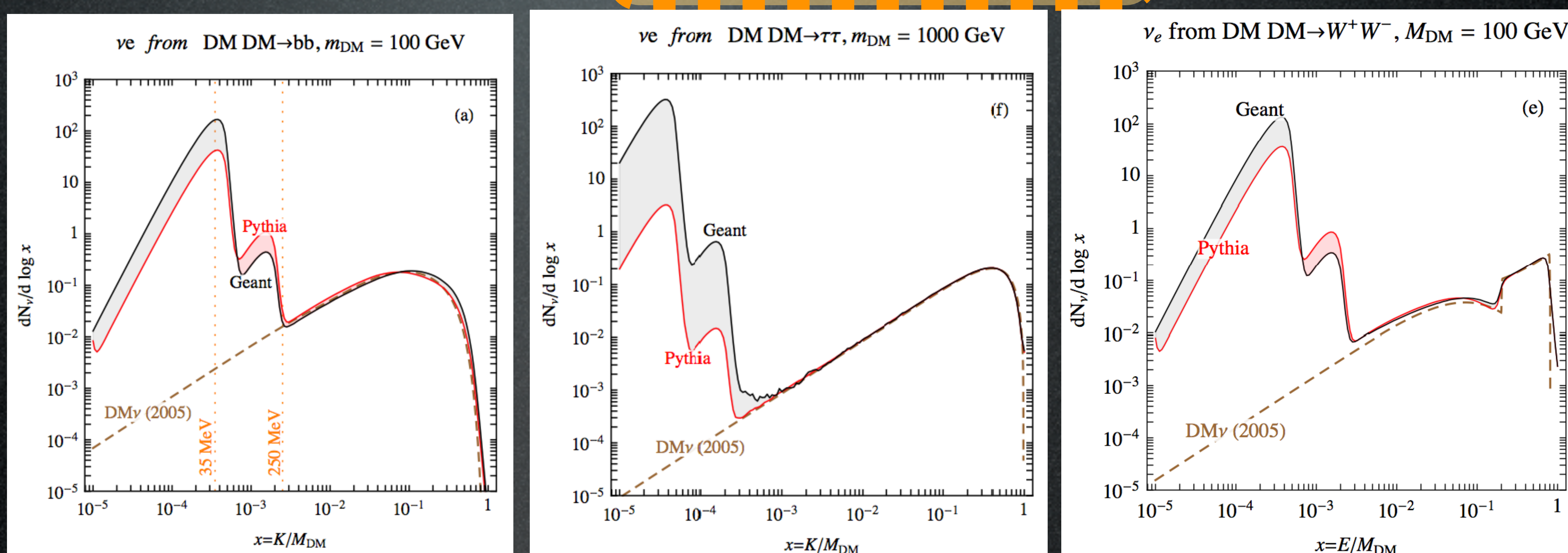
Effects of the medium:

1) light hadrons (π, K...) and leptons (μ) are stopped and decay at rest

2) heavy hadrons/leptons lose some energy before decaying

ID with neutrinos





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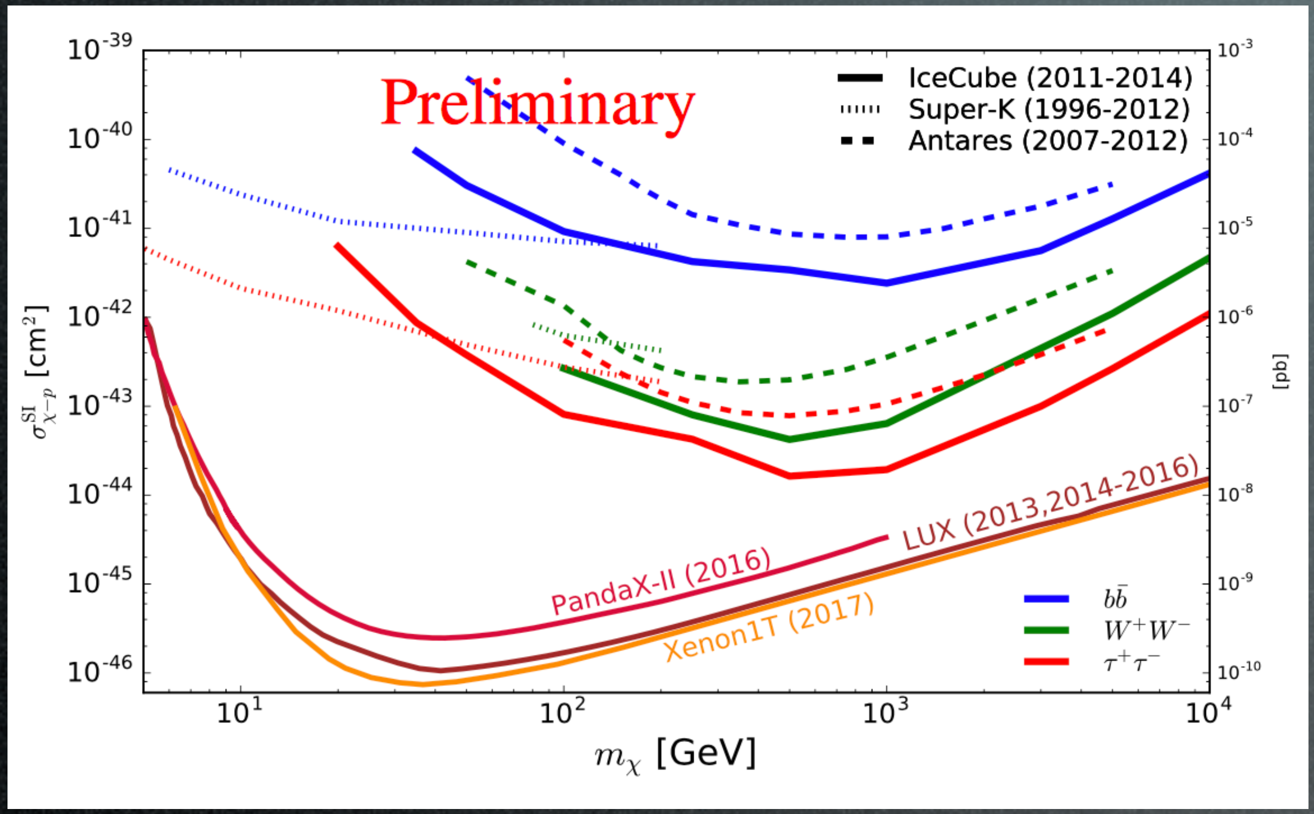
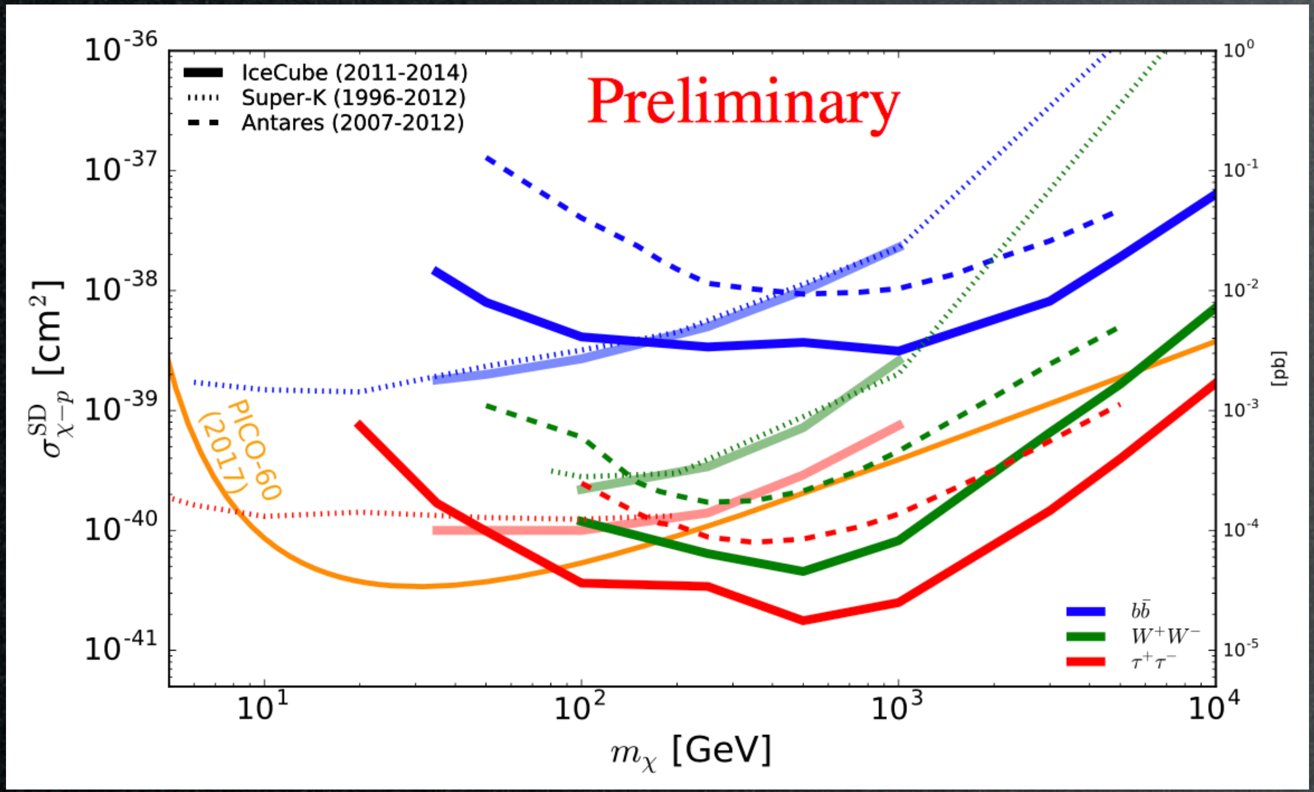
ICECUBE, Antares & SuperKamiokande



Subdominant constraints



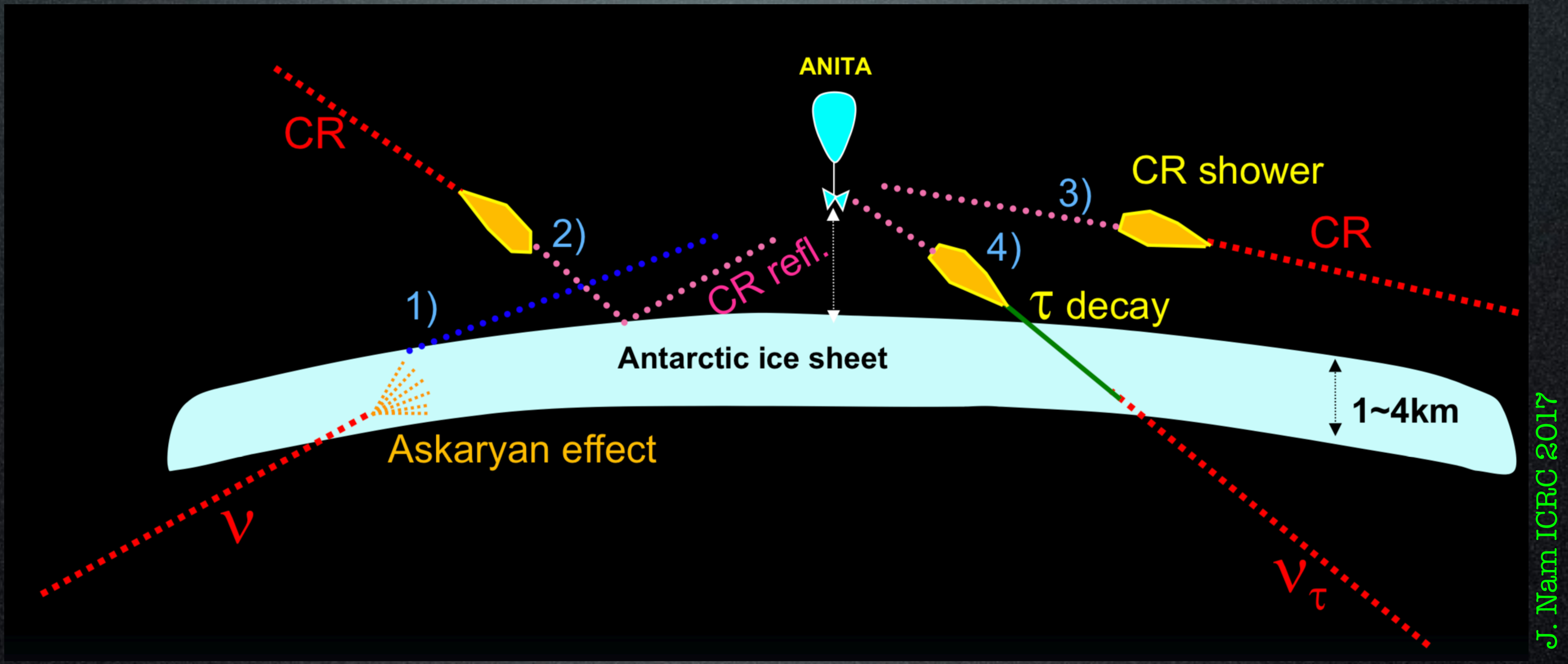
Competitive constraints

status as of ICRC 2017,
C. Rott, rapporteur talk,
1712.00666

ANITA anomaly?

ANITA anomaly?

Who? ANtarctic Impulsive Transient Antenna
balloon-borne radio antenna, 4 1mo flights >2006



NB Askaryan is Cherenkov for **uncharged** particles;
(radio) emission is from a shower of charged particles

ANITA anomaly?

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What? 2 very energetic upgoing air showers
consistent with τ decay (type 4 above)
from a CC of a τ neutrino

NB *focus is on taus because electrons and muons do not even manage to emerge from ice*

event, flight	3985267, ANITA-I	15717147, ANITA-III
date, time	2006-12-28,00:33:20UTC	2014-12-20,08:33:22.5UTC
Lat., Lon.[(1)]	-82.6559, 17.2842	-81.39856, 129.01626
Altitude	2.56 km	2.75 km
Ice depth	3.53 km	3.22 km
El., Az.	$-27.4 \pm 0.3°$, $159.62 \pm 0.7°$	$-35.0 \pm 0.3°$, $61.41 \pm 0.7°$
RA, Dec[(2)]	282.14064, +20.33043	50.78203, +38.65498
$E_{shower}^{(3)}$	0.6 ± 0.4 EeV	$0.56^{+0.3}_{-0.2}$ EeV

[1] Latitude, Longitude of the estimated ground position of the event.
[2] Sky coordinates projected from event arrival angles at ANITA.
[3] For upward shower initiation at or near ice surface.

ANITA coll. 1803.05088

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from a CC of a τ neutrino

How? EeV τ-neutrinos cannot cross the Earth

transmission probability $\simeq 10^{-6}$, even including τ-regeneration

ANITA anomaly?

Who? ANtarctic Impulsive Transient Antenna
balloon-borne radio antenna, 4 1mo flights >2006

What? 2 very energetic upgoing air showers
consistent with τ decay

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Very heavy DM
($m \gtrsim 10^9$ GeV)
decaying to RH ν
which convert to τ
in Earth
Heurtier+ 1902.04584

Very heavy DM
($m \gtrsim 10^{10}$ GeV)
decaying to some χ
(weaker than ν)
which do Askaryan
in ice (type 1 above)
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complexifications
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Or? some other BSM
(long-lived particles, staus,
leptoquarks, sphalerons,
sterile nus...) several 2018+

exp misunderstanding
could they be type 2 events after all?
Shoemaker+ 1905.02846

DM detection

direct detection

production at colliders

indirect

γ from annihil in galactic center or halo and from secondary emission

Fermi, ICT, radio telescopes...

e^+ from annihil in galactic halo or center

PAMELA, Fermi, HESS, AMS,balloons...

$\bar{p}$ from annihil in galactic halo or center

$\bar{d}$ from annihil in galactic halo or center

GAPS, AMS

$\nu, \bar{\nu}$ from annihil in massive bodies

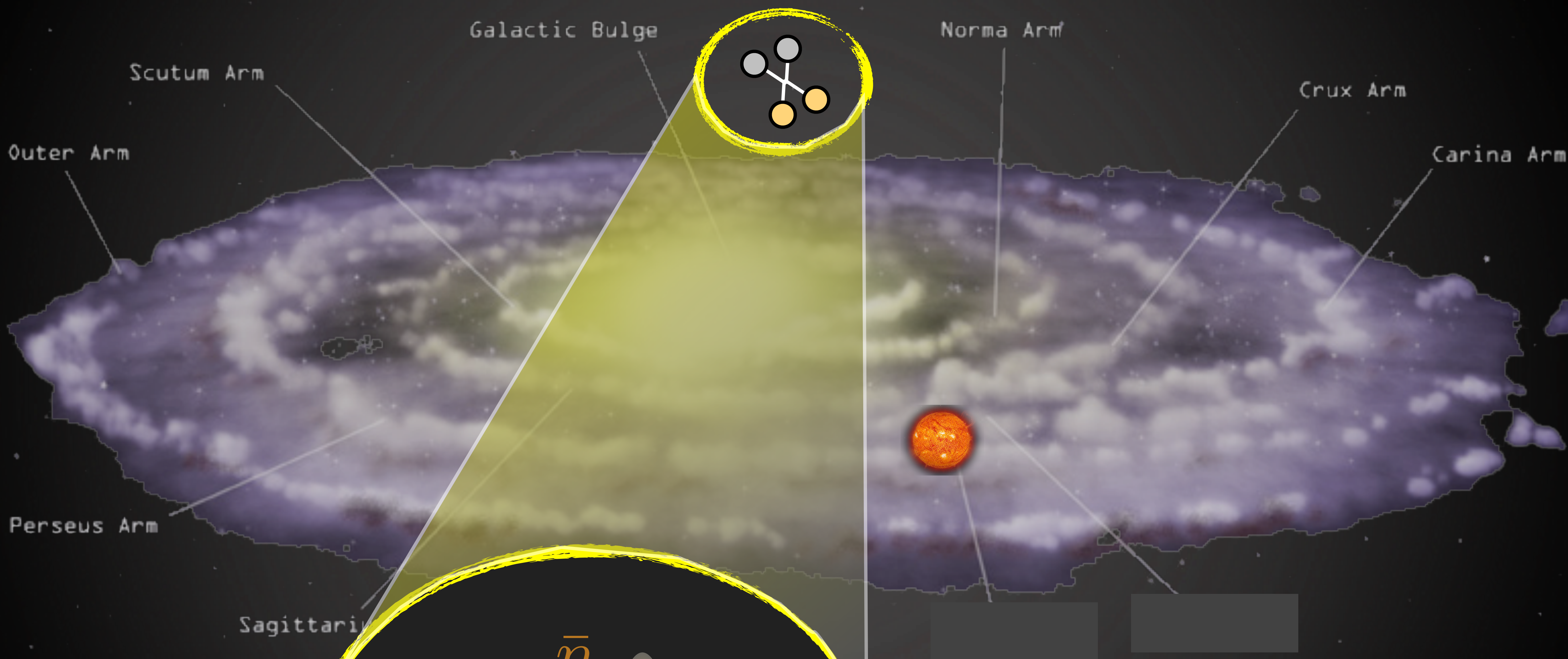
SK, Icecube, Km3Net

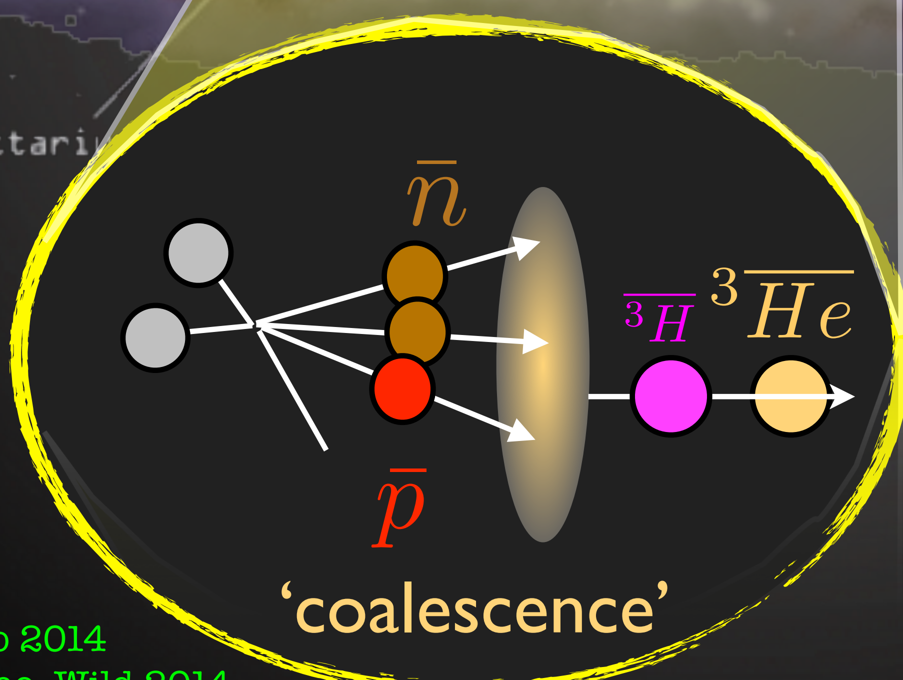
$\overline{He}$ from annihil in galactic halo or center

AMS?

Indirect Detection

$\overline{He}$ from DM annihilations in halo





event-by-event
with Pythia

$$|\vec{k}_1 - \vec{k}_2| \leq p_0$$
$$|\vec{k}_1 - \vec{k}_3| \leq p_0$$
$$|\vec{k}_2 - \vec{k}_3| \leq p_0$$

a bit arbitrary…

coalescence momentum $p_0 = 195$ MeV

Cirelli, Fornengo, Vittino, Taoso 2014
Carlson, Linden, Ibarra, Profumo, Wild 2014

Indirect Detection

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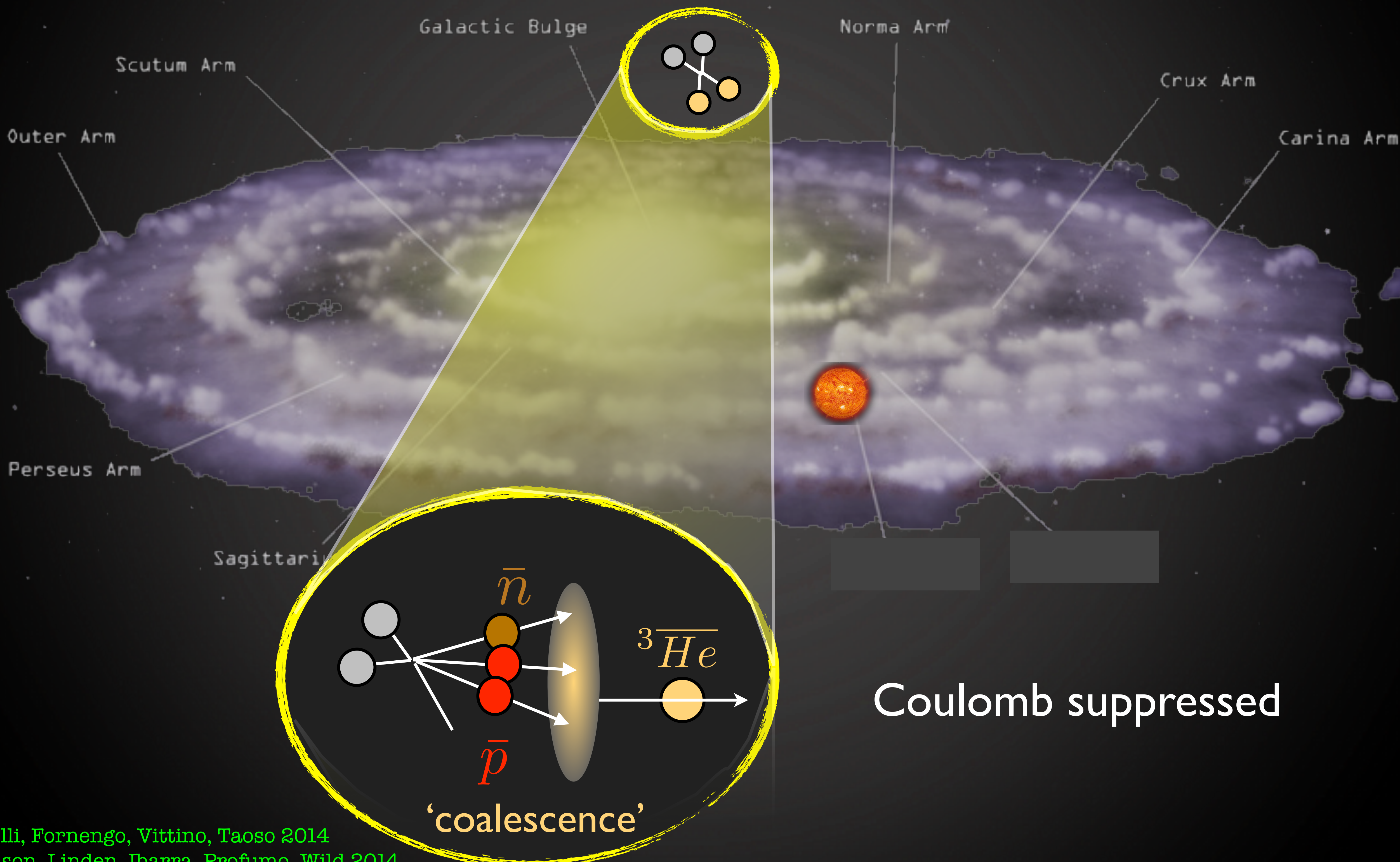

Coulomb suppressed

Cirelli, Fornengo, Vittino, Taoso 2014
Carlson, Linden, Ibarra, Profumo, Wild 2014

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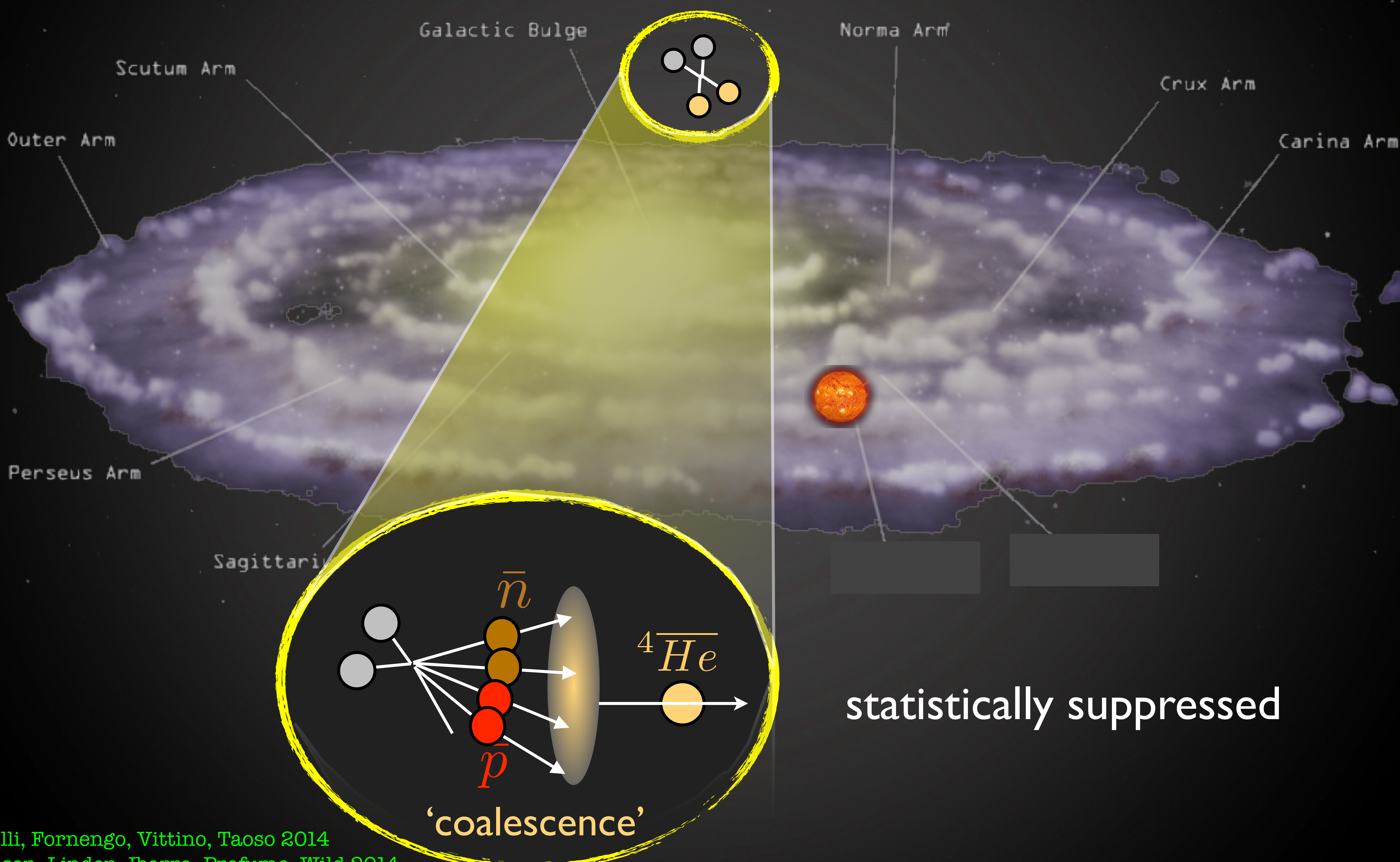

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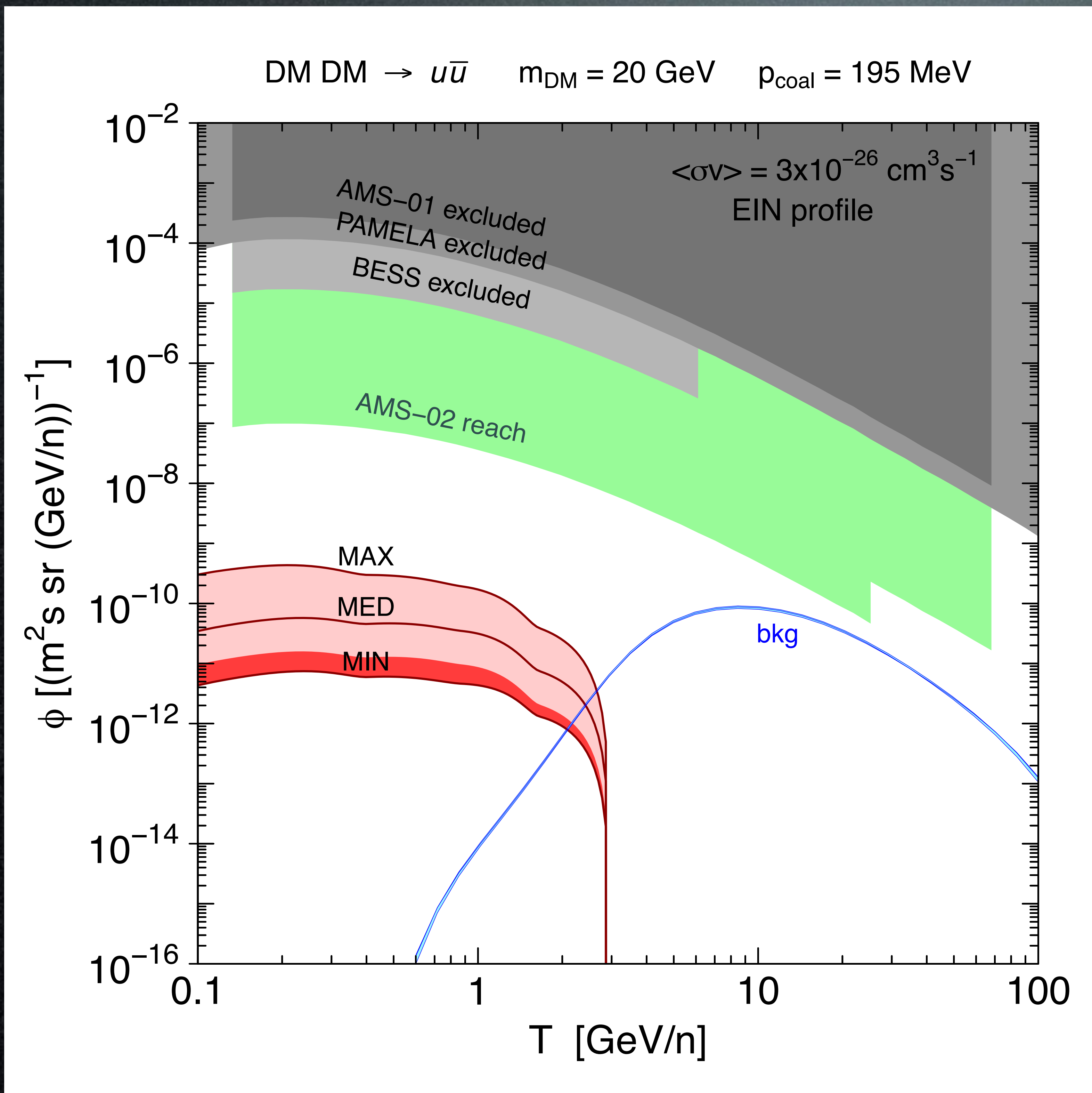
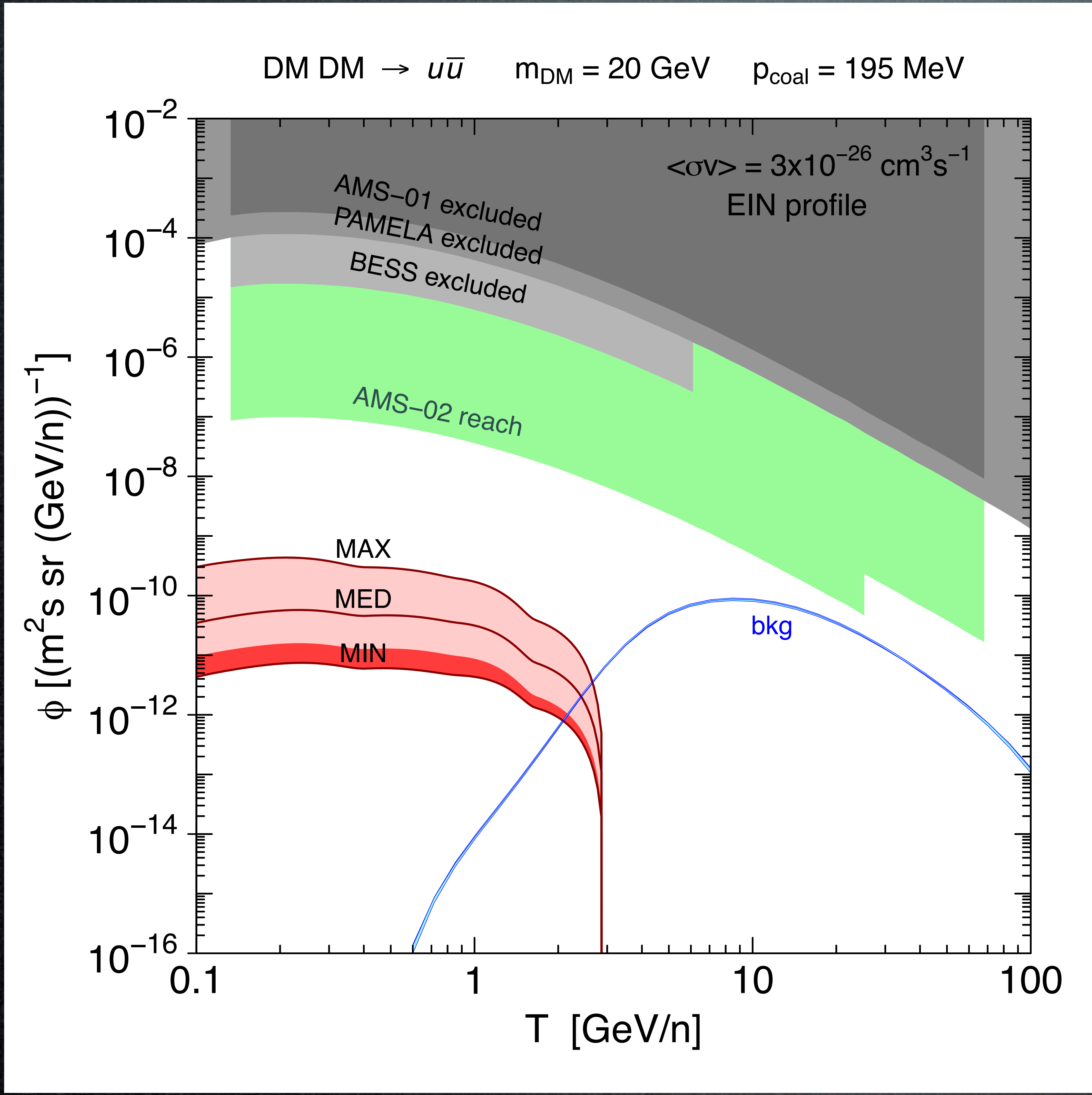


all

consistent with antiproton bounds

Indirect Detection

$\overline{He}$ from DM annihilations in halo



Indirect Detection

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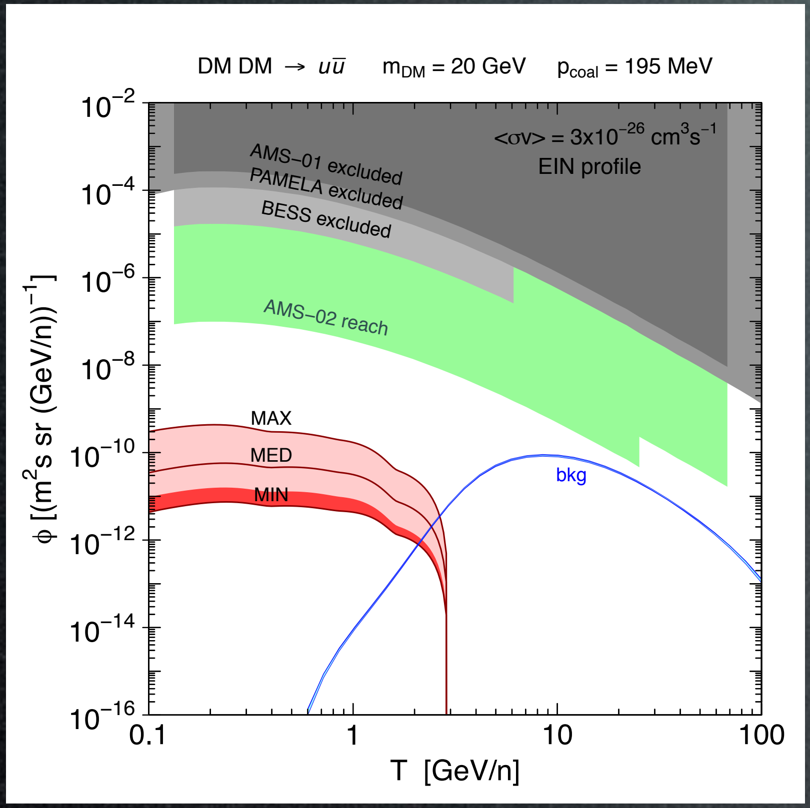

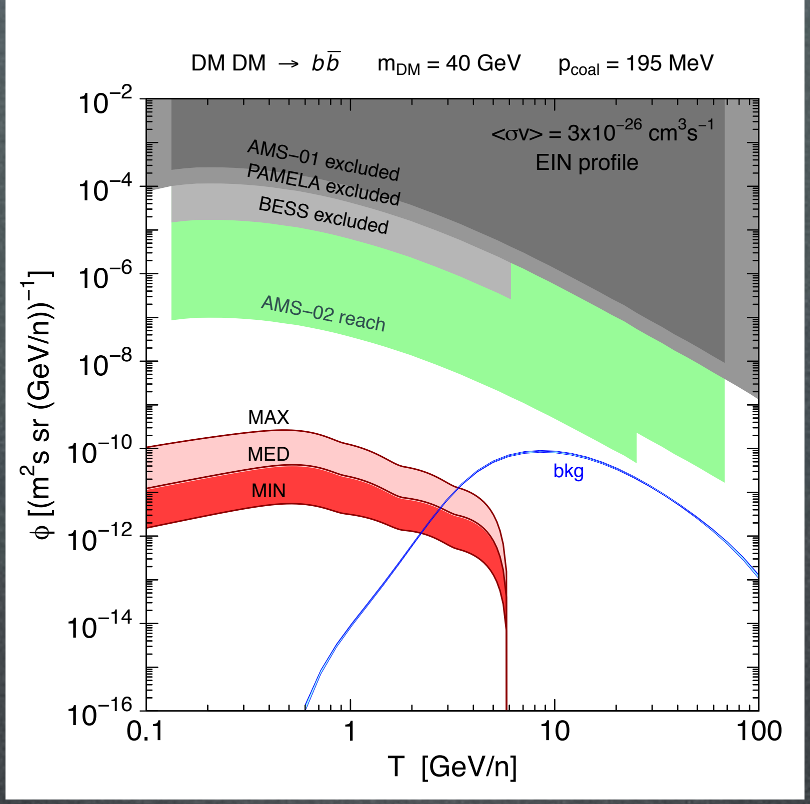

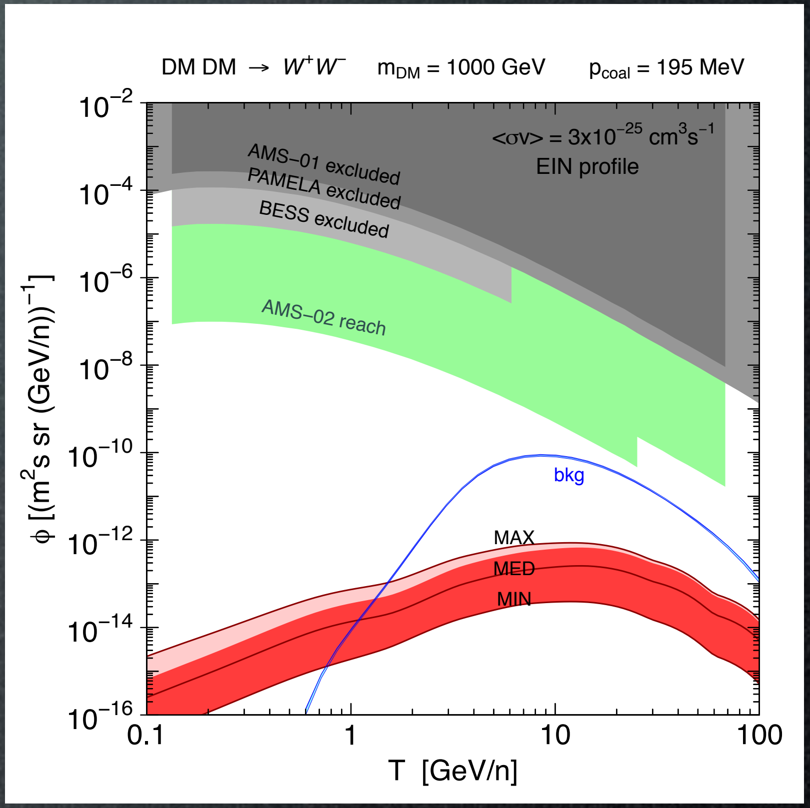

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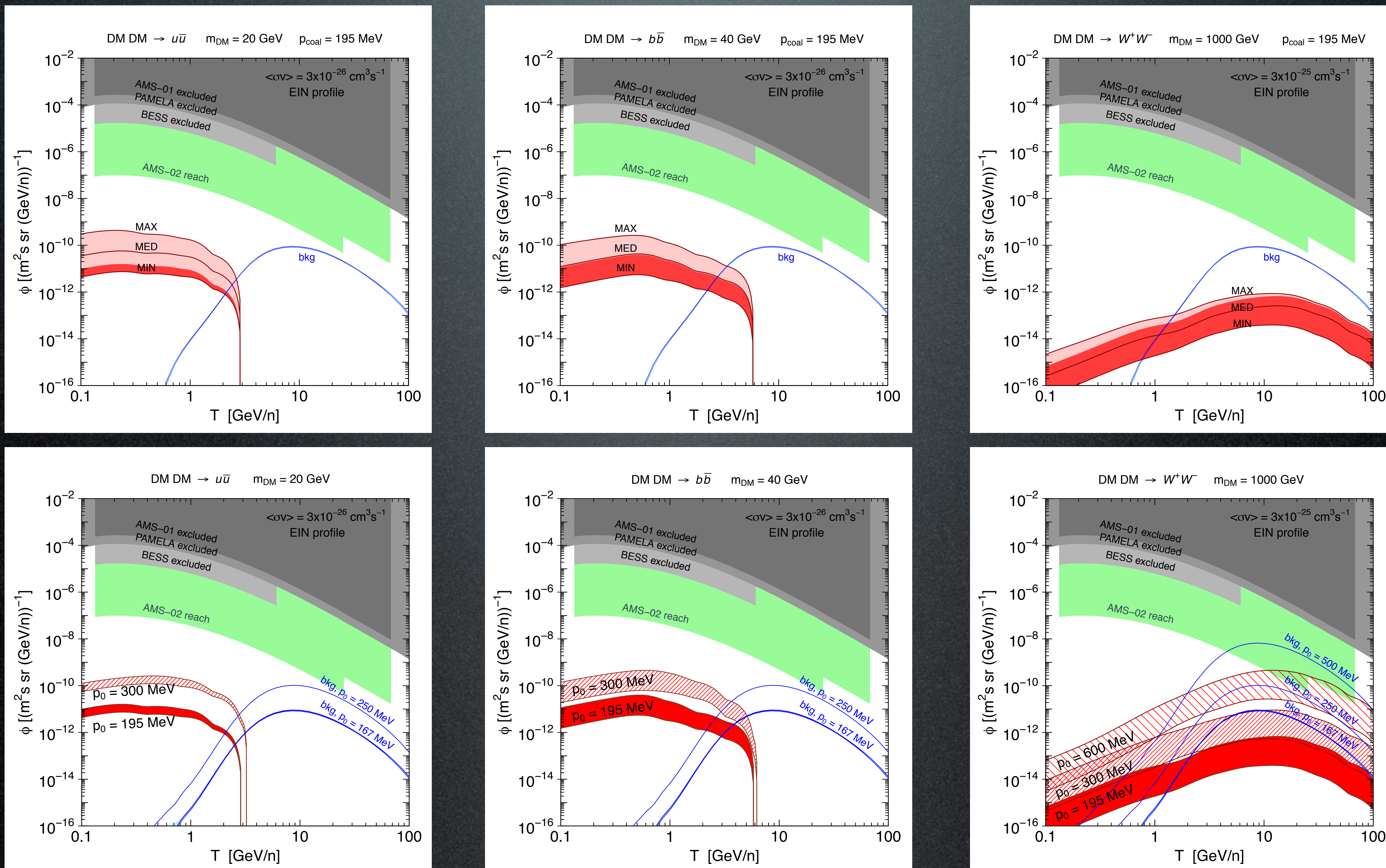

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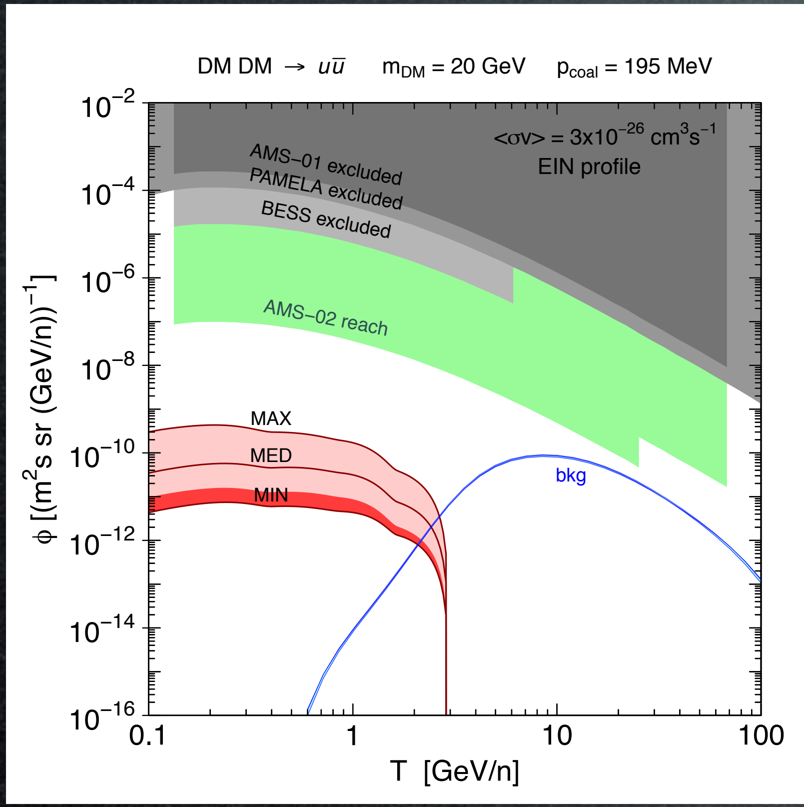
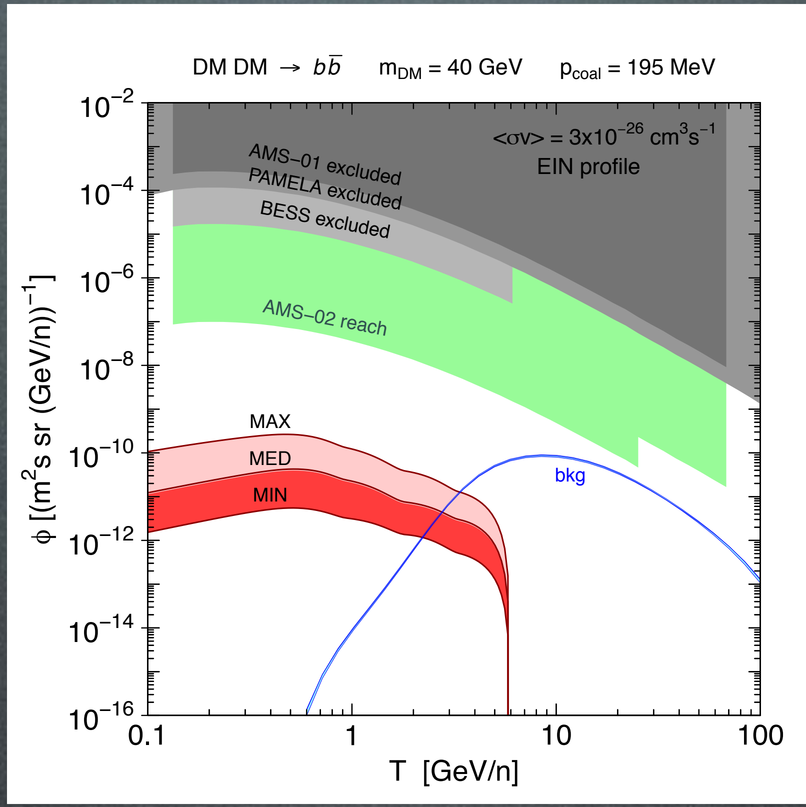
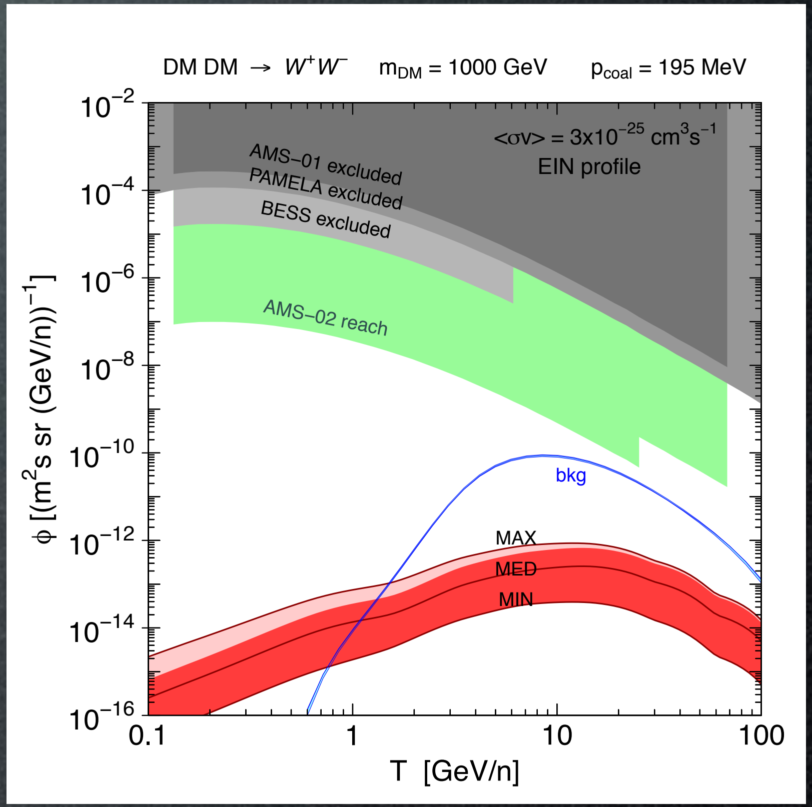
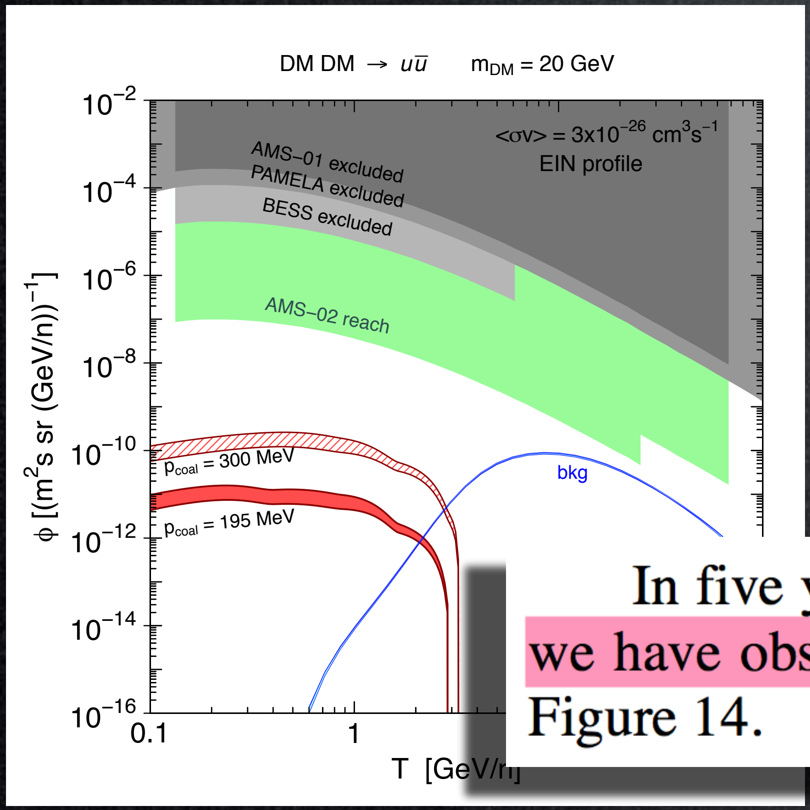
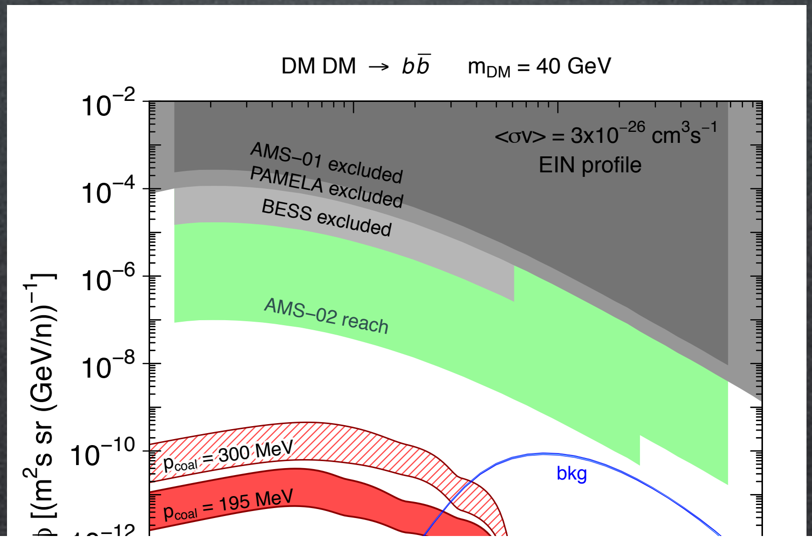
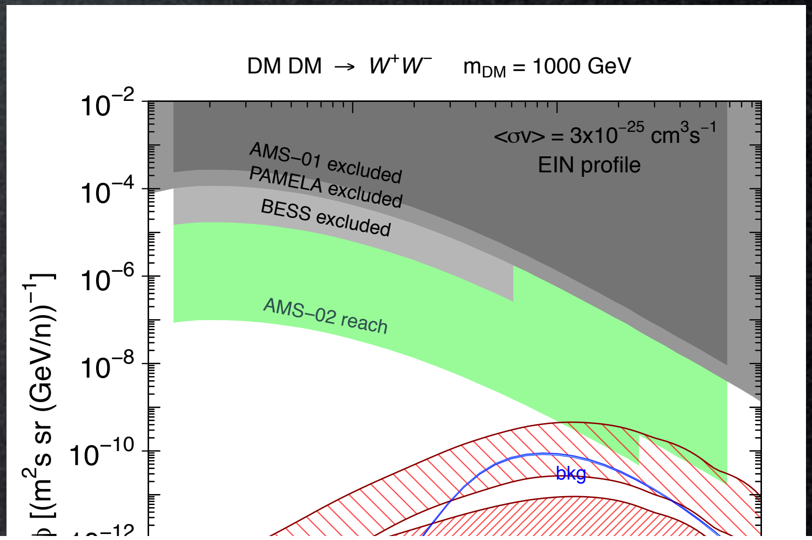










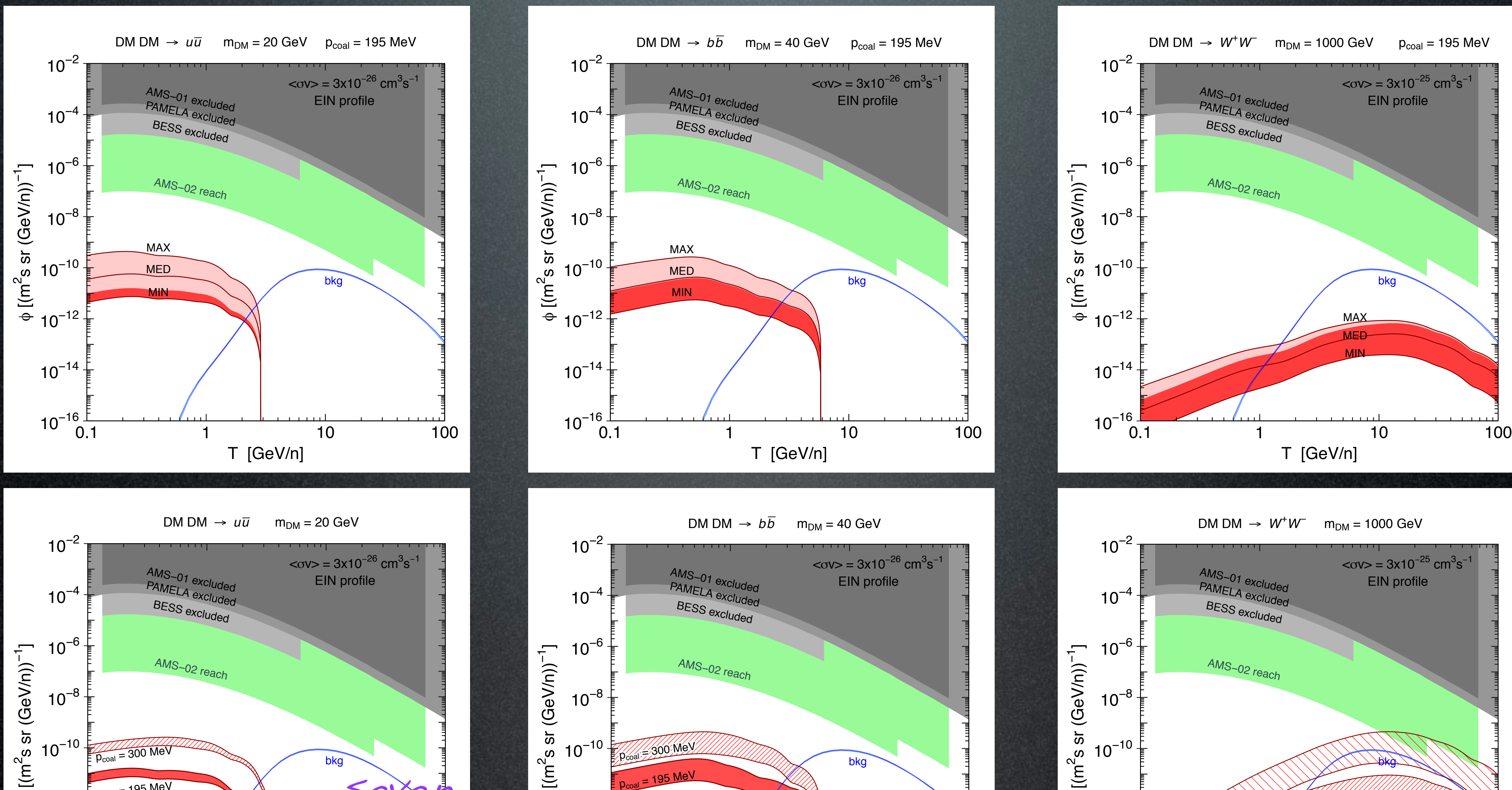


In five years, AMS has collected 3.7 billion helium events (charge Z = +2). To date we have observed a few Z = -2 events with mass around ^{3}He. An event is displayed in Figure 14.

S.Ting - AMS-02 press release - December 2016

Indirect Detection

$\overline{He}$ from DM annihilations in halo



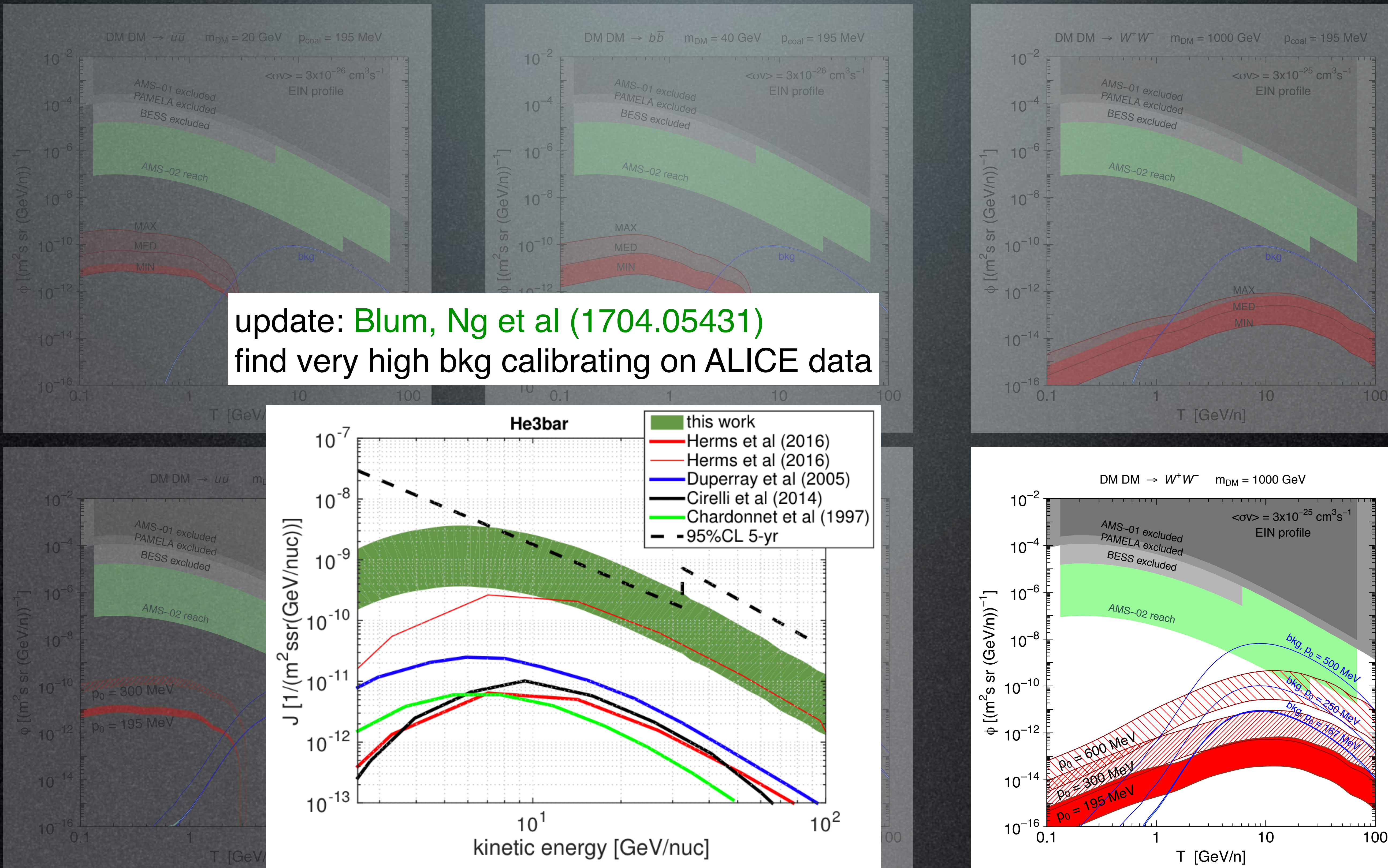
In ~~five~~ seven years, AMS has collected 3.7 billion helium events (charge Z = +2). To date we have observed a ~~few~~ eight Z = -2 events with mass around ^{3}He. An event is displayed in Figure 14. (two of which ^{4}He)

S.Ting - AMS-02 colloquium @ CERN - May 2018

Indirect Detection

$\overline{He}$ from DM annihilations in halo

update: Blum, Ng et al (1704.05431)
find very high bkg calibrating on ALICE data



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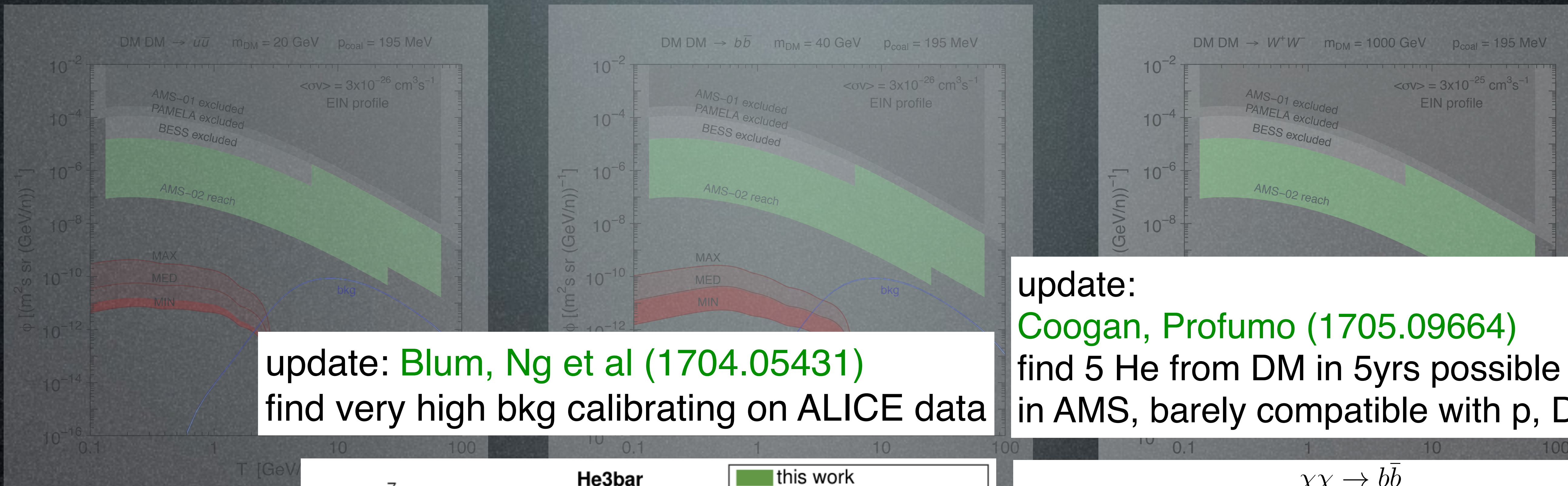

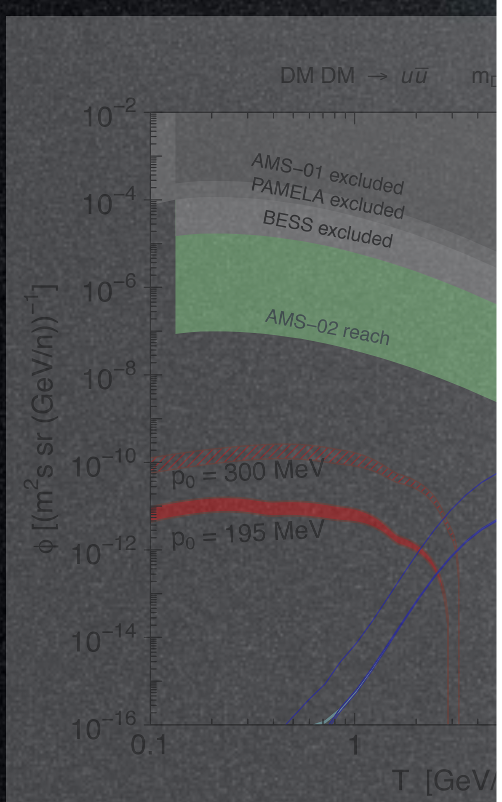

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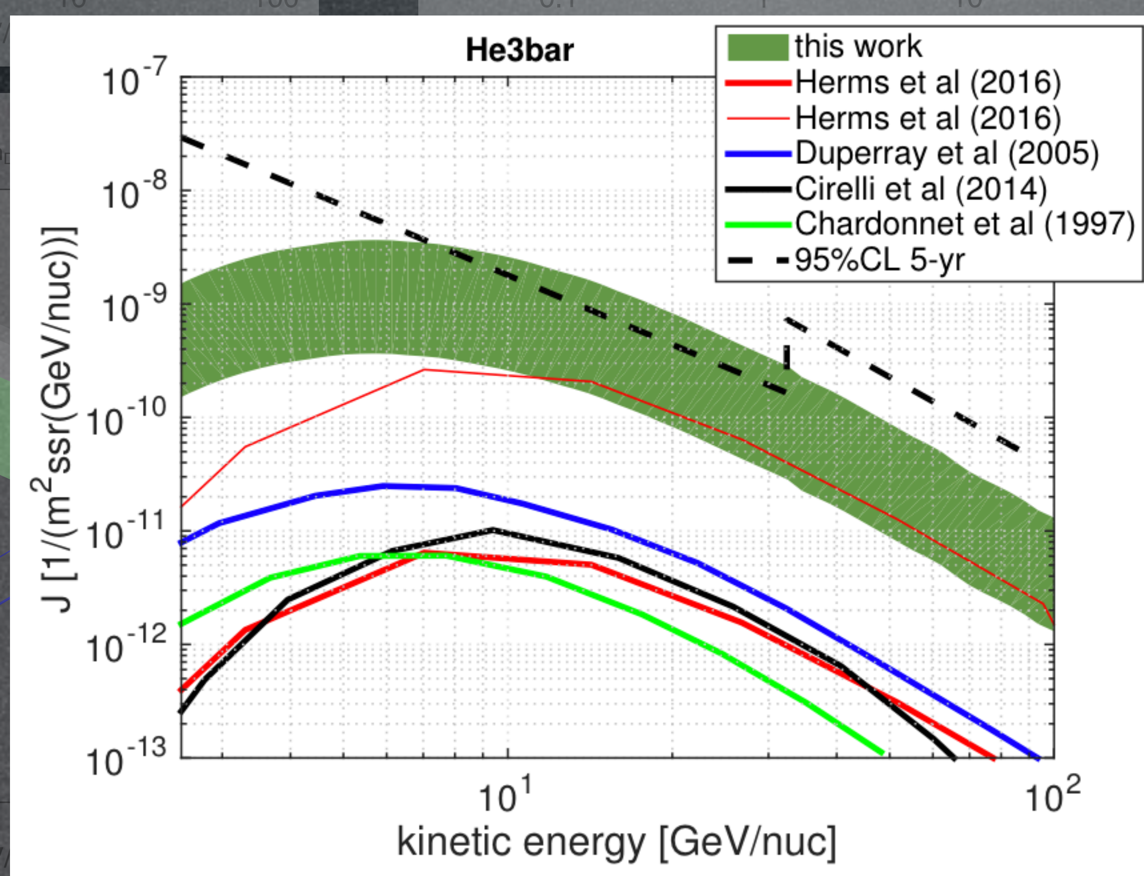

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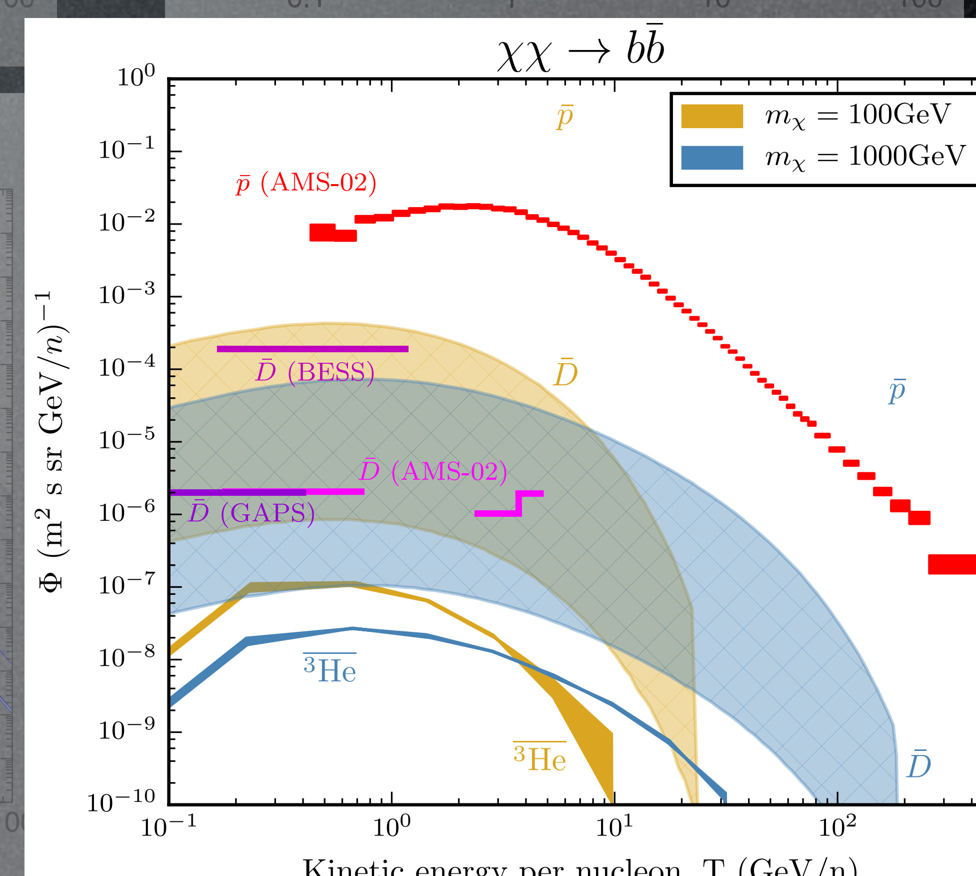
Coogan, Profumo (1705.09664)
find 5 He from DM in 5yrs possible
in AMS, barely compatible with p, D



Indirect Detection

$\overline{He}$ from DM annihilations in halo

alternative: Poulin, Salati, Cholis, Kamionkowski, Silk (1808.08961)

anti-He from anti-clouds or anti-stars!

however: strong constraints from gamma-rays, CMB etc
need exotic (anti-)BBN to have right isotopic ratios…

also: Heck, Rajaraman (1906.01667):
$\overline{\text{He}}$ from decay of exotic Φ carrying negative baryon number (but very fine tuned or killed by antiprotons)

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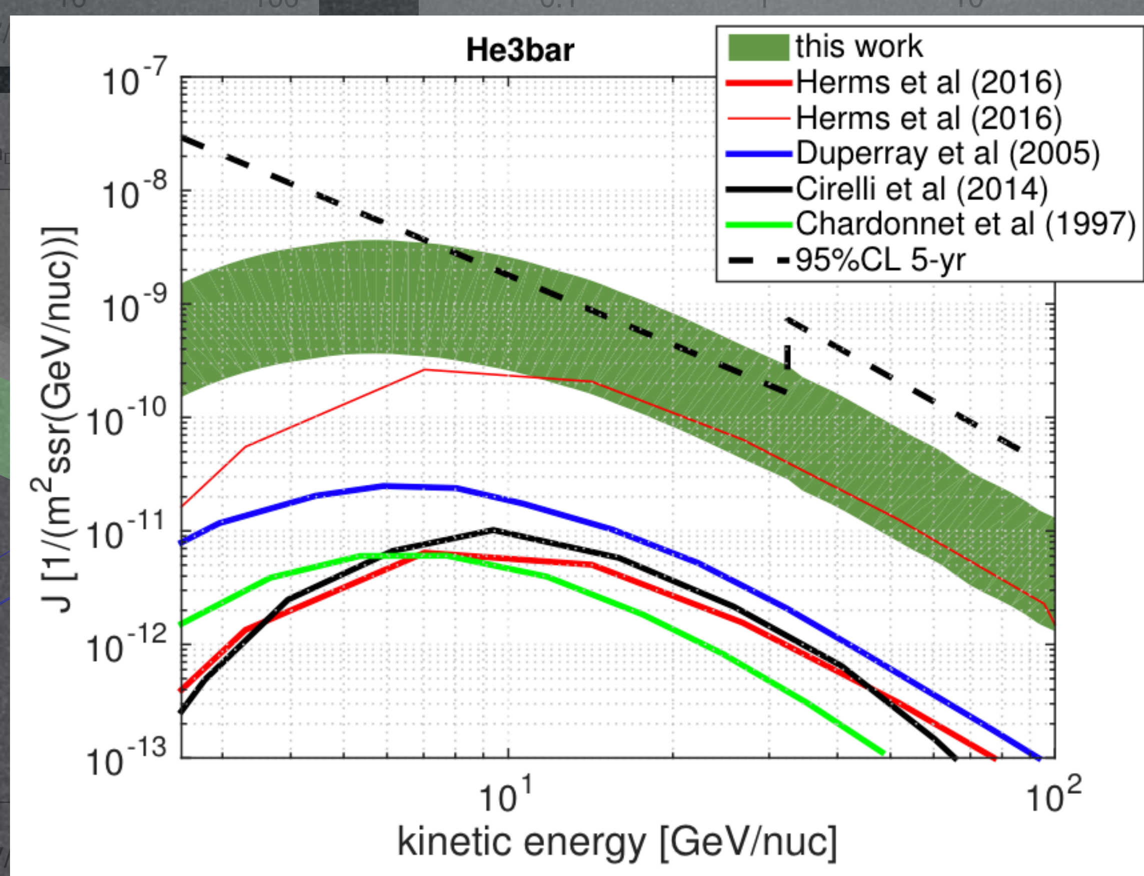

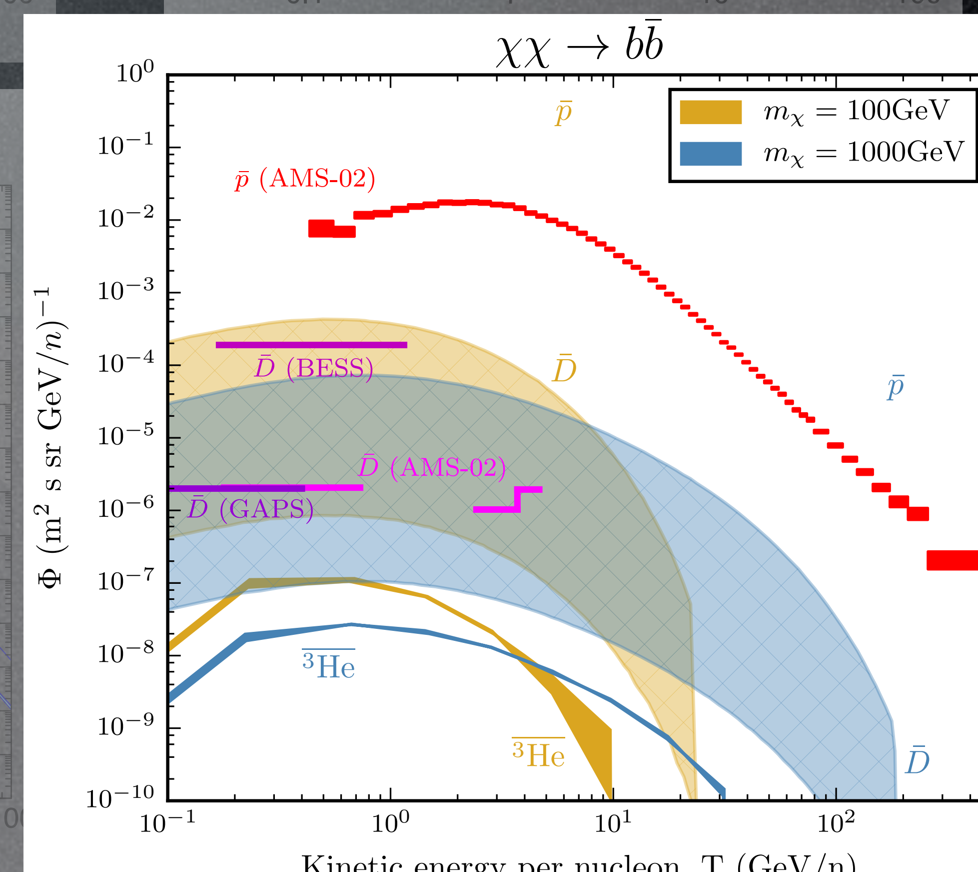

Comparing all bounds

updated from M. Cirelli 1511.02031; M. Cirelli, A. Strumia, J. Zupan to appear

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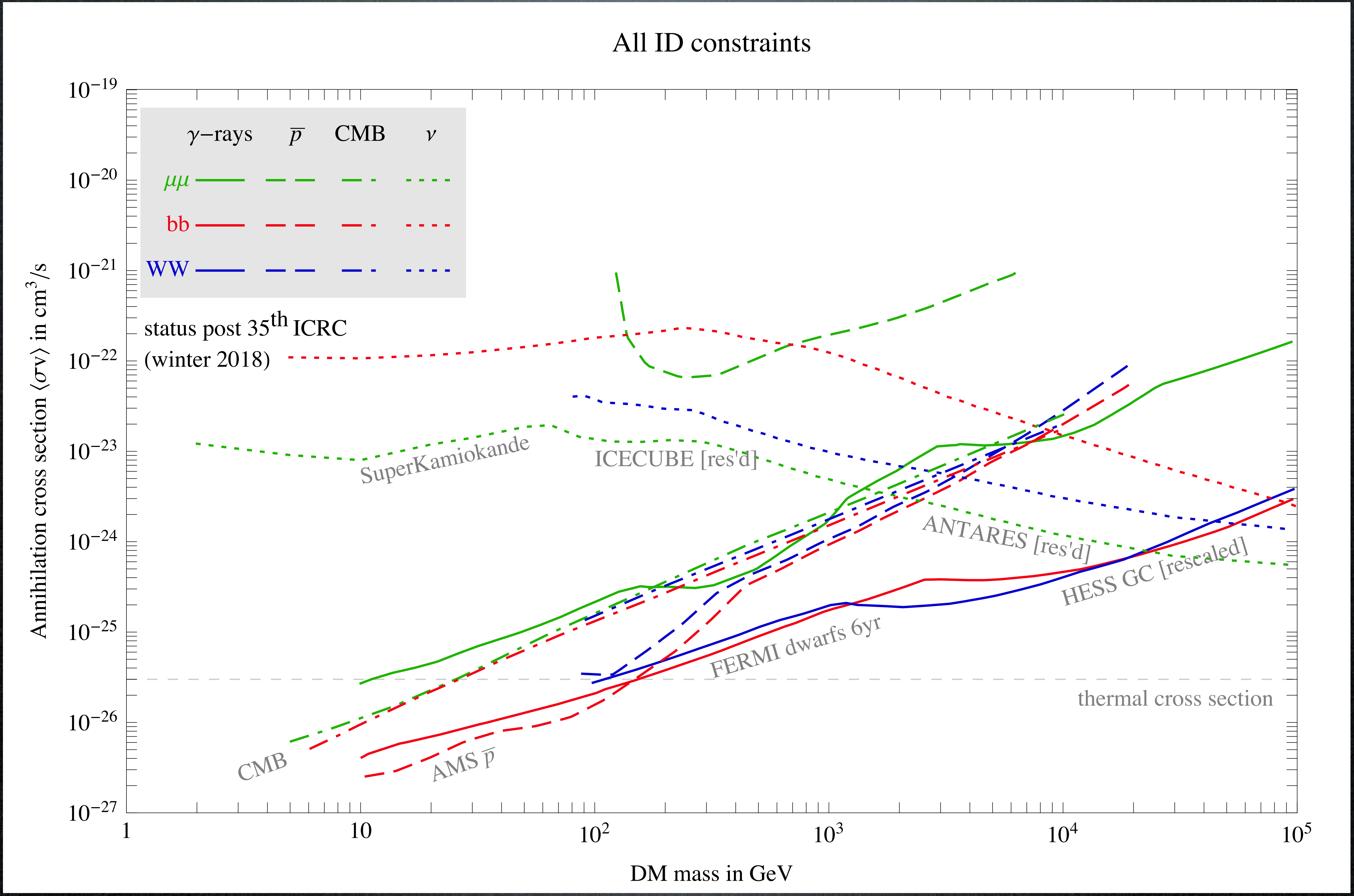

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in $\bar{p}$: still large uncertainties
in $\bar{d}$: challenging flux
in $\overline{He}$: hopeless? who knows
in ν: challenging detection
in γ: astrophysical background

Solution:

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Solution:

- multimessenger
- switch-off astrophysics